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# Optimization and Realization of Quadruple-ridge Flared Horn with New Spline-defined Profiles as a High-efficiency Feed for Reflectors over 4.6 – 24 GHz

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**Abstract**—In this paper, we present a new optimization and realization of a quadruple-ridge flared horn (QRFH) as a feed for reflector antennas of the Square Kilometre Array (SKA) project. The QRFH has been numerically optimized by spline-defined profiles for both the ridges and the horn sidewall, with a conical cavity in the back-short for maximum aperture efficiency. The final aperture efficiency better than (for horizontal polarization) or around (for vertical polarization) 78% over 4.6–8 GHz, 70% over 8–15 GHz, 65% over 15–20 GHz and 60% over 20–24 GHz in the SKA offset-Gregorian dual-reflector antenna has been achieved. The realization of the horn has been carried out carefully by applying several new mechanical design solutions in order to guarantee the accurate positioning of the ridges, the feeding pins and a good electrical contact. The measured  $S_{11}$  is mostly better than -10 dB and the predicted aperture efficiency based on the measured far field patterns agrees well with the simulated result. System performance, such as the sensitivity and system noise temperature, are also estimated and presented.

**Index Terms**—quadruple-ridge flared horn, wideband feed, spine-defined profiles, aperture efficiency, sensitivity.

## I. INTRODUCTION

The Square Kilometre Array (SKA) [1] is an international project for the next-generation radio astronomical observations at meter- to centimeter-wavelength. It will be an interferometric array with a total collecting area of about one million square meters, which in its final stage will provide a sensitivity more than 50 times higher than the current largest interferometers in the world. The SKA will be constructed in two distinct phases, SKA1 and SKA2, the former being a subset of the latter. The SKA1 is now in its pre-construction phase. The Advanced Instrumentation Programme (AIP) is to enhance the SKA1 baseline and demonstrate an important technology prototype of direct relevance to SKA2 [2]. Sweden (Onsala

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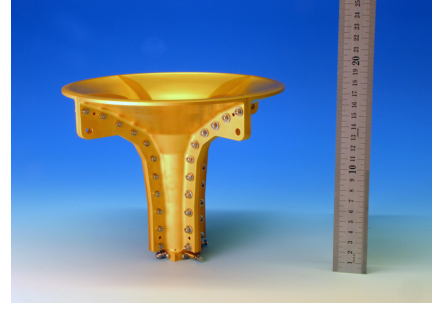


Fig. 1. Prototype of the spline-profiled quadruple-ridge flared horn for the SKA Band-B feed.

Space Observatory and Chalmers University of Technology) committed the technical research and design of the Band-B (4.6–24 GHz) feed for AIP.

Several ultra-wideband (UWB) feed technologies have been developed during the last decade. Among them the Eleven feed [3], [4] and the quadruple-ridge flared horn (QRFH) [5]–[7] are the most representative ones. The challenges in designing and realizing a Band-B feed include: 1) very challenging requirements for both aperture efficiency and sensitivity over a 5:1 bandwidth; 2) high frequency up to 24 GHz, therefore small dimensions and tight tolerances for manufacturing and assembling. In terms of the integration ability with LNA (Low Noise Amplifier), the feasibility for cryogenic operation and the flexibility of the beam shape to match with the SKA offset optics, a QRFH feed has advantages over an Eleven feed as the former offers standard single-ended 50-Ω output ports, all metal structure and the freedom of controlling the beamwidth via changing its profiles.

The novelty of this work is the combination of existing and available ideas demonstrated previously in [5], [8], [9] to improve the QRFH feed. Contributions of this work include: 1) new spline-defined profiles for both the ridges and the horn sidewall of the QRFH, with the aim to achieve the very high aperture efficiency as well as sensitivity over a 5:1 bandwidth; 2) the realization of a 4.6–24 GHz QRFH by implementing an edge-extrude mechanical technique to ensure good conductive contact between the ridges and the horn sidewall.

## II. DESIGN AND ANALYSIS

The SKA dish antenna has a shaped offset-Gregorian optics, with a projected diameter of 15 m for the main reflector and 5.8 m for the sub-reflector with subtended angle of  $\theta_c=58^\circ$  and a shaped extension on the bottom side to reduce the ground spillover noise [10]. Specification for  $\epsilon_{ap}$  and  $A_{eff}/T_{sys}$  is listed in Table I. This requirement, if roughly estimated into the scenario of a primary-fed paraboloid, requires the feed to possess an  $\epsilon_{ap}$  of about 10% lower [7], [11], i.e. 68% over 4.6–8 GHz, 60% over 8–15 GHz, 55% over 15–20 GHz and 50% over 20–24 GHz, which, to the best knowledge of the authors, have not been very well achieved by any existing 5:1 or 6:1 QRFHs [7], [12], [13].

The shapes of the horn sidewall and ridges are of great importance to the performance of QRFHs and have conventionally been expressed by analytical profiles [7]. However,

TABLE I  
THE SKA REQUIREMENT FOR  $\epsilon_{ap}$  AND  $A_{eff}/T_{sys}$  OVER BAND B

frequency	4.6–8 GHz	8–15 GHz	15–20 GHz	20–24 GHz
$\epsilon_{ap}$	78%	70%	65%	60%

frequency	4.6–13.8 GHz	13.8–20 GHz	20–24 GHz
$A_{eff}/T_{sys}$	$6.1 \text{ m}^2/\text{K}$	$4.7 \text{ m}^2/\text{K}$	$3.5 \text{ m}^2/\text{K}$

QRFHs with such profiles suffer from a strong beamwidth variation over frequency (especially in H-plane) as well as the bad cross-polarization level (in D-plane, the maximum relative cross-polar level is only -6 dB [7]), which makes the illumination efficiency  $\epsilon_{ill}$  drop towards higher frequencies and results in a bad polarization efficiency  $\epsilon_{pol}$  and BOR<sub>1</sub> efficiency  $\epsilon_{BOR1}$ . A lot of efforts were made to improve these sub-efficiencies, such as the utilization of quadaxial feeding [9], [12] and other kinds of aperture shapes [8], which make  $\epsilon_{ill}$  higher and more flat, but mostly at the cost of a worse spillover efficiency  $\epsilon_{sp}$  (barely up to 95% [7], [12]). In other words,  $\epsilon_{ill} \cdot \epsilon_{sp}$  is still not good enough to meet the Band-B requirement. The quadaxial feeding managed to improve the relative cross-polar level to -9 dB, but this differential technique needs additional baluns and hence more insertion loss would be introduced. Differential LNAs feeding can be another alternative, but it is normally more expensive and exhibits higher noise factors.

Spline-defined profile already showed very promising beam performance in [14] and is also adopted for the Band-B feed [15]. It is more flexible than the analytical profiles and enables a highly customizable optimization on the feed structure. After the optimization, the new spline-profiled Band-B feed has fulfilled and been very close to fulfill the SKA requirement for aperture efficiency for horizontal and vertical polarizations respectively, while still keeping the spillover noise temperature to an acceptable level.

#### A. The Spline-defined Profiles of Band-B Feed

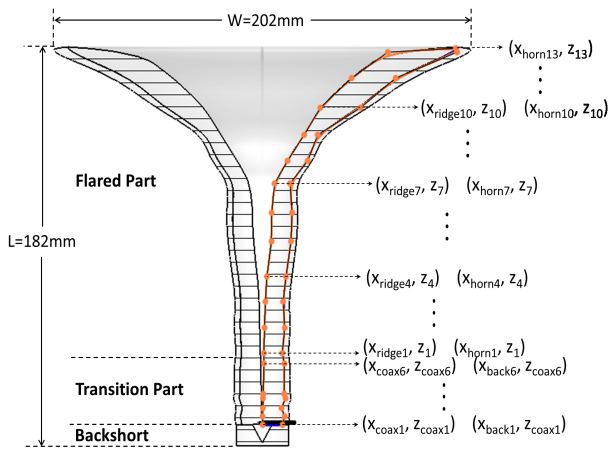


Fig. 2. Feed geometry: the spline-defined profiles defined by discrete points.

The proposed QRFH geometry is shown in Fig. 2, where the spline curve function in CST is utilized to construct the

ridges and horn sidewall. In our design, 13 points ( $x_{horn1}, z_1$ ), ( $x_{horn2}, z_2$ ), ( $x_{horn3}, z_3$ ), ... ( $x_{horn13}, z_{13}$ ) are chosen for the horn profile. 13 points with the same z-coordinates ( $x_{ridge1}, z_1$ ), ( $x_{ridge2}, z_2$ ), ( $x_{ridge3}, z_3$ ), ... ( $x_{horn13}, z_{13}$ ) (the last point is shared with the horn profile) are chosen for the ridge profile. These 38 parameters in total form the flared part of the feed, as is shown in Fig. 2. The transition part is to have impedance transformation from the flared part to a standard 50- $\Omega$  coaxial cable, where 6 points ( $x_{coax1}, z_{coax1}$ ), ( $x_{coax2}, z_{coax2}$ ), ... ( $x_{coax6}, z_{coax6}$ ) are defined for the inner profile of the ridges, and 6 points with the same z-coordinates ( $x_{back1}, z_{coax1}$ ), ( $x_{back2}, z_{coax2}$ ), ... ( $x_{back6}, z_{coax6}$ ) defined for the horn sidewall profile, 12 points as illustrated in Fig. 2. A conical-shaped backshort cavity, which possesses a good wideband performance, was utilized and defined by two parameters (radius and depth). The four ridges are made of metal pieces with thickness of 1.5 mm, a large chamfering angle of  $\theta=55.31^\circ$  is chosen to make the gap width between opposite ridges small enough. In total, there are 60 parameters in this spline-profiled QRFH feed.

~~The optimization procedure was divided into three stages: the flared part, the transition part and the whole feed.~~ A time-efficient Matlab-CST optimization pipeline [7], [16] was established to search for the best performance with the Genetic Algorithm (GA). Please note that aperture efficiency  $\epsilon_{ap}$  was defined as the goal function and it was calculated in a primary-fed paraboloid using the method described in [17], [18]. Through intensive investigations, it was found that for optimizations in a primary-fed paraboloid, a smaller half-subtended angle of  $\theta_c=51^\circ$  is more suitable than  $\theta_c=58^\circ$ , as  $\theta_c=51^\circ$  provides a better control of the spillover efficiency  $\epsilon_{sp}$  and shows a better correlation with the corresponding aperture efficiency in the SKA dish. Over Band B, the  $\epsilon_{ap}$  in a primary-fed paraboloid with  $\theta_c=51^\circ$  would be constantly  $\sim 10\%$  lower than the final  $\epsilon_{ap,SKA}$  in the SKA dish, see Fig. 4.

~~Approximately 5000 different sets of parameters in total have been investigated.~~ After optimization, the whole feed measures 202 mm in aperture diameter and 182 mm in length.

#### B. Aperture Efficiency and Radiation Patterns

The final optimized feed shows an aperture efficiency  $\epsilon_{ap}$  better than 68% over 4.6–8 GHz, 60% over 8–15 GHz, 58% over 15–20 GHz and 55% over 20–24 GHz in a primary-fed paraboloid ( $\theta_c=51^\circ$ ), as shown in Fig. 4. The corresponding aperture efficiency in the SKA offset dish (the black-dashed curve)  $\epsilon_{ap,SKA}$ , also shown in the figure, demonstrates a very similar trend with a constant  $\sim 10\%$  margin to  $\epsilon_{ap}$ . The illumination efficiency  $\epsilon_{ill}$  is higher than 65% over the entire bandwidth and exhibits a relatively flat trend, which implies nearly constant beamwidths over frequency. Meanwhile, the spillover efficiency  $\epsilon_{sp}$  is kept higher than 90% over 5–24 GHz and even exceeds 95% towards mid-band and high-band, which provides a good management of the spillover noise temperature. The polarization efficiency  $\epsilon_{pol}$  is better than 95% over most of the band and falls down to slightly above 90% only around 11 GHz. This is a significant improvement compared to the analytical-profiled QRFHs'  $\epsilon_{pol}$  of only about

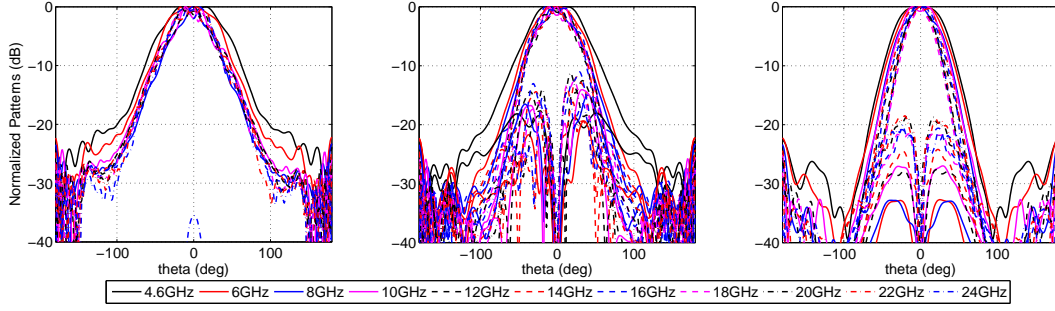


Fig. 3. Simulated radiation patterns in E-, D- and H-planes, both co-polar and cross-polar components are shown.

90% across the bandwidth [13]. The BOR<sub>1</sub> efficiency  $\epsilon_{\text{BOR}_1}$  is around 90% over the band, which is 5~10% higher than that of the QRFHs using analytical profiles [7].

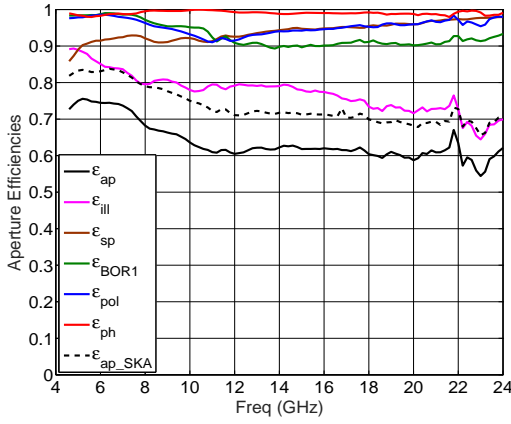


Fig. 4. Simulated aperture efficiency and the sub-efficiencies of the optimized spline-profiled QRFH feed, calculated in a primary-fed paraboloid with half-subtended angle of  $\theta_c=51^\circ$ .

Fig. 3 shows the simulated radiation patterns in E- ( $\phi=0^\circ$ ), D- ( $\phi=45^\circ$ ) and H- ( $\phi=90^\circ$ ) planes, normalized co-polar as well as relative cross-polar. The co-polar patterns exhibit very good Gaussian shapes with sidelobe and backlobe levels well below -20 dB. The beam profiles are nearly constant with frequency, even in H-plane the fluctuation of nominal 10 dB beamwidth is only about  $2 \times 20^\circ$ , an improvement on the beam rotational symmetry from the existed QRFH's H-plane beamwidth fluctuation of  $2 \times (30 \sim 35)^\circ$  [7], [13]. The cross-polar level is most significant in D-plane for BOR and quasi-BOR antennas and the QRFH shows a relative cross-polar level below -11 dB in D-plane, another improvement from the analytical-profiled QRFHs, which is about -6 dB in [7], [13].

### III. MECHANICAL REALIZATION

The final electrical design was realized by dividing the whole feed into four pieces of ridges, four sidewall quad-pieces, a backshort cavity and two feed pins, see Fig. 5(a) and Fig. 5(c), with each ridge being clamped between two adjacent sidewall quad-pieces, see Fig. 5(b). The ridges and sidewall quad-pieces were manufactured by Computer Numerical Control (CNC) machines of Aluminum Alloy (AlSi1MgMn) which is cheap and also well-performed in vacuum and

cryostat. On top of the metal surface, a gold layer of  $3 \mu\text{m}$  thick was plated to reduce thermal loss and prevent oxidation. All the small components (feed-pins, etc) were fabricated by using a watchmaker lathe.

Fig. 6 shows the simulated surface current on the inner sidewall of the horn at 24 GHz, where the transverse component of the current reaches its maximum at the jointing between the ridges and the sidewall, so any gaps between them will change the current distribution significantly, which will result in degraded reflection coefficient and radiation performance, and even leakage. A special edge-extrude technique is used to guarantee the good electrical contact: the sidewall quad-pieces were designed and manufactured with an edge extrusion with a height of 0.1 mm and a width of 0.2 mm, see Fig. 5(e). Thus the sharp edge extrusion will indent the surface on both

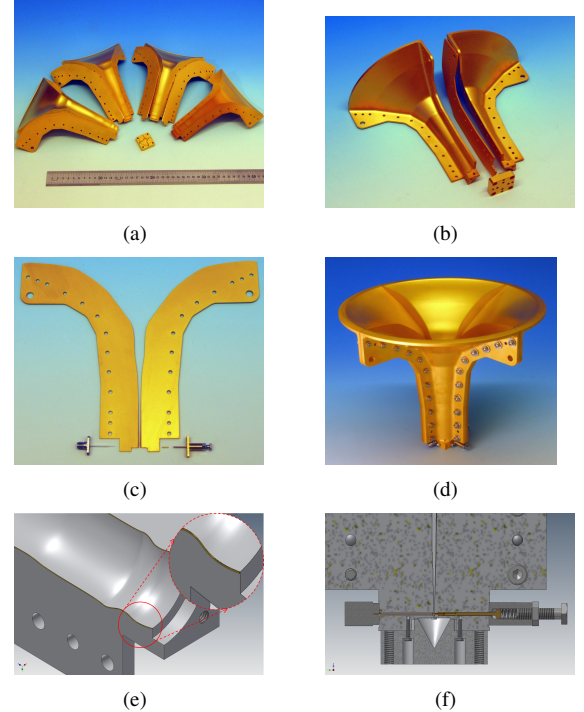


Fig. 5. Mechanical design of the Band-B feed: (a) manufactured major components; (b) zoom-in of the inner edge on one sidewall quad-piece (marked in yellow); (c) mechanical design around the feeding part; (d) each ridge is clamped between two quad-pieces during assembly; (e) assembly of the feeding pin; (f) the Band-B feed after assembly.



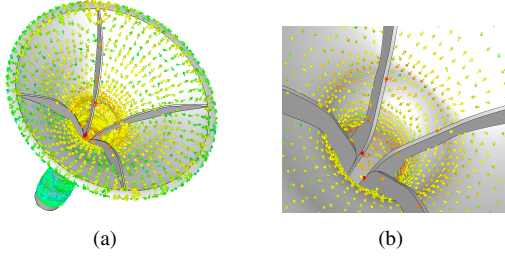


Fig. 6. Surface current distribution at 24 GHz (in magnitude, based on CST simulation): (a) surface currents on the whole feed; (b) surface currents around the radiation region.

sides of each ridge as the ridge material is a bit softer and will act like a punching die. With the golden layer this would guarantee an extremely good electrical contact after tightening the two adjacent sidewall quad-pieces with screws.

A coaxial cable with an inner conductor diameter of  $\phi=0.28$  mm was chosen as the feeding pin. The inner conductor was inserted into the central hole of a spring-loaded cylinder stud embedded within the opposite ridge, which locks the pin by a chuck, see Fig. 5(f). This design secures the clamp of the pin and at the same time allows some small movements of the material due to heating- and cooling-extension.

#### IV. MEASURED PERFORMANCE

##### A. Scattering Parameters

The simulated and measured scattering parameters are shown in Fig. 7. The measured  $S_{11}$  and  $S_{22}$  are mostly below -10 dB across the bandwidth except for around 10.6 GHz and 23.2 GHz. At low frequencies, the measured results follow the same trend with simulation, while at high frequencies the measured reflection coefficient is higher than the simulation for about 5 dB. The measured port isolation is around -25 dB over the bandwidth. The inconsistency between measurement and simulation is caused mainly by a mechanical error, that on one ridge the feeding hole was not center-drilled due to a drift of the drill, the corresponding feed pin therefore has to be slightly bent in order to fit in.

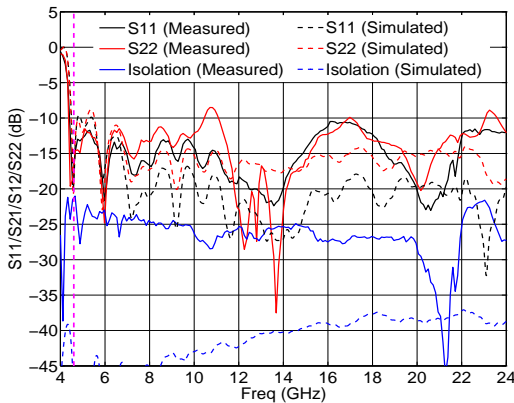


Fig. 7. Measured and simulated S-parameters of the prototype feed.

##### B. Pattern Performance

The feed radiation patterns were measured by the Antenna Measurement Team at Yebes Observatory in Spain. A spherical near-field measurement was conducted (Fig. 8) and five sets of Open Ended Waveguide (OEWG) probes (WR-229, WR-159, WR-112, WR-75 and WR-42) were utilized to cover the entire Band B. The measurement frequency step is 0.2 GHz, but there was a gap between 15 and 15.6 GHz that can't be covered by the aforementioned probes.

The measured radiation patterns show great consistence with simulations. Fig. 8 shows the measured and simulated 10 dB beamwidths in E-, D- and H-planes as well as the maximum relative cross-polar level (XPI) in D-plane. The measured relative cross-polar level in D-plane exhibits an expected performance of below -11 dB. Fig. 8 indicates the good beamwidth stability and therefore the rotational symmetry of radiation patterns has been improved.

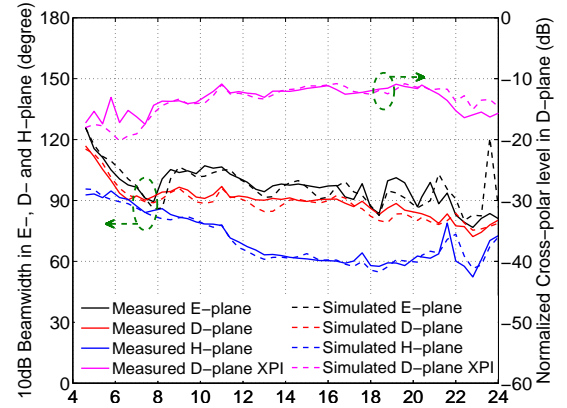


Fig. 8. Measured and simulated 10 dB beamwidth in E-, D- and H-planes, as well as the maximum relative cross-polar level (XPI) in D-plane.

##### C. Predicted System Performance in the SKA Dish

With the measured radiation patterns and estimated receiver noise, both aperture efficiency and sensitivity ( $A_{\text{eff}}/T_{\text{sys}}$ ) in the SKA shaped offset-Gregorian dual-reflector antenna (geometry shown in Fig. 9) were analyzed afterwards, using a System Simulator code written in Matlab [19], which models the total antenna beam pattern through physical optics and physical theory of diffraction calculations in GRASP. The effective area  $A_{\text{eff}}$  is calculated from the theoretical aperture

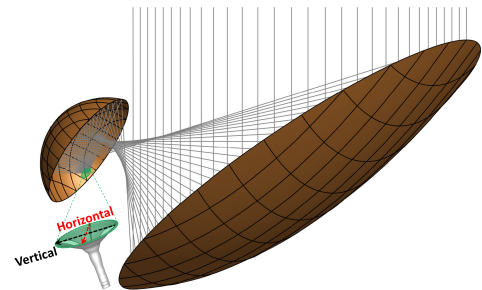


Fig. 9. The SKA dish geometry: an offset-Gregorian dual-reflector antenna.

efficiency that is derived from the antenna directivity. A sky noise temperature model [20] for all elevations, specific to the conditions of the suggested SKA-mid site (altitude 1000 m), is weighted with the antenna beam pattern and integrated over all elevations of the sky and the hot ground to calculate the total antenna noise  $T_A$ . This sky noise temperature model mainly includes cosmic microwave background (CMB,  $\sim 2.7$  K), the atmospheric noise and quite weak galactic emission component away from the galaxy plane. Please note that over Band B there is a significant increase in atmospheric noise around the water vapor absorption line at 22.2 GHz, as indicated by Fig. 2 and Fig. 3 in [20].

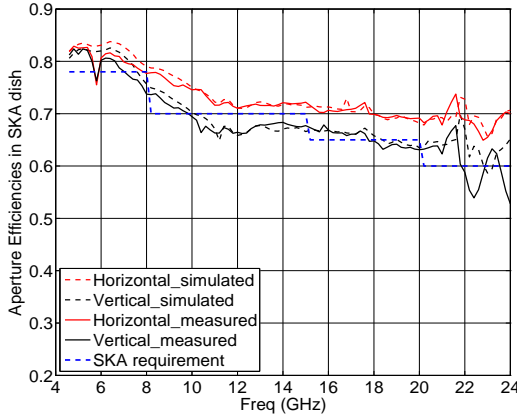


Fig. 10. Predicted aperture efficiencies  $\epsilon_{ap}$  in the SKA dish, based on simulated and measured radiation patterns, with a blue-dashed line representing the SKA requirement.

Fig. 10 shows the aperture efficiencies in the SKA dish, based on simulated patterns and measured patterns respectively, with the blue-dashed line representing the SKA requirement. The horizontal polarization has clearly fulfilled the requirement, with its  $\epsilon_{ap}$  better than 70% across most of the bandwidth, and the measurement agrees quite well with simulation except for the 5.8 GHz-dip and some minor reduction of  $\sim 1\%$  at lower band (4.6–10 GHz). As for the vertical polarization, the measurement exhibits the same trend with simulation but drops to some extent over 20–24 GHz, this is partly due to the worse performance caused by the bent feeding pin in this polarization, as well as the polarization discrepancy that is caused by the offset optics of the SKA dish. Though not fully qualified, the vertical polarization is still close to the SKA requirement, and an aperture efficiency of better than 53% could be obtained over the entire bandwidth.

The estimated system noise temperature  $T_{sys}$  is shown in Fig. 11. It consists of the receiver noise temperature  $T_{REC}$  and the antenna noise temperature  $T_A$ .  $T_{REC}$  includes the feed noise temperature  $T_{feed}$  (in the cryostat) and LNA noise temperature  $T_{LNA}$  (the ohmic losses of coaxial cable and possible directional couplers for calibration between the feed and the receiver were not included). 2 K is estimated for  $T_{feed}$  which covers the noise contributions from the feed window, the insertion loss of the feed and the mismatch loss between feed and 1st LNA. The LNA was designed by the Gothenburg-based company LNF (Low Noise Factory AB) as part of the WBSPF development program and the measured  $T_{LNA}$

at ambient temperature of 20 K was used for estimation. The antenna noise temperature  $T_A$  is composed of the sky temperature  $T_{sky}$  which is picked up by the main beam of the telescope, and the spillover noise temperature  $T_{sp}$  which is contributed from the hot ground and would be picked up through the sidelobes and backlobes of the telescope. Both  $T_{sky}$  and  $T_{sp}$  were calculated by the System Simulator. In Fig. 11, the receiver noise  $T_{REC}$  is relatively flat and less than 10 K over the bandwidth, the antenna noise temperature  $T_A$  is comparable with  $T_{REC}$  at low-band and mid-band, but due to the strong water vapor absorption line around 22.2 GHz,  $T_A$  increases dramatically towards high-band, which is the main reason for the bad performance of  $A_{eff}/T_{sys}$  at higher frequencies in Fig. 12.

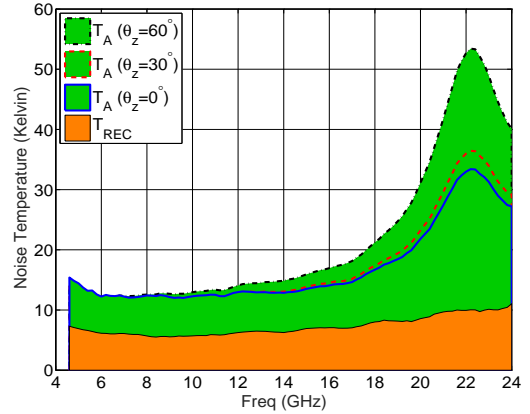


Fig. 11. Estimated system noise temperature over Band B, including the receiver noise  $T_{REC}$  and antenna noise  $T_A$ .

The system sensitivities ( $A_{eff}/T_{sys}$ ) in the SKA dish for different zenith angles  $\theta_z$ , are given in Fig. 12. For  $\theta_z=0^\circ$ , the discrepancy between two polarizations is most severe, especially at the low-band, which is caused by the asymmetry of the radiation patterns and the offset optics of the SKA dish. However, as the zenith angle increases the discrepancy between two polarizations decreases. For  $\theta_z=0^\circ$  and  $\theta_z=30^\circ$ , the sensitivity have fulfilled the SKA requirement over 4.6–21 GHz, but over 21–24 GHz the vertical polarization fails to achieve the requirement. For  $\theta_z=60^\circ$ , both polarizations are performing well up until 18.25 GHz but unqualified towards higher frequencies, the worst performance occurs around 22.5 GHz and it has a difference of about  $1.5 \text{ m}^2/\text{K}$  from the SKA requirement. The biggest problem at high frequencies here is caused by the water-vapor absorption line at 22.2 GHz, which is very severe in terms of sky-noise rather than spillover noise. This noise will be inevitably picked up through the main beam of the telescope. So, to achieve the sensitivity requirement at high frequencies, we would need a very high aperture efficiency but that would be very difficult to achieve for a wideband feed.

## V. CONCLUSION

This paper presents the optimization and realization of a quadruple-ridge flared horn with new spline-defined profiles designed for the Band-B feed and the system performance

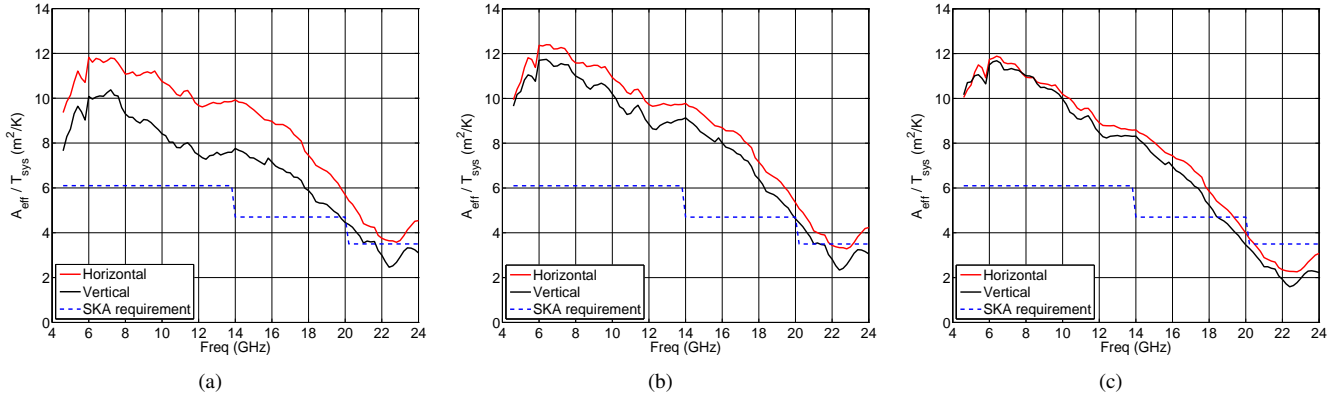


Fig. 12. Predicted system sensitivities of both polarizations for different zenith angle: (a)  $\theta_z=0^\circ$ ; (b)  $\theta_z=30^\circ$  and (c)  $\theta_z=60^\circ$ .

analysis in the SKA dish. The feed exhibits a good impedance matching and beam performance over more than 5:1 bandwidth. Its sub-efficiencies and radiation patterns show that by adopting spline-defined profiles, more constant beamwidth and more symmetrical radiation patterns have been obtained, hence improved  $\epsilon_{\text{BOR}_1}$ ,  $\epsilon_{\text{pol}}$  and  $\epsilon_{\text{ill}} \cdot \epsilon_{\text{sp}}$  can be achieved. An edge-extrude mechanical technique was implemented to realize a good contact between ridges and horn sidewall, so as to guarantee a high operational frequency. Measured results show reasonable agreement with simulations, predicted aperture efficiency in the SKA dish based on the measured farfield patterns have fulfilled the SKA requirement for its horizontal polarization, with a polarization discrepancy that makes the vertical polarization unqualified over part of the bandwidth. This discrepancy is caused by the asymmetrical optics of the SKA offset dish which are not fully optical-matched with the symmetrically-designed Band-B feed, and thus was ignored during our primary-fed-parabolic optimization. As for the predicted  $A_{\text{eff}}/T_{\text{sys}}$ , a large margin above the requirement over low- and mid-band was achieved, while at high-band it slightly fails to satisfy the requirement due to the strong water-vapor absorption line around 22.2 GHz.

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