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# The ALMA Survey of 70 $\mu\text{m}$ Dark High-mass Clumps in Early Stages (ASHES).

## XI. Statistical Study of Early Fragmentation

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### Abstract

Fragmentation during the early stages of high-mass star formation is crucial for understanding the formation of high-mass clusters. We investigated fragmentation within 39 high-mass star-forming clumps as part of the Atacama Large Millimeter/submillimeter Array Survey of 70  $\mu\text{m}$  Dark High-mass Clumps in Early Stages (ASHES) survey. Considering projection effects, we have estimated core separations for 839 cores identified from the continuum emission and found mean values between 0.08 and 0.32 pc within each clump. We find compatibility of the observed core separations and masses with the thermal Jeans length and mass, respectively. We also present subclump structures revealed by the 7 m array continuum emission. Comparison of the Jeans parameters using clump and subclump densities with the separation and masses of gravitationally bound cores suggests that they can be explained by clump fragmentation, implying the simultaneous formation of subclumps and cores within rather than a step-by-step hierarchical fragmentation. The number of cores in each clump positively correlates with the clump surface density and the number expected from the thermal Jeans fragmentation. We also find that the higher the fraction of protostellar cores, the larger the dynamic range of the core mass, implying that the cores are growing in mass as the clump evolves. The ASHES sample exhibits various fragmentation patterns: aligned, scattered, clustered, and subclustered. Using the  $Q$ -parameter, which can help distinguish between centrally condensed and subclustered spatial core distributions, we finally find that in the early evolutionary stages of high-mass star formation, cores tend to follow a subclustered distribution.

*Unified Astronomy Thesaurus concepts:* Star formation (1569); Infrared dark clouds (787)

*Supporting material:* machine-readable tables

### 1. Introduction

High-mass stars ( $>8 M_{\odot}$ ) are mostly formed in clusters (Lada & Lada 2003). Investigating the fragmentation process of molecular clouds provides insights into the formation process of these stellar cluster members through molecular cloud contraction and has implications for the stellar initial mass function (Motte et al. 2022). In particular, infrared dark clouds (IRDCs) are thought to be the best place to study the initial phase of high-mass stellar clusters prior to being significantly disturbed by feedback processes (e.g., outflows, stellar winds, and radiation; Rathborne et al. 2006; Chambers et al. 2009; Sanhueza et al. 2010, 2012; Tan et al. 2013; Rosen et al. 2020). Among IRDCs, those that are 70  $\mu\text{m}$  dark are likely the most pristine, with no evidence of star formation in the IR (e.g., Sanhueza et al. 2013; Tan et al. 2013; Guzmán et al. 2015; Sanhueza et al. 2017; Contreras et al. 2018). Stars form in gas condensations where gravity dominates over any supporting

mechanism, such as turbulence, pressure, and magnetic fields. To understand what dominates core formation, separations and masses of cores have been compared with those expected from Jeans instabilities (Jeans 1902). Considering a uniform, infinite, isothermal medium at rest with a density  $\rho$  and pressure  $P_0 = \rho c_s^2$ , the mean separation of fragments expected from the gravitational collapse, called the thermal Jeans length, is defined as

$$\lambda_J^{\text{th}} = c_s (\pi / (G\rho))^{1/2}, \quad (1)$$

where  $c_s = (k_B T / \mu m_H)^{1/2}$  is the isothermal sound speed,  $G$  is the gravitational constant,  $k_B$  is the Boltzmann constant, and  $m_H$  is the mass of the hydrogen atom.  $\text{H}_2$  and He govern the thermal velocity dispersion, and the mean molecular weight per free particle  $\mu$  can be set to be 2.37 (Kauffmann et al. 2008), which is calculated from the cosmic abundance ratios. The mass of a sphere associated with the Jeans length is called the Jeans mass and is expressed as

$$M_J^{\text{th}} = \frac{4\pi\rho}{3} \left( \frac{\lambda_J^{\text{th}}}{2} \right)^3 = \frac{\pi^{5/2}}{6} \frac{c_s^3}{\sqrt{G^3\rho}}. \quad (2)$$



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If we include the impact of nonthermal motions as well as the thermal velocity dispersion, the Jeans length and mass are called turbulent Jeans length ( $\lambda_J^{\text{turb}}$ ) and mass ( $M_J^{\text{turb}}$ ), respectively.

The recently observed core separations in IRDCs or high-mass star-forming regions are mostly comparable to the thermal Jeans length of clumps rather than to the turbulent Jeans length (e.g., Beuther et al. 2015, 2018; Palau et al. 2018; Liu et al. 2019; Sanhueza et al. 2019; Lu et al. 2020; Beuther et al. 2021; Ishihara et al. 2024). For core masses, observations of IRDCs using the Atacama Large Millimeter/submillimeter Array (ALMA) revealed that a large portion of cores have low to intermediate mass ( $\lesssim 30 M_\odot$ ), which also preferentially exhibit thermal Jeans fragmentation (e.g., Sanhueza et al. 2019), although less sensitive observations at lower resolution with the ALMA 7 m array or the Submillimeter Array (SMA) have found most massive cores comparable to the turbulent Jeans mass (greater than a few solar mass; Zhang et al. 2009, 2015). Only a few ALMA studies of more evolved high-mass star-forming regions affected by feedback from massive stars show less effective fragmentation, favoring turbulent fragmentation (Rebolledo et al. 2020; Jiao et al. 2023). Some recent studies, on the other hand, propose hierarchical fragmentation (i.e., from clump to subclump and subclump to cores; Palau et al. 2018; Pokhrel et al. 2018; Svoboda et al. 2019; Rosen et al. 2020; Zhang et al. 2021) based on two different spatial resolution observations or double peak of core separation distribution. It should be noted that most of these studies are case studies (limited sample) or single-pointing observations (limited spatial area). The general trend of core separations and core masses remains unclear, and a statistical study offers the opportunity to conclusively answer the following questions: Are average core properties explained by thermal Jeans fragmentation of clumps? Is clump fragmentation hierarchical?

We use 839 cores identified from continuum emission of thirty-nine  $70 \mu\text{m}$  dark IRDC clumps using the dendrogram algorithm in the Atacama Large Millimeter/submillimeter Array Survey of  $70 \mu\text{m}$  Dark High-mass Clumps in Early Stages (ASHES) survey (Sanhueza et al. 2019; Morii et al. 2023). This is the largest sample of cores embedded in IRDCs and is thought to be the best sample for a statistical study of cores in the very early phase of evolution to date. Morii et al. (2023) revealed that the majority of the ASHES clumps only host low- to intermediate-mass cores, implying the need for core growth, and that core mass segregation does not clearly appear at such very early stages, although cores are likely segregated in terms of their density. This paper is outlined as follows: Section 2 describes the observation setups and data reduction process. Section 3 summarizes the analysis to measure core separations and estimate core masses. We compare observed core properties (separations and masses) with thermal and turbulent Jeans parameters and with Jeans parameters estimated from clump and subclump densities, and discuss the hierarchical fragmentation and diversity found within our sample in Section 4. Our conclusions are given in Section 5.

## 2. Observations and Data Reduction

We have used observations from the ASHES survey, which was carried out with ALMA in Band 6 ( $\sim 224 \text{ GHz}$ ;  $\sim 1.34 \text{ mm}$ ) through three cycles: Cycle 3 (2015.1.01539.S, PI: P.

Sanhueza), Cycle 5 (2017.1.00716.S, PI: P. Sanhueza), and Cycle 6 (2018.1.00192.S, PI: P. Sanhueza). The data was taken with the main 12 m array and the Atacama Compact Array, including both the 7 m array and total power (TP). Targets are the thirty-nine  $70 \mu\text{m}$  dark IRDC clumps with the potential for high-mass star formation (see Table 1 in Morii et al. 2023). All of the IRDC clumps were covered by Nyquist-sampled 10-pointing and three-pointing mosaics with the 12 m array and the 7 m array, respectively. The mosaicked area corresponds to  $0.97 \text{ arcmin}^2$  within 20% of the primary beam response, equivalent to the effective field of view (FOV) of  $\sim 1'$  per target. The detailed observation setups, such as on-source time and maximum recoverable scale for all sources, are summarized in Table 2 of Morii et al. (2023).

Our spectral setup includes several molecular lines that have been used in the series of ASHES works: outflow tracers (e.g., CO  $J = 2-1$  and SiO  $J = 5-4$ ; Li et al. 2020; Morii et al. 2021; Tafuya et al. 2021), dense gas tracers (e.g.,  $\text{N}_2\text{D}^+$   $J = 3-2$ , DCN  $J = 3-2$ , and  $\text{DCO}^+$   $J = 3-2$ ; Sakai et al. 2022), and shock or warm gas tracers (e.g.,  $\text{H}_2\text{CO}$   $J_{K_a, K_c} = 3_{2,2}-2_{2,1}$ ,  $\text{H}_2\text{CO}$   $J_{K_a, K_c} = 3_{2,1}-2_{2,0}$ ,  $\text{CH}_3\text{OH}$   $J_K = 4_2-3_1$ , and  $\text{HC}_3\text{N}$   $J = 24-23$ ; Izumi et al. 2024). The velocity resolution of CO,  $\text{CH}_3\text{OH}$ ,  $\text{H}_2\text{CO}$ , and  $\text{HC}_3\text{N}$  is  $\sim 1.3 \text{ km s}^{-1}$ , and that of other molecules is  $\sim 0.17 \text{ km s}^{-1}$ . A detailed description of lines is summarized in Morii et al. (2021).

Data reduction was carried out using the CASA software package versions 4.5.3, 4.6, 4.7, and 5.4.0 for calibration and 5.4.0 and 5.6.0 for imaging (CASA Team et al. 2022). Continuum images were produced by averaging line-free channels. The effective bandwidth for continuum emission was  $\sim 3.7 \text{ GHz}$ . After subtracting continuum emission, we combined the 12 m array data with the 7 m array data using the CASA task `concat`, and then they were cleaned together. Additionally, we also produced continuum images only from the 7 m array data following the same procedure. In this work, we only used TP data of the  $\text{C}^{18}\text{O}$  ( $J = 2-1$ ) line to estimate the velocity dispersion of the clumps because TP antennas do not provide continuum emission. Some of the  $\text{C}^{18}\text{O}$  ( $J = 2-1$ ) data has already been presented in the ASHES pilot survey (Sabatini et al. 2022). We used `TCLEAN` with Briggs's robust weighting of 0.5 to the visibilities and an imaging option of `MULTI-SCALE` with scales of 0, 5, 15, and 25 times the pixel size, considering the extended components of IRDCs. The average  $1\sigma$  rms noise level of the combined image is  $\sim 0.094 \text{ mJy beam}^{-1}$ , with a beam size of  $\sim 1''.2$  (Morii et al. 2023). The rms noise level of the 7 m array image is summarized in the second column of Table 1.

For molecular lines, we used the automatic cleaning algorithm for imaging data cubes, `YCLEAN` (Contreras et al. 2018; Contreras 2018) to `CLEAN` each spectral window with custom-made masks. We adopted a Briggs's robust weighting of 2.0 (natural weighting) to improve the signal-to-noise (S/N) ratio. The average synthesized beam size is  $\sim 1''.4$ . The average rms noise levels are  $\sim 0.03 \text{ K}$  for cubes with a velocity resolution of  $1.3 \text{ km s}^{-1}$  and  $\sim 0.09 \text{ K}$  for cubes with a velocity resolution of  $0.17 \text{ km s}^{-1}$ . All images have  $512 \times 512$  pixels with a pixel size of  $0''.2$ , and all images shown in this paper are the ALMA 12 and 7 m combined before the primary beam correction, while all measured fluxes are derived from the combined data corrected for the primary beam attenuation.

**Table 1**  
Clump Information and Jeans Parameters

Clump Name	Rms Noise <sup>a</sup> (mJy beam <sup>-1</sup> )	$\lambda_{J,cl}^{th}$ (pc)	$M_{J,cl}^{th}$ ( $M_{\odot}$ )	$\sigma^b$ (km s <sup>-1</sup> )	$\lambda_{J,cl}^{tu}$ (pc)	$M_{J,cl}^{tu}$ ( $M_{\odot}$ )	$n(\text{bound core})^c$	$f(\text{proto})^d$	$Q^e$
G010.991-00.082	0.88	0.08	1.3	1.1	0.46	230	14	0.29	0.78
G014.492-00.139	1.90	0.06	1.0	1.8	0.53	660	21	0.71	0.81
G015.203-00.441	0.90	0.11	2.9	1.0	0.43	170	...	...	0.78
G016.974-00.222	0.33	0.11	1.8	1.2	0.64	370	3	0.67	0.77
G018.801-00.297	0.44	0.11	1.9	1.3	0.68	460	...	...	0.68
G018.931-00.029	0.59	0.19	5.3	1.5	1.07	920	1	0.0	0.79
G022.253+00.032	0.31	0.11	2.1	1.0	0.51	190	1	1.0	0.86
G022.692-00.452	0.46	0.19	4.7	1.2	0.97	570	4	0.75	0.77
G023.477+00.114	1.00	0.08	1.4	1.3	0.46	300	7	0.75	0.71
G024.010+00.489	0.65	0.06	1.1	1.0	0.31	130	...	...	0.78
G024.524-00.139	0.53	0.13	2.2	1.5	0.90	780	9	0.91	0.73
G025.163-00.304	0.58	0.11	1.9	1.2	0.67	400	12	0.57	0.67
G028.273-00.167	0.70	0.08	1.2	1.6	0.68	650	8	0.23	0.66
G028.541-00.237	0.50	0.13	2.5	1.4	0.88	700	8	0.78	0.72
G028.564-00.236	1.20	0.07	1.3	1.9	0.69	980	...	...	0.81
G028.927+00.394	0.40	0.15	3.1	1.0	0.67	270	...	...	0.87
G030.704+00.104	0.47	0.18	3.6	1.5	1.17	1000	...	...	0.73
G030.913+00.719	0.40	0.09	1.4	0.9	0.37	110	6	0.67	0.91
G033.331-00.531	0.26	0.21	4.0	2.0	1.84	2740	...	...	0.68
G034.133+00.076	0.41	0.19	4.1	1.2	1.04	640	4	0.4	0.81
G034.169+00.089	0.32	0.19	4.4	1.0	0.76	310	2	0.5	0.77
G034.739-00.119	0.42	0.10	1.7	1.2	0.60	360	17	0.63	0.75
G036.666-00.114	0.39	0.10	1.8	0.9	0.43	160	7	0.71	0.82
G305.794-00.096	0.61	0.08	1.7	1.3	0.45	300	21	0.5	0.78
G327.116-00.294	0.71	0.11	2.2	1.4	0.70	520	2	0.75	0.71
G331.372-00.116	0.73	0.16	2.9	1.8	1.31	1670	3	0.14	0.71
G332.969-00.029	0.45	0.15	2.6	1.1	0.86	450	3	0.2	0.67
G333.016-00.751	0.55	0.17	4.1	2.1	1.49	2570	1	0.0	0.83
G333.481-00.224	0.65	0.13	3.3	1.2	0.65	400	12	0.36	0.76
G333.524-00.269	1.4	0.08	2.4	1.5	0.49	450	19	0.77	0.84
G337.342-00.119	0.36	0.19	3.7	2.8	2.35	6920	1	0.0	0.75
G337.541-00.082	0.82	0.09	1.4	1.1	0.48	230	11	0.46	0.71
G340.179-00.242	0.42	0.18	3.4	1.9	1.56	2190	0	0.0	0.82
G340.222-00.167	0.75	0.11	2.3	1.0	0.52	220	5	0.0	0.81
G340.232-00.146	1.1	0.14	2.5	2.1	1.28	2170	5	0.4	0.7
G340.398-00.396	0.59	0.13	2.4	1.8	1.08	1330	10	0.1	0.79
G341.039-00.114	0.72	0.13	2.5	1.1	0.63	290	16	0.37	0.82
G343.489-00.416	0.60	0.10	1.3	1.0	0.50	190	...	...	0.71
G345.114-00.199	0.50	0.08	1.1	1.1	0.43	210	...	...	0.75

**Notes.**

<sup>a</sup> The rms noise level of the 7 m array continuum image used to identify subclumps.

<sup>b</sup> Velocity dispersion ( $\sigma$ ) was obtained by the fitting of the line profile of C<sup>18</sup>O ( $J=2-1$ ) averaged within the clump with a 1D Gaussian, which was observed by TP.

<sup>c</sup> Number of bound cores within  $r = 0.45$  pc (see Section 4.3).

<sup>d</sup> The fraction of protostellar core to all bound cores in each clump.

<sup>e</sup> The parameter describes how centrally concentrated the core spatial distribution is. See Section 4.4 for more details.

### 3. Results

#### 3.1. Core Sample

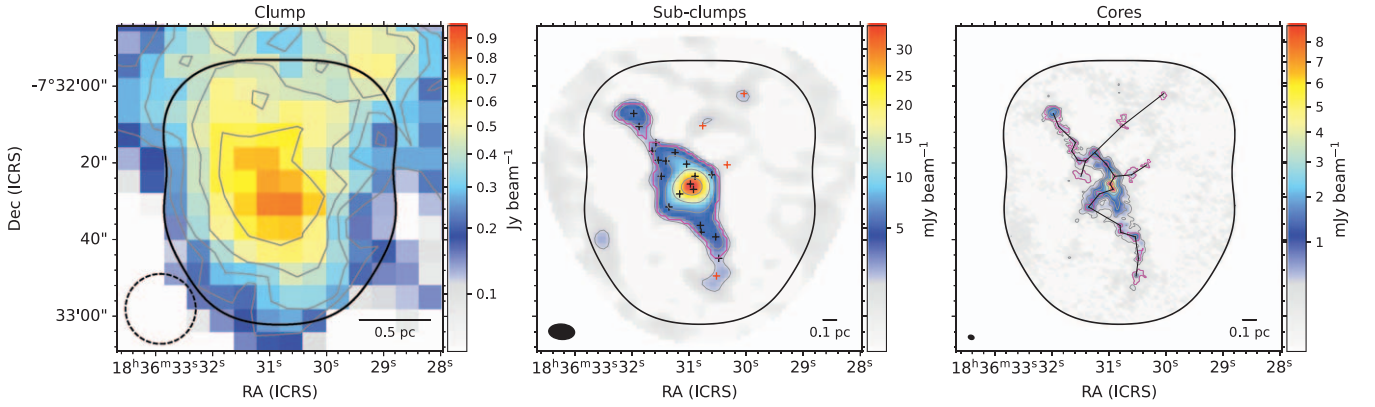
Using the dendrogram technique (Rosolowsky et al. 2008), Morii et al. (2023) identified cores using the 1.3 mm continuum images. They set a minimum value,  $F_{\min}$ , as  $2.5\sigma$ , a minimum significance to separate them,  $\delta$ , as  $1.0\sigma$ , and the minimum number of pixels to be contained in the smallest individual structure,  $S_{\min}$ , as the half-pixel numbers of the beam area. Here,  $\sigma$  is the rms noise level of the continuum image. Since these parameters are optimistic values, they applied the additional constraint to the flux density to exclude suspicious structures. They have excluded cores with a flux density smaller than  $3.5\sigma$ . Additionally, they eliminated cores at the edge of the FOV. The total number of identified `leaf`

structures is 839 from 39 clumps (see an example in Figure 1). For the analysis carried out in this work, we have adopted the core masses from Morii et al. (2023).

#### 3.2. Classification

We classified cores into three evolutionary categories and another three categories in terms of their gravitational stability. As star formation signatures, we followed Sanhueza et al. (2019) and adopted the detection of molecular outflows traced by CO  $J=2-1$  and SiO  $J=5-4$ , and the detection of warm gas tracers (H<sub>2</sub>CO  $J_{K_a,K_c}=3_{2,2}-2_{2,1}$ , H<sub>2</sub>CO  $J_{K_a,K_c}=3_{2,1}-2_{2,0}$ , CH<sub>3</sub>OH  $J_K=4_2-3_1$ , and HC<sub>3</sub>N  $J=24-23$ ), which all have an upper-state energy higher than 45 K. We judged cores associated with outflows as outflow cores, and any detection





**Figure 1.** Continuum images for G024.524-00.139 obtained by (left) a single-dish telescope (Schuller et al. 2009), (middle) ALMA 7 m array, and (right) ALMA 12 and 7 m arrays. The beam sizes of each image are shown as a black circle or ellipse at the bottom left corner. (Left) Gray-solid contours represent  $3 \times 2^{\text{n}}\sigma$  ( $n = 1, 2, 3, \dots$ ), where  $\sigma = 71 \text{ mJy beam}^{-1}$  is the rms noise level. The black contour represents the FOV of ALMA observations. (Middle) Gray-solid contours represent  $3 \times 2^{\text{n}}\sigma$  ( $n = 1, 2, 3, \dots$ ), with  $\sigma = 0.53 \text{ mJy beam}^{-1}$ . The magenta thick contour represents leaf structures identified by the dendrogram algorithm. Black and red crosses show core peak positions included in subclumps and not-included ones, respectively. The black ellipse in the bottom left corner represents the synthesized beam size. The black line indicates the spatial scale in the bottom right corner. (Right) Black segments show the outcome from the MST, which corresponds to the set of straight lines that connect cores in a way that minimizes the sum of the lengths. Gray contour levels are the same with the middle panel but with  $\sigma = 0.095 \text{ mJy beam}^{-1}$ . The magenta thick contour represents leaf structures identified by the dendrogram algorithm.

of warm line emission without outflow detections as warm cores. In turn, cores without any detections of outflow and warm gas tracers are classified as prestellar core candidates. This is consistent with the classification of the pilot ASHES survey (Li et al. 2023). In the end, we had 514 prestellar core candidates and 325 protostellar cores (222 warm cores and 103 outflow cores).

To determine the gravitational state of cores, we calculated the virial parameter ( $\alpha$ ), and the ratio of the virial mass ( $M_{\text{vir}}$ ) to core mass ( $M_{\text{core}}$ ). The virial mass was calculated as follows:

$$M_{\text{vir}} = \frac{5}{a\beta} \frac{\sigma_{\text{tot}}^2 R}{G}, \quad (3)$$

where  $\sigma_{\text{tot}}^2 = \sigma_{\text{th}}^2 + \sigma_{\text{nt}}^2$  is the velocity dispersion,  $R$  is the core radius,  $a = (1 - b/3)/(1 - 2b/5)$  is the correction factor for a power-law density profile ( $\rho \propto R^{-b}$ ), and  $\beta = (\arcsine)/e$  is the geometry factor (see Fall & Frenk 1983; Li et al. 2013, for a detailed derivation). Here, we adopt a typical density profile index  $b = 1.6$  (e.g., Beuther et al. 2002; Palau et al. 2014), and  $\beta = 1.2$  (Li et al. 2023). The thermal velocity dispersion and the nonthermal velocity dispersion are given by  $\sigma_{\text{th}}^2 = \frac{kT}{\mu m_{\text{H}}}$  and  $\sigma_{\text{nt}}^2 = \sigma_{\text{obs}}^2 - \frac{kT}{m_{\text{obs}}}$ , respectively. We assumed that the nonthermal component is independent of the molecular tracer and that  $\sigma_{\text{obs}}$  is the observed velocity dispersion.  $m_{\text{obs}}$  is the molecular weight of the molecule, here  $30m_{\text{H}}$ , corresponding to  $\text{DCO}^+$  and  $\text{N}_2\text{D}^+$ .

The identified cores at early stages of evolution are in dense ( $>10^5 \text{ cm}^{-3}$ ) and cold ( $<20 \text{ K}$ ) environments where CO depletion and a high-level deuteration occur (e.g., Caselli et al. 2002; Sabatini et al. 2020; Redaelli et al. 2021, 2022; Sabatini et al. 2022; Sakai et al. 2022; Sabatini et al. 2023). The pilot survey revealed that  $\text{N}_2\text{D}^+ J = 3-2$  and  $\text{DCO}^+ J = 3-2$  succeeded in tracing such quiescent dense gas, excellent lines to measure the gas velocity dispersion minimizing the contribution from more diffuse intraclump gas and gas related to protostellar activity like outflows (Sakai et al. 2022). We fitted the core-averaged spectrum of the  $\text{N}_2\text{D}^+$  and  $\text{DCO}^+$  taken from the 12 and 7 m arrays combined data with a single

Gaussian. We judged whether the emission was detected with an S/N larger than 3. Following the pilot survey (Li et al. 2022), to increase the S/N of the weak line emission, the core-averaged spectrum is spectrally smoothed over two native channels prior to the Gaussian fittings, if it shows a marginal  $\sim 3\sigma$  confidence in the native spectral resolution. Table 2 summarizes the fitting results, the peak intensity, velocity center, and total velocity dispersion for each line emission. Following the discussion in Li et al. (2022), both lines trace the same physical location (dense gas), and if both lines are detected in a core, we generally use  $\sigma_{\text{obs}}$  measured from  $\text{DCO}^+$ , except for a few cases. We adopt the  $\sigma_{\text{obs}}$  value if either line is detected. As a result, we obtained  $\sigma_{\text{obs}}$  information for 492 cores among 839 cores. Following Li et al. (2023), we consider cores with  $\alpha < 2$  to be gravitationally bound cores and cores with  $\alpha > 2$  to be unbound, which might be transient objects if we ignore additional support of magnetic fields or external pressure. For cores without the detection of  $\text{DCO}^+$  and  $\text{N}_2\text{D}^+$ , we classified them as nondetection cores. The derived  $\alpha$  ranges from 0.06 to  $\sim 10$  with a median value of 1.2. Among 492 cores, 340 cores are classified as bound cores. The detailed results for each core's virial analysis will be presented in a future paper.

### 3.3. Core Separation

We used the minimum spanning tree (MST) method developed by Barrow et al. (1985) to measure core separation. MST connects structures (cores in this case), minimizing the sum of the length and determining a set of straight lines. The right panel of Figure 1 shows an example of the minimum core separation (edge length determined by MST, hereafter  $\delta_{\text{sep}}$ ) as black segments. This separation is the projected separation, and the real separation could be equal or longer.<sup>11</sup> Taking the average of the projection effect, we divide the measured separation by a correction factor of  $\pi/4$  (see Ishihara et al. 2024 for the details). The mean  $\delta_{\text{sep}}$  after correction in each clump ranges from 0.08–0.32 pc, with a median of 0.15 pc.

<sup>11</sup> The real separation can be expressed as the projected separation divided by  $\cos i$ , where  $i$  is the inclination angle from the plane of the sky.

**Table 2**  
Summary of  $\text{N}_2\text{D}^+$  and  $\text{DCO}^+$  Gaussian Fitting Results

Clump Name	Core Name	$\text{N}_2\text{D}^+$			$\text{DCO}^+$		
		$T_a$ (K)	$v_{\text{LSR}}$ ( $\text{km s}^{-1}$ )	$\sigma_{\text{obs}}$ ( $\text{km s}^{-1}$ )	$T_a$ (K)	$v_{\text{LSR}}$ ( $\text{km s}^{-1}$ )	$\sigma_{\text{obs}}$ ( $\text{km s}^{-1}$ )
G010.991-00.082	ALMA1	0.28 (0.04)	29.71 (0.06)	0.34 (0.06)	0.35 (0.03)	29.74 (0.03)	0.29 (0.03)
G010.991-00.082	ALMA2	0.30 (0.04)	29.88 (0.06)	0.38 (0.06)	...	...	...
G010.991-00.082	ALMA3	0.23 (0.03)	29.98 (0.1)	0.67 (0.1)	0.37 (0.04)	29.65 (0.06)	0.48 (0.06)
G010.991-00.082	ALMA4	0.18 (0.03)	30.11 (0.09)	0.48 (0.09)	...	...	...
G010.991-00.082	ALMA5	0.41 (0.07)	29.73 (0.05)	0.27 (0.05)	0.45 (0.06)	29.85 (0.04)	0.25 (0.04)
G010.991-00.082	ALMA6	...	...	...	...	...	...
G010.991-00.082	ALMA7	0.44 (0.06)	29.46 (0.05)	0.34 (0.05)	1.25 (0.06)	29.56 (0.01)	0.21 (0.01)
G010.991-00.082	ALMA8	0.27 (0.04)	30.01 (0.05)	0.3 (0.05)	0.32 (0.05)	29.81 (0.06)	0.31 (0.06)
G010.991-00.082	ALMA9	0.48 (0.06)	29.5 (0.06)	0.43 (0.06)	0.52 (0.06)	29.52 (0.05)	0.38 (0.05)
G010.991-00.082	ALMA10	0.26 (0.05)	30.14 (0.04)	0.19 (0.04)	0.28 (0.04)	30.45 (0.05)	0.28 (0.05)

**Note.** The corresponding uncertainty is given in parentheses. Dashes denote no available data.

(This table is available in its entirety in machine-readable form.)

**Table 3**  
Subclump Properties Obtained from a Dendrogram

Clump Name	Subclump No.	R.A.	Decl.	$\text{FWHM}_{\text{maj}} \times \text{FWHM}_{\text{min}}$ ("×")	Peak Intensity ( $\text{mJy beam}^{-1}$ )	Flux Density ( $\text{mJy}$ )
G010.991-00.082	1	18 <sup>h</sup> 10 <sup>m</sup> 06 <sup>s</sup> .72	−19°27′46″.69	12.04 × 4.26	12.73	23.51
G010.991-00.082	2	18 <sup>h</sup> 10 <sup>m</sup> 08 <sup>s</sup> .20	−19°28′17″.69	6.26 × 3.37	12.67	14.17
G010.991-00.082	3	18 <sup>h</sup> 10 <sup>m</sup> 07 <sup>s</sup> .42	−19°28′02″.69	9.66 × 6.28	12.49	39.73
G014.492-00.139	1	18 <sup>h</sup> 17 <sup>m</sup> 22 <sup>s</sup> .44	−16°25′00″.89	6.79 × 5.72	56.71	110.46
G014.492-00.139	2	18 <sup>h</sup> 17 <sup>m</sup> 21 <sup>s</sup> .54	−16°25′02″.89	5.42 × 3.22	44.68	41.71
G014.492-00.139	3	18 <sup>h</sup> 17 <sup>m</sup> 22 <sup>s</sup> .23	−16°25′30″.89	5.88 × 3.47	24.25	26.03

(This table is available in its entirety in machine-readable form.)

In addition to the minimum core separation, which is the distance between a pair identified by the MST method, we also estimate the nearest separation for each core, which is the distance to the nearest fragments. The average of nearest core separation after correction in each clump ranges from 0.07–0.29 pc, with a median of 0.12 pc.

### 3.4. Subclump Identification

Inside clumps, cores are not always uniformly distributed but consist of subclusters. To study the fragmentation properties of such subclusters, we applied the dendrogram technique for the continuum images produced by data only taken by the 7 m array. We used the same parameters as those used for the core identification but with the rms noise level measured in the 7 m array continuum images (Table 1). For example, the middle panel of Figure 1 shows the 7 m array continuum image for G024.524-00.139, and magenta contours represent the identified leaf structures. We define these leaves as subclumps. From all clumps, we identified from one to seven subclumps per clump, 135 subclumps in total. Table 3 gives the peak position, beam-convolved size (major and minor full width half-maximum; FWHM), peak intensity, and flux density of subclumps identified by the dendrogram algorithm (the properties for all subclumps are summarized in a machine-readable table).

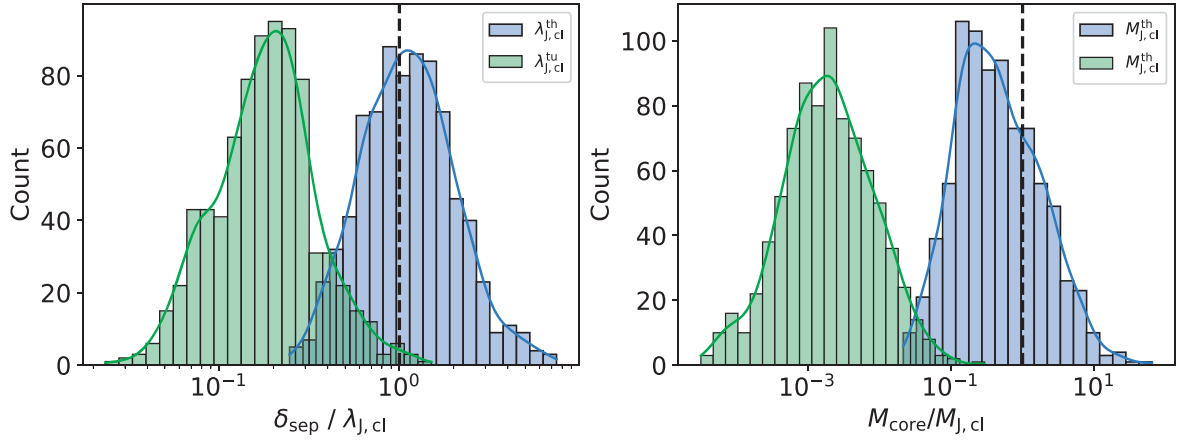
Masses of subclumps are estimated from the flux density of 1.3 mm continuum emission assuming optically thin conditions

by

$$M_{\text{core}} = \mathbb{R} \frac{d^2 F_\nu}{\kappa_\nu B_\nu(T_{\text{dust}})}, \quad (4)$$

where  $\mathbb{R} = 100$  is the gas-to-dust mass ratio,  $\kappa_\nu$  is the dust absorption coefficient,  $d$  is the heliocentric distance associated with each ASHES clump, and  $B_\nu$  is the Planck function for a dust temperature  $T_{\text{dust}}$ . We adopt  $\kappa_\nu$  of  $0.9 \text{ cm}^2 \text{ g}^{-1}$  from the dust coagulation model of the MRN (Mathis et al. 1977) distribution with thin ice mantles at a number density of  $10^6 \text{ cm}^{-3}$  computed by Ossenkopf & Henning (1994). We used the clump dust temperature listed in Mori et al. (2023). As the ASHES clumps are in their early stages, the free-free contamination, especially at 1.3 mm, is negligible. Here, we assume that the continuum emission comes only from dust emission.

The surface density,  $\Sigma$ , and the molecular volume density,  $n(\text{H}_2)$ , were estimated assuming a uniform spherical density distribution as  $\Sigma = M_{\text{subcl}}/\pi R_{\text{subcl}}^2$  and  $n(\text{H}_2) = M_{\text{subcl}}/\bar{m}_{\text{H}_2}(4\pi R_{\text{subcl}}^3/3)$ , where  $R_{\text{subcl}}$  is half of the geometric mean of the FWHM (Table 3). In this paper, we defined cores overlapped with the subclump as members of the subclump (black crosses in the middle panel of Figure 1). The number of cores in each subclump varies from zero to more than 10 (the maximum is 23). The estimated physical parameters and the number of gravitationally bound cores are summarized in Table 4.



**Figure 2.** Core separation ( $\delta_{\text{sep}}$ ) and core masses normalized with Jeans length and Jeans mass of clumps, respectively. Blue and green bins represent the case of thermal or turbulent Jeans fragmentation, respectively. The thick solid lines represent kernel density distribution. The vertical lines correspond to the ratio of unity.

**Table 4**  
Subclump Physical Parameters

Clump Name	Subclump Name	$M_{\text{subcl}}$ ( $M_{\odot}$ )	$R_{\text{subcl}}$ (pc)	$\Sigma$ ( $\text{g cm}^{-2}$ )	$n(\text{H}_2)$ ( $\times 10^5 \text{ cm}^{-3}$ )	$\lambda_{J, \text{subcl}}^{\text{th}}$ (pc)	$M_{J, \text{subcl}}^{\text{th}}$ ( $M_{\odot}$ )	$N(\text{bound core})$
G010.991-00.082	1	15.0	0.064	0.24	4.45	0.048	0.76	3
G010.991-00.082	2	9.0	0.041	0.35	2.56	0.032	0.5	0
G010.991-00.082	3	25.3	0.07	0.34	1.95	0.041	0.66	8
G014.492-00.139	1	69.4	0.059	1.33	4.45	0.02	0.35	5
G014.492-00.139	2	26.2	0.039	1.12	2.56	0.018	0.31	3
G014.492-00.139	3	16.4	0.043	0.60	1.95	0.026	0.45	0

**Note.** The radius is calculated from the geometric mean of the FWHM divided by 2.

(This table is available in its entirety in machine-readable form.)

### 3.5. Jeans Length and Jeans Mass

We estimated the thermal Jeans length and mass of clumps and subclumps from Equations (1) and (2) with their density and dust temperature. The clump density is estimated from the flux density of the continuum emission obtained by the Atacama Pathfinder Experiment Telescope Large Area Survey of the galaxy (Schuller et al. 2009), and summarized in Table 1 of Morii et al. (2023). One example of the continuum image is shown in the left panel of Figure 1. The derived  $\lambda_{J, \text{cl}}^{\text{th}}$  ranges from 0.06–0.22 pc, with a median of 0.12 pc, and  $M_{J, \text{cl}}^{\text{th}}$  ranges from 1.1–5.5  $M_{\odot}$ , with a median of 2.4  $M_{\odot}$ . Using the subclump density in Table 4, the subclumps’ thermal Jeans length ( $\lambda_{J, \text{subcl}}^{\text{th}}$ ) and mass ( $M_{J, \text{subcl}}^{\text{th}}$ ) can be derived. The density of subclumps is about 1 order of magnitude larger than clumps, and the estimated  $\lambda_{J, \text{subcl}}^{\text{th}}$  is a few times smaller than  $\lambda_{J, \text{cl}}^{\text{th}}$ , ranging from 0.018–0.12 pc, with a median of 0.05 pc.  $M_{J, \text{subcl}}^{\text{th}}$  ranges from 0.3–3.5  $M_{\odot}$ , with a median of 1.1  $M_{\odot}$ .

We also estimate the turbulent Jeans parameters of clumps. We applied a 1D Gaussian fitting toward the line profile of  $\text{C}^{18}\text{O } J=2-1$  averaged within the clump, which was observed by TP. The obtained velocity dispersion ( $\sigma_{\text{obs}}$ ) is summarized in Table 1. The nonthermal velocity dispersion is given by  $\sigma_{\text{nt}}^2 = \sigma_{\text{obs}}^2 - \frac{kT}{m_{\text{obs}}}$  with  $m_{\text{obs}} = 28m_{\text{H}}$ . By replacing  $c_s$  with  $\sqrt{c_s^2 + \sigma_{\text{nt}}^2}$  in Equation (1), the turbulent Jeans length ( $\lambda_{J, \text{cl}}^{\text{tu}}$ ) can be estimated. It is about five times larger than the thermal Jeans length ranging from 0.31–2.35 pc, and the turbulent Jeans masses ( $M_{J, \text{cl}}^{\text{tu}}$ ) are estimated to be in the range of  $10^2$  to

$7 \times 10^3 M_{\odot}$ . All estimated values for each clump are summarized in Table 1.

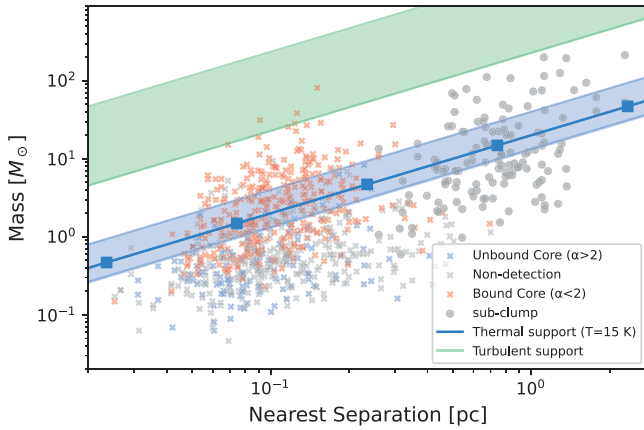
Figure 2 shows the histogram of the ratios of the observed separation and mass divided by Jeans length and mass of clumps, respectively, in logarithmic scale. The top two panels show the case for the thermal Jeans fragmentation ( $\delta_{\text{sep}}/\lambda_{J, \text{cl}}^{\text{th}}$  and  $M_{\text{core}}/M_{J, \text{cl}}^{\text{th}}$ ), while the bottom two panels are the cases of turbulent Jeans fragmentation ( $\lambda_{J, \text{cl}}^{\text{tu}}$  and  $M_{J, \text{cl}}^{\text{tu}}$ ). The  $\delta_{\text{sep}}/\lambda_{J, \text{cl}}^{\text{th}}$  distributions peak at around 1–2. For mass, the distribution has a large variation across more than 2 orders of magnitude, and more than half of the cores have lower mass than the thermal Jeans mass of clumps. The peak is around 0.2, but there is a secondary peak around unity. For the turbulent Jeans fragmentation, both ratios of  $\delta_{\text{sep}}/\lambda_{J, \text{cl}}^{\text{tu}}$  and  $M_{\text{core}}/M_{J, \text{cl}}^{\text{tu}}$  show their peaks at values much smaller than unity.

## 4. Discussion

In this section, we aim to characterize the fragmentation at early stages of high-mass star formation found in the ASHES sample by comparing core separation and mass with thermal and turbulent Jeans parameters, and by comparing clump and subclump Jeans parameters to study hierarchical fragmentation.

### 4.1. Thermal versus Turbulent Jeans Fragmentation

What dominates the fragmentation is one key question in star formation that has been investigated in some previous studies (e.g., Zhang et al. 2009; Palau et al. 2013, 2015;



**Figure 3.** The relation between mass and nearest separation distance. The crosses represent cores identified in the ASHES survey (Mori et al. 2023), and the gray circles correspond to subclumps. The three different colors of the crosses indicate the gravitational states of cores; bound cores (red), unbound cores (blue), and cores without detections of dense gas tracers such as  $\text{N}_2\text{D}^+$  and  $\text{DCO}^+$  (gray). The blue line shows thermal Jeans fragmentation with  $T = 15$  K and  $n(\text{H}_2) = [10^2, 10^6] \text{ cm}^{-3}$ , and the blue-shaded region corresponds to the same density range but with  $T = [10, 30]$  K. The five squares correspond to  $n(\text{H}_2)$  of  $10^2$ ,  $10^3$ ,  $10^4$ ,  $10^5$ , and  $10^6 \text{ cm}^{-3}$  from right to left. The green-shaded region shows turbulent Jeans fragmentation to the same density and temperature range but  $\sigma = [0.8, 2.7] \text{ km s}^{-1}$ .

Sanhueza et al. 2017; Beuther et al. 2018; Sanhueza et al. 2019; Beuther et al. 2024). We have the largest sample in IRDCs thanks to the mosaicked high-spatial resolution and high-sensitivity ALMA observations. IRDCs are thought to be the best for answering this question in the very early phase of high-mass star and cluster formation.

To visualize this clearly, Figure 3 displays the mass–nearest separation relation, as in Wang et al. (2014); following this work, for this figure, we use the nearest separation between cores rather than the minimum separation defined in MST. In this figure, blue- and green-shaded regions show what is expected from thermal Jeans fragmentation and turbulent Jeans fragmentation, respectively. Our identified cores and subclumps are denoted as crosses and circles, respectively. The blue line shows thermal Jeans fragmentation with  $T = 15$  K (mean temperature of the ASHES sample) varying the density from  $10^2$  to  $10^6 \text{ cm}^{-3}$ . The blue-shaded region shows the same density range but with  $T$  from 10–30 K. The green-shaded region shows turbulent Jeans fragmentation to the same density and temperature range but varying the velocity dispersion from 0.8–2.7  $\text{km s}^{-1}$ . These ranges cover the ASHES sample properties (see Table 1 of Mori et al. 2023). The majority of cores (99.6%) and all subclumps are plotted below the green area, and on average, both cores and subclumps are in the thermal Jeans fragmentation regime, while the variation of core masses is large, about 2 orders of magnitude. Especially gravitationally bound cores (red crosses) are located around blue-shaded regions, preferring thermal Jeans fragmentation.

This trend can also be seen in Figure 2. The  $\delta_{\text{sep}}/\lambda_{\text{J,cl}}^{\text{th}}$  distribution takes a peak around unity, and both the mean and median values of  $M_{\text{core}}/M_{\text{J,cl}}^{\text{th}}$  are around unity (1.28 and 0.393, respectively). On the contrary, the Jeans length and mass in the turbulence-dominated case are both much larger than  $\lambda_{\text{J,cl}}^{\text{th}}$  and  $M_{\text{J,cl}}^{\text{th}}$ , and the ratios become much smaller than unity (see the bottom panels). Thus, based on the estimated masses and separations, we conclude that core formation found in the ASHES sample (and likely in IRDCs in general) is regulated

by thermal Jeans fragmentation rather than turbulent fragmentation.

This conclusion, based on a large ALMA sample uniformly analyzed, consolidate earlier findings in IRDCs and active high-mass star-forming regions (e.g., Palau et al. 2013; Beuther et al. 2015; Palau et al. 2015; Liu et al. 2017; Beuther et al. 2018; Liu et al. 2019; Sanhueza et al. 2019; Svoboda et al. 2019; Lu et al. 2020; Beuther et al. 2021; Saha et al. 2022). However, this result disagrees with some observations that suggest the importance of turbulent or magnetic field support (e.g., Zhang et al. 2009; Pillai et al. 2011; Wang et al. 2011; Zhang & Wang 2011; Wang et al. 2014; Zhang et al. 2015; Li et al. 2019; Rebollo et al. 2020). One significant difference is the image fidelity reached by our observations (i.e., higher spatial resolution, mass sensitivity, and better UV coverage). For example, we revealed further fragments and succeeded in identifying about three times more cores than previous studies using SMA or the Plateau de Bure Interferometer (Beuther et al. 2013, G023.477+00.114; Lu et al. 2015, G028.564-00.236; Sanhueza et al. 2017, G028.273-00.167; Li et al. 2019, G014.492-00.139; Pillai et al. 2019, G010.991-00.082). Additionally, our mosaicked observations covered a larger FOV than in single-pointing observations. That is why we succeeded in resolving cores and identifying more cores with lower masses than in previous studies.

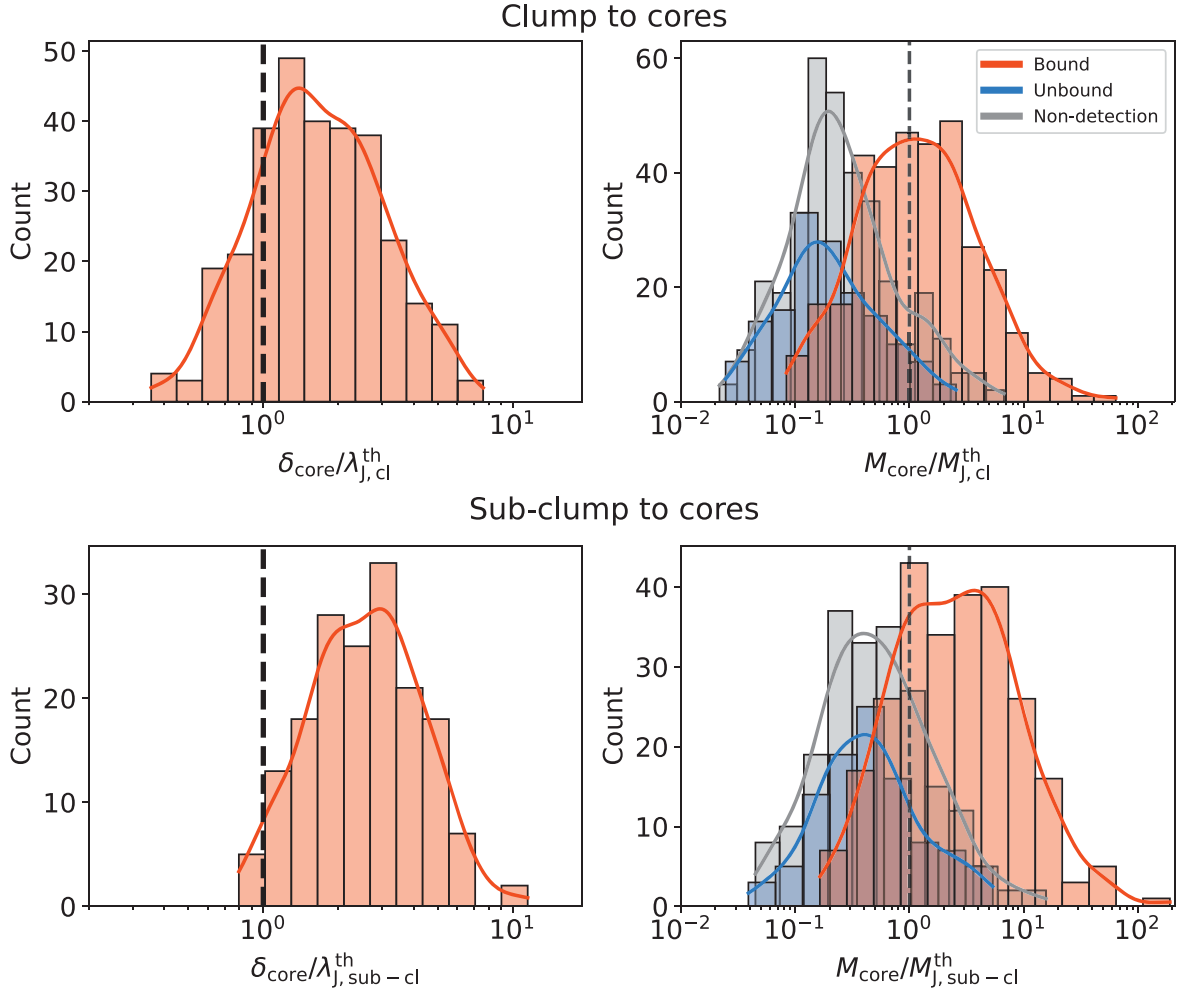
Other ALMA observations with similar setups like the ASHES survey still suggest turbulence-dominated fragmentation (e.g., Rebollo et al. 2020; Xu et al. 2023), but most are not in a quiet/early stage. They focus on more evolved, active high-mass star-forming regions where some feedback from the newborn stars can suppress the fragmentation by inducing additional turbulence and warming up the surrounding gas (e.g., Krumholz & McKee 2008).

#### 4.2. Hierarchical Fragmentation

We found that the overall properties of identified cores are consistent with thermal Jeans fragmentation, as shown in Figure 3. However, as Figure 1 shows, cores are embedded in subclump structures inside clumps. Next, we address whether cores form from the fragmentation of clumps or dense substructures (here, we refer to them as subclumps) by using clumps’ and subclumps’ Jeans parameters. In this section, we use only gravitationally bound cores that will likely form stars. Core separations are recalculated only using bound core positions.

Figure 4 compares the fragmentation of clumps and subclumps with the observed core properties. The top two panels show the ratios of separation and masses normalized by thermal Jeans parameters estimated by using clump density, and the bottom two panels show the case of using subclump density. For the bottom panels, only cores inside subclumps are considered. The left panels present the ratios of separation of bound cores normalized by thermal Jeans length. Their distribution peaks are located around 1–2 and 2–3 for the top and bottom, respectively. The right panels display the histogram of the mass ratios, and bound cores are highlighted in red. The peak is just around unity for clump fragmentation. Although the bottom panel shows no clear single peak, it has a broad peak around 1–5. These histograms also show that unbound cores and cores without the detection of dense gas tracers generally have masses smaller than the Jeans mass of





**Figure 4.** Core separation ( $\delta_{\text{sep}}$ ) and core masses normalized with thermal Jeans length and masses, respectively. The top three panels show ratios normalized by thermal Jeans lengths and masses of clumps, and the bottom three show the case of subclump fragmentation. Only cores inside subclumps are used for this analysis. The thick solid lines represent kernel density distribution. The vertical lines correspond to the ratio of unity. (Left) Core separations of bound cores normalized by Jeans length. (Right) Core masses normalized by Jeans mass. Gravitationally bound and unbound are colored red and blue, respectively (cores without detections of dense gas tracers are in gray).

clumps. Overall, the observed core properties favor the fragmentation from clumps rather than subclumps.

Hierarchical fragmentation is expected in one of the theoretical scenarios for high-mass star formation, called the global hierarchical collapse scenario (Vázquez-Semadeni et al. 2019). Such hierarchical fragmentation has been reported in some previous studies observing high-mass starless clump candidates (Svoboda et al. 2019; Zhang et al. 2021), IRDCs (Wang et al. 2011, 2014), and also OMC-1S (Palau et al. 2018). Our analysis also revealed subclumps, intermediate structures connecting clumps and cores, but the core properties cannot be explained well by subclump fragmentation. This implies that a step-by-step fragmentation (clump to subclump to core) is unlikely.

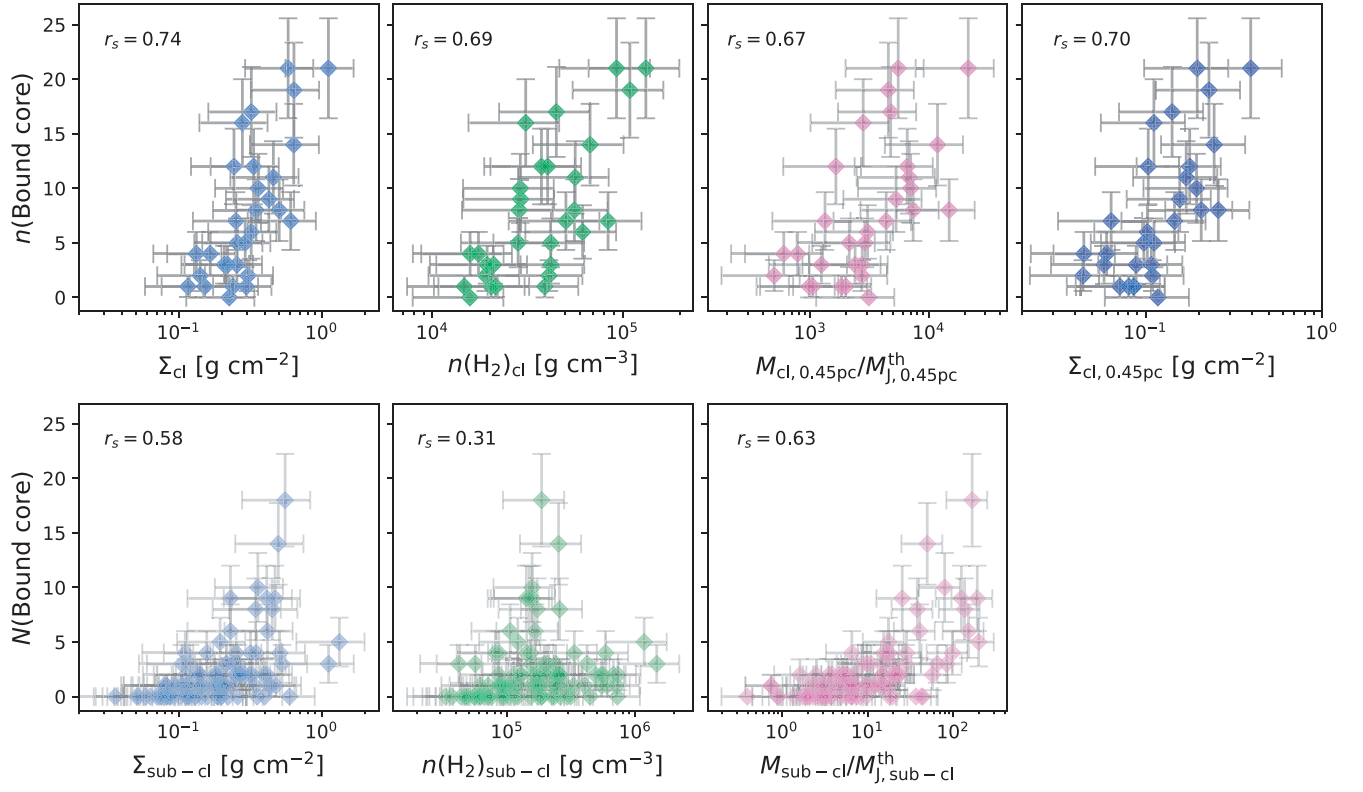
We also investigated the fragmentation from clump to subclump by comparing Jeans parameters with subclump separations and masses. We find no clear peak in the distributions due to the small number of statistics. However, thermal Jeans fragmentation may still be favored over turbulent fragmentation, as shown in Figure 3. We note that the subclumps locate (see Figure 3) in the area at which the parental clumps should have densities of  $\sim 10^3 \text{ cm}^{-3}$ , inconsistent with observations. This could be explained if

subclump structures form as a result of clump fragmentation into cores. Subclumps could be the ensemble of cores that evolve as cores evolve, not in a step-by-step hierarchical fragmentation, but rather in a simultaneous formation process.

We point out that this simultaneous formation of subclumps and cores, in which the core properties are not determined by subclumps, may only be valid at the very early stages of high-mass star formation traced in the ASHES survey. Later, in more active high-mass star-forming regions, once subsequent fragmentation occurs, core properties may be explained differently.

#### 4.3. Level of Fragmentation and Clump Properties

Jeans fragmentation invokes the idea that the level of fragmentation or the number of cores depends on clump density. We investigated the correlation of clump density or Jeans number (the ratio of clump mass to Jeans mass) with the number of bound cores,  $n(\text{bound core})$ . To calculate the  $n(\text{bound core})$ , we count the number of bound cores within the same physical area for all clumps and impose a mass threshold to reduce the effect of having different distances and sensitivities for each clump. We have limited the sample for



**Figure 5.** (Top) The number of bound cores above  $0.41 M_{\odot}$  located around the clump peak ( $r < 0.45$  pc) and  $n(\text{bound core})$ , as a function of clump surface density  $\Sigma_{\text{cl}}$ , clump mean number density  $n(\text{H}_2)_{\text{cl}}$ , Jeans number, the ratio of  $M_{\text{cl},0.45 \text{ pc}}$  to thermal Jeans mass ( $M_{\text{J},0.45 \text{ pc}}^{\text{th}}$ ), and clump surface density with a limited area of  $\Sigma_{\text{cl},0.45 \text{ pc}}$  for the selected 30 clumps. Spearman's rank correlation coefficients are denoted inside each panel. All  $p$ -values are much smaller than 0.01. (Bottom) The number of bound cores above  $0.41 M_{\odot}$  in subclumps,  $N(\text{bound core})$ , as a function of subclump surface density  $\Sigma_{\text{sub-cl}}$ , mean number density  $n(\text{H}_2)_{\text{sub-cl}}$ , and Jeans number ( $M_{\text{sub-cl}}/M_{\text{J},\text{sub-cl}}^{\text{th}}$ ).

this discussion, excluding clumps that are located too close ( $< 3.5$  kpc) and too far ( $> 5.5$  kpc), and two more with the worst mass sensitivity ( $> 0.45 M_{\odot}$ ). As a result, the 30 clumps remaining are located between 3.5 and 5.5 kpc and have a mass sensitivity between  $0.086$  and  $0.41 M_{\odot}$ . We count cores within a circle with a radius of  $0.45$  pc centered on the mean positions of cores. The circle size almost corresponds to the FOV of the closest clump. In addition, for all clumps, we impose a mass threshold of  $0.41 M_{\odot}$ , which corresponds to the worst mass sensitivity among the 30 clumps. The measured  $n(\text{bound cores})$  are listed in Table 1).

The top-left two panels of Figure 5 show moderate to strong correlations of  $n(\text{bound core})$  with surface density  $\Sigma_{\text{cl}}$  and mean clump number density  $n(\text{H}_2)_{\text{cl}}$ . This indicates that denser clumps produce a larger number of cores (higher fragmentation level). We found a stronger correlation between  $n(\text{bound core})$  and  $\Sigma_{\text{cl}}$  with a Spearman's rank correlation coefficient of  $r_s = 0.74$ , while that is  $r_s = 0.69$  for  $n_{\text{cl}}$ . Both  $p$ -values are much smaller than 0.01. The less scattered plot of  $n(\text{bound core})$ – $\Sigma_{\text{cl}}$  implies that the clump surface density is the best indicator of the level of fragmentation: the higher the clump surface density, the more likely it is to have a larger number of cores. To confirm that these correlations do not result from the codependence on the distance, we recalculated the clump mass ( $M_{\text{cl},0.45 \text{ pc}}$ ) and surface density ( $\Sigma_{\text{cl},0.45 \text{ pc}}$ ) using flux within the same physical area ( $r = 0.45$  pc). The right panel displays the strong correlation between  $\Sigma_{\text{cl},0.45 \text{ pc}}$  and  $n(\text{bound core})$ . We also compared  $n(\text{bound core})$  with the Jeans numbers ( $M_{\text{cl},0.45 \text{ pc}}/M_{\text{J},0.45 \text{ pc}}^{\text{th}}$ ), which is proportional to

$\Sigma_{\text{cl},0.45 \text{ pc}}^3/n(\text{H}_2)_{\text{cl},0.45 \text{ pc}}^{3/2}$ . It is confirmed that the measured  $n(\text{bound core})$  has a strong correlation with the number of cores expected from thermal Jeans fragmentation ( $r_s = 0.70$  and  $p$ -value  $< 0.01$ ).

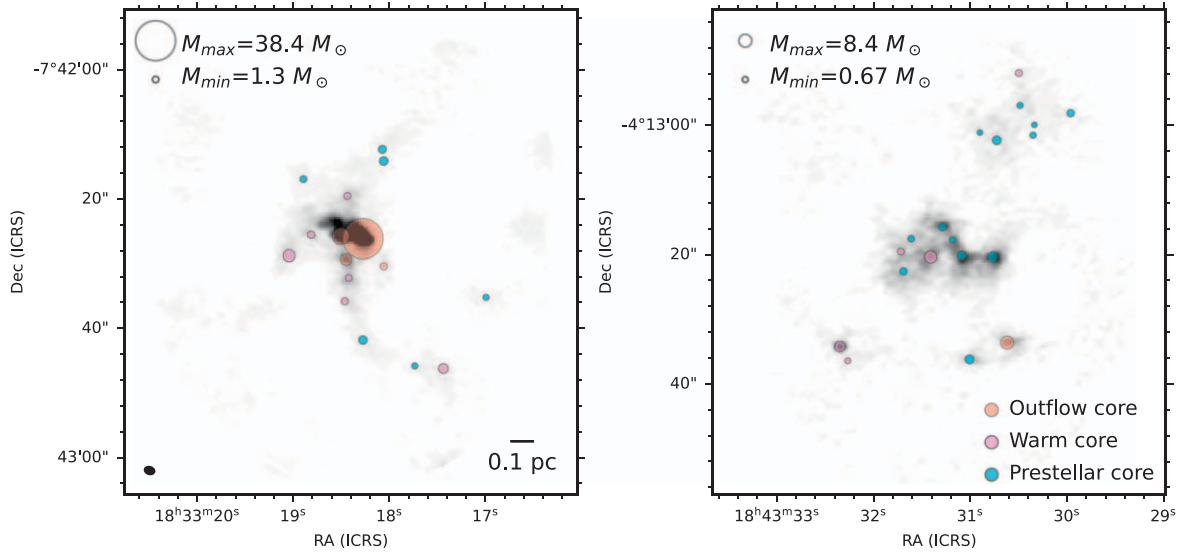
These correlations are still found in the case of subclump to core fragmentation, but in this case, the coefficients are relatively weaker. The bottom three panels also imply the strong or moderate correlation between the number of bound cores in each subclump and subclump surface density, volume density, and Jeans number, although relatively weaker than the clump case. It suggests that subclumps are also involved with core formation.

To summarize, we have revealed that a higher level of fragmentation (or higher core number density) can be expected from a region with a higher surface density, and it is consistent with Jeans fragmentation. The tight correlation between the level of fragmentation and the clump/cloud surface density has been observed in more evolved star-forming regions as well, indicating that this correlation begins early on in IRDCs and prevails during the evolution of high-mass star-forming regions (e.g., Palau et al. 2014; Sokol et al. 2019).

#### 4.4. Fragmentation Diversity

The ASHES survey also reveals a diversity in both the range of core masses per clump (mass dynamic range) and fragmentation patterns.

Morii et al. (2023) reported that most clumps host low- to intermediate-mass cores. However, the dynamic range in core masses varies from clump to clump. For example, the left panel



**Figure 6.** ALMA 1.3 mm continuum image of (left) G024.010+00.489 and (right) G028.273-00.167. The circle size represents the core mass, and the position is centered at the continuum peak of each core. The three different colors have the same meanings as in Figure 3.

in Figure 6 shows some (relatively) massive cores surrounded by some low-mass cores, and the right panel shows a cluster of low-mass cores with a small dynamic range in mass. We find that such a mass dynamic range correlates with the fraction of protostellar cores (cores with outflow or warm line detection). Figure 7 shows the fraction of bound protostellar cores to the total number of bound cores as a function of the maximum and minimum mass difference ( $M_{\max} - M_{\min}$ ), the ratio ( $M_{\max}/M_{\min}$ ), the standard deviation of core masses ( $\sigma M$ ), and interquartile range. The fraction of protostellar cores is summarized in Table 1. The  $x$ -axis in the top two panels is derived from the maximum and minimum core masses and can be significantly affected by a single peculiar object if it exists. The interquartile range, the difference between the 25th percentile and 75th percentile, is on the contrary, less affected by the most massive object. The Spearman's rank correlation coefficients are  $r_s \sim 0.5$ , and  $p$ -values are less than 0.01 except for the case of  $M_{\max}/M_{\min}$ . One outlier in these plots is G340.232-00.146, shown as an unfilled marker. As discussed in Sanhueza et al. (2019), in this peculiar clump, the most massive core is rather large with a radius of  $\sim 10^4$  au, and more fragmented structures are expected in higher angular resolution observations from visual inspection of the continuum image. The correlation coefficients displayed in Figure 7 become higher if we exclude them, as shown in the values in parenthesis (bottom, right of each panel). These plots imply that if we take the fraction of protostellar cores as an indicator of cluster evolution, the dynamic range of mass increases with cluster evolution. This correlation can be interpreted as clumped accretion onto cores. Clumps would initially fragment into cores of similar mass (small dynamic mass range), and as time goes on, some cores grow in mass more than others, resulting in a larger difference in mass range (large dynamic mass range).

Among 39 ASHES targets, some clumps show aligned fragmentation, as shown in the left panel of Figure 6, some show concentrated, and some show spread core distributions with several subclumps (e.g., the right panel of Figure 6). We calculated the  $Q$ -parameter to investigate the cluster members' distribution. The  $Q$ -parameter was defined by Cartwright &

Whitworth (2004) as

$$Q = \frac{\bar{m}}{\bar{s}}. \quad (5)$$

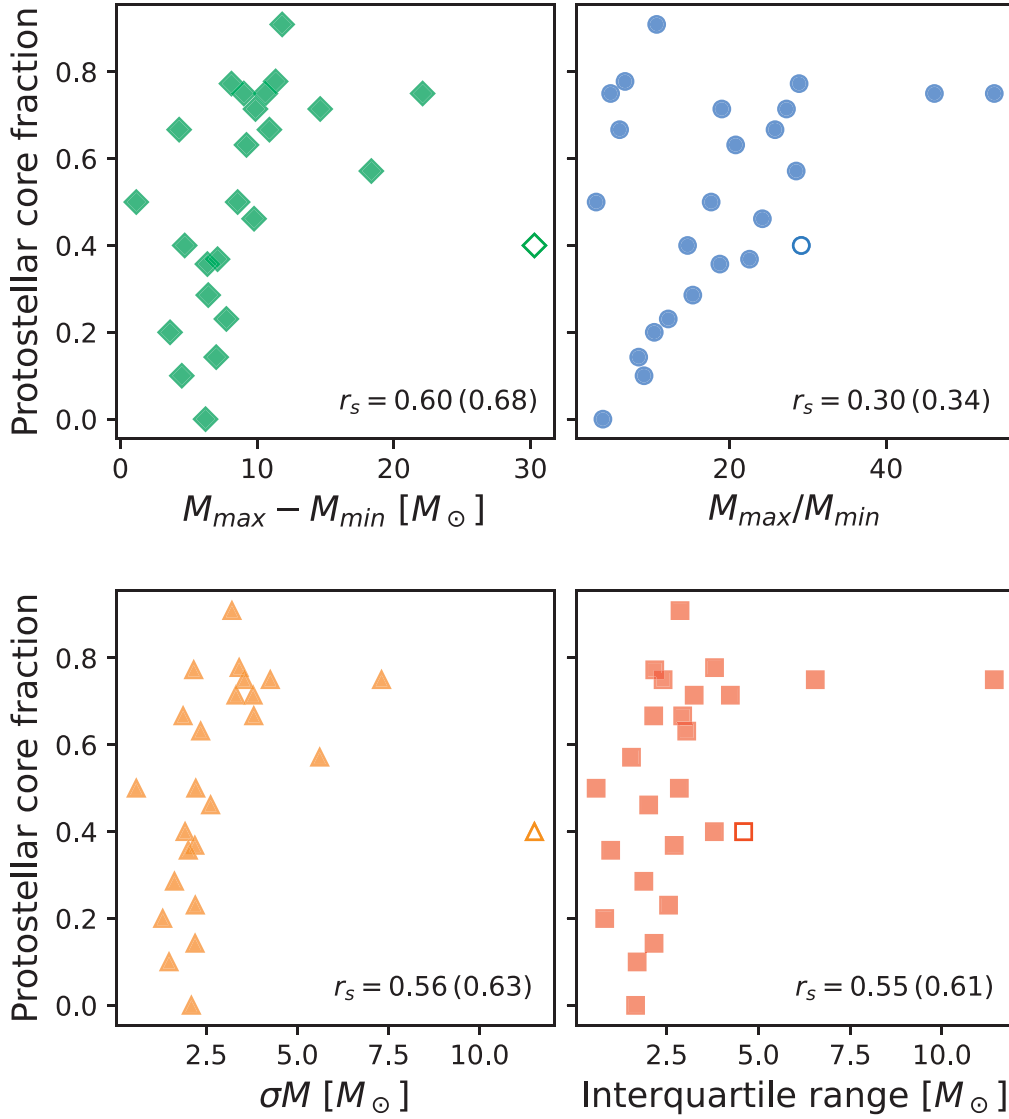
The term  $\bar{m}$  is the normalized mean edge length of MST and is defined as

$$\bar{m} = \frac{\sum_{i=1}^{N_c-1} L_i}{N_c - 1} \times \frac{N_c}{(N_c A)^{1/2}} = \frac{\sum_{i=1}^{N_c} L_i}{(N_c A)^{1/2}}, \quad (6)$$

where  $N_c$  is the number of cores in the region,  $\sum_{i=1}^{N_c} L_i$  is the total length of all the lines MST connected, hereafter *edges*, and  $A$  is the area of the cluster and estimated by  $A = \pi R_{cl}^2$ . Here, the radius of the cluster  $R_{cl}$  is defined as the distance from the mean position of cores to the farthest core position. The second term in the first equation of Equation (6) is the factor used to normalize the mean edge length of cores (the first term is defined as  $l_{MST}$ ) having different areas ( $A$ ) and/or different numbers of cores ( $N_c$ ). The term  $\bar{s}$  is the ratio of the mean core separation to the cluster radius ( $R_{cl}$ ). Here, the core separation is different from the minimum core separation ( $\delta_{sep}$ ), and it is the core separation within the region, not only considering the minimum separation. Now both  $\bar{m}$  and  $\bar{s}$  are independent of the number of cores in the cluster-forming clump.

For clusters with a smooth radial density gradient ( $n \propto r^{-\alpha}$ ),  $Q$  increases from  $\sim 0.8$  to 1.5 as the degree of concentration increases from  $\alpha = 0$  to 2.9, and for subclustering clusters  $Q$  becomes smaller than 0.8 (Cartwright & Whitworth 2004). The parameters estimated for the ASHES sample range from 0.6–0.9, and their average is 0.76. Some clumps indicate a uniform distribution of cores ( $Q \sim 0.8$ ), but most (70%) prefer subclustering ( $Q < 0.8$ ). Clumps with aligned fragmentation have a  $Q$ -parameter of 0.7–0.8. However, we note that, as can be inferred from the definition of the  $Q$ -parameter, this parameter is unable to identify aligned fragmentation. Clumps with several subclumps or showing spread fragmentation have  $Q < 0.8$ .

The origin of such variation in the fragmentation pattern is not yet clear from the current data (e.g., clump mass, density, virial parameter, protostellar, or core fraction), and further



**Figure 7.** Correlation plots between the fraction of the protostellar core and the mass dynamic range (the difference between the most massive core and the least massive core and the ratio of the most massive core to the least massive core, the standard deviation, and the interquartile range) for each clump. An additional five clumps with less than one bound core are excluded from the sample in Figure 5. The unfilled marker corresponds to G340.232-00.146 (see the main text). Spearman's rank correlation coefficients are given within each panel, and the values in parentheses are the case that the unfilled point is excluded.  $p$ -values are all less than 0.01 except in the top-right panel (0.1).

information on the magnetic field or clump-scale properties such as large-scale gas dynamics seems to be necessary. For example, Tang et al. (2019) suggest that the balance among the magnetic field, turbulence, and gravity determines the core fragmentation pattern, such as no fragmentation, aligned fragmentation, and clustered fragmentation. It should be noted that we find no clear correlation between the dynamic range in mass and the fragmentation pattern; both aligned fragmentation and spread fragmentation show large and small mass dispersion.

#### 4.5. Early Fragmentation Picture

We have revealed that the observed mean core separation and masses are comparable to the thermal Jeans lengths and masses, respectively, and much smaller than turbulent Jeans parameters. This implies that turbulence is not a dominant source characterizing core formation. This is consistent with the study by Traficante et al. (2020), indicating that gravity

dominates over turbulence once the regions become dense (e.g.,  $\Sigma > 0.1 \text{ g cm}^{-2}$ ). ASHES targets are all dense with a surface density larger than  $0.1 \text{ g cm}^{-2}$ . Compared with other ALMA studies in high-mass star-forming regions, in which core properties are better explained by turbulent Jeans fragmentation (e.g., Rebolledo et al. 2020; Xu et al. 2023), the ASHES sample contains  $70 \mu\text{m}$  dark, cold regions not affected by feedback mechanism from massive stars. Thus, our finding implies that the initial fragmentation in massive clumps, prior to the changes due to gravitational accretion and some feedback effects, is described by thermal Jeans fragmentation.

However, our sample still contains super-Jeans cores with a mass more than 10 times  $M_{J,\text{cl}}^{\text{th}}$ . Those cores may have grown in mass by acquiring gas from the surrounding environment. Infall rates in the range of  $10^{-4}$  to  $10^{-3}$  have been measured in two ASHES targets (Contreras et al. 2018; Redaelli et al. 2022), allowing the cores to quickly grow in mass in a freefall time. Alternatively, the magnetic field may play a role in suppressing fragmentation, as suggested by theoretical



studies (Hennebelle & Teyssier 2008; Commerçon et al. 2011). The theoretical prediction that the magnetic field can inhibit fragmentation in high-mass star-forming regions has garnered support from observational studies in more evolved high-mass star-forming regions. Dust polarization emission from IR-bright sources has highlighted the significant role magnetic fields play in exerting pressure on the medium, from clump to core scale, effectively curbing fragmentation during gravitational collapse (Zhang et al. 2014; Hull & Zhang 2019). Supporting evidence for fragmentation suppression due to magnetic fields has also been presented by Frau et al. (2014). In addition, Das et al. (2021) showcased the impact of the magnetic field on reducing fragment numbers, while Palau et al. (2021) reported a tentative correlation between fragment quantity and the mass-to-flux ratio among massive dense cores, as suggested by theoretical and numerical works. Future observations of dust polarization toward such subclumps or massive cores would verify this effect on suppressing fragmentation. Studying how significantly the magnetic field contributes to the fragmentation is also important to the understanding of the diversity of fragmentation patterns seen in Figure 6. Although at relatively smaller scales, these points are one of the goals of the Magnetic Fields in Massive Star-forming Regions (MagMaR; Cortés et al. 2021; Fernández-López et al. 2021; Sanhueza et al. 2021) survey once the whole survey sample is analyzed.

We found substructures inside clumps using the 7 m array data, which is consistent with the measured  $Q$  of 0.6–0.8, implying that the initial core distribution is not yet so concentrated but rather subclustered. These substructures are located in the thermal Jeans fragmentation-dominated regime, as shown in Figure 3, as well as cores. We therefore investigated whether cores are directly formed from such subclumps rather than from clumps. We find no evidence that subclump fragmentation is the preferred mechanism to explain the observed core properties. The comparison of the number of cores or the degree of fragmentation with the properties of clumps or subclumps also suggests a stronger link between clumps and cores. A possible picture is that subclumps and cores are simultaneously formed from clumps, and core properties are determined from clump properties.

It should be noted that such subclumps are likely to contribute to density segregation since denser cores are generally embedded in subclumps. If they can grow by more effective gas feeding, this may later lead to mass segregation, as discussed in Morii et al. (2023). Gas feeding or gravitational collapse of clumps and subclumps would increase the mass dynamic range, as can be seen in Figure 7. Xu et al. (2024) suggest that the gravitational concentration or gas accretion toward the center of mass would result in the appearance of mass segregation and in the increase of the  $Q$ -parameter. Comparing our results with more evolved clusters would confirm this hypothesis.

## 5. Conclusions

We have studied the fragmentation properties in 39 clumps as a part of the ASHES survey, which aims to characterize the very early phase of high-mass star formation. Using the 839 cores identified in the continuum images, we compared their masses and separations with Jeans parameters. We have obtained the following conclusions:

1. The mean core separation measured by the MST method ranges from 0.08–0.32 pc in each region, and core masses range from 0.05–81  $M_{\odot}$ . The core mass and core separation are explained by thermal Jeans fragmentation ruling out turbulent Jeans fragmentation at the very early stages of high-mass star formation.
2. Comparing the Jeans parameters of clumps and subclumps with the observed core properties, core properties, especially for bound cores, are likely determined from clumps. We interpret this as a simultaneous formation of subclumps and cores within clumps.
3. The level of fragmentation or the number of cores within each clump shows a strong correlation with the Jeans number, the ratio of clump mass to Jeans mass, implying that early core formation can be described with thermal Jeans fragmentation. It also has a strong correlation with clump surface density.
4. Furthermore, our sample shows the diversity of fragmentation in terms of mass dynamic range and spatial distribution. The correlation between the protostellar core fraction and the mass dynamic range is likely a sign of the clump-fed accretion scenarios. We have revealed aligned, spread, clustered, and subclustered fragmentation patterns, and the measured  $Q$ -parameter also implies that the early fragmentation seen in ASHES fields is not centrally concentrated.





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*Facility:* ALMA.

*Software:* CASA versions 4.5.3, 4.6, v4.7, 5.4, and 5.6 (CASA Team et al. 2022).

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