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Validation and atmospheric exploration of the sub-Neptune TOI-2136b around a nearby M3 dwarf

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ABSTRACT

Context. The NASA space telescope TESS is currently in the extended mission of its all-sky search for new transiting planets. Of the thousands of candidates that TESS is expected to deliver, transiting planets orbiting nearby M dwarfs are particularly interesting targets since they provide a great opportunity to characterize their atmospheres by transmission spectroscopy.

Aims. We aim to validate and characterize the new sub-Neptune-sized planet candidate TOI-2136.01 orbiting a nearby M dwarf ($d = 33.36 \pm 0.02$ pc, $T_{\text{eff}} = 3373 \pm 108$ K) with an orbital period of 7.852 days.

Methods. We use TESS data, ground-based multicolor photometry, and radial velocity measurements with the InfraRed Doppler (IRD) instrument on the Subaru Telescope to validate the planetary nature of TOI-2136.01, and estimate the stellar and planetary parameters. We also conduct high-resolution transmission spectroscopy to search for helium in its atmosphere.

Results. We confirm that TOI-2136.01 (now named TOI-2136b) is a bona fide planet with a planetary radius of $R_p = 2.20 \pm 0.07 R_{\oplus}$ and a mass of $M_p = 4.7^{+3.1}_{-2.6} M_{\oplus}$. We also search for helium 10830 Å absorption lines and place an upper limit on the equivalent width of < 7.8 mÅ and on the absorption signal of $< 1.44\%$ with 95% confidence.

Conclusions. TOI-2136b is a sub-Neptune transiting a nearby and bright star ($J = 10.8$ mag), and is a potentially hycean planet, which is a new class of habitable planets with large oceans under a H₂-rich atmosphere, making it an excellent target for atmospheric studies to understand the formation, evolution, and habitability of the small planets.

Key words. planets and satellites: individual: TOI-2136b – planets and satellites: detection – planets and satellites: atmospheres – techniques: photometric – techniques: radial velocities

1. Introduction

Small planets with radii between 1 and $4 R_{\oplus}$ are extremely common in the Milky Way, but do not exist in our solar system. NASA's *Kepler* space telescope revealed that this population has a bimodal distribution that separates the planets into super-Earth ($1-1.5 R_{\oplus}$) and sub-Neptune ($2-4 R_{\oplus}$) populations (Fulton et al. 2017) for planets with $P < 100$ days. This “radius valley” is consistent with the transition from rocky to non-rocky planets and is predicted by photoevaporation models (e.g., Owen & Wu 2013, 2017; Jin et al. 2014; Lopez & Rice 2018; Mordasini 2020) and/or core-powered mass-loss models (Ginzburg et al. 2016, 2018; Gupta & Schlichting 2019, 2021).

Furthermore, the location of the radius valley is revealed to be dependent on the planetary period, insolation, and stellar type (e.g., Van Eylen et al. 2018; Martinez et al. 2019; Wu 2019). The position of the radius gap in the radius-period distribution plane around Sun-like stars decreases with the orbital period following a power law $R_{\text{gap}} \propto P^{-0.11}$, which is consistent with photoevaporation and core-powered mass-loss model predictions (Martinez et al. 2019). On the other hand, it has been claimed that

the position of the radius valley around late-type stars increases with the orbital period following a power law $R_{\text{gap}} \propto P^{+0.058}$, which is consistent with gas-poor formation models (Cloutier & Menou 2020). However, a recent study found that the correlation between the position of the radius gap and orbital period is negative around M-type stars whose masses are above $0.5 M_{\odot}$ ($R_{\text{gap}} \propto P^{-0.12}$), and this correlation does not seem to vary with the stellar type (Petigura et al. 2022). To further explore the nature and origins of the radius valley around M dwarfs, it is important to discover new small planets and to measure their masses as well as radii, in order to have a more complete sample.

Sub-Neptune-sized planets are likely to have H₂O-dominated ices and/or fluids (water worlds) in addition to a rocky core enveloped in a H₂-He gas (e.g., Zeng et al. 2019). Recent studies indicated the possibility that sub-Neptune-sized planets could have large oceans with habitable conditions under H₂-rich atmospheres (Hycean worlds; Madhusudhan et al. 2021). The direct observation of their atmospheres helps our understanding of the composition of the planets.

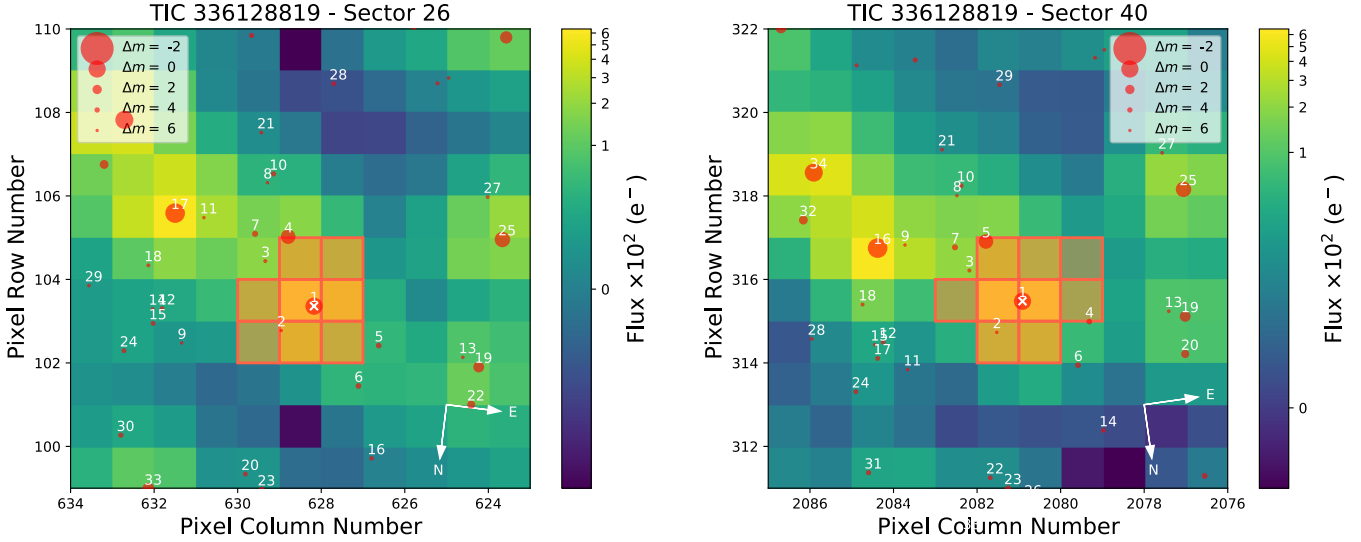


Fig. 1. TESS target pixel file image of TOI-2136 observed in Sector 26 (*left*) and Sector 40 (*right*). The pixels in red show the photometric aperture used by TESS to create the PDCSAP light curves. The position of nearby stars and their magnitudes are represented by the red circles. This image was produced using `tpfp1otter` (Aller et al. 2020).

The infrared helium triplet is a good indicator of a primary atmosphere accreted from the protoplanetary nebula (Elkins-Tanton & Seager 2008). Previous studies have detected the He I 10830 Å absorption lines in the atmosphere of about a half-dozen Jupiter- and Neptune-sized planets (e.g., Spake et al. 2018; Nortmann et al. 2018; Allart et al. 2018), but only in three sub-Neptunes: GJ3470b (Palle et al. 2020), GJ1214b (Orell-Miquel et al. 2022), and TOI 560.01 (Zhang et al. 2022). Other studies of sub-Neptunes have been able to place only upper limits on He I absorption, but they still help to constrain the mass-loss rate (e.g., Krishnamurthy et al. 2021).

Here we report the discovery of TOI-2136b, a sub-Neptune-sized planet around an M dwarf, using photometry from NASA’s Transiting Exoplanet Survey Satellite (TESS; Ricker et al. 2015) and ground-based follow-up observations. We also search for helium in its atmosphere with high-resolution transmission spectroscopy. The paper is organized as follows. In Sects. 2 and 3, we introduce the observation of TESS and ground-based telescope. In Sects. 4, 5, 6, and 7, we describe the analysis and results for stellar parameters, planet parameters, and transmission spectroscopy. We discuss the physical properties in Sect. 8, and we summarize this work in Sect. 9.

2. TESS photometry

TIC 336128819 (TOI-2136) was observed with a two-minute cadence by TESS in Sector 26 from UT 9 June 2020 to 4 July 2020, and in Sector 40 from UT 24 June 2021 to 23 July 2021. For each sector there are 0.919 days downtime to download data and 23.95 days of science data in total. For both TESS sectors, the target star was positioned on charge-coupled device (CCD) 1 Camera 1. Figure 1 shows the TESS images around TOI-2136 for Sectors 26 and 40.

TESS raw images were processed by the Science Processing Operations Center (SPOC) at NASA Ames Research Center. After the data were reduced by the SPOC pipeline (Jenkins et al. 2016), a preliminary TESS light curve was produced using simple aperture photometry (SAP; Morris et al. 2020) and instrumental systematic effects were removed using the presearch data

conditioning (PDC) pipeline module (Smith et al. 2012; Stumpe et al. 2014). A transit search (Jenkins 2002; Jenkins et al. 2020) was carried out on the TESS light curves and a signal of ~ 7.8 days was detected. After the transit signal was validated by the TESS team (Twicken et al. 2018), it was assigned a TESS object of interest number (TOI-2136.01) and announced to the community.

3. Follow-up observations

3.1. MuSCATs

We observed a total of five full transit events with the Multicolor Simultaneous Camera for studying Atmospheres of Transiting exoplanets (MuSCAT; Narita et al. 2015), MuSCAT2 (Narita et al. 2019) and MuSCAT3 (Narita et al. 2020). MuSCAT is mounted on the 1.88 m telescope of the National Astronomical Observatory of Japan (NAOJ) in Okayama, Japan. MuSCAT2 is mounted on the 1.52 m Telescopio Carlos Sánchez (TCS) at the Teide Observatory, in the Canary Islands, Spain. MuSCAT3 is mounted on the 2 m Faulkes Telescope North at the Haleakala Observatory, in Maui, Hawai’i. These instruments have g , r , i and z_s bands (only g , r , and z_s bands for MuSCAT). MuSCAT and MuSCAT2 are equipped with 1024×1024 pixel CCDs, and MuSCAT3 is equipped with 2048×2048 pixel CCDs. MuSCAT provides a field of view (FOV) of 6.1×6.1 arcmin² with a pixel scale of 0.36 arcsec per pixel. MuSCAT2 provides a field of view of 7.4×7.4 arcmin² with a pixel scale of 0.44 arcsec per pixel. MuSCAT3 provides a field of view of 9.1×9.1 arcmin² with a pixel scale of 0.27 arcsec per pixel.

The first transit was observed by MuSCAT3 on UT 12 June 2021, and the exposure times were set to 60, 26, 23, and 20 s for g , r , i , and z_s , respectively. The fourth transit was observed by MuSCAT on UT 23 September 2021, and the exposure times were set to 30 and 30 s for r and z_s ; for that night the g -band CCD of MuSCAT did not work. The rest of the transits were observed by MuSCAT2 on UT 22 and 30 August 2021, and 24 October 2021. For the observations of 22 August 2021, the exposure times were set to 30, 15, and 15 s for g , i , and z_s , respectively; for that

Table 1. Ground-based photometric observations.

Date (UT)	Telescope/instrument	Telescope size	Filters
2021-06-12	Faulkes/MuSCAT3	2.0 m	g, r, i, z_s
2021-08-22	TCS/MuSCAT2	1.5 m	i, z_s
2021-08-30	TCS/MuSCAT2	1.5 m	g, i, z_s
2021-09-23	Okayama/MuSCAT	1.88 m	r, z_s
2021-10-08	LCO/Sinistro	1.0 m	i, z
2021-10-24	TCS/MuSCAT2	1.5 m	g, r, i, z_s

night the r -band CCD of MuSCAT2 did not work and we also discarded the g -band data due to the large photometric scatter present in the light curve. For the MuSCAT2 observations of 30 August 2021, the exposure times were the same for each band as the night of 22 August 2021; the r -band CCD did not work that night also, but the g -band data were much better and we decided to include them in the analysis. For the night of 24 October 2021, the exposure times were set to 30, 20, 15, and 15 s for the g , r , i , and z_s bands, respectively. We used a fixed aperture radius of 4.3 arcsec for each band for the MuSCAT photometry, and 10.9 arcsec for each band for the MuSCAT2 photometry. For the MuSCAT3 photometry, we used fixed aperture radii of 3.7, 4.3, 4.3, and 4.8 arcsec for g , r , i , and z_s , respectively.

3.2. LCO 1 m/Sinistro

We observed one almost-full transit of TOI-2136b on UT 9 October 2021 simultaneously with two 1 m telescopes of Las Cumbres Observatory (LCO) at the McDonald Observatory in the USA. Each telescope is equipped with a single-band imager Sinistro, which has a $4k \times 4k$ CCD with a pixel scale of $0''.389 \text{ pixel}^{-1}$ and an FOV of $26'.5 \times 26'.5$. We selected the i -band filter for one telescope and the Z -band filter for the other. We set the exposure time at 30 s with the full-frame readout mode for both telescopes. The telescopes were slightly defocused (by 0.2 mm) to avoid detector saturation. The obtained raw images were processed by the BANZAI pipeline (McCully et al. 2018) to perform dark-image and flat-field corrections. Aperture photometry was then performed using a custom pipeline (Fukui et al. 2011). A summary of the ground-based photometric observations, facilities, instruments, and filters used in this work is presented in Table 1.

3.3. Subaru/IRD

We obtained a total of 38 high-resolution spectra of TOI-2136 from UT 27 September 2020 to 27 October 2021 with the InfraRed Doppler (IRD) instrument on the Subaru 8.2 m telescope under the Subaru-IRD TESS intensive follow-up program (ID: S20B-088I, S21B-118I). The IRD is a fiber-fed, near-infrared (NIR) spectrometer covering from 970 to 1750 nm with a spectral resolution of 70 000 (Tamura et al. 2012; Kotani et al. 2018). The exposure times were set to 750–1800 s depending on the sky condition. We also observed a telluric standard star (A0 or A1 star) to correct for telluric lines in the template spectrum, for the radial velocity (RV) analysis.

Raw IRD data were reduced with IRAF (Tody 1993) and our custom code to correct for the bias pattern of the detectors, and also to perform wavelength calibration with the Th–Ar lamp and laser-frequency comb (Kuzuhara et al. 2018; Hirano et al. 2020). The reduced one-dimensional spectra have signal-to-noise ratios

(S/Ns) of 40–94 per pixel around 1000 nm. We measured the precision RV for each frame with the RV analysis pipeline for the IRD described in Hirano et al. (2020). The typical RV internal errors are $3\text{--}5 \text{ m s}^{-1}$. The RV individual measurements are given in Table A.1.

The IRD data also include the observation of a full transit of TOI-2136b on 27 September 2020, with the goal of searching for an excess helium absorption by the planetary atmosphere. We obtained 11 consecutive spectra with a 900 s exposure time from UT 05:01 to 07:51, out of which five frames are out-of-transit spectra and six frames are in-transit spectra. The S/Ns of the extracted spectra are 50 – 63 per pixel around 1000 nm, and the airmass varied from 1.05 to 1.33 during the observations.

3.4. Gemini North/Alopeke

Exoplanet parameters, such as the radius determination, depend on the transit depth as well as the assumption that the star is single. A spatially close companion (bound or line-of-sight) can create a false-positive transit signal, or cause “third-light” flux, leading to an underestimation of the planet radius (Ciardi et al. 2015). Such companion stars can even cause nondetections of small planets residing within the same exoplanetary system (Lester et al. 2021). Nearly 50% of FGK-type stars are binary or multiple star systems (Matson et al. 2018), providing crucial information toward our understanding of exoplanetary formation, dynamics, and evolution (Howell et al. 2021). Thus, to search for close-in ($<1''$) companions unresolved in TESS images or other ground-based follow-up observations, we obtained high-resolution imaging speckle observations of TOI-2136.

TOI-2136 was observed on UT 17 October 2021 using the ‘Alopeke speckle instrument on the Gemini North 8-m telescope (Scott et al. 2021). ‘Alopeke provides simultaneous speckle imaging in two bands (562 nm and 832 nm) with output data products including a reconstructed image with robust contrast limits on companion detections (e.g., Howell et al. 2016). Five sets of $1000 \times 0.06 \text{ s}$ exposures were collected for TOI-2136 and subjected to Fourier analysis in our standard reduction pipeline (see Howell et al. 2011). Figure 2 shows our final contrast curves and the 832 nm reconstructed speckle image. We find that TOI-2136 is a single star with no companion dimmer than four to seven magnitudes above that of the target star from the diffraction limit (20 mas) out to $1.2''$. At the distance of TOI-2136 ($d = 33 \text{ pc}$), these angular limits correspond to spatial limits of 0.7–40 au.

4. Stellar properties

4.1. Stellar parameters

We determined the effective temperature T_{eff} and metallicity of the host star from the template spectrum, which was constructed from the IRD spectra. The analysis is based on a line-by-line comparison of equivalent widths (EWs) between the observed spectra and synthetic spectra, following the procedure described in Sect. 3.1.2 of Hirano et al. (2021). The synthetic spectra were calculated by a one-dimensional LTE spectral synthesis code, which was based on the same assumptions as the model atmosphere program of Tsuji (1978), with the interpolated grid of MARCS model atmosphere (Gustafsson et al. 2008). Through spectral analysis, a surface gravity $\log g$ value of 4.88 ± 0.0008 was adopted from the TESS Input Catalog (TIC) version 8

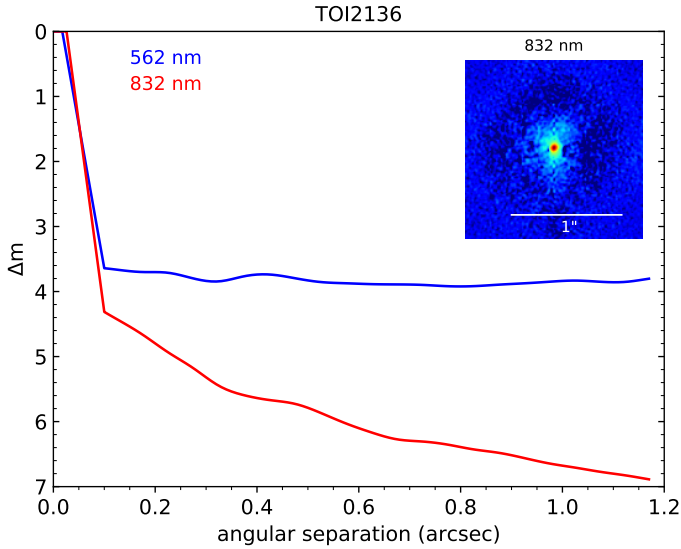


Fig. 2. Gemini/Alopeke 5-sigma contrast curves in 562 nm (blue) and 832 nm (red). The inset figure is the 832 nm reconstructed speckle image.

(Stassun et al. 2019), and the microturbulent velocity was fixed at $0.5 \pm 0.5 \text{ km s}^{-1}$ for simplicity.

The selected 47 FeH molecular lines in the Wing-Ford band at 990–1020 nm were used for the T_{eff} estimation. Ishikawa et al. (2022) provides the selection criteria and the list of the 47 lines. The original data of these line transitions are available on the MARCS¹ web page. We measured the EW of each FeH line by fitting the Gaussian profile to find the T_{eff} at which the synthetic spectra best reproduce the EW by an iterative search.

For the stellar metallicity, we used the analysis tool developed by Ishikawa et al. (2020). We used the atomic absorption lines of Na, Mg, Ca, Ti, Cr, Mn, Fe, and Sr based on the selection criteria of: (1) not suffering from blending of other absorption lines, (2) sensitive to elemental abundances, and (3) continuum level can be reasonably determined. The original data of the atomic lines are taken from the Vienna Atomic Line Database (VALD; Kupka et al. 1999; Ryabchikova et al. 2015). The EWs were measured by fitting the synthetic spectra on a line-by-line basis. We searched for an elemental abundance until the synthetic EW matched the observed one for each line, and took the average for all the lines to estimate $[X/H]$ for an element X. The $[M/H]$ is calculated as an error-weighted average of abundances of all the eight elements.

As a first step, we estimated T_{eff} assuming the solar metallicity. The average of the T_{eff} estimates for the 47 lines was adopted as the best estimate. We temporarily determined the elemental abundances using the T_{eff} estimate and, as a second step, we redetermined T_{eff} adopting the iron abundance $[Fe/H]$ as the metallicity of the atmospheric model grid. Its uncertainty was given as the sum of the statistical error (8 K), which was calculated by dividing the standard deviation of the 47 lines by the square root of the number of lines (σ/\sqrt{N}), and the possible systematic error (100 K). We then adopted the resulting T_{eff} to finally determined the elemental abundances. The procedure up to this point allowed the results of T_{eff} and abundances to converge well within the measurement errors. Consequently, we

Table 2. Stellar parameters.

Parameter	Value	Reference
TIC	336128819	
α (J2000)	18:44:42.37	(a)
δ (J2000)	36:33:47.36	(a)
$\mu_{\alpha} \cos \delta$ (mas yr ⁻¹)	-33.809 ± 0.017	(a)
μ_{δ} (mas yr ⁻¹)	177.053 ± 0.020	(a)
Parallax (mas)	29.9756 ± 0.0169	(a)
<i>Gaia</i>	12.9462 ± 0.0005	(a)
TESS	11.7371 ± 0.0074	(b)
<i>V</i> mag	14.32 ± 0.2	(c)
<i>J</i> mag	10.184 ± 0.024	(d)
<i>H</i> mag	9.604 ± 0.028	(d)
<i>K</i> mag	9.343 ± 0.022	(d)
Distance (pc)	33.361 ± 0.019	This work
T_{eff} (K)	3373 ± 108	This work
<i>U</i> (km s ⁻¹)	-35.20 ± 0.13	This work
<i>V</i> (km s ⁻¹)	-19.61 ± 0.29	This work
<i>W</i> (km s ⁻¹)	6.05 ± 0.10	This work
$[Fe/H]$ (dex)	0.02 ± 0.14	This work
$[M/H]$ (dex)	0.06 ± 0.06	This work
$\log g$ (cgs)	4.881 ± 0.026	This work
M_* (M_{\odot})	0.3272 ± 0.0082	This work
R_* (R_{\odot})	0.3440 ± 0.0099	This work
ρ_* (ρ_{\odot})	$8.04^{+0.74}_{-0.66}$	This work
L_* (L_{\odot})	$0.0137^{+0.0019}_{-0.0017}$	This work
P_{rot} (days)	82.56 ± 0.45	This work

References. (a) *Gaia* EDR3 (Riello et al. 2021), (b) Stassun et al. (2019), (c) Lépine & Shara (2005), (d) Cutri et al. (2003).

obtained $T_{\text{eff}} = 3373 \pm 108 \text{ K}$, $[Fe/H] = 0.02 \pm 0.14 \text{ dex}$, and $[M/H] = 0.06 \pm 0.06 \text{ dex}$.

Based on these parameters, as well as the literature values in Table 2, we estimated the other physical parameters, including the stellar radius, which is required to infer the planet radius. Making use of the empirical relations for M dwarfs by Mann et al. (2015, 2019), we derived the stellar mass M_* , radius R_* , and density ρ_* with their uncertainties estimated by the Monte Carlo approach using Gaussian distributions for the *Gaia* parallax, 2MASS magnitudes, and $[Fe/H]$. The stellar luminosity L_* was also derived from the Stefan–Boltzmann law using T_{eff} in Table 2.

4.2. Stellar rotation

We obtained publicly available photometry of TOI-2136 from the MEarth project (Berta et al. 2012) and the Zwicky Transient Facility (ZTF, Masci et al. 2019). The MEarth project uses eight 40 cm telescopes equipped with RG715 filters and it mostly focuses on observing M stars to search for transits (Nutzman & Charbonneau 2008). The project has two observing sites in both hemispheres. One site is located at the Fred Lawrence Whipple Observatory, on Mount Hopkins in Arizona, in the USA. The other is located at the Cerro Tololo Inter-American Observatory (CTIO) in Chile. From the MEarth project, we obtained TOI-2136 (LSPMJ1844+3633) photometry from the 2011 to 2020 north target light curves.

The ZTF uses a 47 squared degree camera mounted on the Palomar 48-inch Schmidt Telescope to search and study transient and variable objects in the sky. Using its large observing area,

¹ <https://marcs.astro.uu.se/>

ZTF is capable of scanning the northern hemisphere sky every two nights. From ZTF we obtained g - and r -band photometry of TOI-2136 from the ZTF data release 8 (DR8).

Visual inspection of the MEarth and ZTF data shows that TOI-2136 presented periodic photometric variations (see Fig. 3). These variations can be produced by inhomogeneities in the stellar surface (e.g., spots, plages) that appear and disappear out of sight as the star rotates. We performed a joint fit to the MEarth and ZTF photometry to establish the rotation period of the star. We used the package for Gaussian Processes *Celerite* (Foreman-Mackey et al. 2017) and fit the photometry with the following kernel:

$$k_{ij \text{ Phot}} = \frac{B}{2+C} e^{-|t_i - t_j|/L} \left[\cos\left(\frac{2\pi|t_i - t_j|}{P_{\text{rot}}}\right) + (1+C) \right], \quad (1)$$

where $|t_i - t_j|$ is the difference between two epochs or observations, B , C , and L are positive constants, and P_{rot} is the stellar rotational period (see Foreman-Mackey et al. 2017 for details). For the joint fit, each data set had the constants B , C , and L as free parameters, but shared a common rotational period. The prior functions and parameter limits used in the joint fit are in Table B.1. After global optimization of a log likelihood function, we sampled the posterior distribution of the kernel used to fit the photometry with *emcee* (Foreman-Mackey et al. 2013), using 80 chains and 20 000 iterations. The final parameter values (median and 1σ uncertainties) were estimated from the posterior distribution.

Figure 3 shows the MEarth and ZTF photometry, and the best fitted model using the kernel described in Eq. (1). The generalized Lomb–Scargle (GLS, Zechmeister & Kürster 2009) periodograms for each photometric data set are shown in the right panel. The three data sets present peaks in their respective periodograms for periods between 70 and 90 days. The stellar rotational period found by the best fitted model is $P_{\text{rot}} = 82.56 \pm 0.45$ days. This rotation period is in agreement with Newton et al. (2016), which found a rotational period of $P_{\text{rot}} = 82.97$ days for this star, using MEarth data.

Engle & Guinan (2018) investigated the relationship between the stellar age and rotation of M dwarfs. Using their equation, the stellar age of TOI-2136 is calculated as ~ 5 Gyr. We also estimated the stellar age from its correlation with the Galactic space velocities (U , V , W). We computed the UVW velocities using the *Gaia* Early Data Release (EDR) 3 information and the absolute RV measured from the IRD spectra using the procedure described in Johnson & Soderblom (1987; Table 2). Following Hirano et al. (2021), we obtained an age of $6.2^{+3.7}_{-3.1}$ Gyr when the age prior based on the Geneva-Copenhagen Survey catalog (Casagrande et al. 2011) is imposed, and $6.3^{+5.0}_{-4.1}$ Gyr for the uniform age prior ($0 < \text{age} < 14$ Gyr). These results indicate that the stellar rotation and UVW velocities do not constrain the age of the star.

5. Methods

5.1. Transit light curve modeling

For the transit and RV modeling, we followed the procedure described in Murgas et al. (2021). The light curves were modeled using the python package *PyTransit*² (Parviainen 2015). To fit the transit events analyzed here, we used *PyTransit*'s Mandel & Agol (2002) analytic models, assuming a quadratic limb-darkening law.

² <https://github.com/hpparvi/PyTransit>

Before performing the joint fit of all the data available, the ground-based transit observations were fitted simultaneously using a common transit model and for each night the systematic effects were accounted for independently. To model the instrumental noise, we used a linear model with one free parameter dependent on the airmass of the target, two free parameters dependent on the star position in the detector (X - and Y -axis), and a term dependent on the full width at half maximum (FWHM) of the point spread function (PSF) as a proxy for seeing variations. For the final joint fit we used the LCO, MuSCAT, MuSCAT2 and MuSCAT3 detrended light curves (i.e., without systematic effects).

For the joint fit, the common transit free parameters used in the transit modeling were the planet-to-star radius ratio R_p/R_* , the central time of the transit T_c , the stellar density ρ_* , and the transit impact parameter b . A constant orbital period (i.e., assuming no transit timing variations) was used as a common parameter for the transit light curves and RV measurements. The passband dependent quadratic limb-darkening (LD) coefficients u_1 and u_2 were set free. During the fitting process, we converted the LD coefficients (u_1, u_2) to the parametrization proposed by Kipping (2013), (q_1, q_2) with $q_i \in [0, 1]$. The fitted LD coefficients were weighted against the predicted coefficients delivered by the Python Limb Darkening Toolkit³ (*ldtk*, Parviainen & Aigrain 2015), computed using the TOI-2136 stellar parameters from Table 2.

To model the stellar variability and residual systematic noise present in the TESS data, we used Gaussian processes (GPs; e.g., Rasmussen & Williams 2010, Gibson et al. 2012, Ambikasaran et al. 2015). For the GPs used in the TESS data, we chose the Matern 3/2 kernel implementation of *Celerite* (Foreman-Mackey et al. 2017):

$$k_{ij \text{ TESS}} = c_1^2 \left(1 + \frac{\sqrt{3}|t_i - t_j|}{\tau_1} \right) \exp\left(-\frac{\sqrt{3}|t_i - t_j|}{\tau_1}\right), \quad (2)$$

where $|t_i - t_j|$ is the difference between two epochs or observations, c_1 is the amplitude of the flux variation, and τ_1 is a characteristic timescale. For each TESS data set available, the constants c_1 and τ_1 were set as free parameters in the fit.

5.2. Radial velocity modeling

The RV data were modeled using the Radial Velocity Modeling Toolkit *RadVel*⁴ (Fulton et al. 2018). The free parameters used to model the RVs were: the planet-induced RV semiamplitude (K_{RV}), the host star systemic velocity (γ_0), the orbital eccentricity and argument of the periastron (using the parametrization of $\sqrt{e} \cos \omega$, $\sqrt{e} \sin \omega$), and the instrumental RV jitter ($\sigma_{\text{RV jitter}}$). The orbital period and central transit time were also set free, but were taken to be global parameters in common with the light curves. To account for systematic noise present in the RV time series, we used GPs with an exponential squared kernel (i.e., a Gaussian kernel):

$$k_{ij \text{ RV}} = c_2^2 \exp\left(-\frac{(t_i - t_j)^2}{\tau_2^2}\right), \quad (3)$$

where $t_i - t_j$ is the difference between two epochs or observations, c_2 is the amplitude of the exponential squared kernel, and

³ <https://github.com/hpparvi/ldtk>

⁴ <https://github.com/California-Planet-Search/radvel>

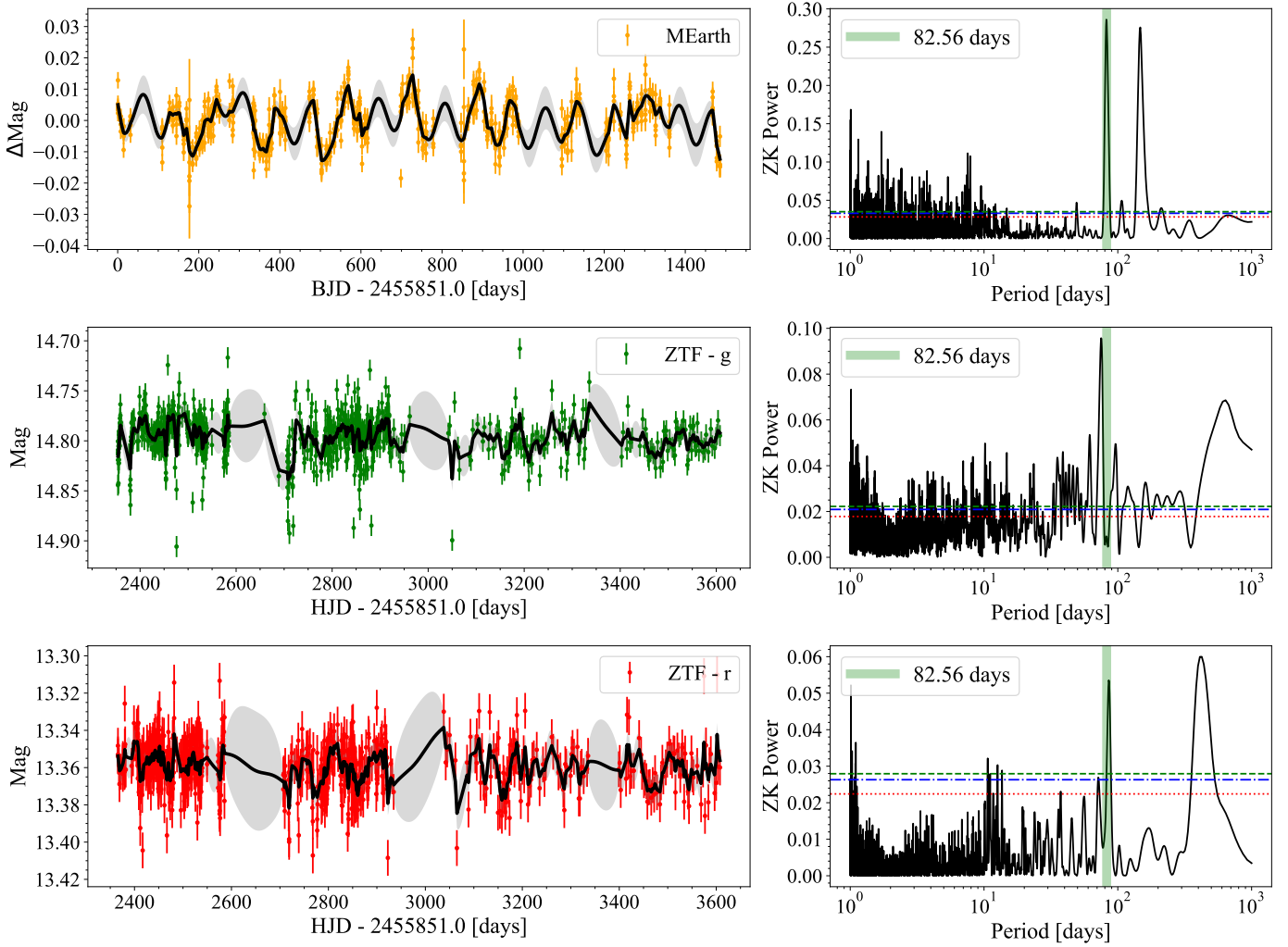


Fig. 3. Ground-based long-term photometric observations of TOI-2136 using MEarth and ZTF, and their respective GLS periodograms (Zechmeister & Kürster 2009) for each data set. The stellar photometric variability is detected with a stronger signal in MEarth data. Fitting a kernel with a periodic term to all the data sets using Gaussian processes, we find a stellar rotation period of 82.56 ± 0.45 days.

τ_2 is a characteristic timescale. The constants c_2 and τ_2 were set as free parameters.

We tested how the fitted parameter values changed by fitting a single-planet Keplerian to the data without taking into account the red noise in the RV measurements, and by modeling the systematics using GPs. We then computed the model comparison metrics Bayesian information criterion (BIC, Schwarz 1978) and Akaike information criterion (AIC, Akaike 1974), defined by

$$\text{BIC} = k \ln n - 2 \ln \mathcal{L}_{\max}, \quad (4)$$

and

$$\text{AIC} = 2k - 2 \ln \mathcal{L}_{\max}. \quad (5)$$

Where n is the number of observed data points, k is the number of fitted parameters, and \mathcal{L}_{\max} is the maximized likelihood. We found that the RV model that includes GPs is slightly preferred, although the difference between the criterion values for the RV model including GPs and without GPs (ΔBIC and ΔAIC) is not large enough to select the single Keplerian model plus the GPs model according to the criteria of Raftery (1995) ($\Delta\text{BIC} < 2.0$ and $\Delta\text{AIC} < 2.5$).

5.3. Joint fit

For the joint fit of all the data available, we employed a Bayesian approach. We performed an uninformative transit search in the TESS time series (from Sectors 26 and 40) using Transit Least Squares (TLS, Hippke & Heller 2019). From TLS we obtained an estimate of the period and epoch of the central time of the transit with their respective uncertainties; we used these values as priors for the joint fit. Then, we implemented a Markov chain Monte Carlo (MCMC) to explore the posterior distribution of the parameters using emcee (Foreman-Mackey et al. 2013). With emcee we evaluated a likelihood plus a prior function for several iterations. The likelihood function was the sum of the log likelihood for each transit time series plus a log likelihood for the star LD coefficients and the log likelihood of the RV observations. A total of 26 free parameters were sampled in the joint fit. The fitting process started with the optimization of the global likelihood with PyDE⁵, and once the optimization process ended we used the optimal parameter distribution to start the MCMC. The MCMC consisted of 200 chains and ran for 2000 iterations as a burn-in, and the main MCMC ran for 8000 iterations.

⁵ <https://github.com/hpparvi/PyDE>

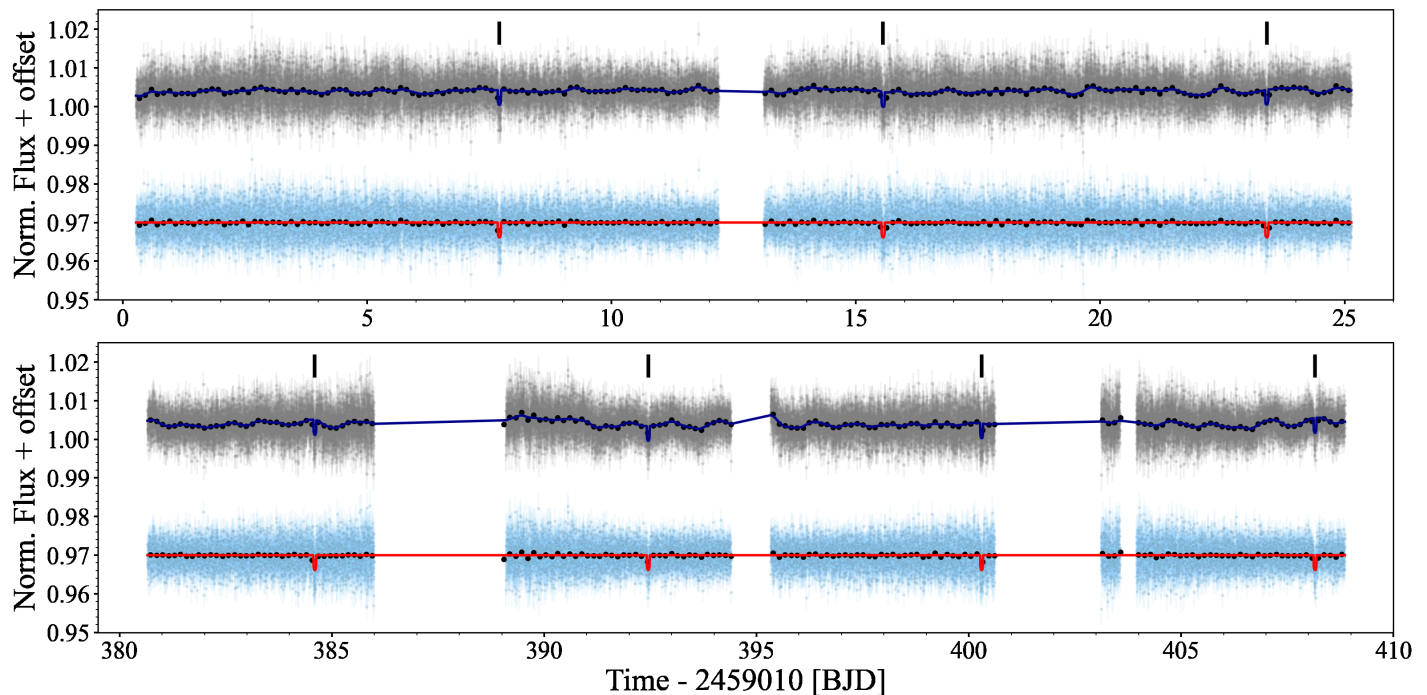


Fig. 4. TOI-2136 light curves of TESS Sector 26 (*top panel*) and Sector 40 (*bottom panel*). The black points represent TESS binned photometry. The gray points are the TESS PDCSAP data. The best fitting transit model including systematic effects is shown in dark blue. Light-blue points are the TESS data after correcting the systematic effects. The best fitting transit model is shown in red. Transit events of TOI-2136b are marked with black vertical lines.

The final parameter values and their respective uncertainties were computed from the posterior distributions, adopting the median as parameter values and using the 1σ levels of the distribution as uncertainties. The prior functions and parameter limits used in the joint fit are in Table C.1. Figure C.1 shows the correlation plots of the fitted orbital parameters, excluding the limb-darkening coefficients and parameters related to the modeling of red noise present in the data.

6. Results

We present the fitted and derived parameters for the joint fit of the data with and without the use of GPs to model the systematics present in the RV measurements (Table 3). For the case of the fit without GPs to model the red noise, the RV jitter parameter increased by a factor of two when compared to the GP modeling results; an expected result since there was no component to model the red noise in this approach. For the non-GP approach, we find an upper limit for the planetary mass of $8.3 M_{\oplus}$ while for the fit including GPs in the RV modeling, we find $M_p < 9.9 M_{\oplus}$. As discussed in Sect. 5.2, the fits performed with and without using GPs for the RV measurements are not statistically different according to the difference between the model selection metrics estimated using the BIC and AIC definitions. We choose to adopt the results of the fit using GPs, since this approach takes into account the red noise in the data and the results provide a more conservative upper mass limit for the mass of the planet.

Figure 4 presents the photometric light curves from TESS Sectors 26 and 40, and the best fitting model (including systematic effects) found by the joint fit. Figure 5 shows the detrended and phase-folded TESS light curves for both TESS sectors analyzed in this work. Figure 6 shows the multicolor ground-based transit follow-up observations of TOI-2136b used in the joint fit.

Finally, Fig. 7 presents the IRD RV measurements and the best fitting model.

Our joint fit finds that TOI-2136b has a period of $P = 7.851925 \pm 0.000016$ days and an orbital eccentricity consistent with 0. We also determined that TOI-2136b has a radius of $R_p = 2.20 \pm 0.07 R_{\oplus}$ and an equilibrium temperature of $T_{\text{eq}} = 378 \pm 13$ K. We note TOI-2136 light curves of TESS Sector 26 have a nontrivial bias by sky background noise correction. As this bias is about 1.54% in the data of Sector 26, it is only about 0.19% in our results, which is below the size of our error bars.

From Fig. 7, it is clear that the RV measurements present some variability, the origin of which it is not clear. Looking at the long-term photometric follow-up (see Sect. 4.2) and the TESS light curves, the star seems to present a low level of activity in the form of flares that could affect the precision of the RV measurements. To be conservative, we decided to fit a simple single-planet model to the RV data and assume that the rest of the signal is attributable to systematic noise modeled by the GPs. Figure 8 presents the posterior distribution for the planetary mass. The upper-limit (95% confidence level) for the mass of TOI-2136b is $9.9 M_{\oplus}$. The planetary mass estimate from the joint fit is of $M_p = 4.7^{+3.1}_{-2.6} M_{\oplus}$.

The left panel in Fig. 9 shows the median mass and radius values of TOI-2136b derived in this work compared to known transiting exoplanets with precisely determined bulk properties. The parameters for the known exoplanets were taken from TEPcat database (Southworth 2011). We only show planets having mass measurements with a precision better than 30%. In the panel, we show theoretical mass-radius relationships for planets with five different bulk compositions, including bare rocky planets with an Earth-like composition (32.5% Fe plus 67.5% MgSiO_3), 100% water worlds, and Earth-like rocky

Table 3. Planet parameters.

Parameter	Joint fit – no GPs for RVs	Joint fit with GPs
Fitted orbital and transit parameters		
R_p/R_*	$0.0588^{+0.0006}_{-0.0006}$	$0.0588^{+0.0006}_{-0.0006}$
T_c (BJD)	$2459214.00325^{+0.00043}_{-0.00040}$	$2459214.00322^{+0.00045}_{-0.00042}$
P (days)	$7.851925^{+0.000015}_{-0.000016}$	$7.851925^{+0.000016}_{-0.000016}$
ρ_* (g cm ⁻³)	$11.33^{+0.69}_{-0.67}$	$11.38^{+0.70}_{-0.71}$
b	$0.463^{+0.063}_{-0.072}$	$0.462^{+0.067}_{-0.078}$
$\sqrt{e} \cos(\omega)$	$0.070^{+0.284}_{-0.270}$	$-0.086^{+0.290}_{-0.251}$
$\sqrt{e} \sin(\omega)$	$-0.024^{+0.155}_{-0.143}$	$-0.029^{+0.160}_{-0.145}$
$\gamma_0 - \langle \gamma_0 \rangle$ (m s ⁻¹)	$-1.89^{+1.23}_{-1.22}$	$0.10^{+2.53}_{-2.49}$
K (m s ⁻¹)	$2.53^{+1.92}_{-1.48}$	$3.29^{+2.14}_{-1.80}$
σ_{RV} (m s ⁻¹)	$5.89^{+1.18}_{-1.08}$	$2.49^{+1.63}_{-1.56}$
Derived orbital parameters		
e	$0.07^{+0.11}_{-0.05}$	$0.07^{+0.09}_{-0.05}$
ω (deg)	$-8.32^{+120.09}_{-113.43}$	$-18.24^{+160.32}_{-140.29}$
a/R_*	$33.30^{+0.66}_{-0.67}$	$33.35^{+0.67}_{-0.71}$
i (deg)	$89.20^{+0.11}_{-0.09}$	$89.20^{+0.12}_{-0.09}$
Fitted LD coefficients		
u_1 TESS	$0.29^{+0.02}_{-0.02}$	$0.28^{+0.02}_{-0.02}$
u_2 TESS	$0.26^{+0.03}_{-0.04}$	$0.26^{+0.04}_{-0.04}$
$u_1 g$	$0.52^{+0.04}_{-0.04}$	$0.52^{+0.03}_{-0.03}$
$u_2 g$	$0.30^{+0.06}_{-0.06}$	$0.29^{+0.05}_{-0.05}$
$u_1 r$	$0.54^{+0.03}_{-0.03}$	$0.54^{+0.03}_{-0.03}$
$u_2 r$	$0.22^{+0.06}_{-0.05}$	$0.22^{+0.06}_{-0.05}$
$u_1 i$	$0.34^{+0.02}_{-0.01}$	$0.34^{+0.02}_{-0.02}$
$u_2 i$	$0.29^{+0.03}_{-0.03}$	$0.29^{+0.04}_{-0.04}$
$u_1 z_s$	$0.25^{+0.01}_{-0.01}$	$0.24^{+0.02}_{-0.02}$
$u_2 z_s$	$0.29^{+0.03}_{-0.03}$	$0.30^{+0.03}_{-0.03}$
Fitted GP parameters		
$\log(c_1)$ TESS S26	$-7.56^{+0.10}_{-0.10}$	$-7.54^{+0.11}_{-0.10}$
$\log(\tau_1)$ TESS S26	$-1.61^{+0.18}_{-0.18}$	$-1.60^{+0.18}_{-0.18}$
$\log(c_1)$ TESS S40	$-7.13^{+0.12}_{-0.11}$	$-7.13^{+0.12}_{-0.11}$
$\log(\tau_1)$ TESS S40	$-1.21^{+0.16}_{-0.16}$	$-1.21^{+0.16}_{-0.16}$
c_2 IRD	–	$6.78^{+2.63}_{-1.88}$
τ_2 IRD	–	$13.72^{+21.52}_{-6.69}$
Derived planet parameters		
R_p (R_\oplus)	2.21 ± 0.07	2.20 ± 0.07
M_p (M_\oplus)	$3.7^{+2.8}_{-2.2}$	$4.7^{+3.1}_{-2.6}$
	< 8.3 (95%)	< 9.9 (95%)
ρ_p (g cm ⁻³)	$1.91^{+1.46}_{-1.13}$	$2.39^{+1.61}_{-1.36}$
g_p (m s ⁻²)	$7.5^{+5.7}_{-4.4}$	$9.4^{+6.3}_{-5.3}$
a (au)	0.0533 ± 0.0015	0.0533 ± 0.0015
T_{eq} (K)	378 ± 13	378 ± 13
$\langle F_p \rangle$ (10^5 W m ⁻²)	0.066 ± 0.010	0.066 ± 0.010
S_p (S_\oplus)	4.83 ± 0.73	4.82 ± 0.72

Notes. T_{eq} computed assuming a Bond albedo of 0.3.

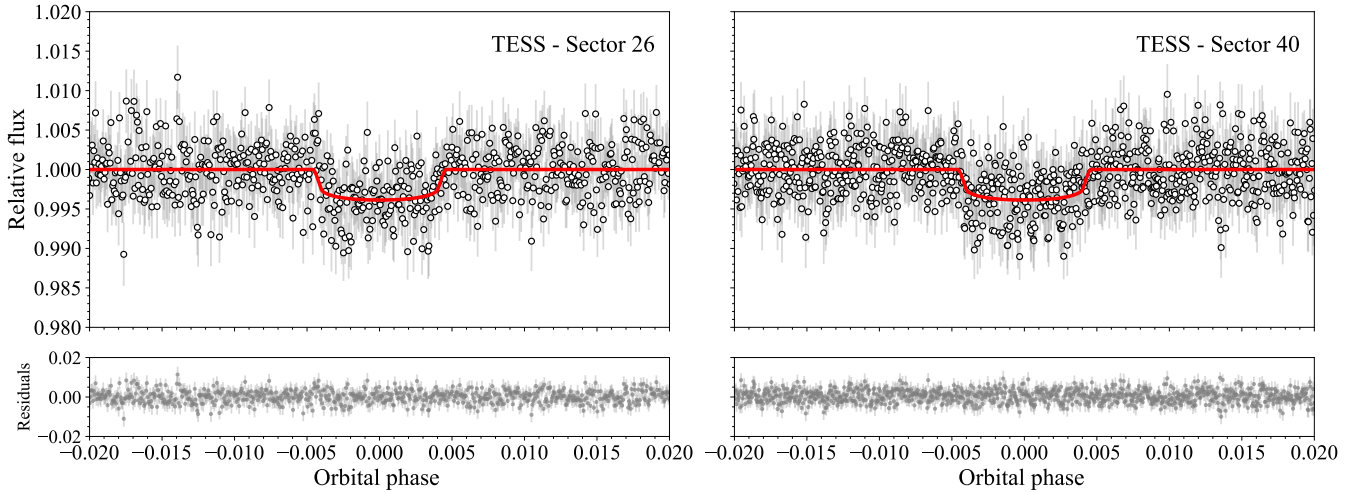


Fig. 5. TESS sector 26 and 40 folded light curves, and the best fitting model. The photometric variability present in both TESS sectors have been removed.

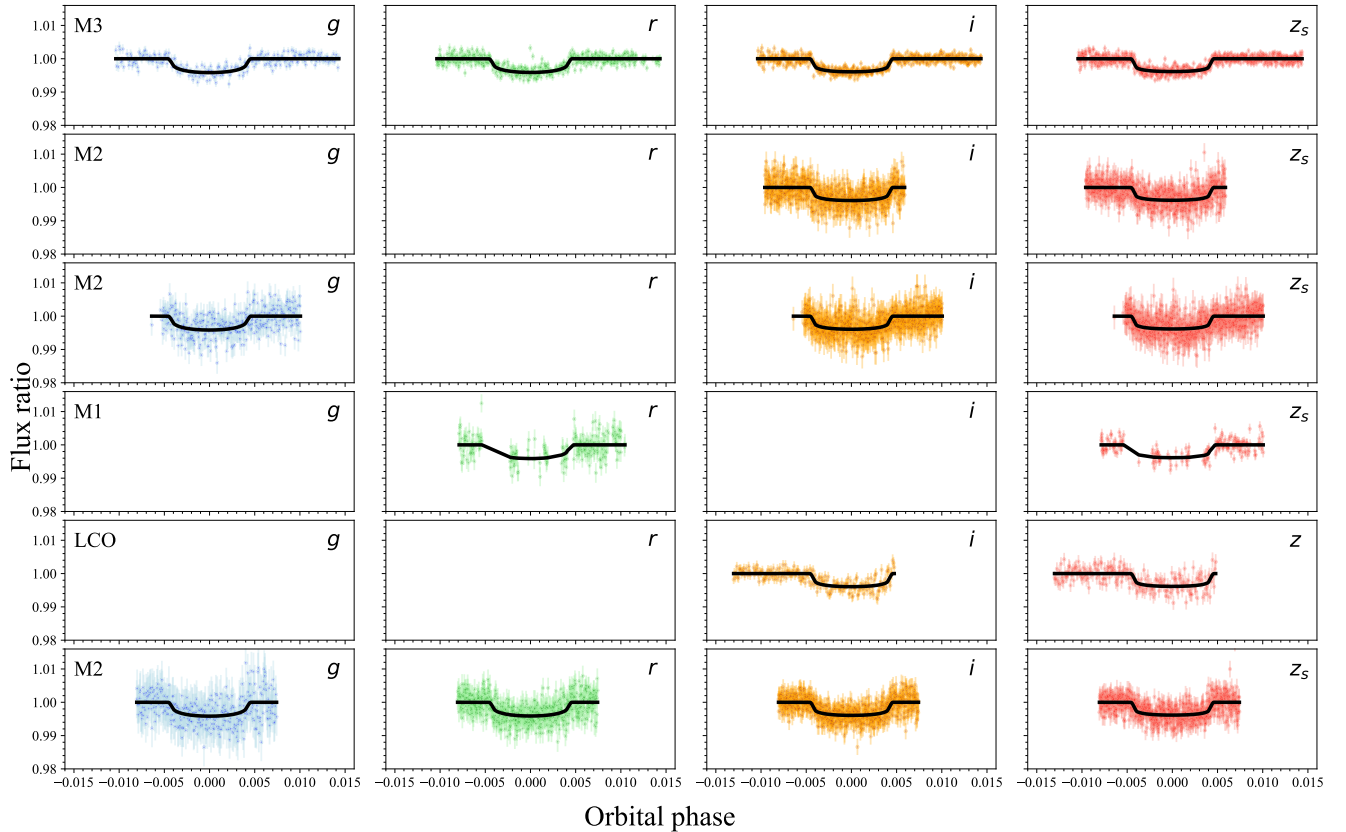


Fig. 6. Ground-based transit observations of TOI-2136b. The figure shows the photometry obtained with MuSCAT (M1), MuSCAT2 (M2), MuSCAT3 (M3), and LCO facilities, and the best fitting model (black line). The systematic effects presented in each data set were fitted and subtracted before performing the joint fit. The dates for each observation are in Table 1. The oldest observations are in the *top* row, and the most recent are in *bottom* row.

cores with 0.1%, 1%, and 5% H+He envelopes. For bare rocky planets and water worlds, we take the data provided by Zeng et al. (2019) and use the intermediate value of the radii for 300 K and 500 K at a given mass. For the envelopes, we numerically integrate convective structure with Saumon et al. (1995)'s adiabats for a 75% H + 25% He mixture up to 1 kbar. Above this we use the analytic solutions from Matsui & Abe (1986)

for radiative atmospheres with equilibrium temperature of 400 K (see Kurosaki & Ikoma 2017, for details). Since the mass measurement of TOI-2136b presents a large uncertainty, when compared to models, TOI-2136b could be a low-density planet ($\rho_p \sim 1.0 \text{ g cm}^{-3}$) with a significant gaseous envelope in its low mass limit, or a rocky planet with an Earth-like composition and a 1% H₂ gaseous envelope in its upper mass limit.

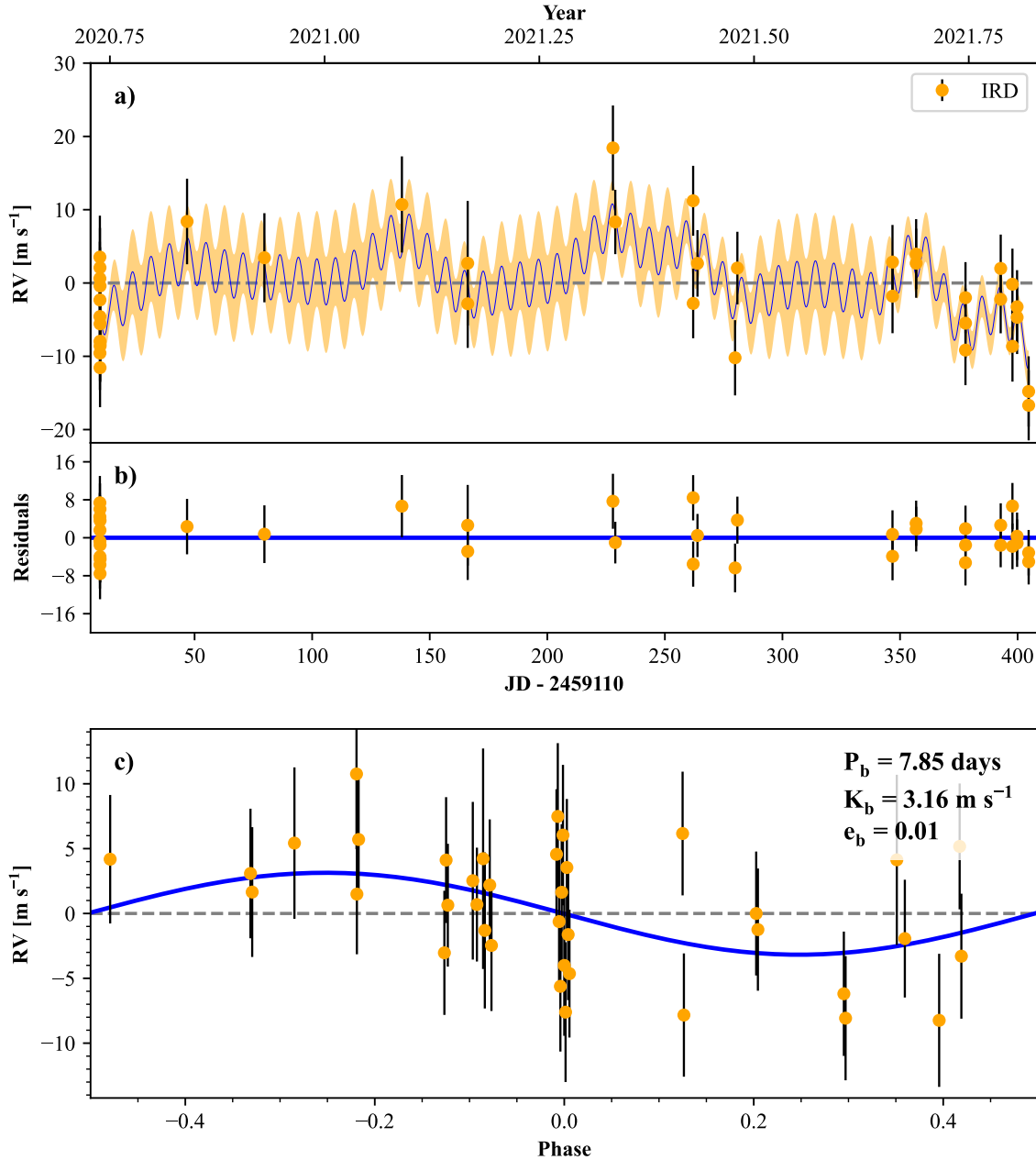


Fig. 7. Radial velocity measurements of TOI-2136 taken with the IRD. The *top panel* presents the time series and best fitting model with the use of GPs to model the red noise (the fit without GPs is presented in Fig. C.2). The *middle panel* presents the residuals of the fit. The *bottom panel* shows the RV measurements in phase.

To evaluate the potential for follow-up observations with the aim of detecting the presence of an atmosphere in TOI-2136b, we computed the transmission spectroscopy metric (TSM) of this planet following [Kempton et al. \(2018\)](#). This metric is an estimate of the S/N achievable with the Near Infrared Imager and Slitless Spectrograph (NIRISS) instrument on the *James Webb* Space Telescope (*JWST*, [Gardner et al. 2006](#)), using ten hours of observation time. For the median mass value of TOI-2136b, we find a TSM of ~ 93 , and using the low and upper 1σ mass limits, we find TSM values of 212 (for $M_{p,low} = 2.04 M_{\oplus}$) and 56 (for $M_{p,up} = 7.75 M_{\oplus}$). We compared the TSM obtained with the median mass value with those of previously discovered exoplanets that have a radius smaller than $4 R_{\oplus}$ (see Fig. 9, right panel). We find that TOI-2136b has a relatively large TSM when compared with other planets with radii around $2 R_{\oplus}$ and a similar

J-band magnitude, and it also has an almost comparable TSM to hyccean planets (labeled in Fig. 9, right panel), making it a very suitable target for *JWST* atmospheric exploration.

7. Transmission spectroscopy

TOI-2136b likely has a H/He gaseous envelope from the mass-radius relationship. Using the IRD high-resolution spectra taken during transit, it is possible to search for an extended atmosphere of H/He. In this particular case focusing on the helium infrared triplet at 10829.08, 10830.25, and 10830.34 Å in air.

Around the He absorption lines, there are OH⁻ emission lines and telluric H₂O absorption lines that sometimes contaminate a possible planetary absorption feature ([Nortmann et al. 2018](#); [Salz et al. 2018](#); [Alonso-Floriano et al. 2019](#);

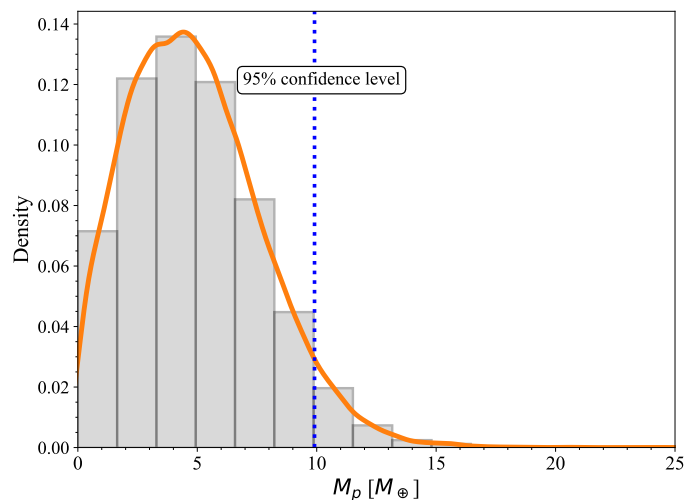


Fig. 8. Planetary mass posterior distribution. We can place an upper limit on the mass of TOI-2136b of $9.9 M_{\oplus}$ (95% confidence level, dashed blue line).

Palle et al. 2020). To obtain the transmission spectrum, we corrected for these telluric absorption lines, as well as for the stellar RV shift and the planet RV. To obtain an OH^- emission model for our data, we fitted the strongest OH^- emission line with a Gaussian to estimate the central wavelength, and we compared it with previous works (Palle et al. 2020; Orell-Miquel et al. 2022) to estimate the central wavelength of other OH^- emission lines. The region of the strongest emission was then masked.

The telluric H_2O absorption lines were removed from the observed spectrum using a telluric standard star, and the correlation between line depth and airmass, following Kawauchi et al. (2018). We used a telluric standard star observed on 29 September 2020 as there were no observations of a telluric standard star on the same day. We created a telluric template spectrum by fitting the spectrum of the telluric standard star with the first spectrum of our transit time series. Figure 10 shows that, with this correction, the telluric absorption lines are eliminated from the spectra down to the noise level.

The stellar system is shifted due to the Earth’s rotation and revolution. These stellar RVs in the barycentric reference frame were estimated from RA, Dec and observational time (UT) with IRAF task `rvcorrect`. We shifted the stellar spectra and corrected them (see the left panel of Fig. 11). However, these spectra were not aligned with the model stellar spectra at air and vacuum wavelength.

To correct the wavelength, we estimated the additional stellar RV shift from the model spectra with least-squares deconvolution (LSD) method (Donati et al. 1997). Using this method, we created a line profile of the observed spectra in velocity space, convolved with the delta functions obtained from the center wavelength and the strength of atomic and molecular lines created in the model stellar atmosphere (Watanabe et al. 2020; Collier Cameron et al. 2010). We used the observed spectra in three orders around He lines ($10681.55\text{--}10979.75 \text{ \AA}$), and the 426 absorption lines, assuming the stellar temperature and $\log g$ from VALD.

After we created a line profile of each frame, we fitted it with a flat line and a Gaussian to estimate the shifted velocity from the model. We obtained a value of -52.6 km s^{-1} of the shifted velocity, and we shifted the observed spectrum by that value

(see right panel of Fig. 11). As a result, the observed spectra are consistent with the model spectra in air wavelength.

After the above correction, we created the template stellar spectrum to combine all out-of-transit spectra (five frames) and divided each frame by it. The planet signals were shifted in wavelength by planet orbital motion. We calculated the RV of the planet at each frame with Eq. (8) of Khalafinejad et al. (2017). The RV of TOI-2136b changes from $+1.7 \text{ km s}^{-1}$ to -1.5 km s^{-1} during transit. We shifted each frame by the calculated velocity and combined all frames during transit to create the final transmission spectrum.

Figure 12 shows the final transmission spectrum of TOI-2136 over the He I triplet spectral region. A possible He I signal at the triplet wavelengths (10829.081 , 10830.25 and, 10830.34 \AA) is not apparent in the data, although some spectral features fall near the expected wavelengths. In the presence of an atmospheric outflow caused by stellar wind, the He I absorption lines could be blue-shifted, but red-shifted helium absorption has been detected in HAT-P-32b (Czesla et al. 2022) and TOI 560.01 (Zhang et al. 2022). Therefore, in order to estimate the detectability of He I, we searched for the helium absorption line in a large range around the theoretical center of wavelength, and assumed that the largest absorption feature in this region of the spectra could be associated with the He triplet.

To further explore this spectral feature at $\sim 10830.8 \text{ \AA}$ (which is red-shifted by $\sim 15 \text{ km s}^{-1}$ from the expected wavelength of the strongest triplet line), we fitted a Gaussian with the central wavelength, depth, sigma, and baseline as free parameters, using the MCMC algorithm `emcee`. We used a normal prior for the central wavelength between 10829.21 and 10831.38 \AA , which correspond to about $\pm 30 \text{ km s}^{-1}$ from the center of the stronger He I lines, and a uniform prior for sigma between 0 and 0.344, which corresponds to the thermal broadening at a temperature of 30 000 K. We obtained a statistically nonsignificant signal with a depth of $2.2 \pm 1.1\%$, a sigma of 0.08 ± 0.06 and an EW of $4.3 \pm 2.0 \text{ m\AA}$ (Fig. 12). However, this spectral feature is only detected at the 2σ level, and cannot be identified as significant or planetary in origin. Thus, our data only allow us to place upper limits to the presence of He I in TOI-2136b’s atmosphere. Any possible He I signal would have an EW $< 7.8 \text{ m\AA}$, with 95% confidence, and an absorption signal $< 1.44\%$, with 95% confidence.

8. Discussion

In previous sections, we validated the planetary nature of TOI-2136.01. Figure 13 shows its location in a period–radius diagram, together with all other known planets around M-type stars whose radius measurements were better than 8% precise. In the figure, the black line indicates the radius valley measured around low-mass stars (Cloutier & Menou 2020), which is consistent with the gas-poor formation model (Lopez & Rice 2018). The dashed black line indicates the radius valley measured around Sun-like stars (Martinez et al. 2019), which is consistent with the photoevaporation (Lopez & Rice 2018) and core powered mass-loss models (Gupta & Schlichting 2019). TOI-2136b lies above both radius valleys, so it would be expected to be non-rocky in composition and possess some fraction of volatile elements in its atmosphere. This prediction is consistent with our results in the mass-radius diagram (Fig. 9).

Additionally, thanks to the multiple ground-based photometric observations, we obtained the radius of TOI-2136b with an uncertainty of $\sim 3\%$. Presently, there are only a few

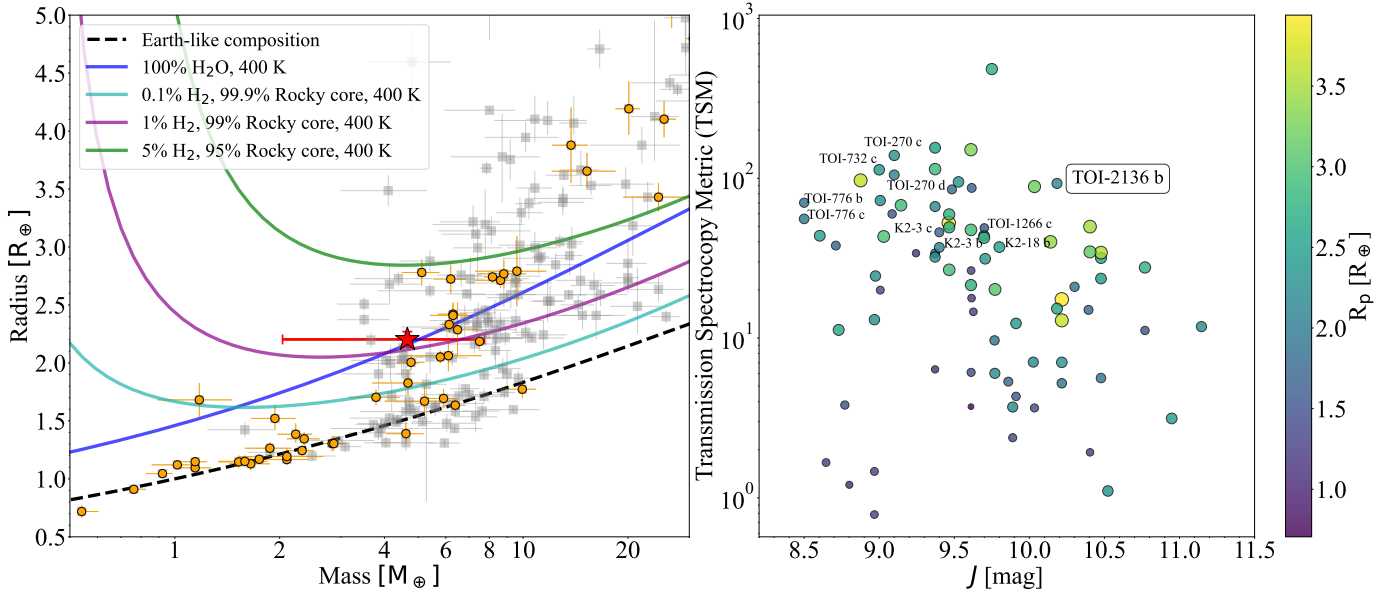


Fig. 9. Mass-radius diagram and TSM for TOI-2136b. *Left panel:* mass and radius values for known exoplanets (data taken from TEPcat database, Southworth 2011). The position of TOI-2136b is marked by the red star. In the figure we only show planets having mass measurements with a precision better than 30%. The lines represent theoretical mass–radius relationships for five different bulk compositions, including bare rocky planets with an Earth-like composition, pure water worlds, and Earth-like rocky cores with 0.1%, 1%, and 5% H+He envelopes with radiative equilibrium temperature of 400 K (see text). The orange points represent planets orbiting M-type stars, and the gray squares are planets around other types of stars. *Right panel:* apparent magnitude in *J*-band vs. the TSM by Kempton et al. (2018) for TOI-2136b and known transiting planets with radii smaller than $4 R_{\oplus}$. The other labeled planets are the hycean candidates.

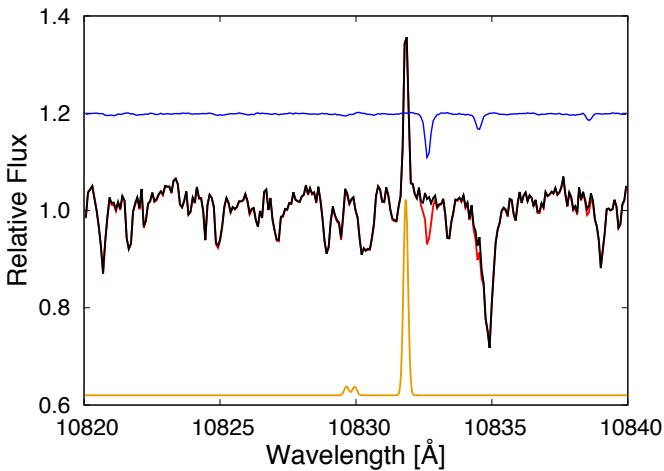


Fig. 10. Example of a single IRD spectrum of TOI-2136 around the helium triplet before (red) and after (black) the telluric correction. Blue and orange lines indicate the telluric absorption lines and OH[−] emission lines. This figure indicates the position of the emission lines and the correction of the telluric absorption lines down to the noise level.

planets measured with such high precision around M-type stars. Among them, LTT3780c ($2.42 \pm 0.10 R_{\oplus}$, $6.29 \pm 0.63 M_{\oplus}$, $2.88 \pm 0.28 S_{\oplus}$; Nowak et al. 2020), L321-32d ($2.13 \pm 0.06 R_{\oplus}$, $4.78 \pm 0.43 M_{\oplus}$, $3.74 \pm 0.61 S_{\oplus}$; Van Eylen et al. 2021), and TOI-776c ($2.02 \pm 0.14 R_{\oplus}$, $5.30 \pm 1.80 M_{\oplus}$, $4.9 \pm 0.2 S_{\oplus}$; Luque et al. 2021) have similar parameters to TOI-2136b. These planets are candidates to be hycean worlds, which are composed of water-rich interiors with massive oceans underlying H₂-rich atmospheres (Madhusudhan et al. 2021). Hycean planets have received attention since a possible biosignature might be

detectable in their atmospheres with a relatively modest amount of JWST observation time.

We also obtained a transmission spectrum of the planet around the He I 10 830 Å absorption lines. In other works, helium absorption is detected on the atmosphere of three sub-Neptunes (TOI-560.01; Zhang et al. 2022, GJ3470b; Palle et al. 2020, and GJ1214b; Orell-Miquel et al. 2022). To compare our results with previous He I detections in the literature, we calculated the equivalent height of the atmosphere normalized by the scale height. This value for TOI-2136b is <114 , adopting a mass of $M_p = 4.7 M_{\oplus}$ and is comparable with the value of TOI 560.01 (Fig. 4 in Orell-Miquel et al. 2022). TOI 560.01 has a detection of He I (Zhang et al. 2022), meaning that TOI-2136b likely has a less extended or a less rich H/He atmosphere than TOI 560.01. For GJ1214b and GJ3470b, the ratio between the equivalent height and atmospheric scale height is ~ 57 for both planets, about half of our upper limit for TOI-2136b. Despite our non-detection of He I on this planet, the atmospheric metrics and its place in the mass-radius space suggest that TOI-2136b is an ideal candidate to search for atmospheric signals using ground-based high-resolution spectroscopic observations taken during a transit event and space-based observations with JWST.

During the submission and revision stage of this manuscript, two studies on the TOI-2136 system were announced: Gan et al. (2022); Beard et al. (2022). Gan et al. (2022) used TESS data, ground-based follow-up observations, and the Spirou instrument on the Canada-France-Hawaii Telescope (CFHT) to obtain RV measurements; the study presents a planetary radius of $2.19 \pm 0.17 R_{\oplus}$ and a mass of $6.4 \pm 2.4 M_{\oplus}$ for the transiting planet. Beard et al. (2022) used TESS data, ground-based follow-up observations, and the Hobby-Eberly Telescope (HET) instrument Habitable-Zone Planet Finder (HPF) instrument to obtain RV measurements. Beard et al. (2022) presents a planetary radius of $2.09 \pm 0.08 R_{\oplus}$ and a mass upper limit of $< 15 M_{\oplus}$, with

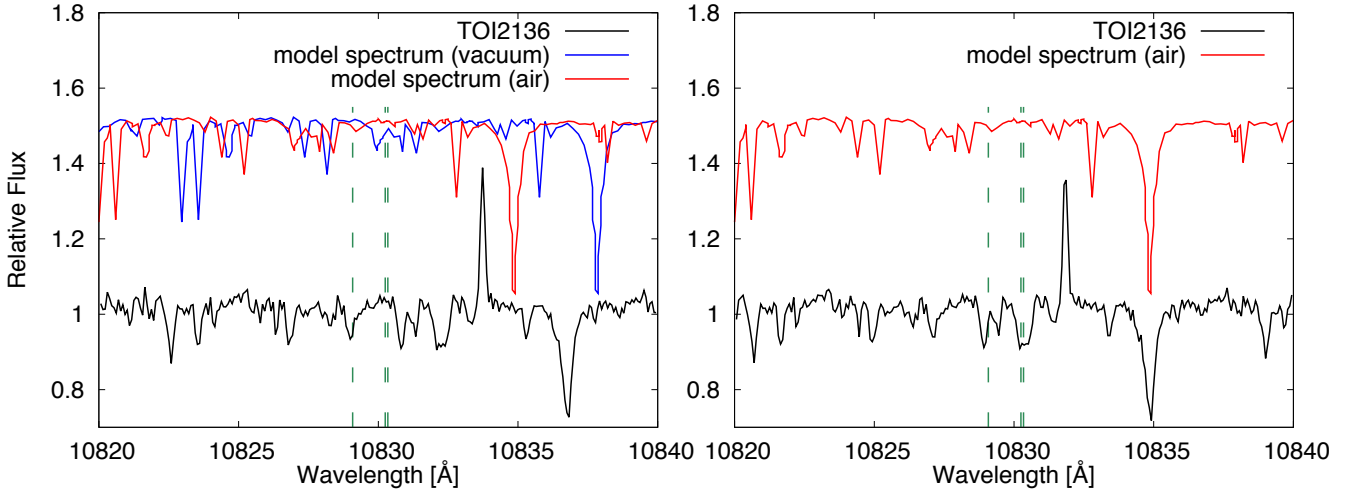


Fig. 11. Example of a single IRD spectrum of TOI-2136b (black) compared with the BT-Settl model (Allard 2014) spectrum in vacuum (blue) and air (red). The *left figure* indicates the raw IRD spectrum and the *right figure* indicates the IRD-spectrum corrected shifted velocity. Vertical green dash lines in the *left and right panels* indicate the center of the predicted He triplet in air.

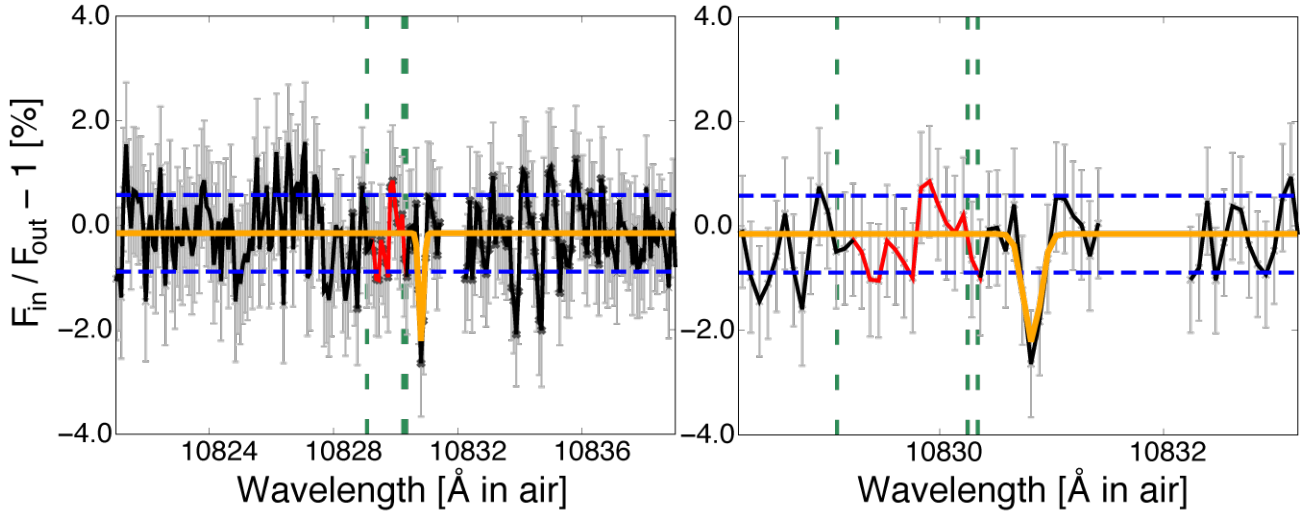


Fig. 12. Transmission spectrum of TOI-2136 around the He triplet. The blank region is where we mask the strongest OH⁻ emission. The red line is contaminated by the other shallower OH⁻ emission lines. The orange line is the best-fitted model with a Gaussian function. Vertical green dashed lines indicate the center of the predicted He triplet in air, and the horizontal blue lines are standard deviations between 10820 and 10840 Å.

a median planet mass of $4.64 M_{\oplus}$. Our results for the planetary mass and radius of TOI-2136b are consistent with both studies (taking into account the uncertainties).

9. Summary

In this work, we reported the follow-up efforts on the planet candidate TOI-2136b, a sub-Neptune planet orbiting around a nearby M3 dwarf. The initial observations made by TESS led to the detection of the transiting planet candidate. Multicolor, ground-based, follow-up photometric observations and RV measurements made it possible to confirm the planetary nature of the candidate and rule out false positives.

Using the IRD instrument on the Subaru 8.2 m telescope, we obtained high-resolution spectra and derived stellar parameters for the planet host star. We find that TOI-2136 is a nearby ($d = 33.361 \pm 0.019$ pc) M dwarf with $T_{\text{eff}} = 3373 \pm 108$ K, $[\text{Fe}/\text{H}] =$

0.02 ± 0.14 dex, and an estimated stellar mass of $M_{*} = 0.3272 \pm 0.0082 M_{\odot}$ and a radius of $R_{*} = 0.3440 \pm 0.0099 R_{\odot}$. We also measured the rotation period of the star using MEarth and ZTF long-term photometric observations, and found a value of $P_{\text{rot}} = 82.56 \pm 0.45$ days.

We analyzed the available transit and RV observations of the planet candidate fitting the data jointly. The transit observations and RV measurements from IRD were fitted simultaneously using an MCMC procedure that included Gaussian processes to model the systematic effects present in the TESS and RV measurements. We find that TOI-2136b has a planetary radius of $R_p = 2.20 \pm 0.07 R_{\oplus}$ and a mass of $M_p = 4.7^{+3.1}_{-2.6} M_{\oplus}$ (with an upper limit on the mass value of $< 9.9 M_{\oplus}$, with 95% confidence). Our model supports a circular orbit ($e = 0.07^{+0.09}_{-0.05}$) with an orbital period of 7.851925 ± 0.000016 days.

Finally, we observed a transit event using high-resolution spectroscopy in an effort to detect the existence of an extended atmosphere around this planet. In particular we focused on the

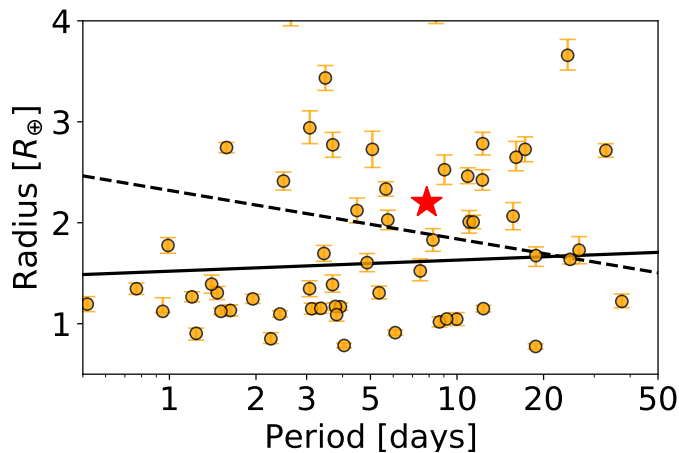


Fig. 13. Period–radius diagram for known exoplanets around M-type stars with radius measurements with a precision better than 8% (orange points). The red star represents TOI-2136b. The black line and the dashed line indicate the radius valley around low-mass stars (Cloutier & Menou 2020) and Sun-like stars (Martinez et al. 2019), respectively.

He I 10 830 Å absorption lines. We found no statistical significant extra absorption around the He I lines that could be attributable to the planet, but we can place an upper limit on the presence of a He-rich extended atmosphere: an absorption depth of $<1.44\%$ and an EW of $<7.8 \text{ mÅ}$, with 95% confidence.

TOI-2136b is a small sub-Neptune planet orbiting a relatively bright ($J = 10.1 \text{ mag}$) M star. Based on the TSM, this planet presents a very valuable opportunity for further atmospheric characterization studies and the search for additional planets in the system.

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Table A.1. Subaru IRD RV measurements

BJD [-2459000]	RV [ms ⁻¹]	σ_{RV} [ms ⁻¹]
119.71575590	0.440	4.250
119.72649620	3.360	4.980
119.73722720	-4.750	4.530
119.74796610	-9.740	4.260
119.75869830	-2.490	4.520
119.76943850	1.920	4.710
119.78016590	-8.130	4.720
119.79090390	-11.730	4.680
119.80164520	-0.580	4.560
119.81237620	-5.750	4.290
119.82310590	-8.760	4.140
156.80336260	8.220	5.190
189.68940840	3.270	5.460
248.16697680	10.540	5.990
276.14592460	2.520	8.070
276.16016820	-3.010	5.390
337.91216630	18.250	5.140
338.91002330	8.150	3.480
372.02252050	11.040	3.950
372.03322450	-2.960	3.930
373.86437420	2.500	3.670
389.85266460	-10.380	4.380
390.83389110	1.860	4.170
456.79522350	2.680	4.300
456.80944320	-1.970	4.290
466.85547440	3.760	3.960
466.86970380	2.520	3.880
487.82708040	-9.330	3.970
487.84130740	-2.180	4.050
487.85553810	-5.660	3.910
502.80575910	-2.400	3.790
502.81998640	1.810	3.770
507.80134910	-0.340	4.050
507.81558570	-8.810	4.010
509.77572120	-3.420	4.220
509.78996010	-4.860	4.230
514.69360020	-14.970	3.980
514.70783790	-16.860	3.970

Table B.1. Stellar rotation period fitted parameters, prior functions, and limits.

Parameter	Value
$\log B$	$\mathcal{U}(-12, 6)$
$\log C$	$\mathcal{U}(-12, 12)$
$\log L$	$\mathcal{U}(-12, 12)$
$\log P_{\text{rot}}$	$\mathcal{N}(1.6, 6.0)$
$\log j_{\text{jitter}}$	$\mathcal{U}(10^{-6}, 10)$

Notes. \mathcal{U} , \mathcal{N} represent Uniform and Normal prior functions, respectively.

Appendix A: Radial velocities

Appendix B: Stellar rotation priors

Appendix C: Light curve and radial velocity joint fit

Table C.1. Global fit parameters, prior functions, and limits.

Parameter	Value
(Fitted orbital and transit parameters priors)	
R_p/R_*	$\mathcal{U}(0.005, 0.35)$
T_c [BJD]	$\mathcal{U}(2459212.5, 2459215.5)$
P [days]	$\mathcal{U}(3.8, 11.8)$
ρ_* [g cm ⁻³]	$\mathcal{N}(11.33, 3.06)$
$b = (a/R_*) \cos(i) \left(\frac{1-e^2}{1+e \sin(\omega)} \right)$	$\mathcal{U}(0, 1)$
$\sqrt{e} \cos(\omega)$	$\mathcal{U}(-1, 1)$
$\sqrt{e} \sin(\omega)$	$\mathcal{U}(-1, 1)$
$\gamma_0 - \langle \gamma_0 \rangle$ [m/s]	$\mathcal{U}(-42.075, 37.925)$
K [m/s]	$\mathcal{U}(0, 110)$
σ_{RV} [m/s]	$\mathcal{U}(0, 10)$
(LD coefficients priors)	
$q_1 = (u_1 + u_2)^2$	$\mathcal{U}(0, 1)$
$q_2 = 0.5u_1/(u_1 + u_2)$	$\mathcal{U}(0, 1)$
(Fitted GP parameters priors)	
$\log(c_1)$ TESS S26	$\mathcal{U}(-8, 6)$
$\log(\tau_1)$ TESS S26	$\mathcal{U}(-2.65, 6)$
$\log(c_1)$ TESS S40	$\mathcal{U}(-8, 6)$
$\log(\tau_1)$ TESS S40	$\mathcal{U}(-2.65, 6)$
c_2 IRD	$\mathcal{U}(0, 100)$
τ_2 IRD	$\mathcal{U}(10^{-3}, 150)$

Notes. \mathcal{U} , \mathcal{N} represent Uniform and Normal prior functions, respectively.

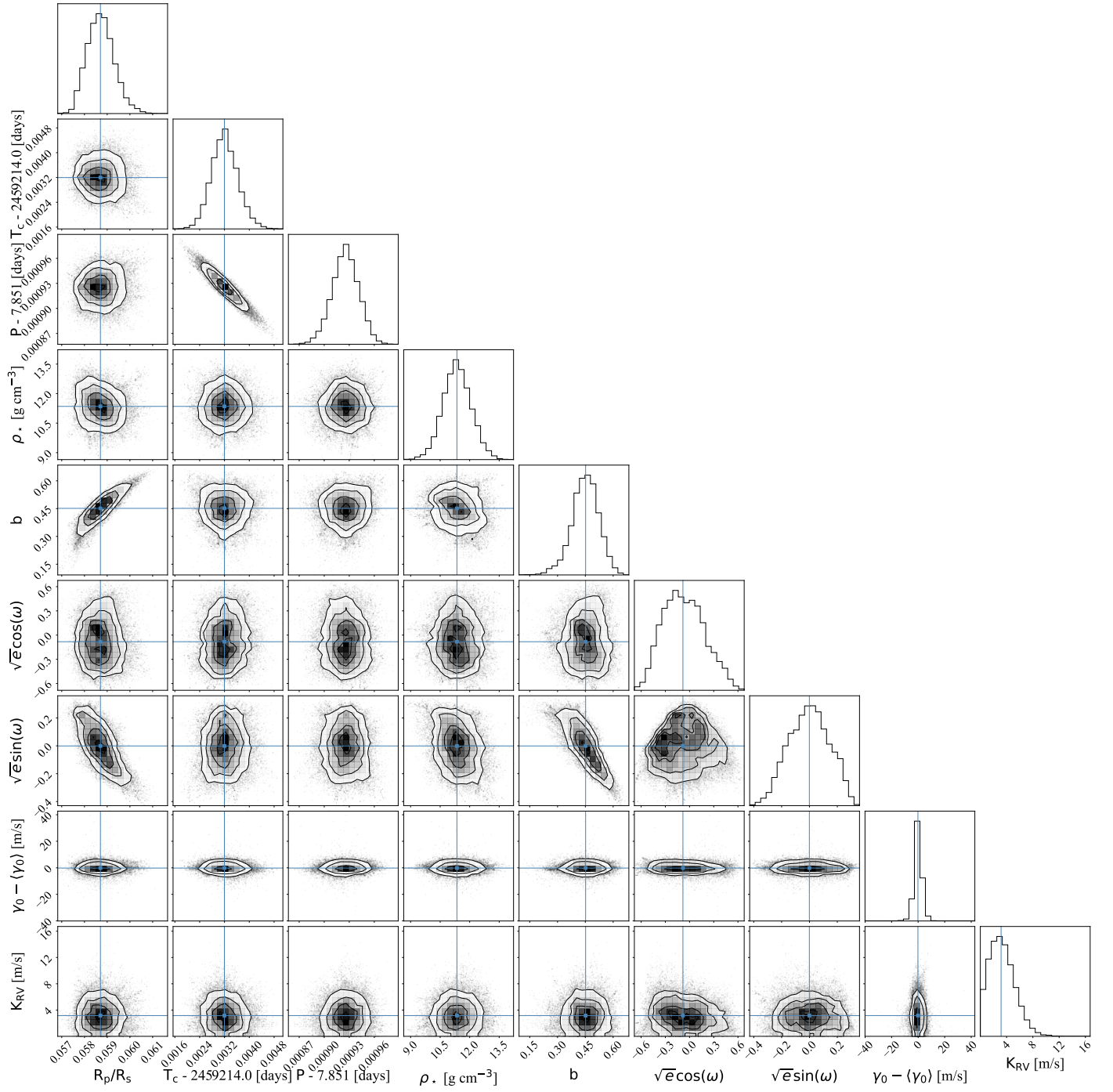


Fig. C.1. Correlation plot for the fitted orbital parameters. Limb-darkening coefficients and systematic effect parameters were intentionally left out for easy viewing. The blue lines give the median values of the parameters.

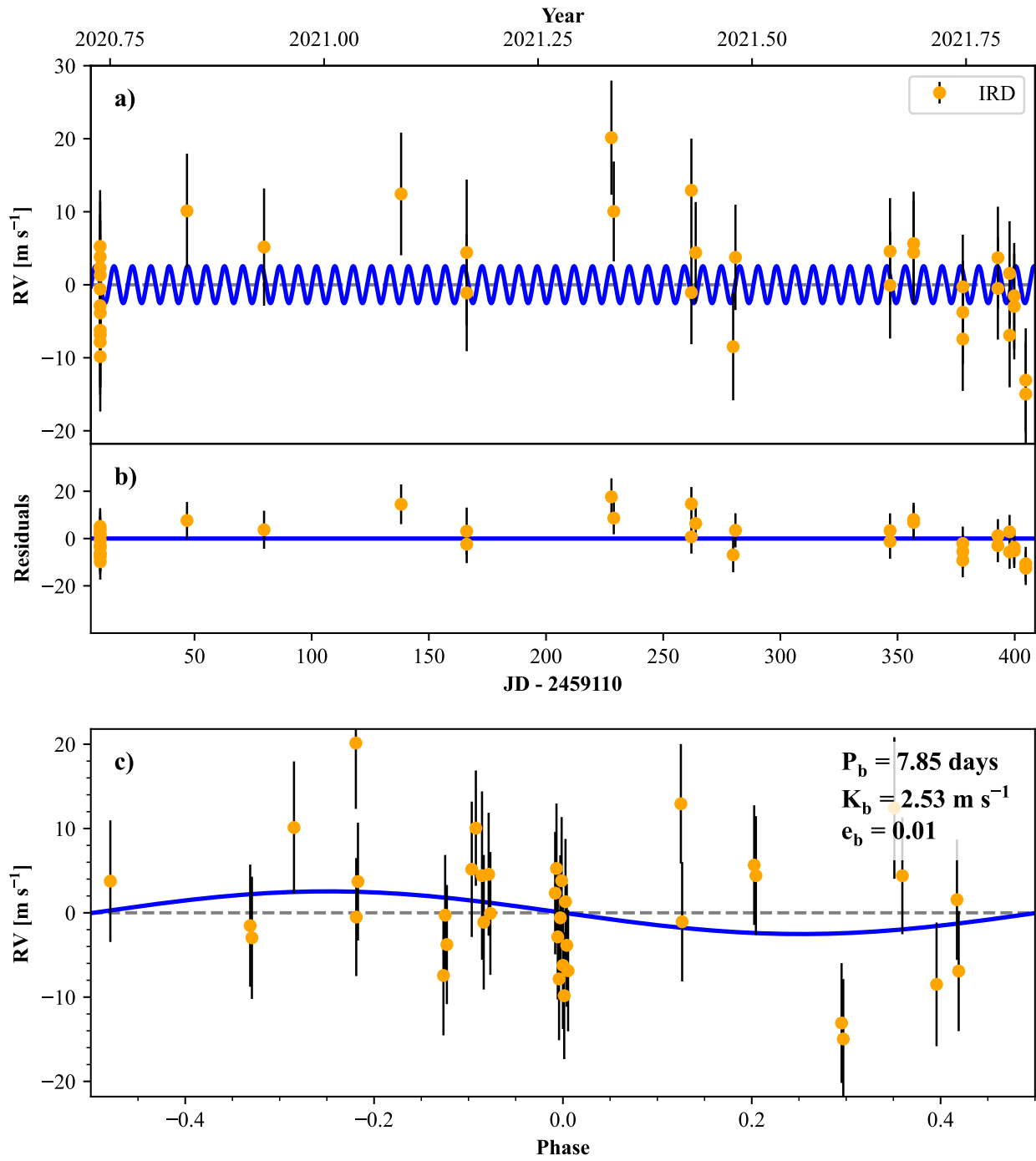


Fig. C.2. RV measurements of TOI-2136 taken with IRD. The top panel presents the time series and best fitting model without GPs. The middle panel presents the residuals of the fit. The bottom panel shows the RV measurements in phase after subtracting the red noise.

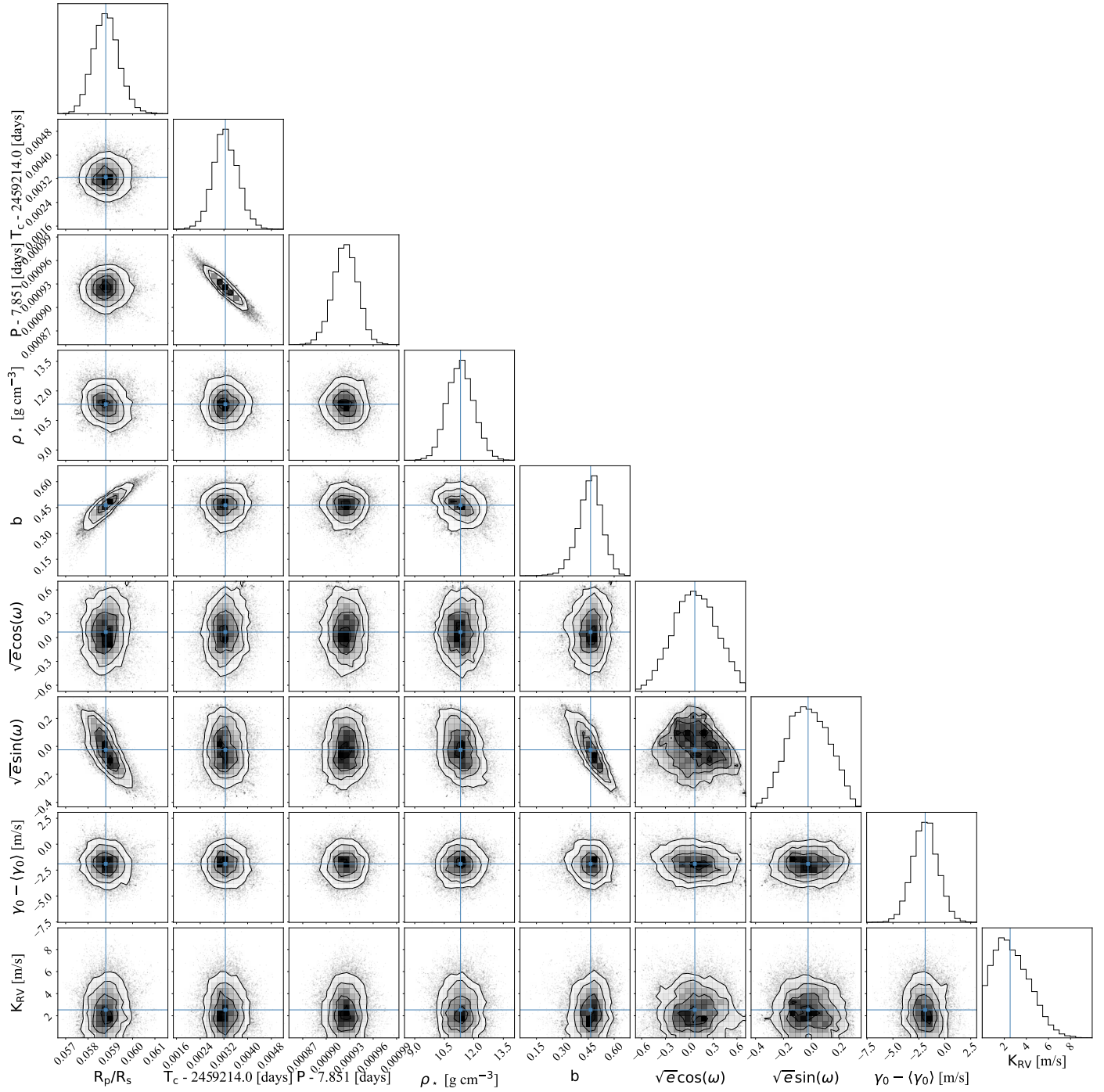


Fig. C.3. Correlation plot for the fitted orbital parameters without GPs. Limb-darkening coefficients and systematic effect parameters were intentionally left out for easy viewing. The blue lines give the median values of the parameters.