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# EVN observations of 6.7 GHz methanol maser polarization in massive star-forming regions

## IV. Magnetic field strength limits and structure for seven additional sources<sup>★</sup>

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### ABSTRACT

**Context.** Magnetohydrodynamical simulations show that the magnetic field can drive molecular outflows during the formation of massive protostars. The best probe to observationally measure both the morphology and the strength of this magnetic field at scales of 10–100 au is maser polarization.

**Aims.** We measure the direction of magnetic fields at milliarcsecond resolution around a sample of massive star-forming regions to determine whether there is a relation between the orientation of the magnetic field and of the outflows. In addition, by estimating the magnetic field strength via the Zeeman splitting measurements, the role of magnetic field in the dynamics of the massive star-forming region is investigated.

**Methods.** We selected a flux-limited sample of 31 massive star-forming regions to perform a statistical analysis of the magnetic field properties with respect to the molecular outflows characteristics. We report the linearly and circularly polarized emission of 6.7 GHz CH<sub>3</sub>OH masers towards seven massive star-forming regions of the total sample with the European VLBI Network. The sources are: G23.44–0.18, G25.83–0.18, G25.71–0.04, G28.31–0.39, G28.83–0.25, G29.96–0.02, and G43.80–0.13.

**Results.** We identified a total of 219 CH<sub>3</sub>OH maser features, 47 and 2 of which showed linearly and circularly polarized emission, respectively. We measured well-ordered linear polarization vectors around all the massive young stellar objects and Zeeman splitting towards G25.71–0.04 and G28.83–0.25. Thanks to recent theoretical results, we were able to provide lower limits to the magnetic field strength from our Zeeman splitting measurements.

**Conclusions.** We further confirm (based on ~80% of the total flux-limited sample) that the magnetic field on scales of 10–100 au is preferentially oriented along the outflow axes. The estimated magnetic field strength of  $|B_{||}| > 61$  mG and  $>21$  mG towards G25.71–0.04 and G28.83–0.25, respectively, indicates that it dominates the dynamics of the gas in both regions.

**Key words.** masers – magnetic fields – polarization – stars: formation

## 1. Introduction

Several theoretical and observational efforts are advancing our understanding of the formation of high-mass stars ( $M > 8 M_{\odot}$ ). In the last twenty years several models were developed, among which we mention the Core Accretion model (e.g., McKee & Tan 2003), the Competitive Accretion model (e.g., Bonnell et al. 2001), and the hybrid model of these two models (e.g., Tan et al. 2014 and references therein). From a theoretical point of view, intensive simulation campaigns have been carried out by several authors in the last decade (e.g., Krumholz et al. 2009; Peters et al. 2010, 2011; Hennebelle et al. 2011; Klessen et al. 2011; Seifried et al. 2011,

2012, 2015; Klassen et al. 2012, 2014, 2016; Kuiper et al. 2011, 2015, 2016). The simulations are mainly focused on understanding the role of feedback before and after the formation of cores (e.g., Myers et al. 2013, 2014) and young stellar objects (YSOs; e.g., Peters et al. 2012; Kuiper et al. 2016), and the effects of more than two feedback mechanisms are rarely taken into consideration within the same simulation (for more details see Tan et al. 2014).

Magnetohydrodynamical (MHD) simulations show that jets and molecular outflows are common in massive YSOs and that they have an important role in the formation process (e.g., Tan et al. 2014; Matsushita et al. 2018). For instance, Banerjee & Pudritz (2007) showed that early outflows can reduce the radiation pressure allowing the further growth of the protostar. The outflows are found to be initially poorly collimated and became more collimated only when a nearly Keplerian disk is formed and fast

<sup>★</sup> Tables A.1–A.7 are only available at the CDS via anonymous ftp to [cdsarc.u-strasbg.fr](http://cdsarc.u-strasbg.fr) (130.79.128.5) or via <http://cdsarc.u-strasbg.fr/viz-bin/qcat?J/A+A/623/A130>.

jets are generated (Seifried et al. 2012). The mass ejection rate through a MHD outflow and the accretion rate through an accretion disk are comparable only when the magnetic energy ( $E_B$ ) of the initial core is comparable to the gravitational energy ( $E_G$ ; Matsushita et al. 2017). Furthermore, when the magnetic field is strong ( $\mu < 5^1$ ) the outflow is slow and poorly collimated (Seifried et al. 2012) and its structure is determined by the large-scale geometry of the magnetic field lines (Matsushita et al. 2017). When the magnetic field is weak ( $\mu > 10$ ), in addition to leading to the fragmentation of the parental cloud core, the outflow is fast and well collimated (Hennebelle et al. 2011), and it might even disappear during the formation of the YSO, leaving the scene to a magnetically supported toroid-like structure (Matsushita et al. 2017).

It needs to be mentioned that other theoretical studies investigated the possibility that the outflows are driven by other mechanisms than magnetic fields (e.g., Yorke & Richling 2002; Vaidya et al. 2011; Peters et al. 2011; Kuiper et al. 2015). In particular, if the outflows are driven by the ionization feedback they appear to be uncollimated (e.g., Peters et al. 2011). In addition, Peters et al. (2014) also found that stars that form in a common accretion flow tend to have aligned outflows, which can combine to form a collective outflow similar to what is observed in W75N(B) (e.g., Surcis et al. 2014a).

From an observational point of view, the presence of molecular outflows in massive star-forming regions (SFRs) is nowadays a fact (e.g., Tan et al. 2014 and references therein). Also, measuring the morphology of the magnetic field around massive YSOs is now regularly done by using polarized dust emission (on a scale of  $10^3$  au; e.g., Zhang et al. 2014a; Girart et al. 2016) and polarized maser emission (on a scale of tens of au; e.g., Vlemmings et al. 2010; Surcis et al. 2012, 2013, 2015, hereafter Papers I–III; Sanna et al. 2017; Dall’Olio et al. 2017). However, the findings at the two different scales conflict. On the large scale (arcsecond resolution) magnetic fields appear to be randomly distributed with respect to the outflow axis (Zhang et al. 2014b), while on the small scale (milliarcsecond resolution) magnetic fields, estimated from the polarized emission of the 6.7 GHz CH<sub>3</sub>OH masers, are preferentially oriented along the outflow (Paper III). We note that the conclusions reported in Paper III are based on 60% of the flux-limited sample (19 out of 31 sources), whereas those in Zhang et al. (2014b) are based on a total of 21 sources. There are no sources in common between the two samples (Zhang et al. 2014b; Paper III). Nonetheless, the comparison around the same YSO, when possible, of the magnetic field morphology at the two scales has shown consistency (e.g., Surcis et al. 2014b).

Measurements of the magnetic field strength close to YSOs are less common. It has been possible, so far, only through the analysis of the circularly polarized emission of H<sub>2</sub>O masers (e.g., Surcis et al. 2014a; Goddi et al. 2017). Because of the shock-nature of the H<sub>2</sub>O masers, the strength of the magnetic field is estimated in the post-shock compressed gas, even if it can be possible to derive the pre-shock magnetic field strength (e.g., Imai et al. 2003; Vlemmings et al. 2006; Goddi et al. 2017). Although the circularly polarized emission of the 6.7 GHz CH<sub>3</sub>OH maser has been regularly detected (e.g., Papers I–III), no estimates of the magnetic field strength have been possible

due to the unknown Landé g-factors (Vlemmings et al. 2011). Very recently, Lankhaar et al. (2018) theoretically calculated the g-factors for all the CH<sub>3</sub>OH maser transitions making it possible to estimate at least a lower limit of the magnetic field strength from the Zeeman splitting measurements.

Here, in the fourth paper of the series after Papers I–III, we present the results of the next seven observed sources of the flux-limited sample, briefly described in Sect. 2. In Sect. 3, in addition to reporting the observations, we describe the changes made to the adapted full radiative transfer method (FRTM) code for the 6.7 GHz CH<sub>3</sub>OH maser emission. The results are presented in Sect. 4 and discussed in Sect. 5, where we also briefly update the previous statistics (see Paper III).

## 2. Massive star-forming regions

We selected a flux-limited sample of 31 massive SFRs with declination  $> -9^\circ$  and a total CH<sub>3</sub>OH maser single-dish flux density greater than 50 Jy from the 6.7 GHz CH<sub>3</sub>OH maser catalog of Pestalozzi et al. (2005), and that in more recent single-dish observations showed a total flux density  $> 20$  Jy (Vlemmings et al. 2011). We already observed and analyzed 19 of these sources (Vlemmings et al. 2010; Surcis et al. 2009, 2011a, 2014b; Papers I–III). Seven more sources, which are described below in Sects. 2.1–2.7, have been observed at 6.7 GHz with the European VLBI Network<sup>2</sup> (EVN). The last five sources in the sample will be presented in the next paper of the series.

### 2.1. G23.44–0.18

G23.44–0.18 is a high-mass SFR at a heliocentric distance of  $5.88^{+1.37}_{-0.93}$  kpc (Brunthaler et al. 2009) in the Norma arm of our Galaxy (Sanna et al. 2014) with a systemic velocity of  $V_{\text{lsr}}^{\text{CS}(2-1)} = +104.2 \text{ km s}^{-1}$  (Bronfman et al. 1996). The region contains two millimeter dust continuum cores, named MM1 and MM2, separated by  $\sim 14''$  ( $8 \times 10^4$  AU), suggesting the presence of two YSOs prior to forming an ultra-compact H II (UC H II) region (Ren et al. 2011). Furthermore, a strong bipolar CO outflow ( $\text{PA}_{\text{outflow}}^{12\text{CO}} = -40^\circ$ ) originates from MM2 (Ren et al. 2011). The CO outflow consists of a low-velocity component (LVC) and a high-velocity component (HVC) whose blue-shifted ( $+83 \text{ km s}^{-1} < V_{\text{blue}}^{\text{LVC}} < +93 \text{ km s}^{-1}$  and  $+65 \text{ km s}^{-1} < V_{\text{blue}}^{\text{HVC}} < +75 \text{ km s}^{-1}$ ) and red-shifted ( $+113 \text{ km s}^{-1} < V_{\text{red}}^{\text{LVC}} < +123 \text{ km s}^{-1}$  and  $+140 \text{ km s}^{-1} < V_{\text{red}}^{\text{HVC}} < +170 \text{ km s}^{-1}$ ) lobes are oriented northwest and southeast, respectively (Ren et al. 2011). Two groups of 6.7 GHz CH<sub>3</sub>OH masers have been detected around MM1 and MM2 (Walsh et al. 1998; Fujisawa et al. 2014; Breen et al. 2015). A 12 GHz CH<sub>3</sub>OH maser emission has been detected with a velocity coverage consistent with the two groups of 6.7 GHz CH<sub>3</sub>OH masers, while the OH maser emission is likely associated with MM2 (e.g., Breen et al. 2016; Caswell et al. 2013). Vlemmings et al. (2011) measured a Zeeman splitting of the 6.7 GHz CH<sub>3</sub>OH maser of  $0.43 \pm 0.06 \text{ m s}^{-1}$  with the Effelsberg telescope.

### 2.2. G25.83–0.18

G25.83–0.18 is a very young SFR in an evolutionary stage prior to the UC H II region phase at a kinematic distance of  $5.0 \pm$

<sup>1</sup>  $\mu = (M/\Phi)/(M/\Phi)_{\text{crit}}$ , where  $(M/\Phi)$  is the mass-to-magnetic flux ratio and  $(M/\Phi)_{\text{crit}} \approx 0.12/\sqrt{G}$  is the critical value of this ratio, where  $G$  is the gravitational constant (Tomisaka et al. 1988). The critical value indicates the maximum mass supported by the magnetic field (Tomisaka et al. 1988). The stronger the magnetic field, the lower  $\mu$ .

<sup>2</sup> The European VLBI Network is a joint facility of European, Chinese, South African, and other radio astronomy institutes funded by their national research councils.

0.3 kpc (Araya et al. 2008; Andreev et al. 2017). Both 6.7 GHz and 12 GHz  $\text{CH}_3\text{OH}$  maser emissions were detected (e.g., Walsh et al. 1998; Błaszczewicz & Kus 2004; Breen et al. 2015, 2016), which are  $\sim 2''$  north of the 4.8 GHz  $\text{H}_2\text{CO}$  maser detected at the center of an IR dark cloud (Araya et al. 2008). The velocities of all the detected maser species, including the  $\text{H}_2\text{O}$  masers (Breen & Ellingsen 2011), are close to the systemic velocity of the region, i.e.,  $V_{\text{lsr}}^{18\text{CO}} = +93.2 \text{ km s}^{-1}$  (de Villiers et al. 2014). A  $^{13}\text{CO}$  outflow has been measured with the James Clerk Maxwell Telescope (JCMT), and its red-shifted ( $+91.8 \text{ km s}^{-1} < V_{\text{red}}^{13\text{CO}} < +101.8 \text{ km s}^{-1}$ ) and blue-shifted ( $+83.8 \text{ km s}^{-1} < V_{\text{blue}}^{13\text{CO}} < +91.8 \text{ km s}^{-1}$ ) lobes are oriented north–south with a position angle of  $\text{PA}_{\text{outflow}}^{13\text{CO}} \sim +10^\circ$  (de Villiers et al. 2014). No radio continuum emission has been detected towards the  $\text{CH}_3\text{OH}$  maser clumps (Walsh et al. 1998).

A Zeeman splitting of the 6.7 GHz  $\text{CH}_3\text{OH}$  maser emission of  $\Delta V_Z = (0.99 \pm 0.26) \text{ m s}^{-1}$  was measured with the Effelsberg 100 m telescope (Vlemmings et al. 2011).

### 2.3. G25.71–0.04

The massive SFR G25.71–0.04, also known as IRAS 18353–0628, is located at a distance of  $10.1 \pm 0.3 \text{ kpc}$  from the Sun (Green & McClure-Griffiths 2011); it is associated with 6.7 GHz and 12 GHz  $\text{CH}_3\text{OH}$  masers and OH masers (Walsh et al. 1997; Fujisawa et al. 2014; Breen et al. 2015, 2016; Szymczak & Gérard 2004). Neither radio continuum emission nor UC H II region are observed at the position of the  $\text{CH}_3\text{OH}$  maser clump (Walsh et al. 1998), a warm dust sub-millimeter source has instead been detected (Walsh et al. 2003). Its bright sub-mm peak suggests that the maser site is likely to be in a stage of evolution before the UC H II region has been created (Walsh et al. 2003). de Villiers et al. (2014) detected a  $^{13}\text{CO}$  outflow of which the blue-shifted lobe ( $\text{PA}_{\text{outflow,blue}}^{13\text{CO}} = -90^\circ$ ) coincides in position and velocity ( $+92.3 \text{ km s}^{-1} < V_{\text{blue}}^{13\text{CO}} < +101.3 \text{ km s}^{-1}$ ) with the  $\text{CH}_3\text{OH}$  masers (Fujisawa et al. 2014). The red-shifted lobe is instead oriented on the plane of the sky with an angle of  $\text{PA}_{\text{outflow,red}}^{13\text{CO}} = -27^\circ$  and its velocity range is  $+101.3 \text{ km s}^{-1} < V_{\text{red}}^{13\text{CO}} < +103.8 \text{ km s}^{-1}$ .

Vlemmings et al. (2011) detected circularly polarized emission of the 6.7 GHz  $\text{CH}_3\text{OH}$  maser with the 100 m Effelsberg telescope, which provided a Zeeman-splitting of  $0.81 \pm 0.10 \text{ m s}^{-1}$ , though few years earlier this emission was not detectable (Vlemmings 2008).

### 2.4. G28.31–0.39

G28.31–0.39 (also known as IRAS 18416–0420a) is a massive YSO, at a parallax distance of  $4.52^{+0.5}_{-0.4} \text{ kpc}$  (Li et al., in prep.), associated with the UC H II regions field G28.29–0.36 (Thompson et al. 2006). Walsh et al. (1997, 1998) detected 6.7 GHz  $\text{CH}_3\text{OH}$  maser emission coinciding with the center of an east-west sub-millimeter dust-emission (Walsh et al. 2003), but with no evident association with any of the H II regions (Thompson et al. 2006). Some of the  $\text{CH}_3\text{OH}$  maser features showed short-lived bursts suggesting a region of weak and diffuse maser emission, probably unsaturated, located far from the central core structure (Szymczak et al. 2018). A  $^{13}\text{CO}$  outflow with a  $\text{PA}_{\text{outflow}}^{13\text{CO}} = -52^\circ$  was detected at the position of the  $\text{CH}_3\text{OH}$  masers (de Villiers et al. 2014). The velocity range of the blue-shifted lobe ( $+80.4 \text{ km s}^{-1} < V_{\text{blue}}^{13\text{CO}} < +85.9 \text{ km s}^{-1}$ )

is consistent with the velocities of most of the  $\text{CH}_3\text{OH}$  masers (e.g., Walsh et al. 1998; Breen et al. 2015, 2016; Szymczak et al. 2018). The red-shifted lobe has a velocity range of  $+85.9 \text{ km s}^{-1} < V_{\text{blue}}^{13\text{CO}} < +88.9 \text{ km s}^{-1}$  (de Villiers et al. 2014). No  $\text{CH}_3\text{OH}$  maser polarization observations had been conducted until now.

### 2.5. G28.83–0.25

The extended green object (EGO) G28.83–0.25 is located at the edge of the mid-infrared bubble N49 at a kinematic distance of  $4.6 \pm 0.3 \text{ kpc}$  (Churchwell et al. 2006; Cyganowski et al. 2008; Green & McClure-Griffiths 2011). Two faint continuum radio sources have been identified at 3.6 cm with G28.83–0.25, named CM1 and CM2 (Cyganowski et al. 2011). CM2 is coincident with the linearly distributed 6.7 GHz  $\text{CH}_3\text{OH}$  masers ( $\text{PA}_{\text{CH}_3\text{OH}} \approx -45^\circ$ , Cyganowski et al. 2009; Fujisawa et al. 2014) and is surrounded by 44 GHz  $\text{CH}_3\text{OH}$  masers (Cyganowski et al. 2009). No 25 GHz  $\text{CH}_3\text{OH}$  maser emission has been detected (Towner et al. 2017). While the 44 GHz  $\text{CH}_3\text{OH}$  masers are located at the edges of a  $^{13}\text{CO}$ -bipolar outflow whose axis is oriented close to the line of sight, the 6.7 GHz  $\text{CH}_3\text{OH}$  masers coincide with the peak emission of the blue-shifted lobe of the outflow that is about  $5''$  southwest from the peak of the red-shifted lobe (de Villiers et al. 2014). Although the outflow is close to the line of sight, the small misalignment of the two lobes implies an orientation on the plane on the sky of  $\text{PA}_{\text{outflow}}^{13\text{CO}} = -40^\circ$  (de Villiers et al. 2014). The velocities of the lobes are  $+77.4 \text{ km s}^{-1} < V_{\text{blue}}^{13\text{CO}} < +88.4 \text{ km s}^{-1}$  and  $+88.4 \text{ km s}^{-1} < V_{\text{red}}^{13\text{CO}} < +96.4 \text{ km s}^{-1}$  (de Villiers et al. 2014). However, the velocity range of the 6.7 GHz and 12 GHz  $\text{CH}_3\text{OH}$  maser emissions agrees with that of the red-shifted lobe (e.g., Fujisawa et al. 2014; Breen et al. 2015, 2016).

Bayandina et al. (2015) detected both 1665 MHz and 1667 MHz OH masers at the same location, within the uncertainties, of the 6.7 GHz  $\text{CH}_3\text{OH}$  maser emission. By measuring the Zeeman splitting of the OH maser emissions they determined magnetic field strengths of 6.6 mG (1665 MHz OH) and of 5.1 mG (1667 MHz OH). No polarization observations of the  $\text{CH}_3\text{OH}$  maser emissions had been made until now.

### 2.6. G29.96–0.02 (W43 S)

G29.96–0.02 is a well-studied high-mass star-forming cloud (e.g., Cesaroni et al. 1994; De Buizer et al. 2002; Pillai et al. 2011; Beltrán et al. 2013) located in the massive SFR W43-South (W43 S) at a parallax distance of  $5.26^{+0.62}_{-0.50} \text{ kpc}$  (Zhang et al. 2014b). G29.96–0.02 contains a cometary UC H II region and a hot molecular core (HMC) located in front of the cometary arc (e.g., Wood & Churchwell 1989; Olmi et al. 2003; Beuther et al. 2007; Cesaroni et al. 1998, 2017).  $\text{H}_2\text{O}$ , OH, and  $\text{H}_2\text{CO}$  maser emissions (e.g., Hofner & Churchwell 1996; Hoffman et al. 2003; Breen & Ellingsen 2011; Caswell et al. 2013), and several  $\text{CH}_3\text{OH}$  maser lines (Minier et al. 2000, 2002; Breen et al. 2015, 2016) were detected towards the HMC. de Villiers et al. (2014) reported a  $^{13}\text{CO}$  outflow from the HMC with a  $\text{PA} = +50^\circ$ . The velocity range of the blue- (southeast) and red-shifted (northwest) lobes of the CO outflow are  $+92.1 \text{ km s}^{-1} < V_{\text{blue}}^{13\text{CO}} < +97.6 \text{ km s}^{-1}$  and  $+97.6 \text{ km s}^{-1} < V_{\text{red}}^{13\text{CO}} < +106.6 \text{ km s}^{-1}$  (de Villiers et al. 2014). At an angular resolution of 0.2 arcseconds, Cesaroni et al. (2017) were able to detect with the Atacama Large Millimeter/submillimeter Array (ALMA) a SiO bipolar jet ( $\text{PA} = -38^\circ$ )



**Table 1.** Observational details.

Source	Observation date	Calibrator	Polarization angle ( $^{\circ}$ )	Beam size (mas $\times$ mas)	Position Angle ( $^{\circ}$ )	rms ( $\frac{\text{mJy}}{\text{beam}}$ )	$\sigma_{\text{s-n}}^d$ ( $\frac{\text{mJy}}{\text{beam}}$ )	Estimated absolute position using FRMAP			
								$\alpha_{2000}$ ( $^{\text{h}} : ^{\text{m}} : ^{\text{s}}$ )	$\delta_{2000}$ ( $^{\circ} : ' : ''$ )	$\Delta\alpha^e$ (mas)	$\Delta\delta^e$ (mas)
G23.44–0.18	01 March 2014	J2202+4216 <sup>a</sup>	$-33 \pm 4$	$9.6 \times 4.2$	$-26.47$	3	31	+18:34:39.187	$-08:31:25.441$	0.8	5.4
G25.83–0.18	02 March 2014	J2202+4216 <sup>b</sup>	$-33 \pm 5$	$8.7 \times 5.5$	$-8.76$	4	43	+18:39:03.630	$-06:24:11.163$	0.8	6.0
G25.71–0.04	03 March 2014	J2202+4216 <sup>c</sup>	$-33 \pm 4$	$16.5 \times 6.5$	$+4.98$	3	9	+18:38:03.140	$-06:24:15.453$	1.7	12.3
G28.31–0.39	12 June 2014	J2202+4216 <sup>c</sup>	$-33 \pm 4$	$12.9 \times 3.0$	$-44.20$	3	13	+18:44:22.030	$-04:17:38.304$	0.6	6.3
G28.83–0.25	13 June 2014	J2202+4216 <sup>c</sup>	$-33 \pm 4$	$10.5 \times 3.5$	$-36.32$	4	10	+18:44:51.080	$-03:45:48.494$	0.4	6.4
G29.96–0.02	14 June 2014	J2202+4216 <sup>c</sup>	$-33 \pm 4$	$10.1 \times 3.2$	$-39.10$	4	55	+18:46:03.740	$-02:39:22.299$	0.3	7.8
G43.80–0.13	15 June 2014	J2202+4216 <sup>c</sup>	$-33 \pm 4$	$9.9 \times 3.4$	$-30.83$	7	16	+19:11:53.990	$+09:35:50.300$	0.3	1.0

**Notes.** <sup>(a)</sup>Calibrated using 3C286 ( $I = 0.39 \text{ Jy beam}^{-1}$ ,  $P_1 = 4.3\%$ ). <sup>(b)</sup>Calibrated using 3C286 ( $I = 0.49 \text{ Jy beam}^{-1}$ ,  $P_1 = 4.8\%$ ). <sup>(c)</sup>Calibrated using results from G23.44–0.18. <sup>(d)</sup>Self-noise in the maser emission channels (e.g., Sault 2012). When no circularly polarized emission is detected we consider the self-noise of the brightest maser feature. <sup>(e)</sup>Formal errors of the fringe rate mapping.

and a rotating disk perpendicular to it. Both are associated with the HMC. The velocity range of the SiO bipolar jet are  $+83.5 \text{ km s}^{-1} < V_{\text{blue}}^{\text{SiO}} < +91.6 \text{ km s}^{-1}$  and  $+102.4 \text{ km s}^{-1} < V_{\text{red}}^{\text{SiO}} < +110.5 \text{ km s}^{-1}$  (Cesaroni et al. 2017). The 6.7 GHz  $\text{CH}_3\text{OH}$  masers are associated with this system, which likely harbors a massive YSO of  $\sim 10 M_{\odot}$  (Sugiyama et al. 2008; Cesaroni et al. 2017).

A Zeeman-splitting of the 6.7 GHz  $\text{CH}_3\text{OH}$  maser line of  $\Delta V_Z = -0.33 \pm 0.11 \text{ m s}^{-1}$  was measured with the 100 m Effelsberg telescope (Vlemmings et al. 2011).

### 2.7. G43.80–0.13 (OH 43.8–0.1)

G43.80–0.13, better known as OH 43.8–0.1, is a massive star-forming region associated with an infrared source (IRAS 19095+0930) and an UCH II region at a parallax distance of  $6.0^{+0.19}_{-0.18} \text{ kpc}$  (Braz & Epchtein 1983; Kurtz et al. 1994; Wu et al. 2014). OH,  $\text{H}_2\text{O}$ , and  $\text{CH}_3\text{OH}$  masers have been detected at VLBI scale towards the UCH II region within the same velocity range (e.g., Fish et al. 2005; Sarma et al. 2008; Sugiyama et al. 2008). López-Sepulcre et al. (2010) detected an  $\text{HCO}^+$ -outflow oriented NE–SW ( $\text{PA}_{\text{outflow}}^{\text{HCO}^+} = +38^{\circ}$ ) with the blue-shifted lobe ( $+32.5 \text{ km s}^{-1} < V_{\text{blue}}^{\text{HCO}^+} < +38.5 \text{ km s}^{-1}$ ) and the red-shifted lobe ( $+49.0 \text{ km s}^{-1} < V_{\text{red}}^{\text{HCO}^+} < +53.5 \text{ km s}^{-1}$ ) directed towards the northeast and southwest, respectively.

Magnetic field strengths were measured at VLBI scale via Zeeman-splitting of both OH and  $\text{H}_2\text{O}$  masers. These are  $|B^{\text{OH}}| = 3.6 \text{ mG}$  and  $0.3 \text{ mG} < |B_{\parallel}^{\text{H}_2\text{O}}| < 22 \text{ mG}$  (Fish et al. 2005; Sarma et al. 2008). No polarization observations of  $\text{CH}_3\text{OH}$  maser emissions had been performed until now.

## 3. Observations and analysis

The second group of seven massive SFRs was observed in full polarization spectral mode at 6.7 GHz with eight of the EVN antennas (Ef, Jb, On, Mc, Nt, Tr, Wb, and Ys) between March and June 2014; the Medicina antenna (Mc) was not available in June 2014 (program code: ES072). The total observing time was 49 h. We covered a velocity range of  $\sim 100 \text{ km s}^{-1}$  by observing a bandwidth of 2 MHz. The correlation of the data was made with the EVN software correlator (SFXC; Keimpema et al. 2015) at the Joint Institute for VLBI ERIC (JIVE, the Netherlands) by using 2048 channels and generating all four polarization combinations (RR, LL, RL, LR) with a spectral resolution of  $\sim 1 \text{ kHz}$  ( $\sim 0.05 \text{ km s}^{-1}$ ). In Table 1 we present all the observational details. Here, the target sources and the date of the

observations are listed in Cols. 1 and 2, respectively; in Cols. 3 and 4 the polarization calibrators with their polarization angles are given. From Col. 5 to Col. 7 some of the image parameters are listed; in particular, the restoring beam size and corresponding position angle are in Cols. 5 and 6 and the thermal noise in Col. 7. In Col. 8 we also show the self-noise in the maser emission channels (see Paper III for details). Finally, the estimated absolute position of the reference maser and the FRMAP uncertainties are listed from Col. 9 to Col. 12 (see Paper III for details).

The Astronomical Image Processing Software package (AIPS) was used for calibrating and imaging the data. Following the same calibration procedure reported in Papers I–III, the bandpass, the delay, the phase, and the polarization calibration were performed on the calibrators listed in Col. 3 of Table 1. Fringe-fitting and self-calibration were subsequently performed on the brightest maser feature of each SFR that is identified as the reference maser feature in Tables A.1–A.7. The cubes of the four Stokes parameters ( $I$ ,  $Q$ ,  $U$ , and  $V$ ) were imaged using the AIPS task IMAGR. The polarized intensity ( $\text{POLI} = \sqrt{Q^2 + U^2}$ ) and polarization angle ( $\text{POLA} = 1/2 \times \text{atan}(U/Q)$ ) cubes were produced by combining the  $Q$  and  $U$  cubes. During the observations we observed a primary polarization calibrator (J2202+4216; Col. 3 of Table 1) and a well-known polarized calibrator (3C286). For G23.44–0.18 and G25.83–0.18 the signal-to-noise ratio of the 3C286 maps were so good that we were able to calibrate the polarization angle of J2202+4216 by using the calibration of 3C286. As expected, within the errors this was consistent with the constant polarization angle measured between 2005<sup>3</sup> and 2012<sup>4</sup>, i.e.,  $-31^{\circ} \pm 4^{\circ}$ . For the other five sources we assumed that the polarization angle of J2202+4216 did not change from March to June 2014. Therefore, we calibrated the linear polarization angles of the maser features by comparing the linear polarization angle of J2202+4216 measured by us with its angle obtained during the calibration of G23.44–0.18. The formal error on POLA due to the thermal noise is given by  $\sigma_{\text{POLA}} = 0.5(\sigma_{\text{P}}/\text{POLI}) \times (180^{\circ}/\pi)$  (Wardle & Kronberg 1974), where  $\sigma_{\text{P}}$  is the rms error of POLI.

As for Paper III, the observations were not performed in phase-referencing mode and the absolute position of the brightest maser feature of each source was estimated through fringe rate mapping (AIPS task FRMAP). The results and the formal errors of FRMAP are listed from Col. 9 to Col. 12 of Table 1. The phase fluctuations dominated the absolute positional

<sup>3</sup> <http://www.vla.nrao.edu/astro/calib/polar/>

<sup>4</sup> [http://www.aoc.nrao.edu/~smyers/evlapolcal/polcal\\_master.html](http://www.aoc.nrao.edu/~smyers/evlapolcal/polcal_master.html)

uncertainties and, from our experience with other experiments and varying the task parameters, we estimate that the absolute position uncertainties are on the order of a few mas.

The analysis of the polarimetric data followed the procedure given in Paper III and references therein. This consists in (1) identifying the individual CH<sub>3</sub>OH maser features by using the process described in Surcis et al. (2011b); (2) determining the mean linear polarization fraction ( $P_1$ ) and the mean linear polarization angle ( $\chi$ ) across the spectrum of each CH<sub>3</sub>OH maser feature; (3) modeling the total intensity and the linearly polarized spectrum of every maser feature for which we were able to detect linearly polarized emission by using the adapted FRTM code for 6.7 GHz CH<sub>3</sub>OH masers (Paper III and references therein); and (4) measuring the Zeeman splitting by including the results obtained from point (3) for fitting the total intensity and circularly polarized spectra of the corresponding CH<sub>3</sub>OH maser feature. While points (1) and (2) were applied as described and used in Paper III (we refer the reader there for more details), for points (3) and (4) we had to modify the adapted FRTM code as follows:

- a new subroutine for calculating the Clebsch–Gordan coefficients was implemented, this provides more accurate values than previously;
- an error in the part of the code used for quantifying the Zeeman-splitting was corrected. We found that the error led us to overestimate the Zeeman-splitting for the massive YSOs W51-e2, W48, IRAS 06058+2138-IRS1, S255-IR, IRAS 20126+4104, G24.78+0.08, G29.86–0.04, and G213.70–12.6, and to underestimate for W3(OH) (in three-quarters of the cases the corrections are less than a factor of 2 than those previously reported; Papers I–III; Surcis et al. 2014b).

The outputs of the FRTM code are the emerging brightness temperature ( $T_b\Delta\Omega$ ), the intrinsic thermal linewidth ( $\Delta V_i$ ), and the angle between the magnetic field and the maser propagation direction ( $\theta$ ). If  $\theta > \theta_{\text{crit}} = 55^\circ$ , where  $\theta_{\text{crit}}$  is the Van Vleck angle, the magnetic field appears to be perpendicular to the linear polarization vectors; otherwise, it is parallel (Goldreich et al. 1973). The fits and models are performed for a sum of the decay and cross-relaxation rates of  $\Gamma = 1 \text{ s}^{-1}$ . Since the emerging brightness temperature scales linearly with  $\Gamma$  the fitted  $T_b\Delta\Omega$  values can be adjusted by simply scaling according to  $\Gamma$ . Furthermore, when fitting the observed polarized CH<sub>3</sub>OH maser features we restricted our analysis to  $10^6 \text{ K sr} < T_b\Delta\Omega < 10^{11} \text{ K sr}$  and to the most plausible range  $0.5 \text{ km s}^{-1} < \Delta V_i < 2.5 \text{ km s}^{-1}$ , in steps of  $0.05 \text{ km s}^{-1}$ , because it takes a prohibitively long time to fit for smaller  $\Delta V_i$  values.

As in Paper III, we consider a detection of circularly polarized emission to be real only when the detected  $V$  peak flux of a maser feature is both  $V > 5 \text{ rms}$  and  $V > 3 \cdot \sigma_{\text{s.-n.}}$ , where  $\sigma_{\text{s.-n.}}$  is the self-noise<sup>5</sup> produced by the maser (Col. 8 of Table 1; e.g., Sault 2012). From the Zeeman effect theory we know that  $\Delta V_Z = \alpha_Z \cdot B_{\parallel}$ , where  $\Delta V_Z$  is the Zeeman-splitting,  $B_{\parallel}$  is the magnetic field strength along the line of sight, and  $\alpha_Z$  is the Zeeman-splitting coefficient that depends on the Landé g-factor(s). Following Lankhaar et al. (2018), who identify for the 6.7 GHz methanol maser transition the hyperfine transition with the largest Einstein coefficient for stimulated emission, i.e.,  $F = 3 \rightarrow 4$ , as more favored among the eight hyperfine transitions that might contribute to the maser line, we estimated  $B_{\parallel}$  by assuming  $\alpha_Z = -0.051 \text{ km s}^{-1} \text{ G}^{-1}$  ( $\alpha_Z = -1.135 \text{ Hz mG}^{-1}$ ,

Lankhaar et al. 2018). Considering that  $\alpha_Z = \mu_N \cdot g_1$ , where  $\mu_N$  is the nuclear magneton and  $g_1$  is the Landé g-factor, and because  $g_1$  for  $F = 3 \rightarrow 4$  is the largest one among the eight hyperfine transitions, our estimate of  $B_{\parallel}$  is therefore a lower limit. Even in the case of a combination of hyperfine components the derived  $B_{\parallel}$  would be higher (e.g., Lankhaar et al. 2018).

## 4. Results

In Sects. 4.1–4.7 the 6.7 GHz CH<sub>3</sub>OH maser distribution and the polarization results for each of the seven massive SFRs observed with the EVN are reported separately. The lists of all the maser features, with their properties, can be found in Tables A.1–A.7.

### 4.1. G23.44–0.18

We detected a total of 61 CH<sub>3</sub>OH maser features, named G23.E01–G23.E61 in Table A.1, towards both MM1 (27/61) and MM2 (34/61), see Fig. 1, with the strongest maser features associated with MM1. We were able to detect the weak maser features (e.g., G23.E29 and G23.E30) with velocities  $> +111 \text{ km s}^{-1}$  previously detected with the Australia Telescope Compact Array (ATCA) by Breen et al. (2015). The maser features distribution around the two cores are similar to what was observed by Fujisawa et al. (2014). The maser features associated with MM2 are distributed along two lines separated by about 250 mas and with position angles of  $\text{PA}_{\text{MM2}}^{\text{east}} = +30^\circ$  and  $\text{PA}_{\text{MM2}}^{\text{west}} = -24^\circ$  (see right panel of Fig. 1). No velocity gradient is measured along them.

Only linearly polarized emission has been detected, in particular from six maser features associated with MM1 ( $P_1^{\text{MM1}} = 1.1\text{--}6.5\%$ ) and from five associated with MM2 ( $P_1^{\text{MM2}} = 1.4\text{--}8.3\%$ ). The FRTM code (see Sect. 3) was able to fit all of them and the outputs of the code are reported in Cols. 9, 10, and 14 of Table A.1. We note that due to their high  $P_1$  the features G23.E25 and G23.E51 might be partially saturated. The estimated  $\theta$  angles indicate that the magnetic field is perpendicular to the linear polarization vectors.

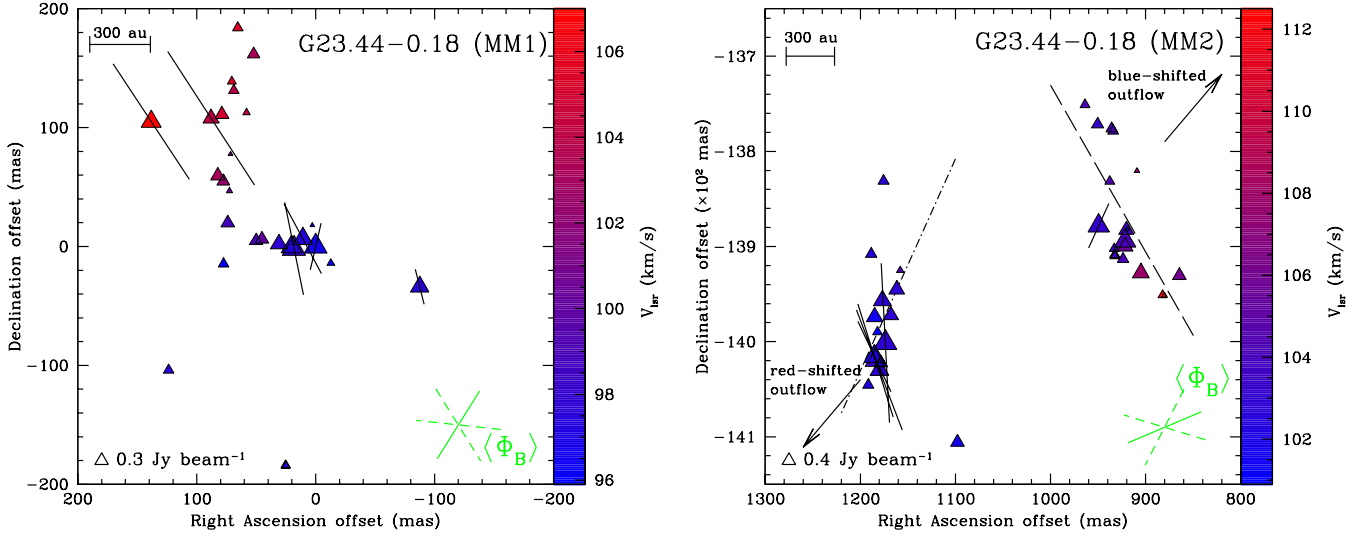
Considering the rms and the  $\sigma_{\text{s.-n.}}$  (see Table 1) for the strongest maser feature (G23.E03) we would have been able to detect circularly polarized emission only if  $P_V > 1\%$ , which is twice the typical fraction ( $P_V \approx 0.5\%$ ; e.g., Paper III).

### 4.2. G25.83–0.18

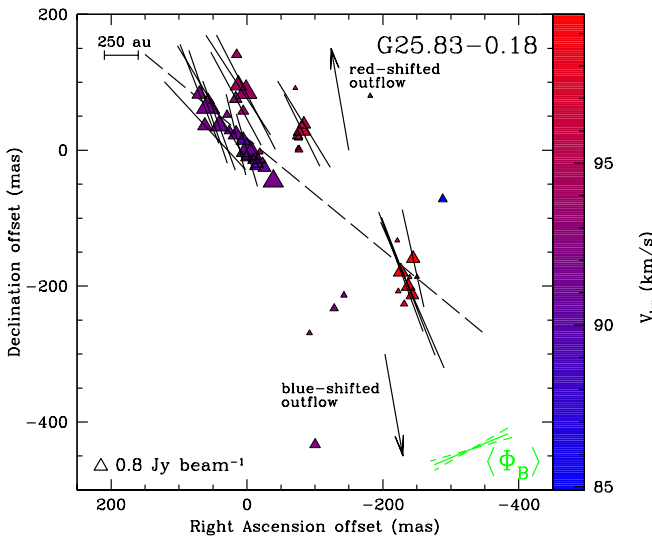
In Table A.2 we report the 46 CH<sub>3</sub>OH maser features detected towards G25.83–0.18 and named G258.E01–G258.E46. The very first VLBI distribution of these maser features is shown in Fig. 2, where we also report the red- and blue-shifted orientation, but not at the actual position, of the <sup>13</sup>CO outflow (de Villiers et al. 2014). The maser features, in particular the northwestern group, are aligned NW–SE ( $\text{PA}_{\text{CH}_3\text{OH}} = +51^\circ \pm 7^\circ$ ) without a clear velocity gradient. Their velocities, except for G258.E01 and G258.E28, are consistent with the red-shifted lobe of the outflow.

Among the seven massive SFRs studied in the present work, G25.83–0.18 has the highest number of linearly polarized features (16 out of 46). The fractional linear polarization ranges from 2.3% to 9.7%, which is one of the highest measured so far. Although the high  $P_1$  indicates that five maser features might be partially saturated, we were able to fit all of the polarized maser features with the FRTM code. From the estimated  $\theta$  angles (Col. 14 of Table A.2) the magnetic field is oriented

<sup>5</sup> The self-noise is high when the power contributed by the astronomical maser is a significant portion of the total received power (Sault 2012).



**Fig. 1.** View of the CH<sub>3</sub>OH maser features detected around G23.44–0.18 MM1 (*left panel*) and MM2 (*right panel*). The reference position is the estimated absolute position from Table 1. Triangles identify CH<sub>3</sub>OH maser features whose side length is scaled logarithmically according to their peak flux density (Table A.1). Maser local standard of rest radial velocities are indicated by color (the assumed velocity of the region is  $V_{\text{lsr}}^{\text{CS}(2-1)} = +104.2 \text{ km s}^{-1}$ , Bronfman et al. 1996).  $0.3 \text{ Jy beam}^{-1}$  and  $0.4 \text{ Jy beam}^{-1}$  symbols are plotted for comparison. The linear polarization vectors, scaled logarithmically according to the polarization fraction  $P_1$  ( $P_1^{\text{MM1}} = 1.1\text{--}6.5\%$  and  $P_1^{\text{MM2}} = 1.4\text{--}8.3\%$ ; see Table A.1), are overplotted. In the bottom right corner the corresponding error-weighted orientation of the magnetic field ( $\langle\Phi_B\rangle$ , see Sect. 5.1) is also shown, the two dashed segments indicate the uncertainty. The two arrows in the *right panel* indicate the direction, but not the actual position, of the red- and blue-shifted lobes of the  $^{12}\text{CO}(2-1)$  outflow associated with G23.44–0.18 (MM2) ( $\text{PA}_{\text{outflow}}^{12\text{CO}} = -40^\circ$ ; Ren et al. 2011). The dashed and dash-dotted lines are the best least-squares linear fit of the western ( $\text{PA}_{\text{CH}_3\text{OH}}^{\text{west}} = +30^\circ \pm 26^\circ$ ) and eastern ( $\text{PA}_{\text{CH}_3\text{OH}}^{\text{east}} = -24^\circ \pm 90^\circ$ ) groups of CH<sub>3</sub>OH maser features in G23.44+0.18 (MM2).



**Fig. 2.** View of the CH<sub>3</sub>OH maser features detected around G25.83–0.18 (Table A.2). Symbols are the same as in Fig. 1. The polarization fraction is in the range  $P_1 = 2.3\text{--}9.7\%$  (Table A.2). The assumed velocity of the YSO is  $V_{\text{lsr}}^{13\text{CO}} = +93.2 \text{ km s}^{-1}$  (de Villiers et al. 2014). The two arrows indicate the direction, but not the actual position, of the red- and blue-shifted lobes of the bipolar outflow ( $\text{PA}_{\text{outflow}}^{13\text{CO}} = +10^\circ$ ; de Villiers et al. 2014). The dashed line is the best least-squares linear fit of the CH<sub>3</sub>OH maser features ( $\text{PA}_{\text{CH}_3\text{OH}} = +51^\circ \pm 7^\circ$ ).

perpendicular to the linear polarization vectors. No circular polarization was measured ( $P_V < 0.9\%$ ).

#### 4.3. G25.71–0.04

We identified 26 CH<sub>3</sub>OH maser features (named G257.E01–G257.E26 in Table A.3). Their complex distribution is shown

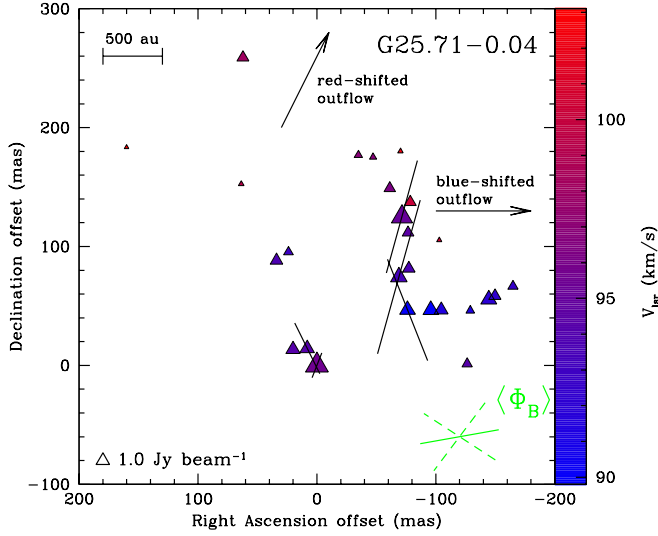
in Fig. 3. Here, the orientation of the red- and blue-shifted lobes of the  $^{13}\text{CO}$  outflow are also displayed (de Villiers et al. 2014). The velocity range of 24 out of 26 CH<sub>3</sub>OH maser features ( $89 \text{ km s}^{-1} < V_{\text{lsr}} < 100 \text{ km s}^{-1}$ ) is within the velocity range of the blue-shifted lobe ( $+92.3 \text{ km s}^{-1} < V_{\text{blue}}^{13\text{CO}} < +101.3 \text{ km s}^{-1}$ ; de Villiers et al. 2014), confirming their association with it.

We detected linearly polarized emission towards five CH<sub>3</sub>OH maser features ( $0.4\% < P_1 < 1.6\%$ ), among which G257.E12 also showed circular polarization ( $P_V = 0.8\%$ ). The error-weighted intrinsic thermal linewidth,  $\langle\Delta V_i\rangle = 2.0^{+0.2}_{-0.4} \text{ km s}^{-1}$ , is one of the largest ever provided by the FRTM code, while the  $T_b\Delta\Omega$  values are within the typical estimated values for the 6.7 GHz CH<sub>3</sub>OH maser emission. Following Paper III we found for the features G257.E13 (offset =  $-71.456 \text{ mas}$ ;  $125.005 \text{ mas}$ ) and G257.E19 ( $0 \text{ mas}$ ;  $0 \text{ mas}$ ) that the magnetic field is parallel to the linear polarization vectors since  $|\theta^+ - 55^\circ| < |\theta^- - 55^\circ|$ , where  $\theta^\pm = \theta \pm \epsilon^\pm$ , and with  $\epsilon^\pm$  the errors associated with  $\theta$ . For all the other maser features the magnetic field is perpendicular. The circularly polarized emission of G257.E12 was fitted with the FRTM code by including the corresponding  $\Delta V_i$  and  $T_b\Delta\Omega$  fit values obtained from the total and linearly polarized intensities. The fitted result is reported in Table A.3 and shown in Fig. 4. The nondetection of circular polarization for the very bright maser features G257.E13 and G257.E19 is likely due to the high  $\sigma_{\text{s-n}}$  ( $55$  and  $97 \text{ mJy beam}^{-1}$ , respectively). Considering the criterion  $3\sigma_{\text{s-n}}$ , we have upper limits of  $P_V < 0.8\%$  and  $< 0.6\%$  for G257.E13 and G257.E19, respectively.

#### 4.4. G28.31–0.39

Towards G28.31–0.39, we detected 13 CH<sub>3</sub>OH maser features (named G283.E01–G283.E13 in Table A.4). The  $V_{\text{lsr}}$  of eleven features ranges from  $79.91 \text{ km s}^{-1}$  to  $83.23 \text{ km s}^{-1}$ ,





**Fig. 3.** View of the  $\text{CH}_3\text{OH}$  maser features detected around G25.71–0.04 (Table A.3). Symbols are the same as in Fig. 1. The polarization fraction is in the range  $P_1 = 0.4\text{--}1.6\%$  (Table A.3). The assumed velocity of the YSO is  $V_{\text{lsr}}^{13\text{CO}} = 101.3 \text{ km s}^{-1}$  (de Villiers et al. 2014). The two arrows indicate the direction, and not the actual position, of the red- and blue-shifted lobe of the bipolar outflow ( $\text{PA}_{\text{outflow,blue}}^{13\text{CO}} = -90^\circ$  and  $\text{PA}_{\text{outflow,red}}^{13\text{CO}} = -27^\circ$ , de Villiers et al. 2014).

in accordance with  $+80.4 \text{ km s}^{-1} < V_{\text{lsr}}^{13\text{CO}} < +85.9 \text{ km s}^{-1}$  (de Villiers et al. 2014). The maser features G283.E12 and G283.E13 have higher velocities,  $V_{\text{lsr}} = 92.72 \text{ km s}^{-1}$  and  $V_{\text{lsr}} = 93.81 \text{ km s}^{-1}$ , respectively, which are close to the highest velocity of the red-shifted lobe of the outflow ( $\sim 89 \text{ km s}^{-1}$ ). In Fig. 5 the complex  $3200 \text{ au} \times 2400 \text{ au}$  distribution of the maser features, which resembles an X, is shown.

The fit with the FRTM code of the two linearly polarized maser features, G283.E03 and G283.E10, provided for both of them that  $\theta > 55^\circ$ , i.e., that the magnetic field is perpendicular to the linear polarization vectors. Circular polarization has not been detected ( $P_V < 0.5\%$ ).

#### 4.5. G28.83–0.25

At VLBI scales we detected 21 6.7 GHz  $\text{CH}_3\text{OH}$  maser features (named G288.E01–G288.E21 in Table A.5) linearly distributed ( $\text{PA}_{\text{CH}_3\text{OH}} = -41^\circ \pm 10^\circ$ ; see Fig. 6) from southwest (the most red-shifted) to northeast (around systemic velocity  $V_{\text{lsr}}^{13\text{CO}} = +88.4 \text{ km s}^{-1}$ ; de Villiers et al. 2014). The maser features at the center of the linear distribution show the most blue-shifted velocities, in accordance with Cyganowski et al. (2009) and Fujisawa et al. (2014). The velocity distribution of the maser features reflects an almost perfect overlap of the red- and blue-shifted lobe emissions of the  $^{13}\text{CO}$  outflow (de Villiers et al. 2014). Therefore, it is difficult to associate the maser features with either the outflow or an accretion disk.

We measured fractional linear polarization between 0.5% and 3.3% from six  $\text{CH}_3\text{OH}$  maser features. For the brightest maser feature G288.E16, by fitting its polarized emission with the FRTM code, we found that  $\theta = 61^{+3}_{-47}$ . This implies that the magnetic field is more likely parallel to the linear polarization vector of G288.E16 (see Paper III). For the other maser features this is instead perpendicular. Furthermore, we detected circularly polarized emission ( $P_V = 0.6\%$ ) from G288.E19 that does not show linearly polarized emission. Hence, to model

the circularly polarized emission we assumed that its emerging brightness temperature is equal to the error-weighted value  $\langle T_b \Delta \Omega \rangle = 9.4 \times 10^8 \text{ K sr}$  of the region; for the intrinsic thermal linewidth we determined that  $\Delta V_i = 1.1 \text{ km s}^{-1}$  is the value that best fits the total intensity emission (see right panel of Fig. 4). The  $\sigma_{\text{s-n}}$  for the bright maser features G288.E16 and G288.E18 are 81 and 64  $\text{mJy beam}^{-1}$ , which imply  $P_V < 0.8\%$  and  $< 1.0\%$ , respectively.

#### 4.6. G29.96–0.02

We detected 34 6.7 GHz  $\text{CH}_3\text{OH}$  maser features (named G29.E01–G29.E34 in Table A.6) towards the rotating disk of G29.96–0.02, distributed perpendicularly to the SiO jet from northeast to southwest ( $\text{PA}_{\text{CH}_3\text{OH}} = +80^\circ \pm 3^\circ$ , Fig. 7). Two of them (G29.E33 and G29.E34), the most northeastern, were previously undetected at VLBI scales (Sugiyama et al. 2008). The velocities of the maser features ranges from  $95.65 \text{ km s}^{-1}$  to  $105.75 \text{ km s}^{-1}$  without an ordered distribution along the major axis of the rotating disk. Even though the association of most of the maser features with the rotating disk is plausible, some of them might be associated with the outflowing gas. This hypothesis can be verified only by measuring their proper motions.

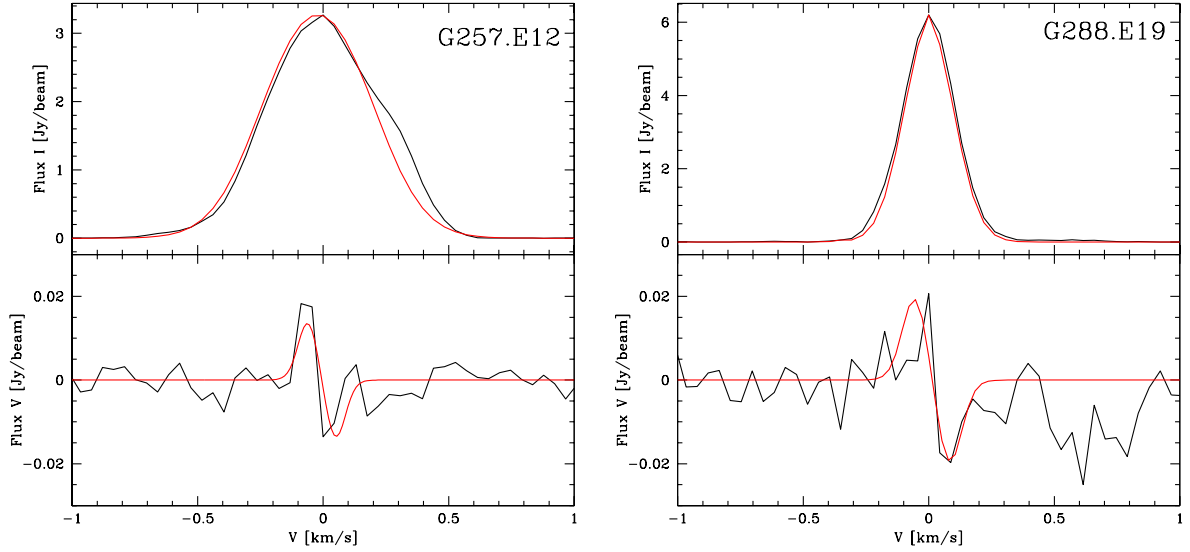
Linearly polarized emission ( $P_1 = 1.4\text{--}5.7\%$ ) was measured towards four maser features grouped two by two: G29.E09 (the brightest) and G29.E12 towards the west and G29.E26 and G29.E30 towards the east. The west group shows  $P_1 > 4\%$  and an error-weighted linear polarization angle of  $\langle \chi \rangle_{\text{west}} = +76^\circ \pm 1^\circ$ , while the east group shows  $P_1 \leq 2\%$  and  $\langle \chi \rangle_{\text{east}} = +49^\circ \pm 6^\circ$ . The total error-weighted linear polarization angle is  $\langle \chi \rangle = +62^\circ \pm 17^\circ$ . The FRTM code properly fits all of them and provided consistent outputs, though the  $T_b \Delta \Omega$  is higher, as expected, for the west group, which might be entering the saturated state. The angle between the magnetic field and the maser propagation direction is greater than  $55^\circ$  for all four maser features indicating that the magnetic field is perpendicular to the linear polarization vectors. No circularly polarized emission was detected ( $P_V < 0.3\%$ ).

#### 4.7. G43.80–0.13

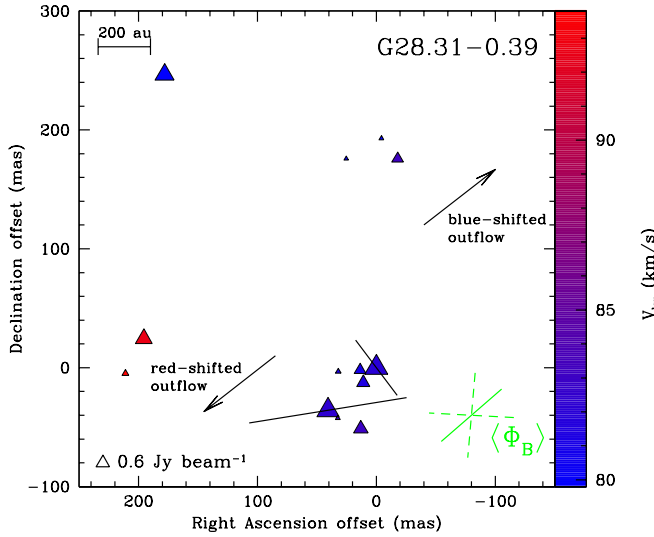
With the EVN we were able to detect twice the number of 6.7 GHz  $\text{CH}_3\text{OH}$  maser features previously detected by Sugiyama et al. (2008). These maser features are listed as G43.E01–G43.E18 in Table A.7, and are shown in Fig. 8 where the perpendicularity of their linear distribution ( $\text{PA}_{\text{CH}_3\text{OH}} = -48^\circ \pm 5^\circ$ ) to the  $\text{HCO}^+$  outflow is clearly seen. The velocity range of the maser features ( $39.5 \text{ km s}^{-1} < V_{\text{lsr}} < 43.2 \text{ km s}^{-1}$ ) is blue-shifted with respect to the systemic velocity of the region ( $V_{\text{lsr}}^{18\text{O}} = +43.9 \text{ km s}^{-1}$ , López-Sepulcre et al. 2010) and it is consistent with the range reported in Sugiyama et al. (2008). In particular, the velocities of the maser features located to the southeast (G43.E15–G43.E18) are closer to the systemic velocity, and the most blue-shifted are located towards the northeast.

Three maser features out of 18 showed linearly polarized emission with  $1\% < P_1 < 4.5\%$ , i.e., they are likely unsaturated. Despite this finding, the FRTM code was able to provide only an upper limit of  $\Delta V_i$  for G43.E04. The estimated  $\theta$  angles indicate that the magnetic field is oriented perpendicular to the linear polarization vectors. No circularly polarized emission was detected, likely due to the weakness of the maser features ( $I < 6 \text{ Jy beam}^{-1}$ ;  $P_V < 0.8\%$ ).

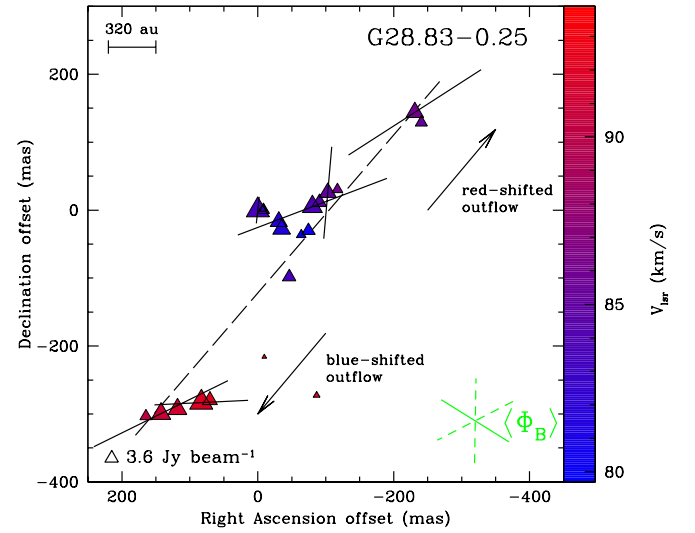




**Fig. 4.** Total intensity ( $I$ , upper panel) and circularly polarized intensity ( $V$ , lower panel) spectra for the  $\text{CH}_3\text{OH}$  maser features named G257.E12 and G288.E19 (see Tables A.3, A.5). The thick red lines are the best-fit models of  $I$  and  $V$  emissions obtained using the adapted FRTM code (see Sect. 3). The maser features were centered on zero velocity.



**Fig. 5.** View of the  $\text{CH}_3\text{OH}$  maser features detected around G28.31–0.39 (Table A.4). Symbols are the same as in Fig. 1. The polarization fraction is in the range  $P_1 = 1.6$ –4.3% (Table A.4). The assumed velocity of the YSO is  $V_{\text{lsr}}^{13\text{CO}} = +85.9 \text{ km s}^{-1}$  (de Villiers et al. 2014). The two arrows indicate the direction, and not the actual position, of the red- and blue-shifted lobe of the bipolar outflow ( $\text{PA}_{\text{outflow}}^{13\text{CO}} = -52^\circ$ , de Villiers et al. 2014).



**Fig. 6.** View of the  $\text{CH}_3\text{OH}$  maser features detected around G28.83–0.25 (Table A.5). Symbols are the same as in Fig. 1. The polarization fraction is in the range  $P_1 = 0.5$ –3.3% (Table A.5). The assumed velocity of the YSO is  $V_{\text{lsr}}^{13\text{CO}} = +88.4 \text{ km s}^{-1}$  (de Villiers et al. 2014). The dashed line is the best least-squares linear fit of the  $\text{CH}_3\text{OH}$  maser features ( $\text{PA}_{\text{CH}_3\text{OH}} = -41^\circ \pm 10^\circ$ ). The two arrows indicate the direction, and not the actual position, of the red- and blue-shifted lobe of the bipolar outflow ( $\text{PA}_{\text{outflow}}^{13\text{CO}} = -40^\circ$ , de Villiers et al. 2014).

## 5. Discussion

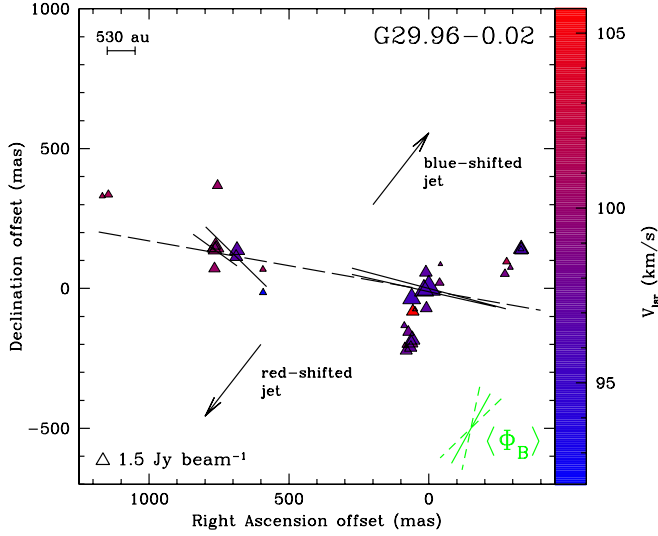
### 5.1. Magnetic field orientations

If the linearly polarized emission passes through a medium where a magnetic field is present before reaching the observer, its linear polarization vector suffers a rotation known as Faraday rotation. In the case of polarized  $\text{CH}_3\text{OH}$  maser emission, two main Faraday rotations can affect the linear polarization vectors that we measure: internal rotation ( $\Phi_i$ ) and foreground Faraday rotation ( $\Phi_f$ ). An analysis of these two components of Faraday rotation was undertaken in the previous papers in the series (Papers I, II, and III); this analysis helps understand why we do not consider them important here. However, in Col. 2 of Table 4 we list  $\Phi_f$  for each source.

The estimated orientation of the magnetic field<sup>6</sup> in the seven massive SFRs under investigation here are separately discussed below.

**G23.44–0.18.** The magnetic fields around MM1 and MM2 are oriented SE–NW with error-weighted orientation of

<sup>6</sup>  $\langle \Phi_B \rangle$  is the mean error-weighted orientation of the magnetic field measured considering all the magnetic field vectors measured in a source. The weights are  $1/e_i$ , where  $e_i$  is the error of the  $i$ th measured vector. The error on  $\langle \Phi_B \rangle$  is the standard deviation. The position angle of the magnetic field vectors are measured with respect to the north, clockwise (negative) and counterclockwise (positive), as the PA of the outflows.



**Fig. 7.** View of the CH<sub>3</sub>OH maser features detected around G29.96–0.02 (Table A.6). Symbols are the same as in Fig. 1. The polarization fraction is in the range  $P_1 = 1.4\text{--}5.7\%$  (Table A.6). The assumed velocity of the YSO is  $V_{\text{lsr}}^{\text{CO}} = +97.6 \text{ km s}^{-1}$  (de Villiers et al. 2014). The dashed line is the best least-squares linear fit of the CH<sub>3</sub>OH maser features ( $\text{PA}_{\text{CH}_3\text{OH}} = +80^\circ \pm 3^\circ$ ). The two arrows indicate the direction, and not the actual position, of the red- and blue-shifted lobe of the bipolar jet ( $\text{PA}_{\text{jet}}^{\text{SiO}} = -38^\circ$ ; Cesaroni et al. 2017).

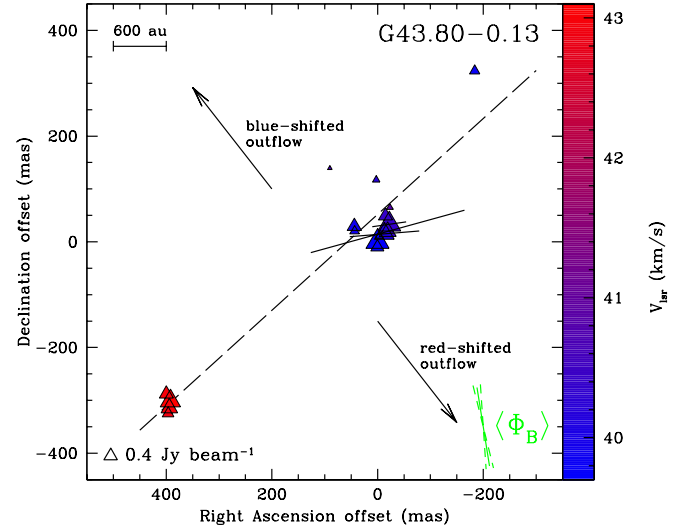
$\langle \Phi_B^{\text{MM1}} \rangle = -32^\circ \pm 64^\circ$  and  $\langle \Phi_B^{\text{MM2}} \rangle = -67^\circ \pm 40^\circ$ . It should be noted that the CO outflow on the plane of the sky is oriented with an angle of  $\text{PA}_{\text{outflow}}^{\text{CO}} = -40^\circ$  and shows an opening angle of  $\sim 30^\circ$  (Ren et al. 2011). However, the velocity range of the masers both in MM1 and in MM2 falls between the velocity ranges of the LVC and HVC CO outflows indicating that the masers are not associated with either of them. The eastern maser group of MM2 shows a linear distribution perpendicular to the CO outflow (see Table 4) suggesting a possible disk structure, though the velocity distribution of the masers does not.

**G25.83–0.18.** The magnetic field is oriented with an angle on the plane of the sky of  $\langle \Phi_B \rangle = -67^\circ \pm 7^\circ$ , which is almost perpendicular both to the linear distribution of the 6.7 GHz CH<sub>3</sub>OH masers ( $\text{PA}_{\text{CH}_3\text{OH}} = +51^\circ \pm 7^\circ$ ) and to the <sup>13</sup>CO outflow ( $\text{PA}_{\text{outflow}}^{\text{CO}} \sim +10^\circ$ , de Villiers et al. 2014).

**G25.71–0.04.** Taking into account the different orientation of the magnetic field with respect to the linear polarization vectors of the maser features, we measured an error-weighted magnetic field orientation of  $\langle \Phi_B \rangle = -80^\circ \pm 43^\circ$ , implying that it is aligned with the blue-shifted lobe of the <sup>13</sup>CO outflow ( $\text{PA}_{\text{outflow}}^{\text{CO}} = -90^\circ$ ; de Villiers et al. 2014). The location and velocities of the CH<sub>3</sub>OH masers also suggest that the masers probe the magnetic field along the blue-shifted lobe of the outflow.

**G28.31–0.39.** We were able to determine the orientation of the magnetic field on the plane of the sky from two blue-shifted CH<sub>3</sub>OH masers. Also in this case, the magnetic field ( $\langle \Phi_B \rangle = -50^\circ \pm 44^\circ$ ) is aligned with the <sup>13</sup>CO outflow ( $\text{PA}_{\text{outflow}}^{\text{CO}} = -52^\circ$ , de Villiers et al. 2014).

**G28.83–0.25.** Considering the different orientation of the magnetic field with respect to the linear polarization vector of G288.E16, the error-weighted orientation of the magnetic field on the plane of the sky is  $\langle \Phi_B \rangle = +58^\circ \pm 59^\circ$ . This is perpendicular to the linear distribution of the masers ( $\text{PA}_{\text{CH}_3\text{OH}} =$



**Fig. 8.** View of the CH<sub>3</sub>OH maser features detected around G43.80–0.13 (Table A.7). Symbols are the same as in Fig. 1. The polarization fraction is in the range  $P_1 = 1.1\text{--}4.4\%$  (Table A.7). The assumed velocity of the YSO is  $V_{\text{lsr}}^{\text{CO}} = +43.9 \text{ km s}^{-1}$  (López-Sepulcre et al. 2010). The dashed line is the best least-squares linear fit of the CH<sub>3</sub>OH maser features ( $\text{PA}_{\text{CH}_3\text{OH}} = -48^\circ \pm 5^\circ$ ). The two arrows indicate the direction, and not the actual position, of the red- and blue-shifted lobe of the bipolar outflow ( $\text{PA}_{\text{outflow}}^{\text{HCO}^+} = +38^\circ$ ; López-Sepulcre et al. 2010).

**Table 2.**  $|B_{\text{crit}}|$  values for the ranges of  $n_{\text{H}_2}$  and  $T_k$ .

$\beta = 1$			
$T_k = 100 \text{ K}$		$T_k = 200 \text{ K}$	
$n_{\text{H}_2} = 10^5 \text{ cm}^{-3}$	$n_{\text{H}_2} = 10^9 \text{ cm}^{-3}$	$n_{\text{H}_2} = 10^5 \text{ cm}^{-3}$	$n_{\text{H}_2} = 10^9 \text{ cm}^{-3}$
0.2 mG	18 mG	0.3 mG	25 mG

$-41^\circ \pm 10^\circ$ ) and to the <sup>13</sup>CO outflow ( $\text{PA}_{\text{CH}_3\text{OH}} = -40^\circ$ ), even though the outflow is almost along the line of sight. Therefore, the interpretation of the morphology of the magnetic field is not straightforward.

**G29.96–0.02.** The magnetic field is oriented on the plane of the sky along the SiO jet ( $\text{PA}_{\text{jet}}^{\text{SiO}} = -38^\circ$ , Cesaroni et al. 2017) with an angle of  $\langle \Phi_B \rangle = -29^\circ \pm 17^\circ$  and perpendicular to the rotating disk (Cesaroni et al. 2017). Considering separately the east and west group ( $\langle \Phi_B \rangle_{\text{east}} = -41^\circ \pm 6^\circ$  and  $\langle \Phi_B \rangle_{\text{west}} = -14^\circ \pm 1^\circ$ ), we note that the morphology of the magnetic field coincides with the morphology of the red-shifted lobe of the SiO jet near to the massive YSO (see Fig. 13 of Cesaroni et al. 2017).

**G43.80–0.13.** The orientation of the magnetic field ( $\langle \Phi_B \rangle = +9^\circ \pm 5^\circ$ ) derived from the masers, all of which belong to the most blue-shifted group, is misaligned by about  $30^\circ$  compared to the HCO<sup>+</sup> outflow ( $\text{PA}_{\text{outflow}}^{\text{HCO}^+} = +38^\circ$ ; López-Sepulcre et al. 2010).

## 5.2. Magnetic field strength

Thanks to the work of Lankhaar et al. (2018), we are able to estimate a lower limit of  $B_{\parallel}$  from the measurements of the Zeeman splitting of the 6.7 GHz CH<sub>3</sub>OH maser by assuming  $\alpha_Z = -0.051 \text{ km s}^{-1} \text{ G}^{-1}$  (Lankhaar et al. 2018). A direct measurement of  $B_{\parallel}$  can indeed be only possible when the contribution of all eight hyperfine transitions to the 6.7 GHz CH<sub>3</sub>OH

**Table 3.** Magnetic fields measurements around massive YSOs determined by observing, with the EVN, the circularly polarized emission of 6.7 GHz CH<sub>3</sub>OH masers.

(1) Source name	(2) Maser ID	(3) Peak flux Density (I) (Jy beam <sup>-1</sup> )	(4) $V_{lsr}$ (km s <sup>-1</sup> )	(5) $\Delta V_i^a$ (km s <sup>-1</sup> )	(6) $T_b \Delta \Omega^a$ (log K sr)	(7) $\theta$ (°)	(8) $P_V$ (%)	(9) $\Delta V_Z$ (m s <sup>-1</sup> )	(10) $ B_{  } ^b$ (mG)	(11) $ B ^c$ (mG)	Ref.
W75N(B)	A4	47.58	+5.82	–	–	–	0.5	+0.80 ± 0.03 <sup>d</sup>	>16 <sup>d</sup>	–	(1)
	A5	39.39	+5.12	–	–	–	0.5	+0.75 ± 0.13 <sup>d</sup>	>15 <sup>d</sup>	–	(1)
	B1	95.38	+7.23	–	–	–	0.4	+0.53 ± 0.04 <sup>d</sup>	>10 <sup>d</sup>	–	(1)
NGC 7538	E02	16.82	–50.49	1.0	9.41	77 <sup>+13</sup> <sub>-11</sub>	1.7	+2.7 ± 0.3 <sup>d</sup>	>53 <sup>d</sup>	>104	(2)
	E26	95.15	–55.92	0.5	11.39	86 <sup>+5</sup> <sub>-19</sub>	1.0	+1.6 ± 0.3 <sup>d</sup>	>31 <sup>d</sup>	>62	(2)
	E48	23.68	–58.03	1.2	9.26	75 <sup>+15</sup> <sub>-35</sub>	1.7	–2.7 ± 0.3 <sup>d</sup>	>53 <sup>d</sup>	>65	(2)
W51-e2	W51E.13	65.78	+57.86	1.7	8.9	77 <sup>+13</sup> <sub>-37</sub>	0.4	–0.9 ± 0.2	>17	>21	(3)
	W51E.14	217.32	+59.26	1.7	8.8	67 <sup>+11</sup> <sub>-42</sub>	0.1	+0.26 ± 0.06	>5	>5	(3)
	W51E.18	27.56	+59.35	–	–	–	1.1	–2.1 ± 0.5 <sup>e</sup>	>42 <sup>e</sup>	>51 <sup>e</sup>	(3)
	W51E.32	8.11	+57.86	0.6	8.8	75 <sup>+14</sup> <sub>-35</sub>	– <sup>f</sup>	– <sup>f</sup>	– <sup>f</sup>	– <sup>f</sup>	(3)
W48	W48.14	294.68	+44.49	1.0	9.5	73 <sup>+16</sup> <sub>-32</sub>	0.7	–1.1 ± 0.2	>22	>27	(3)
W3(OH)	W3OH.11	212.03	–42.60	–	–	–	2.9	–4.5 ± 0.7	>89 <sup>g</sup>	>111 <sup>g</sup>	(3)
	W3OH.22	2051.30	–45.41	1.0	10.0	73 <sup>+10</sup> <sub>-5</sub>	2.1	+4.5 ± 0.7	>88	>182	(3)
	W3OH.25	156.18	–43.74	1.1	8.8	68 <sup>+7</sup> <sub>-45</sub>	2.3	+3.4 ± 0.5	>66	>70	(3)
	W3OH.35	347.43	–42.86	0.9	9.2	71 <sup>+18</sup> <sub>-33</sub>	3.8	+5.8 ± 0.9	>113	>136	(3)
	W3OH.37	110.11	–42.51	1.4	9.2	76 <sup>+13</sup> <sub>-37</sub>	5.0	–10.9 ± 1.6	>213	>260	(3)
	W3OH.40	178.62	–45.14	–	–	–	1.3	–2.0 ± 0.3	>39 <sup>h</sup>	–	(3)
	W3OH.48	19.97	–41.81	1.0	9.0	76 <sup>+11</sup> <sub>-40</sub>	1.5	+8.4 ± 1.8	>165	>195	(3)
IRAS 06058+2138-IRS 1	IRAS06.22	93.29	+10.86	1.2	9.6	73 <sup>+17</sup> <sub>-7</sub>	0.2	+0.4 ± 0.2	>9	>17	(4)
S255-IR	S255.30	10.64	+4.61	1.1	9.5	82 <sup>+8</sup> <sub>-18</sub>	0.3	+0.4 ± 0.1	>9	>16	(4)
IRAS 20126+4104	M05	27.84	–6.10	2.0	8.8	75 <sup>+10</sup> <sub>-43</sub>	0.7	–1.8 ± 0.3	>36	>41	(5)
G24.78+0.08	G24.16	13.70	+110.41	1.7	8.6	82 <sup>+8</sup> <sub>-44</sub>	0.3	–0.6 ± 0.2	>11	>13	(6)
	G24.23	9.79	+107.69	1.9	8.8	67 <sup>+9</sup> <sub>-44</sub>	0.7	+1.7 ± 0.3	>34	>36	(6)
	G24.43	11.48	+114.23	1.6	8.9	88 <sup>+1</sup> <sub>-21</sub>	0.7	–1.5 ± 0.3	>30	>60	(6)
	G24.52	32.97	+113.40	1.8	8.8	73 <sup>+2</sup> <sub>-40</sub>	0.3	–3.7 ± 0.6	>73	>84	(6)
G29.86–0.04	G29.09	19.26	+100.39	1.8	8.6	74 <sup>+15</sup> <sub>-37</sub>	0.6	–1.2 ± 0.2	>24	>28	(6)
G213.70–12.6	G213.15	91.58	+12.57	1.1	9.5	82 <sup>+8</sup> <sub>-17</sub>	0.7	–1.1 ± 0.2	>22	>42	(6)
IRAS 06058+2138-IRS 1	IRAS06.22	93.29	+10.86	1.2	9.6	73 <sup>+17</sup> <sub>-7</sub>	0.2	+0.4 ± 0.2	>9	>17	(4)
S255-IR	S255.30	10.64	+4.61	1.1	9.5	82 <sup>+8</sup> <sub>-18</sub>	0.3	+0.4 ± 0.1	>9	>16	(4)
IRAS 20126+4104	M05	27.84	–6.10	2.0	8.8	75 <sup>+10</sup> <sub>-43</sub>	0.7	–1.8 ± 0.3	>36	>41	(5)
G24.78+0.08	G24.16	13.70	+110.41	1.7	8.6	82 <sup>+8</sup> <sub>-44</sub>	0.3	–0.6 ± 0.2	>11	>13	(6)
	G24.23	9.79	+107.69	1.9	8.8	67 <sup>+9</sup> <sub>-44</sub>	0.7	+1.7 ± 0.3	>34	>36	(6)
	G24.43	11.48	+114.23	1.6	8.9	88 <sup>+1</sup> <sub>-21</sub>	0.7	–1.5 ± 0.3	>30	>60	(6)
	G24.52	32.97	+113.40	1.8	8.8	73 <sup>+2</sup> <sub>-40</sub>	0.3	–3.7 ± 0.6	>73	>84	(6)
G29.86–0.04	G29.09	19.26	+100.39	1.8	8.6	74 <sup>+15</sup> <sub>-37</sub>	0.6	–1.2 ± 0.2	>24	>28	(6)
G213.70–12.6	G213.15	91.58	+12.57	1.1	9.5	82 <sup>+8</sup> <sub>-17</sub>	0.7	–1.1 ± 0.2	>22	>42	(6)

**Notes.** <sup>(a)</sup>Values used to model the circularly polarized emission as evaluated from the linearly polarized emission. <sup>(b)</sup>The lower limit of the magnetic field strength along the line of sight was determined by modeling the circularly polarized emission assuming that the favored hyperfine transition is  $F = 3 \rightarrow 4$  (Lankhaar et al. 2018). The Landé g-factor for this transition is  $g_1 = -1.135 \text{ Hz mG}^{-1}$  (or  $a_Z = -0.051 \text{ km s}^{-1} \text{ G}^{-1}$ ).

<sup>(c)</sup> $B = B_{||} / \cos(\theta + \varepsilon_\theta^-)$ ; considering that  $\theta_{\varepsilon_\theta^+}$ . <sup>(d)</sup>Due to the scarce spectral resolution,  $B_{||}$  was determined from the cross-correlation between the RR and LL spectra, i.e.,  $\frac{B_{||}}{[\text{mG}]} = (\frac{\Delta V_Z}{[\text{m s}^{-1}]}) \cdot (\frac{a_Z}{[\text{km s}^{-1} \text{ G}^{-1}]})^{-1}$ . <sup>(e)</sup>To model the circularly polarized emission we considered the values of  $T_b \Delta \Omega = 1.1 \times 10^9 \text{ K sr}$ ,  $\Delta V_i = 0.7 \text{ km s}^{-1}$ , and  $\theta = 79^{+11}_{-40}^\circ$  measured for W51-e2. <sup>(f)</sup>No  $4\sigma$  detection of the Zeeman splitting has been measured. <sup>(g)</sup>To model the circularly polarized emission we considered the error-weighted values of  $\langle T_b \Delta \Omega \rangle = 1.8 \times 10^9 \text{ K sr}$ ,  $\langle \Delta V_i \rangle = 0.6 \text{ km s}^{-1}$ , and  $\langle \theta \rangle = 78^{+12}_{-37}^\circ$  measured for Group VI in W3(OH). <sup>(h)</sup>To model the circularly polarized emission we considered the values of  $T_b \Delta \Omega = 2.2 \times 10^9 \text{ K sr}$  and  $\Delta V_i = 0.5 \text{ km s}^{-1}$  that best fit the total intensity spectrum.

**References.** (1) Surcis et al. (2009); (2) Surcis et al. (2011a); (3) Paper I; (4) Paper II; (5) Surcis et al. (2014b); (6) Paper III.

maser emission will be determined properly by modeling the pumping mechanism of the maser. However, by knowing the inclination of the magnetic field with respect to the line of sight (i.e., the  $\theta$  angle), we can also estimate a lower limit for

the magnetic field strength ( $B = \frac{B_{||}}{\cos(\theta + \varepsilon_\theta^-)}$ ; considering  $\theta_{\varepsilon_\theta^+}$ ). We were able to measure both  $B_{||}$  and  $B$  towards two of the seven massive YSOs reported in this work: G25.71–0.04 and G28.83–0.25.

**Table 4.** Comparison between position angle of magnetic field, CH<sub>3</sub>OH maser distribution, outflows, and linear polarization angles.

(1) Source	(2) $\Phi_f^a$ ( $^\circ$ )	(3) $\langle\chi\rangle^b$ ( $^\circ$ )	(4) $\langle\Phi_B\rangle^b$ ( $^\circ$ )	(5) $PA_{\text{outflow}}$ ( $^\circ$ )	(6) $PA_{\text{CH}_3\text{OH}}$ ( $^\circ$ )	(7) $\rho^c$	(8) $ PA_{\text{outflow}} - \langle\Phi_B\rangle $ ( $^\circ$ )	(9) $ PA_{\text{CH}_3\text{OH}} - \langle\chi\rangle $ ( $^\circ$ )	(10) $ PA_{\text{CH}_3\text{OH}} - PA_{\text{outflow}} $ ( $^\circ$ )	(11) Ref.
G23.44–0.18 (MM2)	13	+23 $\pm$ 40	–67 $\pm$ 40	–40 $\pm$ 15 <sup>d</sup>	+30 $\pm$ 26 <sup>e</sup> –24 $\pm$ 91 <sup>f</sup>	+0.69 –0.37	27 $\pm$ 43	13 $\pm$ 48 47 $\pm$ 81	70 $\pm$ 30 16 $\pm$ 91	(1)
G25.83–0.18	11	+23 $\pm$ 7	–67 $\pm$ 7	+10 $\pm$ 15 <sup>g</sup>	+51 $\pm$ 7	+0.69	77 $\pm$ 17	28 $\pm$ 10	41 $\pm$ 17	(2)
G25.71–0.04	23	–51 $\pm$ 77	–80 $\pm$ 43 <sup>h</sup>	–90 $\pm$ 15 <sup>g,i</sup>	–71 $\pm$ 9	–0.41	10 $\pm$ 46	20 $\pm$ 78	18 $\pm$ 18	(2)
G28.31–0.39	24	+40 $\pm$ 44	–50 $\pm$ 44	–52 $\pm$ 15 <sup>g</sup>	+85 $\pm$ 22	+0.07	2 $\pm$ 47	45 $\pm$ 49	84 $\pm$ 27 <sup>j</sup>	(2)
G28.83–0.25	10	–45 $\pm$ 34	+58 $\pm$ 59 <sup>h</sup>	–40 $\pm$ 15 <sup>g</sup>	–41 $\pm$ 10	–0.83	82 $\pm$ 61 <sup>j</sup>	4 $\pm$ 35	1 $\pm$ 18	(2)
G29.96–0.02	12	+62 $\pm$ 17	–29 $\pm$ 17	–38 $\pm$ 15 <sup>g</sup>	+80 $\pm$ 3	0.51	9 $\pm$ 23	18 $\pm$ 17	62 $\pm$ 16 <sup>j</sup>	(3)
G43.80–0.13	14	–81 $\pm$ 5	+9 $\pm$ 5	+38 $\pm$ 15 <sup>g</sup>	–48 $\pm$ 5	–0.94	29 $\pm$ 16	33 $\pm$ 7	86 $\pm$ 16	(4)
From Paper III <sup>k</sup>										
IRAS 20126+4104	4	–70 $\pm$ 16	+20 $\pm$ 16	–65 $\pm$ 5	+87 $\pm$ 4	+0.12	85 $\pm$ 17	23 $\pm$ 17	28 $\pm$ 6	(5)
G24.78+0.08-A2	17	–53 $\pm$ 2	+37 $\pm$ 2	–40 $\pm$ 15	–26 $\pm$ 19	–0.77	77 $\pm$ 15	79 $\pm$ 19	66 $\pm$ 24	(5)
G25.65+1.05	7	–80 $\pm$ 8	–23 $\pm$ 51	–15 $\pm$ 15	–49 $\pm$ 7	–0.87	8 $\pm$ 53	31 $\pm$ 11	64 $\pm$ 17	(5)
G29.86–0.04	17	+46 $\pm$ 41	+82 $\pm$ 56	+6 $\pm$ 15	+8 $\pm$ 7	+0.73	76 $\pm$ 58	38 $\pm$ 42	14 $\pm$ 17	(5)
G35.03+0.35	8	–64 $\pm$ 5	+26 $\pm$ 5	+27 $\pm$ 15	–26 $\pm$ 19	–0.77	1 $\pm$ 16	38 $\pm$ 20	53 $\pm$ 24	(5)
G37.43+1.51	4	+90 $\pm$ 3	+90 $\pm$ 3	–4 $\pm$ 15	–64 $\pm$ 5	–0.87	86 $\pm$ 15	26 $\pm$ 6	60 $\pm$ 16	(5)
G174.20–0.08	4	–	–	–40 $\pm$ 15	–63 $\pm$ 16	–0.45	–	–	23 $\pm$ 22	(5)
G213.70–12.6-IRS3	2	+20 $\pm$ 5	–70 $\pm$ 5	+53 $\pm$ 15	+63 $\pm$ 2	+0.95	57 $\pm$ 16	43 $\pm$ 5	10 $\pm$ 15	(5)
Cepheus A	2	–57 $\pm$ 28	+30 $\pm$ 19	+40 $\pm$ 4	–79 $\pm$ 9	–0.34	10 $\pm$ 19	22 $\pm$ 29	61 $\pm$ 10	(5)
W75N-group A	3	–13 $\pm$ 9	+77 $\pm$ 9	+66 $\pm$ 15	+43 $\pm$ 10	+0.96	11 $\pm$ 18	56 $\pm$ 14	23 $\pm$ 18	(5)
NGC7538-IRS1	6	–30 $\pm$ 69	+67 $\pm$ 70	–40 $\pm$ 10	+84 $\pm$ 7	+0.15	73 $\pm$ 71	66 $\pm$ 69	56 $\pm$ 12	(5)
W3(OH)-group II	4	+21 $\pm$ 45	–47 $\pm$ 44	–	–59 $\pm$ 6	–0.84	–	80 $\pm$ 45	–	(5)
W51-e2	12	+33 $\pm$ 16	–60 $\pm$ 21	–50 $\pm$ 20	+57 $\pm$ 8	+0.70	10 $\pm$ 29	24 $\pm$ 18	73 $\pm$ 22	(5)
IRAS18556+0138	5	–2 $\pm$ 11	+88 $\pm$ 11	+58 $\pm$ 23	–40 $\pm$ 2	–0.99	30 $\pm$ 26	42 $\pm$ 11	82 $\pm$ 23	(5)
W48	7	+23 $\pm$ 7	–67 $\pm$ 7	–	+55 $\pm$ 10	+0.70	–	78 $\pm$ 12	–	(5)
IRAS06058+2138-NIRS1	4	+49 $\pm$ 47	–49 $\pm$ 52	–50 $\pm$ 15	+78 $\pm$ 7	+0.64	1 $\pm$ 54	29 $\pm$ 48	52 $\pm$ 17	(5)
IRAS22272+6358A	2	–80 $\pm$ 15	+9 $\pm$ 15	–40 $\pm$ 15	–35 $\pm$ 11	–0.87	49 $\pm$ 21	45 $\pm$ 19	5 $\pm$ 19	(5)
S255-IR	4	+36 $\pm$ 12	–54 $\pm$ 12	+75 $\pm$ 15	–63 $\pm$ 49	–0.11	51 $\pm$ 19	81 $\pm$ 51	42 $\pm$ 51	(5)
S231	4	+28 $\pm$ 49	–62 $\pm$ 49	–47 $\pm$ 5	+28 $\pm$ 8	+0.97	15 $\pm$ 49	0 $\pm$ 50	75 $\pm$ 9	(5)
G291.27–0.70	7	–32 $\pm$ 5	+52 $\pm$ 5	–	–77 $\pm$ 14	–	–	45 $\pm$ 15	–	(5)
G305.21+0.21	9	–51 $\pm$ 14	28 $\pm$ 14	–	+48 $\pm$ 23	–	–	81 $\pm$ 27	–	(5)
G309.92+0.47	12	+2 $\pm$ 56	–75 $\pm$ 56	–	+35 $\pm$ 5	–	–	33 $\pm$ 56	–	(5)
G316.64–0.08	3	–67 $\pm$ 36	+21 $\pm$ 36	–	+34 $\pm$ 29	–	–	79 $\pm$ 46	–	(5)
G335.79+0.17	8	+44 $\pm$ 28	–41 $\pm$ 28	–	–69 $\pm$ 25	–	–	67 $\pm$ 38	–	(5)
G339.88–1.26	7	+77 $\pm$ 24	–12 $\pm$ 24	–	–60 $\pm$ 17	–	–	43 $\pm$ 29	–	(5)
G345.01+1.79	5	+5 $\pm$ 39	–86 $\pm$ 39	–	+74 $\pm$ 4	–	–	69 $\pm$ 39	–	(5)
NGC6334F (central)	5	+77 $\pm$ 20	–13 $\pm$ 20	+30 $\pm$ 15 <sup>g</sup>	–41 $\pm$ 16	–	43 $\pm$ 25	62 $\pm$ 26	71 $\pm$ 41	(5)
NGC6334F (NW)	5	–71 $\pm$ 20	+19 $\pm$ 20	+30 $\pm$ 15 <sup>g</sup>	–80 $\pm$ 38	–	11 $\pm$ 25	9 $\pm$ 43	70 $\pm$ 41 <sup>e</sup>	(5)

**Notes.** <sup>(a)</sup>Foreground Faraday rotation estimated by using Eq. (3) in Paper I. <sup>(b)</sup>Because of the large uncertainties of the estimated  $\Phi_f$ , the angles are not corrected for  $\Phi_f$ . <sup>(c)</sup>Pearson product-moment correlation coefficient  $-1 \leq \rho \leq +1$ ;  $\rho = +1$  ( $\rho = -1$ ) is total positive (negative) correlation,  $\rho = 0$  is no correlation. <sup>(d)</sup>We overestimate the errors by considering half of the opening angle of the outflow. <sup>(e)</sup>Calculated for the western group of masers. <sup>(f)</sup>Calculated for the eastern group of masers. <sup>(g)</sup>We consider an arbitrary conservative error of 15°. <sup>(h)</sup>Before averaging we use the criterion described in Sect. 3 to estimate the orientation of the magnetic field with respect to the linear polarization vectors. <sup>(i)</sup>Since the 6.7 GHz CH<sub>3</sub>OH maser is associated with the blue-shifted lobe of the <sup>13</sup>CO outflow we consider only its PA. <sup>(j)</sup>The differences between the angles are evaluated taking into account that  $PA \equiv PA \pm 180^\circ$ ,  $\langle\chi\rangle \equiv \langle\chi\rangle \pm 180^\circ$ , and  $\langle\Phi_B\rangle \equiv \langle\Phi_B\rangle \pm 180^\circ$ . <sup>(k)</sup>Here we omit all the notes that are already indicated in Table 2 of Paper III.

**References.** (1) Ren et al. (2011); (2) de Villiers et al. (2014); (3) Cesaroni et al. (2017); (4) López-Sepulcre et al. (2010); (5) Paper III and references therein.

**G25.71–0.04.** From the circularly polarized emission of G257.E12 we measured a Zeeman splitting of  $-3.1 \pm 0.7 \text{ m s}^{-1}$ , which implies a magnetic field along the line of sight of  $B_{\parallel} > 61 \text{ mG}$ . By assuming  $\theta_{\min} > 43^\circ$  (see Col. 14 of Table A.3) the 3D magnetic field strength is  $B > 78 \text{ mG}$ .

**G28.83–0.25.** The Zeeman splitting measured by modeling the circularly polarized emission of G288.E19 is  $\Delta V_Z = -1.1 \pm 0.3 \text{ m s}^{-1}$ , for which we have  $B_{\parallel} > 21 \text{ mG}$ . Considering an error-weighted angle of  $\langle\theta\rangle = 81^{+10}_{-34}$  the 3D magnetic field is  $B > 28 \text{ mG}$ . The 3D magnetic field measured from the CH<sub>3</sub>OH maser is four times larger than that measured from the OH maser (Bayandina et al. 2015). From the relation  $|B| \propto n_{\text{H}_2}^{0.47}$  (Crutcher 1999), we determine that the CH<sub>3</sub>OH maser are arising from a gas with a density at least an order of magnitude higher than that of the OH maser. This agrees with the ranges of  $n_{\text{H}_2}$  for the two maser species (e.g., Cragg et al. 2002).

From the estimated  $B$  values we can investigate the importance of the magnetic field in the high-mass star-forming process. If the ratio  $\beta$  between the thermal ( $E_T$ ) and the magnetic energies ( $E_B$ ) is lower than one ( $\beta < 1$ ), the magnetic field dominates in the high-density CH<sub>3</sub>OH maser environment. Following Eq. (11) of Surcis et al. (2011a), we see that the ratio  $\beta$  also depends, in addition to the magnetic field, on the characteristics of the gas where the masers arise, namely on the number density of the gas ( $n_{\text{H}_2}$ ) and on the kinetic temperature of the gas ( $T_k$ ). Cragg et al. (2002, 2005) modeled the Class II CH<sub>3</sub>OH maser emissions and found that the masers arise when  $10^5 \text{ cm}^{-3} < n_{\text{H}_2} < 10^9 \text{ cm}^{-3}$  and  $100 \text{ K} < T_k < 200 \text{ K}$ . Considering  $\beta = 1$  we determined  $|B_{\text{crit}}|$  values for all the possible combinations of the  $n_{\text{H}_2}$  and  $T_k$  extremes (see Table 2). If  $|B| > |B_{\text{crit}}|$  the magnetic field dominates over the thermal motions. In the case of G25.71–0.04 and G28.83–0.25 the magnetic field dominates the dynamics independently of the characteristics of the gas and on the specific dominating hyperfine transition.



**Table 5.** Results of Kolmogorov–Smirnov test.

(1) Angle	(2) $N^a$	(3) $D^b$	(4) $\lambda^c$	(5) $Q_{K-S}(\lambda)^d$
$ \text{PA}_{\text{CH}_3\text{OH}} - \langle\chi\rangle $	35	0.16	0.99	0.28
$ \text{PA}_{\text{CH}_3\text{OH}} - \text{PA}_{\text{outflow}} $	27	0.17	0.91	0.38
$ \text{PA}_{\text{outflow}} - \langle\Phi_B\rangle $	25	0.27	1.42	0.03

**Notes.** <sup>(a)</sup> $N$  is the number of elements considered in the K–S test. <sup>(b)</sup> $D$  is the maximum value of the absolute difference between the data set composed of  $N$  elements and the random distribution. <sup>(c)</sup> $\lambda$  is a parameter given by  $\lambda = (\sqrt{N} + 0.12 + 0.11/\sqrt{N}) \times D$ . <sup>(d)</sup> $Q_{K-S}(\lambda) = 2 \sum_{j=1}^N (-1)^{j-1} e^{-2j^2\lambda^2}$  is the significance level of the K–S test.

In the past, we have reported erroneous values of Zeeman-splitting due to an error in the FRTM code (see Sect. 3). We report the corrected values of Zeeman-splitting with the corresponding lower limits of  $B_{\parallel}$  and  $B$  in Table 3. The varying values of  $|B|$  measured from different maser features within the same source might indicate different gas properties in the massive SFR due either to the association of the  $\text{CH}_3\text{OH}$  masers with different YSOs (e.g., W3(OH) and G24.78+0.08; Papers I and III) or to the different locations of the masers in the associated protostar (e.g., W51-e2 and G24.78+0.08-A1; Papers I and III). Nevertheless, if at least one of the masers detected towards a massive YSO provides  $|B| > 25$  mG (see Table 2) we can assume that the magnetic field dominates there. The only sources for which we cannot determine whether the magnetic field dominates independently of the characteristics of the gas and on the specific dominating hyperfine transition are IRAS 06058+2138-IRS 1 and S255-IR (Paper II).

### 5.3. Updated statistical results

In Paper III, since we were at the midpoint of our project<sup>7</sup>, we updated our first statistical results reported in Paper II. Here, we would like to provide the updated statistical results based on seven more sources with respect to Paper III, i.e., 35 YSOs in total, also including the southern hemisphere sources reported in the literature (Paper II and references therein). Similarly to Papers II and III, we list in Table 4 the sources of the flux-limited sample analyzed so far for which we were able to measure the projection on the plane of the sky of the angles  $|\text{PA}_{\text{outflow}} - \langle\Phi_B\rangle|$ ,  $|\text{PA}_{\text{CH}_3\text{OH}} - \langle\chi\rangle|$ , and  $|\text{PA}_{\text{CH}_3\text{OH}} - \text{PA}_{\text{outflow}}|$ , where  $\text{PA}_{\text{outflow}}$  is the orientation of the large-scale molecular outflow,  $\langle\Phi_B\rangle$  is the error-weighted orientation of the magnetic field,  $\text{PA}_{\text{CH}_3\text{OH}}$  is the orientation of the  $\text{CH}_3\text{OH}$  maser distribution, and  $\langle\chi\rangle$  is the error-weighted value of the linear polarization angles (for more details regarding Table 4, see the table notes and Paper III). In Table 5 we report the results of the Kolmogorov–Smirnov (K–S) test, which is a nonparametric test; here it is used to compare samples of angles ( $|\text{PA}_{\text{CH}_3\text{OH}} - \langle\chi\rangle|$ ,  $|\text{PA}_{\text{CH}_3\text{OH}} - \text{PA}_{\text{outflow}}|$ , and  $|\text{PA}_{\text{outflow}} - \langle\Phi_B\rangle|$ ) with the random probability distribution (see Paper II for more details). Performing the statistical analysis we note that the probability that the angles  $|\text{PA}_{\text{CH}_3\text{OH}} - \langle\chi\rangle|$  are drawn from a random distribution is 28%<sup>8</sup>.

For  $|\text{PA}_{\text{CH}_3\text{OH}} - \text{PA}_{\text{outflow}}|$  and  $|\text{PA}_{\text{outflow}} - \langle\Phi_B\rangle|$  we measured probabilities similar to those reported in Paper III, these are 38%

and 3% instead of 34% and 10% (Paper III), respectively. The updated results reinforce our previous finding: the magnetic field close to the YSO is preferentially oriented along the outflow axis.

## 6. Summary

We observed seven massive SFRs at 6.7 GHz in full polarization spectral mode with the EVN; our aim was to detect the linearly and circularly polarized emission of  $\text{CH}_3\text{OH}$  masers. We detected linearly polarized emission towards all the regions and circularly polarized emission towards G25.71–0.04 and G28.83–0.25. We used the adapted FRTM code to model both the linear and the circular polarization of the masers. In particular, to estimate a lower limit of the magnetic field along the line of sight we assumed that the dominant hyperfine component, the one with the largest Einstein coefficient for stimulated emission, is  $F = 3 \rightarrow 4$  (Lankhaar et al. 2018). By analyzing the linearly polarized emission of the masers we were able to estimate the orientation of the magnetic field around eight massive YSOs (two are located within G23.44–0.18: MM1 and MM2). We found that the magnetic fields are aligned with the outflows ( $|\text{PA}_{\text{outflow}} - \langle\Phi_B\rangle| < 30^\circ$ ) in five YSOs (G23.44–0.18-MM2, G25.71–0.04, G28.31–0.39, G29.96–0.02, G43.80–0.13) and are perpendicular to the outflows in two YSOs (G25.83–0.18 and G28.83–0.25). The estimated magnetic field strengths along the line of sight for G25.71–0.04 and G28.83–0.25 are  $B_{\parallel} > 61$  mG and  $B_{\parallel} > 21$  mG, respectively. The magnetic field seems to dominate the dynamics of the gas in both YSOs.

We further increased the number of sources to 26, which is 80% of the flux-limited sample; the projected angles of the magnetic field and of the outflows of these sources are known. Comparing these angles, we confirm the statistical evidence that the magnetic fields around massive YSOs are preferentially oriented along the molecular outflows. In particular, the probability that the distribution of angles  $|\text{PA}_{\text{outflow}} - \langle\Phi_B\rangle|$  is drawn from a random distribution is lower (3%) than was reported in Paper III (10%).

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<sup>7</sup> To determine whether there is any relation between the morphology of the magnetic field and the ejecting direction of molecular outflow from massive YSOs on a scale of tens of au.

<sup>8</sup> In Paper III we erroneously reported a probability of ~80%.

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## Appendix A: Additional tables

In Tables A.1–A.7 we list the parameters of all the CH<sub>3</sub>OH maser features detected toward the seven massive SFRs observed with the EVN and reported in this work. The tables are organized as follows. In Col. 1 we give the name of the feature, and only in Table A.1 the associated region is reported in Col. 1B. The positions, Cols. 2 and 3, refer to the maser feature used for self-calibration. From Cols. 4 to 6 we give the peak flux density, the LSR velocity ( $V_{\text{LSR}}$ ), and the FWHM ( $\Delta v_{\text{L}}$ ) of the total intensity spectra of the maser features that are obtained using a Gaussian fit. The mean linear polarization fraction ( $P_{\text{L}}$ ) and the mean linear polarization angles ( $\chi$ ) are measured across the spectrum, and are listed in Cols. 7 and 8. The outcomes of the adapted FRTM code are listed in Cols. 9 (intrinsic thermal linewidth), 10 (emerging brightness temperature), and 14 (angle between the magnetic field and the maser propagation direction). The errors were determined by analyzing the full probability distribution function. The value of  $\theta$  in bold indicates that  $|\theta^+ - 55^\circ| < |\theta^- - 55^\circ|$ , i.e., the magnetic field is assumed to be parallel to the linear polarization vector (see Papers I–III). Finally, the circular polarization fraction ( $P_{\text{V}}$ ), the Zeeman splitting ( $\Delta V_{\text{Z}}$ ), and the lower limit of the magnetic field strength along the line of sight ( $B_{\parallel}$ ) determined by fitting the V Stokes spectra by using the best-fitting results ( $\Delta V_i$  and  $T_{\text{b}}\Delta\Omega$ ) and the Landé g-factors calculated by Lankhaar et al. (2018) for the hyperfine transition  $F = 3 \rightarrow 4$  are listed in Cols. 11, 12, and 13. The tables are only available at the CDS.