



## Discovery of two warm mini-Neptunes with contrasting densities orbiting the young K3V star TOI-815

Downloaded from: <https://research.chalmers.se>, 2024-07-05 18:55 UTC

Citation for the original published paper (version of record):

Psaridi, A., Osborn, H., Bouchy, F. et al (2024). Discovery of two warm mini-Neptunes with contrasting densities orbiting the young K3V star TOI-815. *Astronomy and Astrophysics*, 685. <http://dx.doi.org/10.1051/0004-6361/202348678>

N.B. When citing this work, cite the original published paper.

# Discovery of two warm mini-Neptunes with contrasting densities orbiting the young K3V star TOI-815<sup>★,★★</sup>

Angelica Psaridi<sup>1</sup>, Hugh Osborn<sup>2,3</sup>, François Bouchy<sup>1</sup>, Monika Lendl<sup>1</sup>, Léna Parc<sup>1</sup>, Nicolas Billot<sup>1</sup>, Christopher Broeg<sup>2,4</sup>, Sérgio G. Sousa<sup>5</sup>, Vardan Adibekyan<sup>5,6</sup>, Omar Attia<sup>1</sup>, Andrea Bonfanti<sup>7</sup>, Hritam Chakraborty<sup>1</sup>, Karen A. Collins<sup>8</sup>, Jeanne Davoult<sup>2</sup>, Elisa Delgado-Mena<sup>5</sup>, Nolan Grieves<sup>1</sup>, Tristan Guillot<sup>9</sup>, Alexis Heitzmann<sup>1</sup>, Ravit Helled<sup>10</sup>, Coel Hellier<sup>11</sup>, Jon M. Jenkins<sup>13</sup>, Henrik Knierim<sup>10</sup>, Andreas Krenn<sup>7</sup>, Jack J. Lissauer<sup>13</sup>, Rafael Luque<sup>12</sup>, David Rapetti<sup>13,14</sup>, Nuno C. Santos<sup>5,6</sup>, Olga Suárez<sup>9</sup>, Julia Venturini<sup>1</sup>, Francis P. Wilkin<sup>15</sup>, Thomas G. Wilson<sup>16</sup>, Joshua N. Winn<sup>17</sup>, Carl Ziegler<sup>18</sup>, Tiziano Zingales<sup>19,20</sup>, Yann Alibert<sup>2,4</sup>, Alexis Brandeker<sup>21</sup>, Jo Ann Egger<sup>4</sup>, Davide Gandolfi<sup>22</sup>, Matthew J. Hooton<sup>23</sup>, Amy Tuson<sup>24</sup>, Solène Ulmer-Moll<sup>1,4</sup>, Lyu Abe<sup>9</sup>, Romain Allart<sup>25</sup>, Roi Alonso<sup>26,27</sup>, David R. Anderson<sup>16</sup>, Guillem Anglada Escudé<sup>28,29</sup>, Tamas Bárczy<sup>30</sup>, David Barrado<sup>31</sup>, Susana C. C. Barros<sup>5,6</sup>, Wolfgang Baumjohann<sup>7</sup>, Mathias Beck<sup>1</sup>, Thomas Beck<sup>4</sup>, Willy Benz<sup>4,2</sup>, Xavier Bonfils<sup>32</sup>, Luca Borsato<sup>20</sup>, Vincent Bourrier<sup>1</sup>, David R. Ciardi<sup>33</sup>, Andrew Collier Cameron<sup>34</sup>, Sébastien Charnoz<sup>35</sup>, Marion Cointepas<sup>1,32</sup>, Szilárd Csizmadia<sup>36</sup>, Patricio Cubillos<sup>20,7</sup>, Gaspard Lo Curto<sup>37</sup>, Melvyn B. Davies<sup>38</sup>, Tansu Daylan<sup>39</sup>, Magali Deleuil<sup>40</sup>, Adrien Deline<sup>1</sup>, Laetitia Delrez<sup>41,42</sup>, Olivier D. S. Demangeon<sup>5</sup>, Brice-Olivier Demory<sup>2,4</sup>, Caroline Dorn<sup>43</sup>, Xavier Dumusque<sup>1</sup>, David Ehrenreich<sup>1</sup>, Anders Erikson<sup>36</sup>, Alain Lecavelier des Etangs<sup>44</sup>, Diana de Miguel<sup>45</sup>, Andrea Fortier<sup>4,2</sup>, Luca Fossati<sup>7</sup>, Yolanda G. C. Frensch<sup>1</sup>, Malcolm Fridlund<sup>46,47</sup>, Michaël Gillon<sup>41</sup>, Manuel Güdel<sup>48</sup>, Maximilian N. Günther<sup>49</sup>, Janis Hagelberg<sup>1</sup>, Christiane Helling<sup>7,50</sup>, Sergio Hoyer<sup>40</sup>, Kate G. Isaak<sup>49</sup>, Laszlo L. Kiss<sup>51,52</sup>, Kristine W. F. Lam<sup>36</sup>, Jacques Laskar<sup>53</sup>, Baptiste Lavie<sup>1</sup>, Christophe Lovis<sup>1</sup>, Demetrio Magrin<sup>20</sup>, Luca Marafatto<sup>20</sup>, Pierre Maxted<sup>11</sup>, Scott McDermott<sup>54</sup>, Djamel Mékarnia<sup>9</sup>, Christoph Mordasini<sup>4,2</sup>, Felipe Murgas<sup>27</sup>, Valerio Nascimbeni<sup>20</sup>, Louise D. Nielsen<sup>55</sup>, Göran Olofsson<sup>21</sup>, Roland Ottensamer<sup>48</sup>, Isabella Pagano<sup>56</sup>, Enric Pallé<sup>26,27</sup>, Gisbert Peter<sup>36</sup>, Giampaolo Piotto<sup>20,19</sup>, Don Pollacco<sup>16</sup>, Didier Queloz<sup>43,23</sup>, Roberto Ragazzoni<sup>20,19</sup>, Devin Ramos<sup>15</sup>, Nicola Rando<sup>49</sup>, Heike Rauer<sup>36,57,58</sup>, Christian Reimers<sup>48</sup>, Ignasi Ribas<sup>28,29</sup>, Sara Seager<sup>3,59,60</sup>, Damien Ségransan<sup>1</sup>, Gaetano Scandariato<sup>56</sup>, Attila E. Simon<sup>4,2</sup>, Alexis M. S. Smith<sup>36</sup>, Manu Stalport<sup>42,41</sup>, Manfred Steller<sup>7</sup>, Gyula Szabó<sup>61,62</sup>, Nicolas Thomas<sup>4</sup>, Tyler A. Pritchard<sup>63</sup>, Stéphane Udry<sup>1</sup>, Carlos Corral Van Damme<sup>49</sup>, Valérie Van Grootel<sup>42</sup>, Eva Villaver<sup>26,27</sup>, Ingo Walter<sup>36</sup>, Nicholas Walton<sup>24</sup>, Cristilyn N. Watkins<sup>8</sup>, and Richard G. West<sup>16,64</sup>

(Affiliations can be found after the references)

Received 20 November 2023 / Accepted 25 January 2024

## ABSTRACT

We present the discovery and characterization of two warm mini-Neptunes transiting the K3V star TOI-815 in a K–M binary system. Analysis of its spectra and rotation period reveal the star to be young, with an age of  $200^{+400}_{-200}$  Myr. TOI-815b has a 11.2-day period and a radius of  $2.94 \pm 0.05 R_{\oplus}$  with transits observed by TESS, CHEOPS, ASTEP, and LCOGT. The outer planet, TOI-815c, has a radius of  $2.62 \pm 0.10 R_{\oplus}$ , based on observations of three nonconsecutive transits with TESS; targeted CHEOPS photometry and radial velocity follow-up with ESPRESSO were required to confirm the 35-day period. ESPRESSO confirmed the planetary nature of both planets and measured masses of  $7.6 \pm 1.5 M_{\oplus}$  ( $\rho_p = 1.64^{+0.33}_{-0.31} \text{ g cm}^{-3}$ ) and  $23.5 \pm 2.4 M_{\oplus}$  ( $\rho_p = 7.2^{+1.1}_{-1.0} \text{ g cm}^{-3}$ ), respectively. Thus, the planets have very different masses, which is unusual for compact multi-planet systems. Moreover, our statistical analysis of mini-Neptunes orbiting FGK stars suggests that weakly irradiated planets tend to have higher bulk densities compared to those undergoing strong irradiation. This could be ascribed to their cooler atmospheres, which are more compressed and denser. Internal structure modeling of TOI-815b suggests it likely has a H–He atmosphere that constitutes a few percent of the total planet mass, or higher if the planet is assumed to have no water. In contrast, the measured mass and radius of TOI-815c can be explained without invoking any atmosphere, challenging planetary formation theories. Finally, we infer from our measurements that the star is viewed close to pole-on, which implies a spin-orbit misalignment at the  $3\sigma$  level. This emphasizes the peculiarity of the system’s orbital architecture, and probably hints at an eventful dynamical history.

**Key words.** techniques: photometric – techniques: radial velocities – planets and satellites: composition – planets and satellites: detection – planets and satellites: formation – stars: individual: TOI-815

\* The photometric and radial velocity data are available at the CDS via anonymous ftp to [cdsarc.cds.unistra.fr](https://cdsarc.cds.unistra.fr) (130.79.128.5) or via <https://cdsarc.cds.unistra.fr/viz-bin/cat/J/A+A/685/A5>

\*\* This study uses TESS data, CHEOPS data observed as part of guaranteed time observation programs CH\_PR10031 and CH\_PR10048, and ESPRESSO data collected with the ESO 3.6 m telescope under programs 105.20P7.001, 109.23DX.001, and 110.2481.001 (PI: Bouchy).

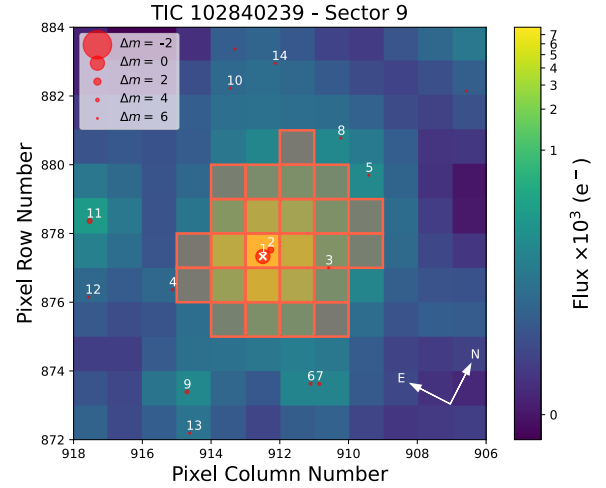
## 1. Introduction

The Transiting Exoplanet Survey Satellite (TESS; Ricker et al. 2015) and the *Kepler* space mission have profoundly transformed our understanding of planets by detecting thousands of exoplanets using the transit method, particularly those with radii below  $4 R_{\oplus}$  and relatively short orbital periods. Fulton et al. (2017), using precise planet radius measurements, identified a clear bimodal distribution in the radii of small planets. It appears that close-orbiting planets predominantly fall into two categories: super-Earths (below  $1.8 R_{\oplus}$ ) or sub-Neptunes ( $2\text{--}4 R_{\oplus}$ ), with very few planets occupying the intermediate range (the so-called radius valley). Four mechanisms have been proposed to explain the origin of these two populations: photoevaporation of the H-He envelopes due to high-energy radiation from the host star (Lopez & Fortney 2013; Owen & Wu 2017; Jin & Mordasini 2018; Rogers & Owen 2021; Attia et al. 2021), atmospheric mass-loss fueled by the cooling luminosity of a planet’s core (Ginzburg et al. 2018; Gupta & Schlichting 2019), a combination of ex situ formation and photoevaporation (Venturini et al. 2020), and an in situ gas-poor formation scenario (Lee et al. 2022). These mechanisms predict different compositions for the mini-Neptunes, ranging from dry planets composed of a H-He envelope atop a rocky core (Owen & Wu 2017; Jin & Mordasini 2018; Gupta & Schlichting 2019; Lee et al. 2022) to water worlds with a thin or no H-He atmosphere (Venturini et al. 2020). The mechanisms also predict distinct relationships between stellar and planetary mass, as well as differing locations of the valley (in the period-radius-mass space) and varying occurrence rates relative to the age of the systems (Berger et al. 2020; Petigura et al. 2022).

Numerous ongoing follow-up programs are dedicated to characterizing the masses of super-Earth and mini-Neptune planets located close to the “radius valley” of short-period orbits. However, there is limited coverage of mini-Neptunes with low stellar irradiation (less than  $30 S_{\oplus}$ ) and long orbital periods (greater than 15 days). We started a radial velocity (RV) follow-up campaign with the Échelle SPectrograph for Rocky Exoplanets and Stable Spectroscopic Observations (ESPRESSO; Pepe et al. 2021) dedicated to the characterizations and precise density measurements of warm mini-Neptune planets that transit FGK dwarfs. Currently, there are known 17 exoplanets in this parameter range with stars brighter than  $V=12$  and well-constrained densities<sup>1</sup> according to the PlanetS catalog<sup>2</sup> (Otegi et al. 2020).

Multi-planet systems provide the opportunity to study planets of different densities and compositions within a single system, and to compare such planets with each other. One observed trend in the diversity of planetary systems is that planets seem to resemble “peas in the pod”, that is to say, multi-transiting systems tend to have planets with similar sizes and geometrically spaced orbital distances (Weiss et al. 2018; Otegi et al. 2022; Mishra et al. 2023). Understanding these trends is crucial since they can elucidate underlying physical processes and constrain formation theories.

In this work we present the discovery of two mini-Neptune planets transiting TOI-815, a  $V = 11.2$  mag K3V star. The innermost planet, designated TOI-815b, has an orbital period of 11.2 days. This planet’s transits were captured by TESS during Sectors 6, 36, and 63, as well as by the CHaracterising



**Fig. 1.** TESS TPF of TOI-815 created with *tpfplotter* (Aller et al. 2020). The orange pixels define the aperture mask used for extracting the photometry. Additionally, the red circles indicate neighboring objects from the *Gaia* DR2 catalog, with the circle size corresponding to the brightness difference compared to the targets (as indicated in the legend). Our target is marked with a white cross. Pixel scale is  $21'' \text{ pixel}^{-1}$ .

ExOPlanet Satellite (CHEOPS; Benz et al. 2021), the Antarctic Search for Transiting Exoplanets (ASTEP) telescope (Guillot et al. 2015), and the Las Cumbres Observatory Global Telescope (LCOGT; Brown et al. 2013). As for the outer planet, TOI-815c, TESS detected three transits during the aforementioned sectors; however, a unique period could not be determined. Additionally, targeted CHEOPS photometry was performed to investigate potential period aliases, leading to the confirmation of a 35-day orbital period for TOI-815c. The planetary nature of both objects has been confirmed via RV follow-up using ESPRESSO.

The paper is structured as follows: in Sect. 2 we provide a detailed description of the space- and ground-based observations with TESS, CHEOPS, ASTEP, and LCOGT, as well as the ESPRESSO RV observations. Section 3 details how we determined the host star parameters by combining high-resolution spectra. In Sect. 4, we present our global photometric and RV analysis and its results. Finally, in Sect. 5 we present a discussion of the system, and we conclude in Sect. 6.

## 2. Observations

### 2.1. TOI-815 in a binary system

Behrard et al. (2022) searched for stellar companions to TESS objects of interest (TOIs; Guerrero et al. 2021) by cross-matching them with the *Gaia* Early Data Release 3 (EDR3) catalog. TOI-815 was identified to be in a binary system with the nearest known neighbor in the TESS Input Catalog Version 8 (TICv8; Stassun et al. 2019), TIC 102840237 (*Gaia* DR3 5415648821874435584) that is located approximately  $359 \text{ au}$  ( $\sim 6''.4$ ) north of TOI-815 (Fig. 1). According to the TICv8 catalog, TIC 102840237 is an M-dwarf with  $T_{\text{eff}} = 3766 \pm 157 \text{ K}$  and is 2.5 magnitudes dimmer than TOI-815 in the *G* band. To compensate for flux contamination in the photometry, we included a dilution factor for the light curves, as presented below.

### 2.2. TESS photometry

The TOI-815 system (TIC 102840239) was observed in TESS Sector 9 (2018 December 12 to 2019 January 06) in 30-min

<sup>1</sup>  $\sigma_M/M \leq 25\%$  and  $\sigma_R/R \leq 8\%$ .

<sup>2</sup> Available on the Data & Analysis Center for Exoplanets (DACE) platform (<https://dace.unige.ch>).

cadence and Sectors 36 (2021 March 07 to 2021 April 01) and 63 (2023 March 10 to 2023 April 06) in 20-s cadence. In Sectors 9 and 36, the target was imaged on CCD 2 of camera 2, and in Sector 63, on CCD 1 of camera 2. The 30-min full frame image light curves were extracted by the MIT Quick Look Pipeline (QLP; Huang et al. 2020), resulting in the simple aperture photometry (SAP) flux and the QLP *Kepler* Spline SAP (KSPSAP; Vanderburg & Johnson 2014) flux. The 2-min and 20-s cadence photometry was produced by the Science Processing Operations Center (SPOC; Jenkins et al. 2016) pipeline at the NASA Ames Research Center, resulting in the SAP (Twicken et al. 2010; Morris et al. 2020) flux and the Presearch Data Conditioning SAP (PDCSAP; Smith et al. 2012; Stumpe et al. 2012, 2014) flux.

On 2019 June 21, the TESS data public website<sup>3</sup> announced the discovery of a 11.2-day TESS TOI (Guerrero et al. 2021), TOI-815b. The SPOC detected the transit signature of TOI-815b in searches of Sector 36, and a search of Sectors 36 and 63 with a noise-compensating matched filter (Jenkins 2002, 2010, 2020), and the signature was fitted with an initial limb-darkened transit model (Li et al. 2019) and passed all the diagnostic tests presented in the Data Validation report (Twicken et al. 2018). In particular, the difference image centroiding test for the multi-sector search located the host star within  $1.4 \pm 3.4$  arcsec of the transit source. Seven transits with a depth of  $1301 \pm 54$  ppm and a duration of  $3.06 \pm 0.43$  h were identified in Sectors 9, 36, and 63 (Fig. A.1). However, TOI-815 showed one additional transit in each sector with a transit depth of  $1004 \pm 73$  ppm and a duration of  $2.53 \pm 0.59$  h, suggesting a second transiting companion. As presented in this work, with follow-up observations using ESPRESSO and CHEOPS, we confirmed the planetary nature and characterized the second planet.

For the analysis of the 30-min cadence data we used the KSP-SAP flux. For the 2-min and 20-s cadence light curves, we used the uncorrected SAP flux and accounted for the systematic noise with the *juliet* package (Espinoza et al. 2019) using Gaussian process (GP) models with an approximate Matérn 3/2 kernel.

In order to check the existence of contaminant sources, we plotted the target pixel file (TPF; Fig. 1) along with the aperture mask used for the SAP flux. The plot was generated using *tpfplotter* (Aller et al. 2020). The apertures used for extracting the light curves in all three sectors were contaminated by TIC 102840237 (Sect. 2.1) and therefore we included a dilution coefficient during the modeling of the TESS light curves. Our TESS light curve analysis is described in more detail in Sect. 4.4.

### 2.3. CHEOPS observations

CHEOPS is a 32 cm aperture ESA telescope dedicated to the characterization of transiting exoplanets. It performs high-precision photometry in a broad optical passband (330–1100 nm). The instrument is optimized for bright stars, making it an ideal facility to observe small-amplitude transits created by low-mass planets (e.g., Delrez et al. 2021; Leleu et al. 2021), or characterize the minute emission signals of hot Jupiters at optical wavelengths (Lendl et al. 2020; Krenn et al. 2023). Among various scientific goals, the CHEOPS space mission is also suited to improve the radius precision of known exoplanets discovered by TESS (Bonfanti et al. 2021a) by observing additional transits and to recover the true period of double transit candidates detected by TESS (e.g., Osborn et al. 2023; Ulmer-Moll et al. 2023; Garai et al. 2023). TOI-815 was jointly observed by two programs within the CHEOPS guaranteed time observation

(GTO): one targeting improved radius precision for multi-planet systems detected by NASA TESS (CHESS; PR0031) and the other targeted on confirming the orbital periods of long-period candidates found by TESS (DUOS; PR0048). The observations had an exposure time of 60 s, typical duration of 5–7 orbits (8.2–11.4 h), and achieved efficiencies between 51 and 73% (e.g., due to Earth occultations, Earth’s stray light contamination and South Atlantic Anomaly crossings). CHEOPS data are automatically processed with an automatic aperture photometry pipeline (Hoyer et al. 2020). We however re-extracted the photometry using point spread function (PSF) photometry with the PIPE package<sup>4</sup> (Brandeker et al. 2022), which, due to the large asymmetric PSF of CHEOPS, typically results in photometry with less influenced by external factors such as companion stars, roll-angle trends, etc. PIPE takes into account the flux contamination from nearby stars by subtracting a synthetic image using the PSF and *Gaia* DR2 parameters for the neighbor stars. Therefore, the dilution factor for the CHEOPS light curves was fixed to one.

### 2.4. Ground-based photometry

The TESS pixel scale is  $\sim 21''$  pixel<sup>-1</sup> and photometric apertures typically extend out to roughly  $1'$ , generally causing multiple stars to blend in the TESS aperture. To determine the true source of transit signals in the TESS data, improve the transit ephemerides, and check the SPOC pipeline transit depth after accounting for the crowding metric, we conducted photometric ground-based follow-up observations with ASTEP and LCOGT of the field around TOI-815 as part of the TESS Follow-up Observing Program (TFOP)<sup>5</sup> Sub Group 1 (Collins 2019). We used the TESS Transit Finder, which is a customized version of the *Tapir* software package (Jensen 2013), to schedule our transit observations.

#### 2.4.1. ASTEP

We observed TOI-815 with ASTEP on 2021 April 04 but could only rule out an ingress of TOI-815b that was, at the time, predicted to occur at BJD 2459309.02. Subsequent analysis (see Sect. 5.5) shows that as a result of transit timing variations (TTVs), the prediction was incorrect, and that the ingress should have started at  $2459309.93 \pm 0.0285$ , later than the end of this first series of observations. On 2022 September 13, using an improved camera system (Dransfield et al. 2022; Schmider et al. 2022) we observed the target again (Fig. 2), with 15 s exposures, at wavelengths ranging from about 700 to 1000 nm (Schmider et al. 2022). The data were processed on site with an IDL-based aperture photometry pipeline (Abe et al. 2013; Mékarnia et al. 2016). We detected an on-time 2.0 ppt transit of TOI-815b, using a 13.9 arcsec target aperture. While the ingress was clearly evident, the egress was affected by sky brightening, with a Sun at about  $12^\circ$  below the horizon at end of the observation sequence.

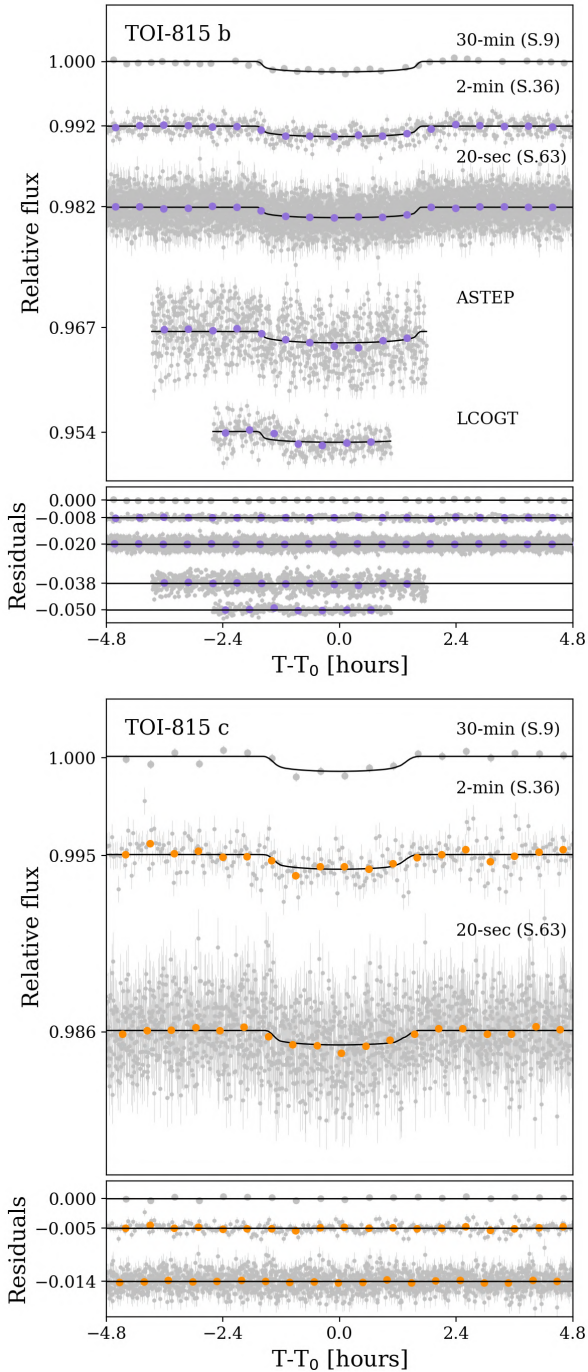
#### 2.4.2. LCOGT

We observed an intended full transit window of TOI-815b on 2021 February 18 (Fig. 2) in Pan-STARRS *z*-short band using the LCOGT 1.0m network node at South Africa Astronomical Observatory (SAAO). However, after updating the ephemeris from TESS Sector 36 data, we determined that the follow-up observations were out of transit by approximately one day.

<sup>4</sup> <https://github.com/alphapsa/pipe>

<sup>5</sup> <https://tess.mit.edu/followup/>

<sup>3</sup> <https://tev.mit.edu/data/>



**Fig. 2.** Phase-folded TESS, ASTEP, and LCOGT light curves of TOI-815b (top panel) and TESS light curves of TOI-815c (bottom panel). The data are binned to 30 min (purple and orange circles).

Therefore, these data are not considered further in this work. We observed a partial transit window of TOI-815b on 2023 March 11 using the same band-pass and LCOGT network node. The images were calibrated by the standard LCOGT BANZAI pipeline (McCully et al. 2018) and differential photometric data were extracted using AstroImageJ (Collins et al. 2017). We used circular photometric apertures with radius  $3''.9$ . The target star was diluted by 0.85% by the neighbor star (Sect. 2.1). The best zero, one, or two detrending vectors were retained if they decreased the Bayesian information criterion (BIC) for a fit by at least two per detrend parameter. We found that a proxy

for sky transparency and the Y-centroid of the target star on the detector detrending pair provided the best improvement to the light curve fit that was justified by the BIC values.

### 2.5. Spectroscopic follow-up with ESPRESSO

We acquired 34 high-resolution spectroscopic observations of TOI-815 using ESPRESSO on the 8.2 m Very Large Telescope (VLT) located in Paranal, Chile. The observations were carried out between 2021 April 04 and 2022 December 04 as part of the observing programs 105.20P7.001, 109.23DX.001 and 110.2481.001 (PI: Bouchy), dedicated to the characterization of warm mini-Neptune transiting exoplanets. The observations have an integration time ranging from 600 to 700 s, a median resolving power of 140,000 using  $2 \times 1$  binning, and a wavelength range of 380–788 nm. The RVs and activity indicators were extracted using version 3.0.0. of the ESPRESSO pipeline, and we computed the RVs by cross-correlating the Echelle spectra with a G2 numerical mask.

The ESPRESSO pipeline was devised to correct the residual atmosphere dispersion, that is, after atmospheric dispersion corrector correction. This reduces significantly the chromatic effects on the data. Out of the total observations, six were identified as unreliable since the flux correction was not applied and therefore were excluded from the global modeling. The average uncertainty of the RV data is  $0.68 \text{ m s}^{-1}$  and the RMS is  $0.74 \text{ m s}^{-1}$ . We report the ESPRESSO RV measurements and their uncertainties, along with the full width at half maximum (FWHM), bisector, contrast, S-index and  $H\alpha$ -index in Table B.1. The excluded data points are noted in the same Table. The RV time series and the phase-folded RVs for TOI-815b and TOI-815c are shown in Fig. 3.

### 2.6. High-resolution imaging

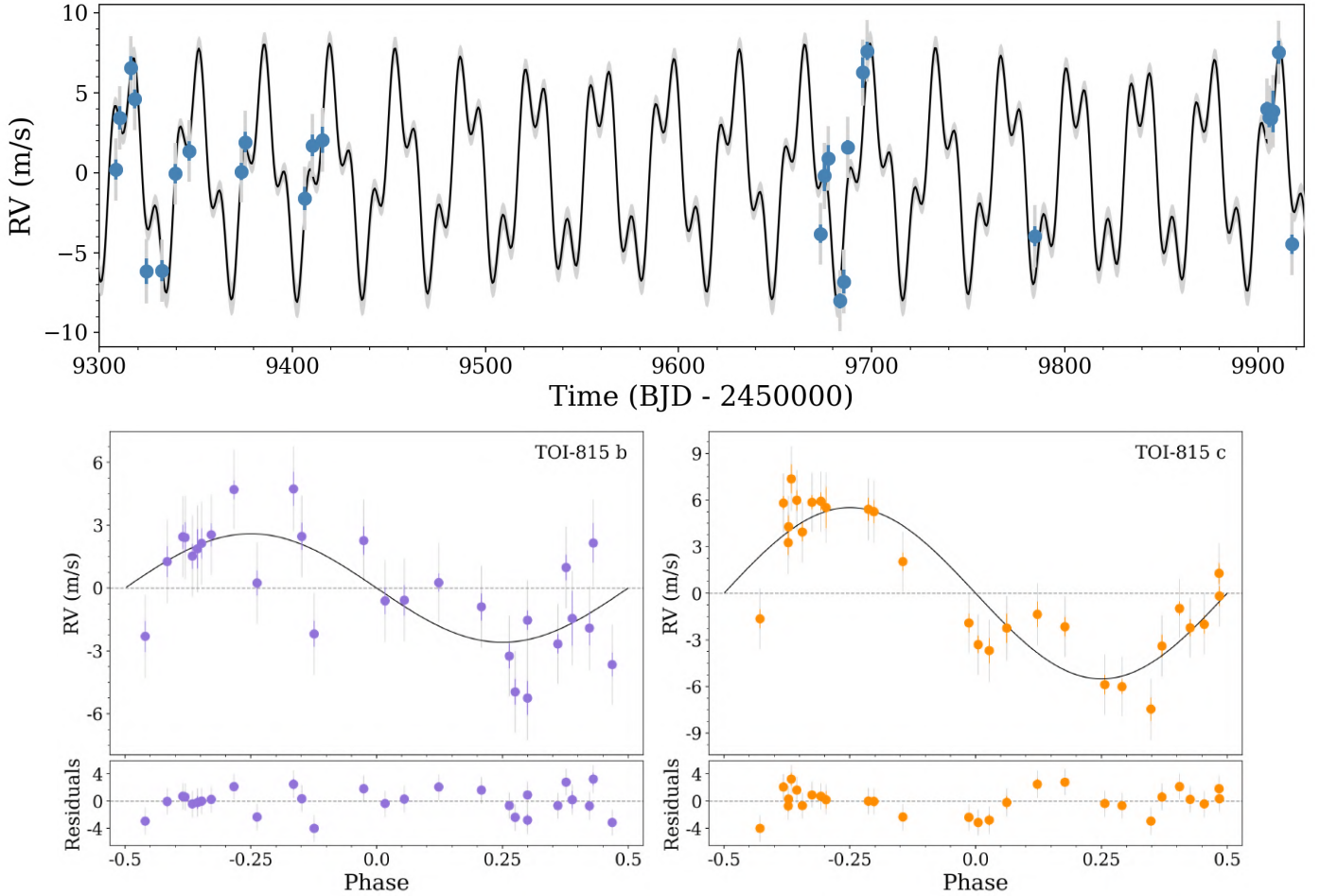
High-angular-resolution imaging is needed to search for nearby sources that can contaminate the TESS photometry, resulting in an underestimated planetary radius, or be the source of astrophysical false positives, such as background eclipsing binaries. We searched for stellar companions to TOI-815 with speckle imaging on the 4.1 m Southern Astrophysical Research (SOAR) telescope (Tokovinin 2018) on 2019 July 14, observing in Cousins  $I$ -band, a similar visible band-pass as TESS. This observation was sensitive to a 5.8-magnitude fainter star at an angular distance of 1 arcsec from the target. More details of the observations within the SOAR TESS survey are available in Ziegler et al. (2020). The  $5\sigma$  detection sensitivity and speckle autocorrelation functions from the observations are shown in Fig. 4. No nearby stars were detected within  $3''$  of TOI-815 in the SOAR observations. The binary star mentioned in Sect. 2.1 has a separation of  $6''.4$ , which places it outside SOAR's field of view; as a result, it was not detected.

## 3. Stellar characterization

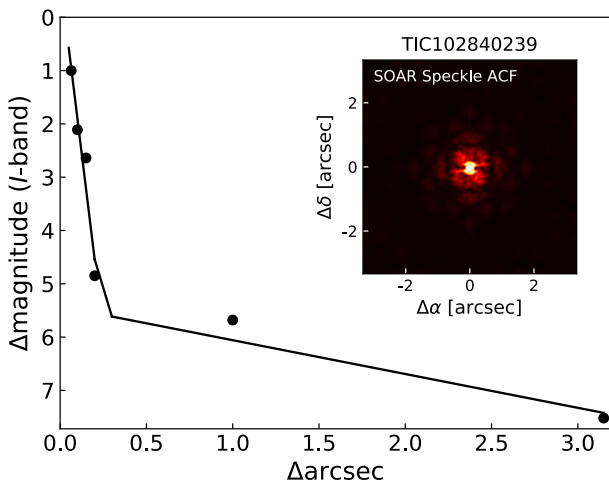
### 3.1. Spectroscopic parameters

The stellar spectroscopic parameters ( $T_{\text{eff}}$ ,  $\log g$ , microturbulence,  $[\text{Fe}/\text{H}]$ ) were estimated using the ARES+MOOG methodology. The methodology is described in detail in Sousa et al. (2021, 2008); Sousa (2014); Santos et al. (2013). For this, we used the latest version of ARES<sup>6</sup>

<sup>6</sup> The last version, ARES v2, can be downloaded at <https://github.com/sousasag/ARES>



**Fig. 3.** Relative RV measurements with ESPRESSO and best-fit models for TOI-815. *Top:* time series of ESPRESSO RV measurements (in blue). The complete inferred model, which comprises signals from two planets along with the activity-induced signal, is represented by a continuous solid line. *Bottom:* phase-folded RVs for TOI-815b (left) and TOI-815c (right). In all plots, the error bars in gray account for the estimated jitter ( $\sigma_{w, \text{ESPRESSO}}$ ).



**Fig. 4.** SOAR speckle imaging ( $5\sigma$  upper limits) of TOI-815 that rules out the presence of stellar companions within  $3''$ . The image inset illustrates the speckle autocorrelation function.

(Sousa et al. 2007, 2015) to consistently measure the equivalent widths of selected iron lines on the combined ESPRESSO spectrum of TOI-815. Because the star has an effective temperature

below 5200 K, we used the appropriate list of iron lines presented in Tsantaki et al. (2013). In this analysis we used a minimization process to find the ionization and excitation equilibrium to converge for the best set of spectroscopic parameters. This process makes use of a grid of Kurucz model atmospheres (Kurucz 1993) and the radiative transfer code MOOG (Snedden 1973). We also derived a more accurate trigonometric surface gravity using recent *Gaia* data following the same procedure as described in Sousa et al. (2021).

Using the stellar atmospheric parameters, we determined the abundances of Mg ( $[\text{Mg}/\text{H}] = -0.03 \pm 0.09$  dex) and Si ( $[\text{Si}/\text{H}] = -0.03 \pm 0.06$  dex), closely following the classical curve-of-growth analysis method described in, for example, Adibekyan et al. (2012, 2015). As for the stellar parameter determination, we used ARES to measure the equivalent widths of the spectral lines of these elements, and used a grid of Kurucz model atmospheres (Kurucz 1993) and the radiative transfer code MOOG (Snedden 1973) to convert the equivalent widths into abundances under assumption of local thermodynamic equilibrium.

Finally, we estimated the projected rotational velocity,  $v \sin i_*$ , by performing spectral synthesis around isolated iron lines in the region of Li line at  $6708 \text{ \AA}$  (discussed in Sect. 3.2). We used the code MOOG and the same model atmosphere used for the abundance determination. The spectral lines are very narrow and, despite the very high resolution of ESPRESSO spectra,

**Table 1.** Stellar parameters for TOI-815.

	TOI-815	Source
Identifiers		
TIC ID	102840239	TICv8
2MASS ID	J10232924-4350059	2MASS
<i>Gaia</i> ID	5415648821879172096	<i>Gaia</i> DR3
TYC	7721-01451-1	<i>Tycho</i> -2
Astrometric parameters		
Right ascension (J2016), $\alpha$	10 <sup>h</sup> 23 <sup>m</sup> 29.25 <sup>s</sup>	<i>Gaia</i> DR3
Declination (J2016), $\delta$	−43° 50′ 05.84″	<i>Gaia</i> DR3
Parallax (mas)	16.825 ± 0.013	<i>Gaia</i> DR3
Distance (pc)	59.437 ± 0.045	<i>Gaia</i> DR3
$\mu_{\text{RA}}$ (mas yr <sup>−1</sup> )	8.898 ± 0.011	<i>Gaia</i> DR3
$\mu_{\text{Dec}}$ (mas yr <sup>−1</sup> )	7.248 ± 0.013	<i>Gaia</i> DR3
Photometric parameters		
TESS (mag)	9.360 ± 0.006	TICv8
<i>B</i> (mag)	11.145 ± 0.061	<i>Tycho</i> -2
<i>V</i> (mag)	10.217 ± 0.005	<i>Tycho</i> -2
<i>G</i> (mag)	9.9398 ± 0.0004	<i>Gaia</i> DR3
<i>J</i> (mag)	8.531 ± 0.024	2MASS
<i>H</i> (mag)	8.066 ± 0.029	2MASS
<i>K</i> (mag)	7.999 ± 0.029	2MASS
Bulk parameters		
Spectral type	K3	Sect. 3.1
$T_{\text{eff}}$ (K)	4869 ± 77	Sect. 3.1
$R_{\star}$ ( $R_{\odot}$ )	0.770 ± 0.009	Sect. 3.2
$M_{\star}$ ( $M_{\odot}$ )	0.776 ± 0.036	Sect. 3.2
[Fe/H] (dex)	−0.09 ± 0.05	Sect. 3.1
[Si/H] (dex)	−0.03 ± 0.06	Sect. 3.1
[Mg/H] (dex)	−0.03 ± 0.09	Sect. 3.1
$\log g_{\star}$ (cm s <sup>−2</sup> )	4.51 ± 0.04	Sect. 3.1
$v \sin i_{\star}$ (km s <sup>−1</sup> )	<1	Sect. 3.1
$P_{\text{rot}}$ (days)	15.3 ± 1.2	Sect. 3.3
$i_{\star}$ (°)	32 <sup>+15</sup> <sub>−23</sub>	Sect. 3.3
Age (Myr)	200 <sup>+400</sup> <sub>−200</sub>	Sect. 3.2
$A(\text{Li})$ (dex)	1.15 ± 0.05	Sect. 3.2

**References:** (1) TICv8 (Stassun et al. 2019); (2) 2MASS (Skrutskie et al. 2006); (3) *Gaia* DR3 (Gaia Collaboration 2021); (4) *Tycho*-2 (Høg et al. 2000).

we are only able to set an upper limit of  $v \sin i_{\star} < 1 \text{ km s}^{-1}$ . All the parameters are presented in Table 1.

### 3.2. Mass, radius, and age

We determined the stellar radius of TOI-815 via a modified Markov chain Monte Carlo (MCMC) infrared flux method (Blackwell & Shallis 1977; Schanche et al. 2020) by constructing spectral energy distributions using stellar atmospheric models from two catalogs (Kurucz 1993; Castelli & Kurucz 2003) and the results of our spectral analysis as priors. By comparing synthetic derived from these spectral energy distribution and the observed broadband photometry in the following bandpasses: *Gaia* *G*,  $G_{\text{BP}}$ , and  $G_{\text{RP}}$ , 2MASS *J*, *H*, and *K*, and WISE *W1* and *W2* (Skrutskie et al. 2006; Wright et al. 2010; Gaia Collaboration 2023) we derived the bolometric flux of TOI-815. This is converted into effective temperature and angular diameter that is subsequently combined with the offset-corrected *Gaia* parallax (Lindgren et al. 2021) to determine the stellar radius. To allow

for model uncertainties to propagate to our stellar radius, we used a Bayesian modeling averaging of the radius posterior distributions produced using the aforementioned stellar catalogs.

We determined the stellar mass  $M_{\star}$  and age  $t_{\star}$  by inputting  $T_{\text{eff}}$ , [Fe/H], and  $R_{\star}$  (Table 1) along with their uncertainties in the isochrone placement algorithm (Bonfanti et al. 2015, 2016). Following interpolation of the input parameters within precomputed grids of PARSEC<sup>7</sup> v1.2S (Marigo et al. 2017) isochrones and evolutionary tracks, we obtained  $M_{\star} = 0.776 \pm 0.036 M_{\odot}$  and  $t_{\star} = 200^{+400}_{-200}$  Myr, where the error bars were computed by adding 4% and 20% in quadrature to the internal uncertainties coming from the interpolation scheme (as justified in Bonfanti et al. 2021a) to account for the isochrone precision. On the one hand, the young age we derived is apparently in disagreement with the low  $v \sin i$  value. On the other hand, an older main sequence star would have a smaller stellar radius differing more than  $3\sigma$  from our  $R_{\star}$  estimate.

To further assess the evolutionary stage of the star, we derived the Li abundance by performing a spectral synthesis (Fig. C.1) using the same model atmospheres and radiative transfer code as to obtain the spectroscopic parameters and abundances (see Delgado Mena et al. 2014, for further details). We obtain a value of  $A(\text{Li})^8 = 1.15$  dex, which is very high for stars with such cool temperatures; their thick convective envelopes allow for the mixing of material into interior regions hot enough to burn Li. As a consequence, Li is depleted at early ages for cool stars. The comparison of this Li abundance to that of stars of similar  $T_{\text{eff}}$  in open clusters also points to a very young age. In particular, the stars around 4900 K in the Hyades (with an age of ~600 Myr) have  $A(\text{Li})$  below 1 dex (see Fig. 5 in Sestito & Randich 2005). Similar values are observed in Praesepe, with an age of ~670 Myr (Cummings et al. 2017). Therefore, TOI-815 should be younger than the Hyades and Praesepe. On the other hand, for clusters of 150–250 Myr (e.g., M35, NGC 6475, and M34) analyzed in Sestito & Randich (2005), Anthony-Twarog et al. (2018), stars with  $T_{\text{eff}} \sim 4900$  K have Li around 1.5 dex. The recent work on M48, a 420 Myr old cluster (Sun et al. 2023) shows that stars with effective temperatures near 4900 K have lower Li values (around 1 dex) than TOI-815, adding an additional constrain for an age younger than ~400 Myr. The Li abundance of TOI-815 is shown with those clusters as comparison in Fig. 5.

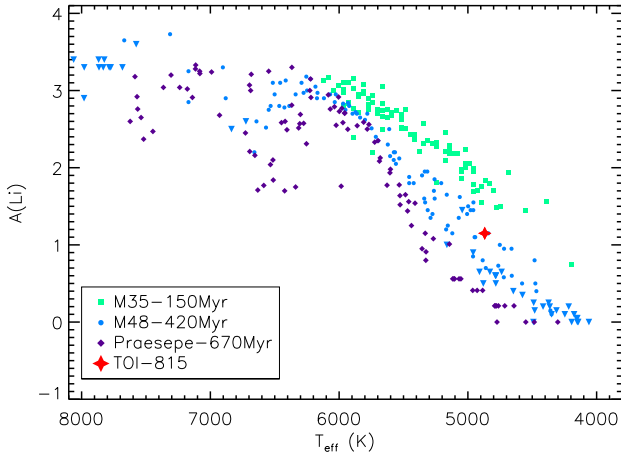
To gain a more precise understanding of TOI-815’s age, we utilized a recently developed kinematic-based method (Maciel et al. 2011; Almeida-Fernandes & Rocha-Pinto 2018) that relies on a higher probability of kinematic disturbances and thus stellar Galactic velocity relative to known kinematic families over the lifetimes of stars (Wielen 1977; Nordström et al. 2004; Casagrande et al. 2011). We compared the Galactic  $U$ ,  $V$ , and  $W$  velocities of TOI-815 to kinematic-age probability distributions bench-marked using a sample of 9000 stars with well-known isochronal ages (Almeida-Fernandes & Rocha-Pinto 2018) and determine a nominal age of ~470 Myr. Finally, the gyrochronological relations of Barnes (2007) and the rotation period of  $P_{\text{rot}} = 15.3 \pm 1.2$  days constrained from the photometry (Sect. 3.3) suggest an age of  $660 \pm 150$  Myr.

### 3.3. Stellar rotation

Using the ESPRESSO RVs, we measured a median  $\log R'_{\text{HK}} = -4.67 \pm 0.03$ , which suggests that TOI-815 is a quite active (see

<sup>7</sup> PAdova and TRieste Stellar Evolutionary Code: <http://stev.oapd.inaf.it/cgi-bin/cmd>

<sup>8</sup>  $A(\text{Li}) = \log[N(\text{Li})/N(\text{H})] + 12$ .



**Fig. 5.** Lithium abundance as a function of  $T_{\text{eff}}$ . The green, blue, and purple points represent the position of targets in the M35 (Sestito & Randich 2005; Anthony-Twarog et al. 2018), M48 (Sun et al. 2023), and Praesepe (Cummings et al. 2017) clusters, respectively. TOI-815 is marked in red.

Fig. 7 in Henry et al. 1996) star. Visual inspection of the raw TESS transit light curves shows a sinusoidal variation that is likely attributed to the stellar activity (Fig. A.1). For our analysis, we used the generalized Lomb–Scargle (GLS; Zechmeister & Kürster 2009) on the TESS SAP photometry of Sector 63 and we identified a significant periodic signal with a period of  $\sim 16.5$  days, which could be associated with the stellar rotation (Fig. D.1). TOI-815 was also observed by the Wide Angle Search for Planets (WASP; Pollacco et al. 2006) transit survey in multiple years between 2006 and 2014, with the observing season in each year spanning  $\sim 120$  nights. A total of 86 000 photometric data points was obtained using 200-mm,  $f/1.8$  Canon lenses until 2012 and 85 mm  $f/1.2$  Canon lenses after that. TOI-815 is by far the brightest star in the extraction aperture. In the years 2008, 2012, and 2013 we detect a significant periodicity with a period of 14.2–15.7 days. In each of those years the estimated false alarm probability (FAP) is below 1%. By combining the TESS Sector 63 and WASP photometric data, we estimate the stellar rotation to be  $15.3 \pm 1.2$  days.

We repeated the frequency analysis for the ESPRESSO RV measurements, as well as for the CaIIH&K activity index ( $\log R'_{\text{HK}}$ ) and the cross-correlation-function asymmetry indicators (FWHM and bisector). The GLS periodogram of the RV measurements shows a strong peak at the orbital frequency of TOI-815c ( $P_{\text{orb,c}} \sim 35$  days) with a FAP  $< 1\%$ . Following the subtraction of the outer planet’s signal, we found a significant peak at the orbital period of TOI-815b ( $P_{\text{orb,b}} \sim 11.2$  days) with a FAP  $\approx 1\%$ . Moreover, the stellar rotation period has no counterpart in the GLS periodograms of the RVs or the activity indicators.

The most straightforward way to estimate the stellar inclination is to assume solid body rotation (e.g., Borucki & Summers 1984; Doyle et al. 1984):

$$\sin i_{\star} = \frac{v \sin i_{\star}}{v} = \frac{v \sin i_{\star}}{2\pi R_{\star}/P_{\text{rot}}}, \quad (1)$$

where  $P_{\text{rot}}$  is the stellar equatorial rotation period. However, directly using Eq. (1) to generate an  $i_{\star}$  distribution is hazardous, because  $v \sin i_{\star}$  and  $P_{\text{rot}}$  are correlated quantities (Masuda & Winn 2020). To avoid such interdependence biases, we used the same procedure as Attia et al. (2023) to derive the stellar

inclination. We fitted for  $v \sin i_{\star}$  with a MCMC (using the emcee package, Foreman-Mackey et al. 2013) based on the following formulation of Eq. (1),

$$v \sin i_{\star} = \frac{2\pi R_{\star}}{P_{\text{rot}}} \sqrt{1 - \cos^2 i_{\star}}, \quad (2)$$

and using a semi-Gaussian distribution for  $v \sin i_{\star}$  (since we only have an upper limit) as a constraint for the computation of the likelihood. We set  $R_{\star}$ ,  $P_{\text{rot}}$ , and  $\cos i_{\star}$  as independent jump parameters with measurement-informed Gaussian priors on  $R_{\star}$  and  $P_{\text{rot}}$  (Table 1), and a uniform prior on  $\cos i_{\star}$  (i.e., assuming an isotropic stellar inclination distribution). The number of walkers and the burn-in phase are adjusted to ensure the convergence of the chains. The ensuing probability distribution function (PDF) yields  $i_{\star} = 32^{+15}_{-23}^{\circ}$ , the posterior estimate corresponding to the median value and the error bars to the 68% highest density interval. This result indicates that the star is spinning faster ( $v = 2.5 \pm 0.2 \text{ km s}^{-1}$ ) but we are observing it close to pole-on.

Additionally, we used the TESS photometric data to constrain the fraction of stellar surface covered with spots. For this, we used the SAGE<sup>9</sup> tool to obtain flux variability models for a given spot configuration. We fitted the observed variability using the same procedure as described in Chakraborty et al. (2023), with uniform priors on latitude, longitude and size of active regions. The wide priors allow spots to be present over the entire stellar surface and the spot sizes vary from small Sun-like spots ( $1^{\circ}$ ) to large M-dwarf-type spots ( $15^{\circ}$ ). To limit the degeneracy on constraining the parameters of active regions, we fixed the stellar inclination (to  $32^{\circ}$ ) and spot contrast in the TESS bandpass (to 0.2691). We iteratively increased the number of spots from 0 to 4 and found that a 1-spot model is most probable and has the lowest BIC. We conclude that at least  $0.26^{+0.20}_{-0.05}\%$  of the stellar surface is covered with spots to match the observed photometric variability. We also estimated the impact of these variations on the radius determination of the planets to be below 0.31%.

## 4. Photometric and RV analysis

### 4.1. TESS-only modeling

As mentioned in Sect. 2.2, visual inspection of the light curves of TOI-815 in Sectors 9 and 36 revealed two similar single transits incompatible with the known  $P = 11.2$  days inner planet. The two transits were separated by a gap of 734.5 days, which initially led to 41 separate period aliases between 17.49 days ( $P_{\text{min}}$ ) and 734.5 days ( $P_{\text{max}}$ ), where period on aliases shorter than this value would have produced a detectable third transit in TESS photometry. In order to determine the most likely periods, we modeled the system using Monotools<sup>10</sup> (Osborn 2022). This uses a period-agnostic fit of the two transit events, combined with priors from period, eccentricity and system stability to determine a marginalized posterior probability for each period. In the case of TOI-815c, we found that the innermost aliases are most likely, with the 23 periods shorter than  $P = 40$  days having probabilities greater than 1% (see Fig. 6).

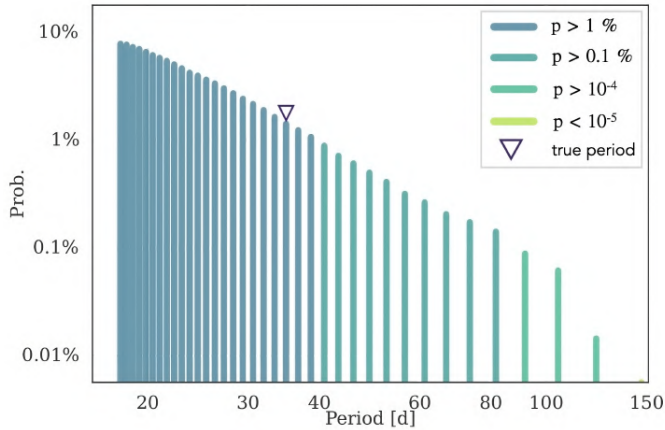
### 4.2. CHEOPS-only modeling

We initially scheduled 13 of the most-likely period aliases of TOI-815c (all  $P < 25$  days) with CHEOPS at varying priorities.

<sup>9</sup> <https://github.com/chakra/sage>

<sup>10</sup> <https://github.com/hposborn/MonoTools>





**Fig. 6.** Marginalized probabilities for each of the period aliases of TOI-815c as calculated by *Monotools* from TESS-data only. The triangle indicates the true period of the planet.

This resulted in five observations in 2022, light curves for which are noted in Table F.1 and can be seen in Fig. F.1. The most important observation occurred on 2022 March 17 (visit 3); with this observation, due to its position exactly 367.25 days after the last transit (i.e.,  $t_1 + P_{\max}/2$ ), we were able to rule out half of all aliases. The CHESS program (PR110048) also observed a single observation of TOI-815b, which showed little to no measurable TTVs. Observations resumed in 2023, with a further seven aliases scheduled, of which five were observed. One such observation, on 2023 February 28 (visit 10), was serendipitously during the transit of planet b, adding a second CHEOPS transit. On 2023 March 19, TESS re-observed a third transit of TOI-815c. However, this occurred exactly 734.5 days after the previously observed transit (i.e.,  $t_2 = t_1 + P_{\max}$ ), and therefore no period information was gained from the observation of this transit. We extracted a quick-look light curve using TICA (Fausnaugh et al. 2020) in order to verify the position of this third transit and updated the ephemeris of scheduled aliases. Thanks to these observations, very few aliases with  $P < 40$  days remained at this point. On 2023 April 23 (visit 11), a transit of TOI-815c at a unique transit time for the 34.97 days alias was finally observed with CHEOPS. Once a transit had been found all other observations were stopped. The CHEOPS light curves of TOI-815b and c are presented in Fig. 7.

#### 4.3. ESPRESSO-only modeling

As discussed in Sect. 2.5, we conducted a spectroscopic follow-up of TOI-815 using ESPRESSO over a period of 20 months. In Fig. D.1, we show that based on these observations we detected TOI-815c with a statistically significant signal at 35 days (FAP < 1%) and marginally detected TOI-815b at 11.2 days (FAP  $\sim$  1%). After fixing the transit ephemerides, the signal of TOI-815c and b was detected at a  $10\sigma$  and  $5\sigma$  level, respectively.

#### 4.4. Joint analysis

We performed an analysis of the planetary system around TOI-815 by combining photometric and RV data. To jointly model the datasets, we used the *juliet* software package, which is a Bayesian framework capable of modeling multi-planet systems. *Juliet* is an openly accessible tool that enables the simultaneous fitting of both transit models (using the *batman*

package, Kreidberg 2015) and RVs (via the *radvel* package, Fulton et al. 2018). It incorporates GPs to model noise (via the *celerite* package; Foreman-Mackey et al. 2017). To explore the parameter space, the code uses nested sampling algorithms with *dynesty* (Speagle 2020) and conducts model comparison through Bayesian evidence (lnZ) using *MultiNest* (Feroz et al. 2009) in conjunction with the *PyMultiNest* (Buchner et al. 2014) Python software package.

*Juliet* employs an efficient parametrization technique for the transit model by fitting the parameters  $r1$  and  $r2$ , which are a combination of the planet-to-star ratio ( $p$ ) and impact parameter ( $b$ ). This approach ensures a uniform exploration of the  $(p, b)$  space. Furthermore, *juliet* estimates the stellar density  $\rho_*$ , which, when combined with the orbital period, determines the scaled semimajor axis ( $a/R_*$ ).

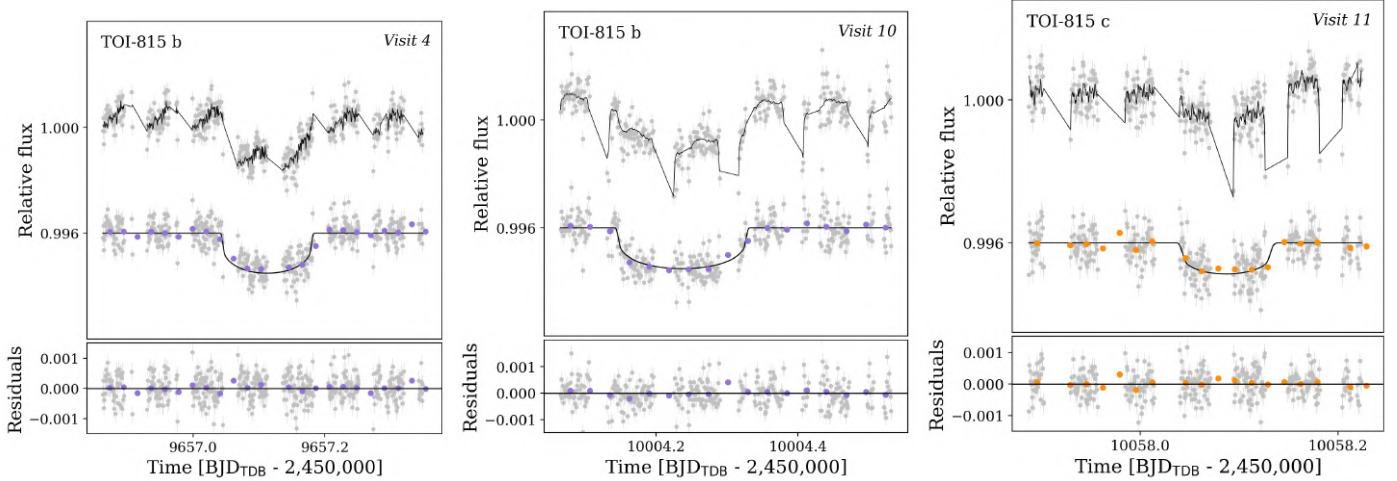
To account for TESS photometric variability, we included a GP for Sectors 36 and 63 using a Matérn 3/2 kernel with hyperparameters amplitude ( $\sigma_{GP}$ ) and timescale ( $\rho_{GP}$ ) as shown in Fig. A.1. The CHEOPS light curves exhibit correlations with certain instrumental parameters like contamination by background stars (bg), centroid movements ( $x, y$ ) and the change in temperature of the telescope tube  $\Delta T$ . Additionally, periodic noise features occur once per CHEOPS orbit due to the satellite's rolling motion around its pointing direction. To account for these periodic instrumental effects, we detrended the light curves by considering the sine or cosine of the roll angle ( $\phi$ ). In our global analysis, we incorporated linear models to account for both instrumental trends and stellar variability. However, it was essential to select only the relevant detrending parameters for each specific CHEOPS observation. To achieve this, we employed the *pycheops* package (Maxted et al. 2022) that provided us with the optimal combination of detrending terms for each visit. The detrending terms for each CHEOPS visit are listed in Table F.1.

The stellar density was fitted using a Gaussian prior based on the stellar spectroscopic analysis conducted in Sect. 3. We derived the quadratic stellar limb-darkening coefficients and their uncertainties for each photometric filter used using the LDCU<sup>11</sup> code (Deline et al. 2022). The LDCU code computes the limb-darkening coefficients and their corresponding uncertainties by considering a set of stellar intensity profiles and accounting for the uncertainties on the stellar parameters. In our global analysis, we utilized the determined limb-darkening coefficients from LDCU as Gaussian priors for all TESS, CHEOPS, LCOGT, and ASTEP data.

The SPOC crowding metric calculated for TOI-815 in Sectors 36 and 63 is  $\sim$ 0.85. This indicates that, according to SPOC modeling after eliminating background noise,  $\sim$ 85% of the flux detected within the photometric aperture originated from the target star. The remaining 15% was attributed to various other sources, with source 2 (TIC 102840237; Sect. 2.1) being particularly significant. To compensate for contamination by these additional sources, we fixed the dilution factor for all the light curves to 0.85.

In order to prevent any potential Lucy-Sweeney bias in the eccentricity measurement (Lucy & Sweeney 1971), we fixed the orbital eccentricity to zero. However, to explore the possibility of noncircular orbits, we ran a separate analysis without any constraints on the eccentricity. With the eccentricity fitted, we found a value of  $e = 0.110^{+0.072}_{-0.065}$  for TOI-815b and  $0.035^{+0.064}_{-0.016}$  for TOI-815c. Therefore, the condition  $e > 2.45\sigma_e$  is not satisfied, which suggests that the RV data are compatible with a circular orbit. Additionally, in Table E.1 we show that the  $3\sigma$  upper limit on

<sup>11</sup> <https://github.com/delinea/LDCU>



**Fig. 7.** CHEOPS light curves of TOI-815b (visits 4 and 10) and TOI-815c (visit 11). The top light curves in the upper panel present the raw flux and the best-fit de-correlation model. The detrending terms are presented in Table F.1. The bottom light curves in the upper panel present the detrended flux, with the best-fit transit model overlaid in black. The residuals of the models are shown in the lower panel. The data are binned to 30 min (purple and orange circles).

**Table 2.** Fitted and derived parameters for TOI-815b and TOI-815c.

Parameter	TOI-815b	TOI-815c
Orbital period, $P_{\text{orb}}$ (days)	$11.197259^{+0.000018}_{-0.000017}$	$34.976145^{+0.000099}_{-0.000097}$
Planet radius, $R_p$ ( $R_{\oplus}$ )	$2.94^{+0.05}_{-0.05}$	$2.62^{+0.10}_{-0.09}$
Planet mass, $M_p$ ( $M_{\oplus}$ )	$7.6^{+1.5}_{-1.4}$	$23.5^{+2.4}_{-2.4}$
Planet density, $\rho_p$ ( $\text{g cm}^{-3}$ )	$1.64^{+0.33}_{-0.31}$	$7.2^{+1.1}_{-1.0}$
Orbital inclination, $i$ (degrees)	$89.36^{+0.17}_{-0.14}$	$89.107^{+0.026}_{-0.026}$
Impact parameter, $b$	$0.279^{+0.058}_{-0.073}$	$0.839^{+0.015}_{-0.016}$
Semi-major axis, $a$ (au)	$0.0903^{+0.0018}_{-0.0019}$	$0.193^{+0.004}_{-0.004}$
Transit duration, $T_{14}$ (h)	$3.379^{+0.023}_{-0.023}$	$2.97^{+0.09}_{-0.10}$
Scaled semi-major axis, $a/R_*$	$25.2^{+0.6}_{-0.6}$	$53.9^{+1.3}_{-1.3}$
Equilibrium temperature <sup>(a)</sup> , $T_{\text{eq}}$ (K)	$686^{+13}_{-14}$	$469^{+9}_{-9}$
Atmospheric scale height <sup>(b)</sup> , $H$ (km)	$288^{+58}_{-54}$	$51^{+7}_{-6}$
Insolation <sup>(c)</sup> , $S_p$ ( $S_{\oplus}$ )	$36.9^{+2.9}_{-2.9}$	$8.03^{+0.64}_{-0.64}$
Restricted Jeans escape param. <sup>(d)</sup> , $\Lambda$	$28.6^{+5.7}_{-5.3}$	$145^{+16}_{-16}$

**Notes.** <sup>(a)</sup>The equilibrium temperature is calculated using  $T_{\text{eq}} = T_{\text{eff}}(1-A)^{1/4} \sqrt{\frac{R_*}{2a}}$ , assuming a Bond albedo of  $A = 0$ . <sup>(b)</sup>The scale height is calculated using  $H = \frac{k_b T_{\text{eq}}}{\mu} \text{g}$  assuming a mean molecular weight of  $2.3 \text{ g mol}^{-1}$ . <sup>(c)</sup>Insolation is calculated using  $S_p = \frac{L_*}{4\pi a^2}$ . <sup>(d)</sup>The restricted Jeans escape parameter is calculated using  $\Lambda = \frac{GM_p m_{\text{H}}}{k_B T_{\text{eq}} R_p}$  (Fossati et al. 2017).

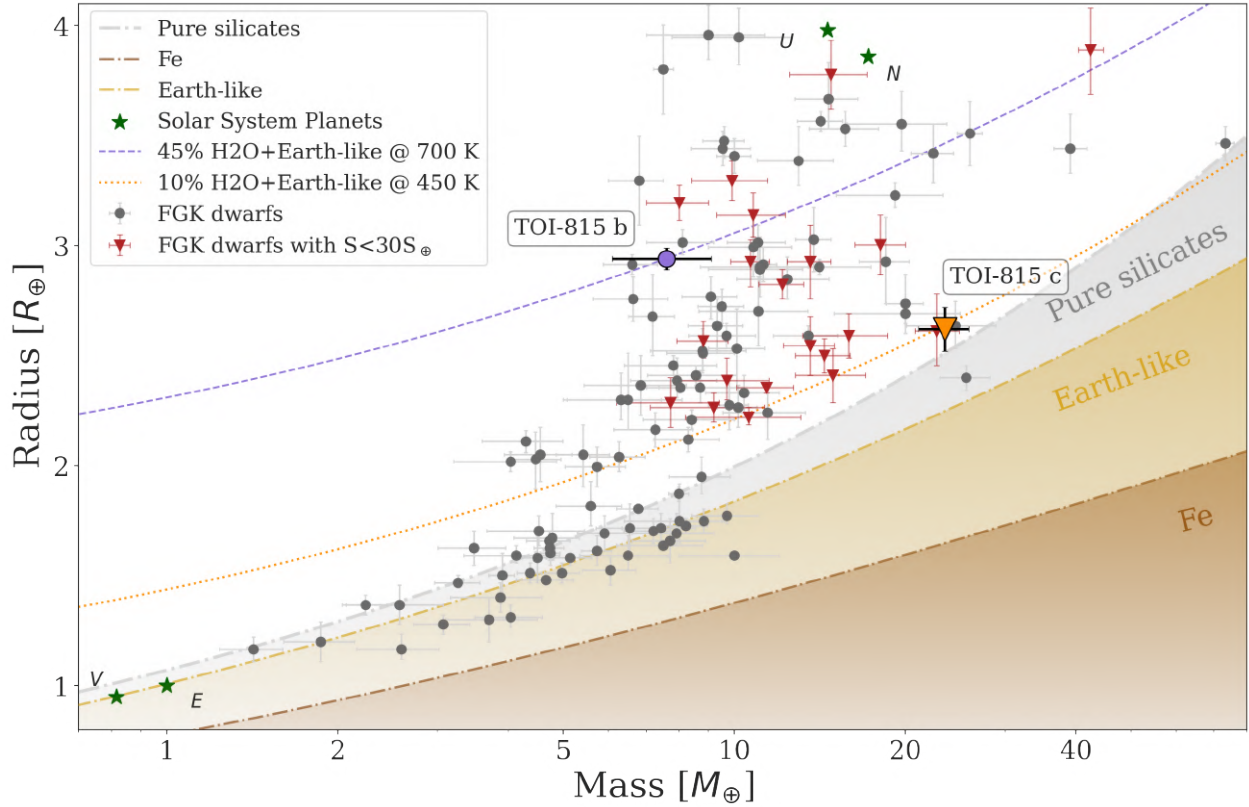
the eccentricity for TOI-815b and TOI-815c is 0.20 and 0.22, respectively. The fitted and derived parameters for TOI-815b and TOI-815c are presented in Table 2.

## 5. Discussion

We report the discovery and characterization of two transiting mini-Neptune planets around TOI-815, a young K3V star. TOI-815b is a 11.2 days period planet with a radius of  $2.94 \pm 0.05 R_{\oplus}$  and a mass of  $7.6 \pm 1.5 M_{\oplus}$ . The outermost planet, TOI-815c, has an orbital period of 35 days, a radius of  $2.62 \pm 0.10 R_{\oplus}$  and a mass of  $23.5 \pm 2.4 M_{\oplus}$ . Based on the stellar parameters

(Table 1), we find that the semimajor axis of TOI-815b and TOI-815c are  $0.0903 \pm 0.0019 \text{ au}$  and  $0.193 \pm 0.004 \text{ au}$ , respectively. Assuming a zero albedo and full heat redistribution ( $f = 1$ ), the equilibrium temperature of the inner planet is  $686^{+13}_{-14} \text{ K}$  and its stellar insolation is  $36.9 \pm 2.9 S_{\oplus}$ . The outer planet has an equilibrium temperature of  $469^{+9}_{-9} \text{ K}$  and receives a total amount of insolation equal to  $8.03 \pm 0.64 S_{\oplus}$ .

Both planets lie within the population of sub-Neptunes that are transiting FGK stars, with 87 well-characterized exoplanet detections<sup>1</sup> (including TOI-815b and TOI-815c) from the PlanetS catalog<sup>2</sup>. TESS has contributed in this sample with 46 discoveries during the past 5 yr (Naponiello et al. 2023). Moreover, TOI-815c adds to the population of warm mini-Neptune-sized planets with low stellar irradiation (less than  $30 S_{\oplus}$ ); with its inclusion, the total number of well-characterized detections is now 21. Figure 8 places TOI-815b and TOI-815c in the mass–radius diagram of small exoplanets ( $2\text{--}4 R_{\oplus}$ ) with precise densities orbiting FGK stars (gray points). The red triangles correspond to exoplanets with stellar irradiation lower than  $30 S_{\oplus}$ . The composition lines of pure silicates (Zeng et al. 2016), Earth-like planets, and pure iron (Dorn et al. 2015) are displayed in gray, light-brown, and brown, respectively. We also plot two other composition lines from Aguichine et al. (2021) as dashed and dotted black lines. For TOI-815b, we plot a line at 700 K, corresponding to its equilibrium temperature, with a composition of 45% of water and the rest as an Earth-like composition (32.5% iron–67.5% silicates). Since TOI-815c is denser and less irradiated, we can see that it lies on the composition lines with 10% water and an Earth-like composition at 450 K, its equilibrium temperature. Finally, we can say that if TOI-815b is rich in water for the model under consideration, its atmosphere is inflated by its irradiation, giving it a large radius and therefore a low density for a quantity of water of around 45%. TOI-815c, on the other hand, would have a lower water content of less than 10%, and its lower irradiation would not have as much influence as its neighbor on its radius, which would explain its higher density. Of course, this remains one possible model, and it is worth remembering that the position of a planet in the mass-radius diagram does not uniquely determine its composition. These planets could also have atmospheres rich in H-He or heavier materials. A more detailed discussion of the internal structures can be found in Sect. 5.3.



**Fig. 8.** Mass-radius diagram of small exoplanets (with radii ranging from 2–4  $R_{\oplus}$ ) with precise densities<sup>1</sup> from the PlanetS catalog<sup>2</sup> orbiting FGK stars. The red triangles correspond to exoplanets with stellar irradiation lower than  $30 S_{\oplus}$ . The composition lines of pure silicates (gray) from Zeng et al. (2016), Earth-like planets (light-brown), and pure iron (brown) from Dorn et al. (2015) are displayed. The purple and orange points represent TOI-815b and c, respectively. Two composition lines (Aguichine et al. 2021) that incorporate both water and terrestrial elements matching the equilibrium temperatures of the two planets are plotted as purple dashed and orange dotted lines.

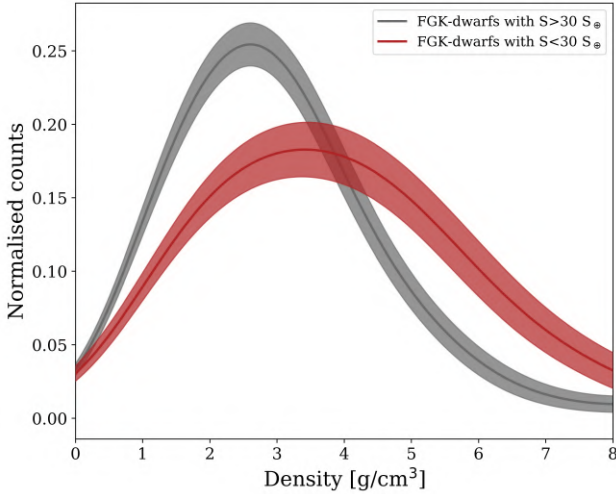
### 5.1. Density discrepancy of lowly and highly irradiated mini-Neptunes

TOI-815c, with a bulk density of  $7.2 \pm 1.1 \text{ g cm}^{-3}$ , becomes the densest well-characterized warm mini-Neptune (2–4  $R_{\oplus}$  and  $\leq 30 S_{\oplus}$ ), with GJ 143b ( $\rho_p = 7.0 \pm 1.6 \text{ g cm}^{-3}$  and  $S_p = 6.0 \pm 0.4 S_{\oplus}$ ; Dragomir et al. 2019) being the second. Also, TOI-815c is the third densest well-characterized mini-Neptune (2–4  $R_{\oplus}$ ), the other two being K2-292b ( $\rho_p = 7.4 \pm 1.6 \text{ g cm}^{-3}$  and  $S_p = 72 \pm 11 S_{\oplus}$ ; Luque et al. 2019) and Kepler-411b ( $\rho_p = 10.3 \pm 1.3 \text{ g cm}^{-3}$  and  $S_p = 196 \pm 9 S_{\oplus}$ ; Sun et al. 2019). We analyzed the differences in bulk density of planets with 2–4  $R_{\oplus}$  and  $\leq 30 S_{\oplus}$  (“warm” sub-Neptunes) and planets of the same size but with  $\geq 30 S_{\oplus}$  (“hot” sub-Neptunes). In the first group, we have 21 planets with a median density of  $3.9 \text{ g cm}^{-3}$ . In the second, we have 65 planets with a median density of  $2.9 \text{ g cm}^{-3}$ . We performed a resampling method to account for uncertainties in density measurements: a large number of new datasets ( $N = 10\,000$  samples) are drawn, which lie within the normal distribution of the density measurements of the planets. For each of these new datasets, the kernel density estimation (KDE) is calculated. Similar to histograms, its aim is to estimate an unknown probability density function from a sample of data but in a more continuous way. We then calculated the mean and standard deviation of all the KDEs obtained. The result for the two groups is shown in Fig. 9. There is evidence suggesting that sub-Neptune-size planets orbiting FGK stars with low stellar irradiation (“warm” sub-Neptunes) exhibit a slightly higher bulk density compared to those exposed to strong irradiation. To test whether the

difference in densities of planets in these two irradiation conditions is statistically significant, we utilized a two-sample Kolmogorov-Smirnov test, which studies whether the two data samples come from the same distribution. With a  $p$ -value of 0.026, we find that the densities of hot and warm sub-Neptunes belong to different distributions with over 95% confidence. The inclusion of TOI-815b and c in the samples plays an important role in the statistical significance of this result, as they reduce the  $p$ -value from 0.051 to 0.026, enabling us to reach the 95% confidence interval. This difference in density could potentially be attributed to the cooler atmospheres of these less irradiated planets, which are more compressed and denser. However, such a pattern could be also linked to observational bias since more massive planets at long periods are easier to characterize using RVs. Therefore, further precise measurements of the mass of similar exoplanets using precise RVs are necessary.

### 5.2. Similarity in the density of multi-planet systems

Well-characterized exoplanets in multi-planet systems are especially interesting since they allow the system’s architecture to be investigated. Otegi et al. (2022) studied the similarity of planets in radii, masses, densities, and period ratios within a multi-planet system and revealed that planets tend to be more similar in radius than in mass. This trend could be linked to the significant influence of the planetary radius on density compared to mass, and consequently, on the overall planetary composition. Interestingly, in their work, they observed a robust correlation



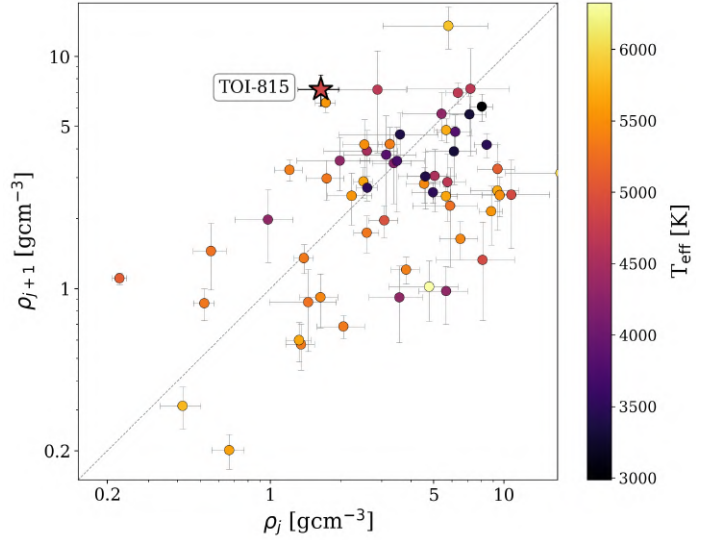
**Fig. 9.** KDE for the two groups of sub-Neptunes ( $2\text{--}4 R_{\oplus}$ ): with an irradiation above  $30 S_{\oplus}$  (gray) and below  $30 S_{\oplus}$  (red). The solid line represents the mean of the KDEs calculated with the 10 000 samples drawn, and the transparent region represents the standard deviation.

between densities of adjacent planets with no clear evidence of bi-modality.

In our study we used the NASA Exoplanet Archive<sup>12</sup> (Akeson et al. 2013) and narrowed down our selection to planets exhibiting mass and radius uncertainties below  $\sigma_M/M \leq 50\%$  and  $\sigma_R/R \leq 16\%$ . For planets with masses characterized by TTVs we used the catalog presented by Hadden & Lithwick (2017) and discarded the less robust (Kepler-25b and c, Marcy et al. 2014; Kepler-18c, Cochran et al. 2011). Our final sample consists of 116 planets within 39 multi-planet systems. In Fig. 10, we plot the density of a planet ( $j$ ) within our sample against the density of the next planet ( $j+1$ ) farther from the star. The points are color-coded by the host effective temperature. The Pearson correlation test resulted in a  $p$ -value = 0.04 that suggests a statistically significant linear correlation between densities. It is worth noting that planets orbiting M-dwarf stars tend to be closer to the 1:1 line compared to FGK stars with a  $p$ -value = 0.03. However, it is important to bear in mind that this observation is based on a limited sample of 22 planets within nine multi-planet systems. We calculated the similarity metric for the TOI-815 system using Eq. (2) of Millholland et al. (2017):

$$\mathcal{D}_X = \sum_{i=1}^{N_{\text{sys}}} \sum_{\substack{j=1 \\ P_j < P_{j+1}}}^{N_{\text{pl}}-1} \left| \log \frac{X_{j+1}}{X_j} \right|, \quad (3)$$

with  $X$  corresponding to the planetary radius or planetary mass, to be  $\mathcal{D}_R = 0.05$  and  $\mathcal{D}_M = 0.49$ . The metrics indicate that the planets in the TOI-815 system are very similar in radii but rather different in mass. Placing the TOI-815 system in Fig. 10, we clearly see that TOI-815 is one of the least uniform systems in density. Similar to TOI-815, the Kepler-107 system (Bonomo et al. 2023) consists of two planets, Kepler-107b and c, with nearly identical radii ( $1.536 \pm 0.025 R_{\oplus}$  and  $1.597 \pm 0.026 R_{\oplus}$ ) but significantly different densities ( $5.8^{+2.7}_{-2.6} \text{ g cm}^{-3}$  and  $13.5^{+2.9}_{-2.8} \text{ g cm}^{-3}$ ). Bonomo et al. (2019) showed that this density discrepancy can be explained by a giant impact on Kepler-107c. As presented in Sect. 5.4, such a mechanism can explain the large



**Fig. 10.** Density of a planet ( $j$ ) against the density of the next planet ( $j+1$ ; farther from the star). The TOI-815 system is indicated with a star symbol.

density of TOI-815c, which suggests that the two planets have formed in a similar way.

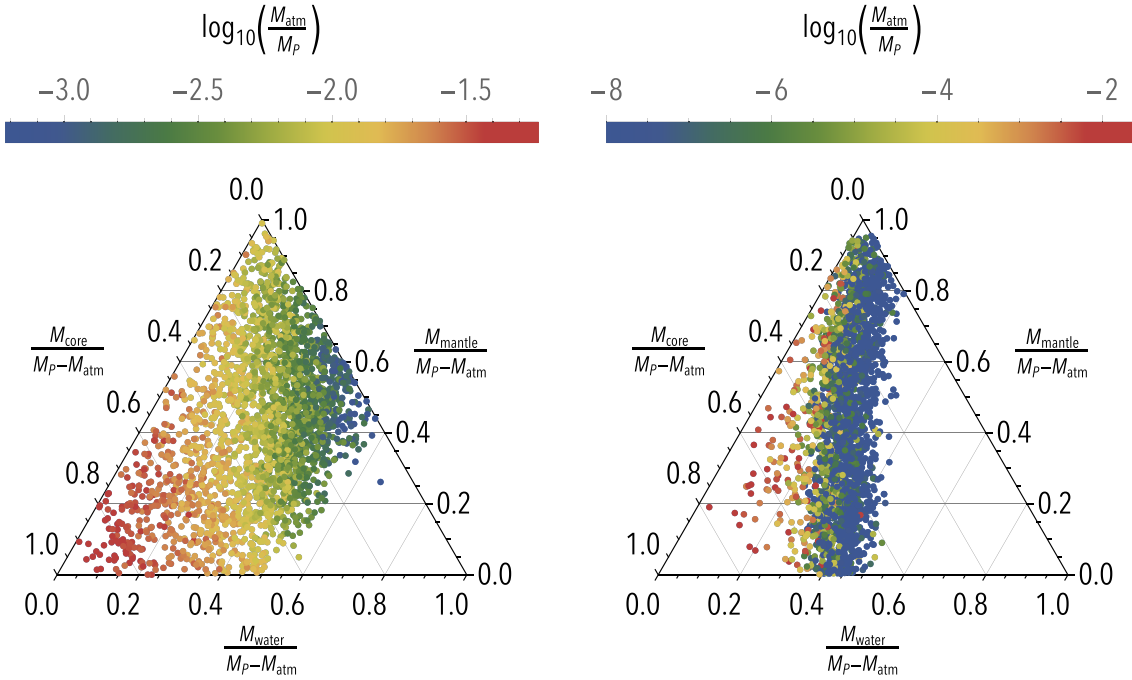
### 5.3. Planetary composition and internal structure

As shown in Fig. 8, the TOI-815 planetary system consists of two exoplanets that are located above an Earth-like composition line in the  $M$ - $R$  diagram. This already indicates that both planets include a considerable amount of volatile materials, and for TOI-815b, a sizable atmosphere.

To explore the range of possible interior structures, we used a nested sampling algorithm (Buchner et al. 2014) on the four-layered interior model (Dorn et al. 2017). This model contains an iron core, a rocky mantle, a water layer, and a H-He atmosphere. For both planets, we ran two bracketing interior models: one where we put no constraints on the layer masses (the free model), and one where we set the water mass fraction to zero in order to put an upper limit on the atmospheric mass fraction (the no-water model). Additionally, for TOI-815c, we included a model where we set the atmospheric mass fraction to zero (the no-atmosphere model). Such a model is not possible for TOI-815b, as its radius is too large to have no atmosphere. For all models, we varied the elemental ratios of [Fe/Si] and [Mg/Si] in the mantle and the planetary age, and we capped the water mass fraction at 50%. As a proxy for the planet's age, we used the age constraints of the host star. The results are summarized in Table 3 and Fig. 11.

As expected, we find that a plethora of interior structures are consistent with the observational constraints of TOI-815b and TOI-815c. These range from water-rich structures with little to no atmosphere in the case of TOI-815c, to structures devoid of water but with a non-negligible H-He atmosphere. However, while the internal structure of TOI-815b is rather poorly constrained, it is clear that the observational constraints can only be satisfied with an H-He atmosphere. This is in alignment with Hakim et al. (2018) who calculated the mass-radius diagram with different core and mantle compositions for planets with Moon-like, Earth-like, or Mercury-like internal structures. Even though they obtained a broad distribution of planetary radii for planets of the same mass, none of the models without an atmosphere can

<sup>12</sup> <https://exoplanetarchive.ipac.caltech.edu/>



**Fig. 11.** Ternary plots of the nested sampling results for the free model of TOI-815b (*left*) and TOI-815c (*right*). The axes are normalized with respect to  $M_p - M_{\text{atm}}$ , and the color of the points indicates the atmospheric mass fraction in logarithmic scale.

**Table 3.** Nested sampling results for the interior structure of the free model (top), the no-water model (middle), and the no-atmosphere model (bottom).

Planet (model)	$M_{\text{core}}/M_p$	$M_{\text{mantle}}/M_p$	$M_{\text{water}}/M_p$	$\log(M_{\text{atm}}/M_p)$
free:				
TOI-815b	$0.29 \pm 0.20$	$0.42 \pm 0.22$	$0.28 \pm 0.13$	$-2.16 \pm 0.40$
TOI-815c	$0.32 \pm 0.16$	$0.44 \pm 0.26$	$0.23 \pm 0.12$	$-7.70 \pm 2.69$
no-water:				
TOI-815b	$0.40 \pm 0.25$	$0.58 \pm 0.26$	0	$-1.64 \pm 0.22$
TOI-815c	$0.30 \pm 0.23$	$0.70 \pm 0.24$	0	$-2.78 \pm 0.75$
no-atmosphere:				
TOI-815c	$0.31 \pm 0.15$	$0.43 \pm 0.25$	$0.26 \pm 0.11$	N/A

reproduce the large radius of TOI-815b. This further supports the conclusion that it is impossible to construct a model of TOI-815b without a H-He atmosphere.

We emphasize that our interior structure model does not include the pollution of the H-He envelope with heavier elements, which leads to contraction of the atmosphere (e.g., Lozovsky et al. 2018). Furthermore, water might not only be present in an ocean layer, but also dissolved within the mantle of the planet (e.g., Dorn & Lichtenberg 2021) or be present in the atmosphere (e.g., Mol Lous et al. 2022). In addition, the assumption of distinct planetary layers might break down at high planetary masses (Helled & Stevenson 2017; Bodenheimer et al. 2018). Nonetheless, in face of the observational constraints currently available, these additions are unlikely to affect the main conclusions of this analysis.

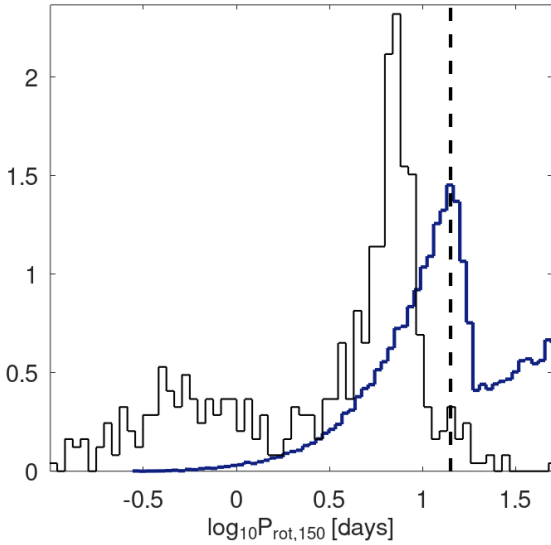
#### 5.4. Planetary atmospheric evolution

To model the atmospheric evolution of the planets in the TOI-815 system and constrain their primordial parameters, we present

modeling conducted using the Planetary Atmospheres and Stellar RoTation RATES (PASTA; Bonfanti et al. 2021b) code, which is a planetary atmospheric evolution code based on the original code presented by Kubyshkina et al. (2019a,b). As stars can be born with different initial rotation periods, their  $L_{\text{XUV}}$  values can spread over a wide range during the first  $\sim 1$  Gyr following protoplanetary disk dispersal (e.g., Tu et al. 2015). To account for that, PASTA simultaneously constrains the evolution of planetary atmospheres and of the stellar rotation rate by combining a model predicting planetary atmospheric escape rates based on hydrodynamic simulations (this has the advantage over other commonly used analytical estimates in that it accounts for both extreme-UV-driven and core-powered mass loss; Kubyshkina et al. 2018), a model of the evolution of the stellar extreme-UV (XUV) flux (Bonfanti et al. 2021b), a model relating planetary parameters and atmospheric mass (Johnstone et al. 2015b), and stellar evolutionary tracks (Choi et al. 2016).

PASTA works under two main assumptions: (1) planet migration did not occur after the dispersal of the protoplanetary disk; and (2) the planets hosted at some point in the past or still host a hydrogen-dominated atmosphere. The free parameters (i.e., subject to uniform priors) are the planetary initial atmospheric mass fractions at the time of the dispersal of the protoplanetary disk ( $f_{\text{atm}}^{\text{start}}$ ), which we assumed occurs at an age of 5 Myr (see for example Alexander et al. 2014; Kimura et al. 2016; Gorti et al. 2016), and the stellar rotation period at 150 Myr, which is used as a proxy for the stellar XUV emission. The code returns constraints on the free parameters and on their uncertainties by implementing the atmospheric evolution algorithm in a Bayesian framework (through the MC3 code of Cubillos et al. 2017), using the system parameters with their uncertainties as input priors. Details of the algorithm can be found in Bonfanti et al. (2021b).

As a proxy for the evolution of the stellar rotation period, and thus XUV emission, Fig. 12 displays the posterior distribution of the stellar rotation period at an age of 150 Myr ( $P_{\text{rot},150}$ ). This distribution is then compared to that of stars member of



**Fig. 12.** Posterior distribution (dark blue) of the stellar rotation rate of TOI-815 after 150 Myr derived by PASTA. The solid black line represents the distribution of the stellar rotation rate of young open cluster stars with masses comparable to that of TOI-815 based on the collection of data provided by Johnstone et al. (2015a). The dashed black line indicates the observed present-day stellar rotation rate.

young open clusters that are of comparable mass, as obtained from Johnstone et al. (2015a). Because the system is still very young ( $200^{+400}_{-200}$  Myr), the posterior distribution of  $P_{\text{rot},150}$  is mostly dominated by the measurement of the present-day rotation rate ( $15.3 \pm 1.2$  days). The distribution is, however, flatter than the prior on  $P_{\text{rot}}$ , because of the large relative uncertainty of the stellar age. The result indicates that the star could not be a fast rotator (i.e.,  $P_{\text{rot},150} < 1$  days), as constrained by the slow present-day rotation period.

Figure 13 shows the posterior distribution of the initial atmospheric mass fraction for TOI-815b (left panel) and TOI-815c (right panel) in comparison to the present-day atmospheric mass fraction. For planet b we find that PASTA clearly prefers evolutionary scenarios in which the evolution of the atmosphere was mostly unaffected by hydrodynamic escape. This indicates that the planet most likely still retains the majority of its primordial atmosphere. However, it must be noted that given the uncertainties of the observed present-day parameters PASTA is not able to completely exclude scenarios with significant escape.

In the case of TOI-815c, PASTA concludes that, independently of the evolution of the stellar rotation rate, the planet has not lost a significant amount of atmosphere via hydrodynamic escape. Given its present-day orbital position and relatively large core mass, it would be expected for the planet to have accreted an extensive hydrogen atmosphere if it had formed within the protoplanetary disk (Ikoma & Hori 2012; Lee & Chiang 2015; Mordasini 2020; Venturini & Helled 2017). However, such an envelope is not observed today and according to PASTA the planet is unable to have lost it due to hydrodynamic escape over the course of its evolution. One possible explanation would be formation within a gas-poor environment (e.g., Lee et al. 2022). Another explanation is that the planet suffered from at least one giant impact toward the end of the disk’s life (Inamdar & Schlichting 2015). This also aligns with the fact that the planets of the system are not in mean motion resonance and the expected spin-orbit misalignment. In both cases the results retrieved by PASTA are compatible with conclusions that can be drawn by the

more simplistic approach of only evaluating the restricted Jeans parameter  $\Lambda$  (Fossati et al. 2017). For TOI-815c the high value of  $\Lambda = 145^{+16}_{-16}$  would imply a stable atmosphere, while for TOI-815b the intermediate value of  $\Lambda = 28.6^{+5.7}_{-5.3}$  would favor a stable atmosphere but not exclude the possibility of a hydrodynamically escaping atmosphere. We note that in principle PASTA is not set up to simulate the atmospheric evolution of planets with mean planetary densities larger than that of Earth. This is due to the fact that the code always assumes for the rocky core of a planet to have a mean density equal to that of the Earth. In practice, PASTA will underestimate the planetary mass of TOI-815c as its observed density is  $1.30 \pm 0.20 \rho_{\oplus}$ . The planetary mass estimate for planet c returned by PASTA is  $M_{\text{P,c}} = 19.9 \pm 1.9 M_{\oplus}$ , which implies a bare rocky-core with a mean planetary density of  $1 \rho_{\oplus}$ . A larger planetary mass, however, would just further hinder hydrodynamic escape, as escaping particles would need even more energy to escape the gravitational well of the planet. Therefore, since hydrodynamic escape affects only marginally the evolution of the planetary atmosphere, the impact of the assumption on the core density has a negligible impact on the results.

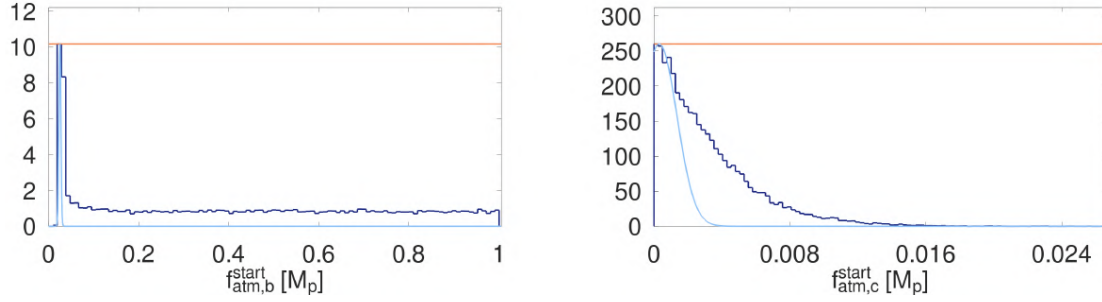
### 5.5. Search for transit timing variations

Although the ratio of the orbital period of TOI-815 and c ( $P_{\text{orb,c}}/P_{\text{orb,b}} \sim eq\ 3.12$ ) lies exterior to the 3:1 resonance, we conducted a TTV analysis to search for additional, non-transiting planets. Using juliet, we performed the analysis by utilizing all available photometric datasets from TESS, CHEOPS, ASTEP, and LCOGT. Instead of fitting a single period ( $P_{\text{orb}}$ ) and time-of-transit center ( $T_0$ ), juliet employs a method that seeks individual transit times. This approach involves fitting each transit independently and determining one transit time for each one, resulting in a more consistent and coherent analysis. Figure 14 illustrates the results of our analysis, which compares the observed transit times with the calculated linear ephemeris derived from all the transits. Notably, no significant variations were detected within the data.

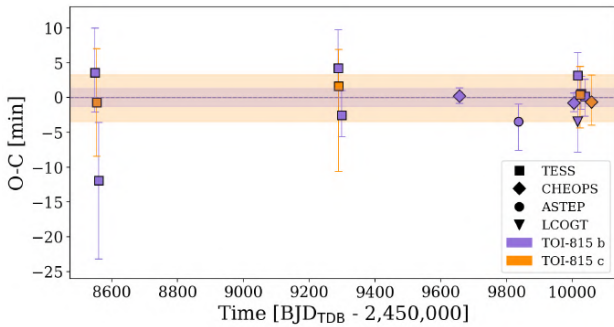
### 5.6. Potential for atmospheric and Rossiter-McLaughlin studies

We used the metric proposed by Kempton et al. (2018) to assess the potential of atmospheric characterization for the planets in the TOI-815 system with the James Webb Space Telescope (JWST; Gardner et al. 2023). To calculate the transmission spectroscopy metric for all exoplanets in the PlanetS catalog<sup>2</sup> with a radii from 2 to 4  $R_{\oplus}$ , we utilized the scale factors provided in Table 1 by Kempton et al. (2018) instead of the suggested value of 0.167 for temperate planets. Notably, the transmission spectroscopy metric values for TOI-815b and c were found to be  $97 \pm 20$  and  $15 \pm 2$ , respectively. Putting these values into context, TOI-815b is located in the top 10% of well-characterized exoplanets in this regime and therefore stands out as an excellent candidate for atmospheric characterization via transmission spectroscopy.

We also explored the possibility of studying the spin-orbit of both planets via Rossiter-McLaughlin (RM) observations (Rossiter 1924; McLaughlin 1924). According to Eq. (40) by Winn (2010), we expect the maximum amplitude of the RM effect to be  $\sim 1.13 \text{ m s}^{-1}$  and  $0.52 \text{ m s}^{-1}$  for TOI-815b and TOI-815c, respectively. Using the classical RM technique, the signals are small and challenging, but the sky-projected spin-orbit angle  $\lambda$  should be detectable using newer methods such as the RM



**Fig. 13.** Posterior distribution (dark blue) for the mass of the planetary atmosphere of TOI-815b (*left*) and TOI-815c (*right*) at the time of the dispersal of the protoplanetary disk. The light blue line represents the distribution of the estimated present-day atmospheric mass fraction derived by PASTA. The orange horizontal line indicates the uninformative prior distribution.



**Fig. 14.** Transit timings for each transit of TOI-815b (purple) and TOI-815c (orange) measured from the TESS and CHEOPS light curves. The purple and orange regions represent the  $1\sigma$  uncertainties of the inner and outer planet ephemeris reported in Table E.1.

effect revolutions technique (Bourrier et al. 2021). Moreover, since the star is seen close to pole-on and since the planets are transiting, the two orbits are probably misaligned. This can be seen using the following equation linking the 3D spin-orbit angle  $\psi$  to other observable quantities (e.g., Winn et al. 2007):

$$\cos \psi = \sin i_* \sin i \cos \lambda + \cos i_* \cos i, \quad (4)$$

yielding a small value for  $\cos \psi$  for any value of  $\lambda$  we measure. The case that would disfavor the orbital misalignment picture the most is if we observe an aligned sky-projected spin-orbit angle  $\lambda \sim 0^\circ$ . Sampling a  $\psi$  distribution from Eq. (4) with a Dirac distribution at  $0^\circ$  for  $\lambda$  and measurement-informed Gaussians for the other parameters, we find  $\psi = 58^{+18}_{-19}$ , which still represents a misaligned orbit at a  $3.2\sigma$  level. Given the young nature of the system, tidal realignment (e.g., Barker & Ogilvie 2010; Attia et al. 2023) may have not had time to occur. Hence, TOI-815 offers us a unique possibility to capture a peculiar orbital architecture, especially given the scarcity of misaligned orbits in multi-planet systems (e.g., Albrecht et al. 2022).

## 6. Conclusions

We have reported the detection and characterization of two sub-Neptune planets that are transiting TOI-815. TOI-815 is a young ( $200^{+400}_{-200}$  Myr) K dwarf in a K–M binary system. The planets were initially detected by the TESS mission and subsequently confirmed with CHEOPS observations as well as precise RV measurements obtained with ESPRESSO, which enabled the determination of their masses. Through the combined use of TESS, CHEOPS, LCOGT, and ASTEP photometry and the

ESPRESSO spectroscopic data, we precisely characterized both planets: TOI-815b has a period of 11.2 days, a radius of  $2.94 R_\oplus$ , and a mass of  $7.6 M_\oplus$ , while TOI-815c has a period of 35 days, a radius of  $2.62 R_\oplus$ , and a mass of  $23.5 M_\oplus$ . The combination of these facilities allowed us to determine the planetary radius and mass with precision better than 2% and 20% for the inner planet and 4% and 10% for the outer planet, respectively.

TOI-815c, with a bulk density of  $7.2 \pm 1.1 \text{ g cm}^{-3}$ , stands as the densest well-characterized warm mini-Neptune. Our study reveals a statistically significant difference between the bulk densities of warm sub-Neptunes and hot sub-Neptunes, with TOI-815b and c being instrumental in establishing a 95% confidence. Further exoplanet detections with precise mass measurements are needed to confirm this trend. Our analysis of the TOI-815 system using similarity metrics demonstrates that while the planets in this system have remarkably similar radii, they exhibit significantly different masses, making TOI-815 one of the least uniform systems in terms of planetary density.

Internal structure modeling that includes four layers (an iron core, a rocky mantle, a water layer, and a H-He atmosphere) suggests that TOI-815b likely has a small H-He atmosphere (on the order of a percent of the planet’s mass) if water is included. If no water is assumed, the atmospheric mass fraction can be an order of magnitude higher. On the contrary, the radius of TOI-815c can be reproduced without any atmosphere. This object is interesting because it is rather massive and yet seems to have only a tiny (or no) atmosphere. This is a challenge for planet formation theories since at these planetary masses strong accretion of H-He is expected (for comparison, Uranus and Neptune have lower masses but consist of  $\sim 10\%$  of H-He). A possible explanation would be for the planet to have lost the majority of its primordial atmosphere. However, our atmospheric evolution model shows that such an extensive mass loss cannot be caused by XUV-driven hydrodynamic escape. Thus, the most likely scenario is that planet c, similar to Kepler-107c (Bonomo et al. 2019), suffered from at least one giant impact (which removed a substantial amount of primordial gas) during the unstable dynamical phase of disk dispersal. This is consistent with the planets of this system not being in resonance. Alternatively, a formation within a gas-poor environment cannot be ruled out either.

Finally, the combination of the stellar rotation period ( $P_{\text{rot}} = 15.3 \pm 1.2$  days) constrained from the WASP and TESS photometry and the very low spectroscopically measured projected rotational velocity ( $v \sin i_* < 1 \text{ km s}^{-1}$ ) suggest that we are observing the star close to pole-on. Given that the planets are on orbits with inclinations close to  $\sim 90^\circ$ , we expect a spin-orbit misalignment at a  $3.2\sigma$  level. Future RM observations will be able to confirm this.

**Acknowledgements.** We thank the Swiss National Science Foundation (SNSF) and the Geneva University for their continuous support to our planet low-mass companion search programs. This work has been carried out within the framework of the National Centre of Competence in Research PlanetS supported by the Swiss National Science Foundation under grants 51NF40\_182901 and 51NF40\_205606. O.A. and V.B. have received funding from the European Research Council (ERC) under the European Union’s Horizon 2020 research and innovation programme (project SPICE DUNE, grant agreement No 947634; grant agreement no. 730890). This material reflects only the authors’ views and the Commission is not liable for any use that may be made of the information contained therein. The authors acknowledge the financial support of the SNSF. This publication makes use of The Data & Analysis Center for Exoplanets (DACE), which is a facility based at the University of Geneva (CH) dedicated to extrasolar planet data visualization, exchange, and analysis. DACE is a platform of NCCR PlanetS and is available at <https://dace.unige.ch>. This paper includes data collected by the TESS mission. Funding for the TESS mission is provided by the NASA Explorer Program. J.S.J. acknowledges support by FONDECYT grant 1201371 and partial support from the ANID Basal project FB210003. Funding for the TESS mission is provided by NASA’s Science Mission Directorate. K.A.C. and C.N.W. acknowledge support from the TESS mission via subaward s3449 from MIT. This research has made use of the Exoplanet Follow-up Observation Program (ExoFOP; DOI: 10.26134/ExoFOP5) website, which is operated by the California Institute of Technology, under contract with the National Aeronautics and Space Administration under the Exoplanet Exploration Program. This paper made use of data collected by the TESS mission and are publicly available from the Mikulski Archive for Space Telescopes (MAST) operated by the Space Telescope Science Institute (STScI). We acknowledge the use of public TESS data from pipelines at the TESS Science Office and at the TESS Science Processing Operations Center. Resources supporting this work were provided by the NASA High-End Computing (HEC) Program through the NASA Advanced Supercomputing (NAS) Division at Ames Research Center for the production of the SPOC data products. CHEOPS is an ESA mission in partnership with Switzerland with important contributions to the payload and the ground segment from Austria, Belgium, France, Germany, Hungary, Italy, Portugal, Spain, Sweden, and the United Kingdom. The CHEOPS Consortium would like to gratefully acknowledge the support received by all the agencies, offices, universities, and industries involved. Their flexibility and willingness to explore new approaches were essential to the success of this mission. CHEOPS data analyzed in this article will be made available in the CHEOPS mission archive (<https://cheops.unige.ch/archive-browser/>). This work makes use of observations from the ASTEP telescope. ASTEP benefited from the support of the French and Italian polar agencies IPEV and PNRA in the framework of the Concordia station program, from OCA, INSU, Idex UCAJEDI (ANR-15-IDEX-01) and ESA through the Science Faculty of the European Space Research and Technology Centre (ESTEC). This work makes use of observations from the LCOGT network. Part of the LCOGT telescope time was granted by NOIRLab through the Mid-Scale Innovations Program (MSIP). MSIP is funded by NSF. Co-funded by the European Union (ERC, FIERCE, 101052347). Views and opinions expressed are however those of the author(s) only and do not necessarily reflect those of the European Union or the European Research Council. Neither the European Union nor the granting authority can be held responsible for them. C.B. acknowledges support from the Swiss Space Office through the ESA PRODEX program. This work has been carried out within the framework of the NCCR PlanetS supported by the Swiss National Science Foundation under grants 51NF40\_182901 and 51NF40\_205606. M.L. acknowledges support of the Swiss National Science Foundation under grant number PCEFP2\_194576. S.G.S. acknowledges support from FCT through FCT contract nr. CEECIND/00826/2018 and POPH/FSE (EC). Y.A.I. acknowledges support from the Swiss National Science Foundation (SNSF) under grant 200020\_192038. A.Br. was supported by the SNSA. D.G. gratefully acknowledges financial support from the CRT foundation under Grant No. 2018.2323 “Gaseous rocky? Unveiling the nature of small worlds”. J.V. acknowledges support from the Swiss National Science Foundation (SNSF) under grant PZ00P2\_208945. M.N.G. is the ESA CHEOPS Project Scientist and Mission Representative, and as such also responsible for the Guest Observers (GO) Programme. M.N.G. does not relay proprietary information between the GO and Guaranteed Time Observation (GTO) Programmes, and does not decide on the definition and target selection of the GTO Programme. K.G.I. is the ESA CHEOPS Project Scientist and is responsible for the ESA CHEOPS Guest Observers Programme. She does not participate in, or contribute to, the definition of the Guaranteed Time Programme of the CHEOPS mission through which observations described in this paper have been taken, nor to any aspect of target selection for the programme. We acknowledge financial support from the Agencia Estatal de Investigación of the Ministerio de Ciencia e Innovación MCIN/AEI/10.13039/501100011033 and the ERDF “A way of making Europe” through projects PID2019-107061GB-C61, PID2019-107061GB-C66, PID2021-125627OB-C31, and PID2021-125627OB-C32, from the Centre of Excellence “Severo Ochoa” award to the Instituto de Astrofísica de Canarias (CEX2019-000920-S), from the Centre of Excellence “María de Maeztu” award

to the Institut de Ciències de l’Espai (CEX2020-001058-M), and from the Generalitat de Catalunya/CERCA programme. S.C.C.B. acknowledges support from FCT through FCT contracts no. IF/01312/2014/CP1215/CT0004. X.B., S.C., D.G., M.F. and J.L. acknowledge their role as ESA-appointed CHEOPS science team members. L.Bo., G.Br., V.Na., I.Pa., G.Pi., R.Ra., G.Sc., V.Si., and T.Zi. acknowledge support from CHEOPS ASI-INAF agreement no. 2019-29-HH.0. A.C.C. acknowledges support from STFC consolidated grant numbers ST/R000824/1 and ST/V000861/1, and UKSA grant number ST/R003203/1. P.E.C. is funded by the Austrian Science Fund (FWF) Erwin Schrodinger Fellowship, program J4595-N. This project was supported by the CNES. The Belgian participation to CHEOPS has been supported by the Belgian Federal Science Policy Office (BELSPO) in the framework of the PRODEX Program, and by the University of Liège through an ARC grant for Concerted Research Actions financed by the Wallonia-Brussels Federation; L.D. is an F.R.S.-FNRS Postdoctoral Researcher. This work was supported by FCT – Fundação para a Ciência e a Tecnologia through national funds and by FEDER through COMPETE2020 – Programa Operacional Competitividade e Internacionalização by these grants: UIDB/04434/2020; UIDP/04434/2020; 2022.06962.PTDC. O.D.S.D. is supported in the form of work contract (DL 57/2016/CP1364/CT0004) funded by national funds through FCT. S.G.S. acknowledges the support from FCT through Investigador FCT contract no. CEECIND/00826/2018 and POPH/FSE (EC). B.-O.D. acknowledges support from the Swiss State Secretariat for Education, Research and Innovation (SERI) under contract number MB22.00046. This project has received funding from the Swiss National Science Foundation for project 200021\_200726. It has also been carried out within the framework of the National Centre of Competence in Research PlanetS supported by the Swiss National Science Foundation under grant 51NF40\_205606. The authors acknowledge the financial support of the SNSF. M.F. and C.M.P. gratefully acknowledge the support of the Swedish National Space Agency (DNR 65/19, 174/18). T.D. acknowledges support by the McDonnell Center for the Space Sciences at Washington University in St. Louis. M.G. is an F.R.S.-FNRS Senior Research Associate. C.He. acknowledges support from the European Union H2020-MSCA-ITN-2019 under Grant Agreement no. 860470 (CHAMELEON). S.H. gratefully acknowledges CNES funding through the grant 837319. K.W.F.L. was supported by Deutsche Forschungsgemeinschaft grants RA714/14-1 within the DFG Schwerpunkt SPP 1992, Exploring the Diversity of Extrasolar Planets. This work was granted access to the HPC resources of MesoPSL financed by the Région Île-de-France and the project Equip@Meso (reference ANR-10-EQPX-29-01) of the programme Investissements d’Avenir supervised by the Agence Nationale pour la Recherche. P.M. acknowledges support from STFC research grant number ST/M001040/1. This work was also partially supported by a grant from the Simons Foundation (PI Queloz, grant number 327127). A. S. acknowledges support from the Swiss Space Office through the ESA PRODEX program. T.Wi. acknowledges support from the UKSA and the University of Warwick. N.A.W. acknowledges UKSA grant ST/R004838/1. V.V.G. is an F.R.S.-FNRS Research Associate. C.D. acknowledges support from the Swiss National Science Foundation under grant TMSGI2\_211313. Gy.M.Sz. acknowledges the support of the Hungarian National Research, Development and Innovation Office (NKFIH) grant K-125015, a PRODEX Experiment Agreement No. 4000137122, the Lendület LP2018-7/2021 grant of the Hungarian Academy of Science and the support of the city of Szombathely. D.R. was supported by NASA under award number NNA16BD14C for NASA Academic Mission Services. R.L. acknowledges funding from University of La Laguna through the Margarita Salas Fellowship from the Spanish Ministry of Universities ref. UNI/551/2021-May 26, and under the EU Next Generation funds. R. A. is a Trotter Postdoctoral Fellow and acknowledges support from the Trotter Family Foundation. This work was supported in part through a grant from the Fonds de Recherche du Québec – Nature et Technologies (FRQNT). This work was funded by the Institut Trotter de Recherche sur les Exoplanètes (iREx). This work made use of `tpfplotter` by J. Lillo-Box (publicly available in [www.github.com/jlillo/tpfplotter](http://www.github.com/jlillo/tpfplotter)), which also made use of the python packages `astropy`, `lightkurve`, `matplotlib` and `numpy`.

## References

- Abe, L., Gonçalves, I., Agabi, A., et al. 2013, *A&A*, 553, A49  
 Adibekyan, V. Z., Sousa, S. G., Santos, N. C., et al. 2012, *A&A*, 545, A32  
 Adibekyan, V., Figueira, P., Santos, N. C., et al. 2015, *A&A*, 583, A94  
 Aguichine, A., Mousis, O., Deleuil, M., & Marçq, E. 2021, *ApJ*, 914, 84  
 Akeson, R. L., Chen, X., Ciardi, D., et al. 2013, *PASP*, 125, 989  
 Albrecht, S. H., Dawson, R. I., & Winn, J. N. 2022, *PASP*, 134, 082001  
 Alexander, R., Pascucci, I., Andrews, S., Armitage, P., & Cieza, L. 2014, in *Protostars and Planets VI*, eds. H. Beuther, R. S. Klessen, C. P. Dullemond, & T. Henning, 475  
 Aller, A., Lillo-Box, J., Jones, D., Miranda, L. F., & Barceló Forteza, S. 2020, *A&A*, 635, A128  
 Almeida-Fernandes, F., & Rocha-Pinto, H. J. 2018, *MNRAS*, 476, 184



- Anthony-Twarog, B. J., Deliyannis, C. P., Harmer, D., et al. 2018, *AJ*, 156, 37
- Attia, O., Bourrier, V., Eggenberger, P., et al. 2021, *A&A*, 647, A40
- Attia, O., Bourrier, V., Delisle, J. B., & Eggenberger, P. 2023, *A&A*, 674, A120
- Barker, A. J., & Ogilvie, G. I. 2010, *MNRAS*, 404, 1849
- Barnes, S. A. 2007, *ApJ*, 669, 1167
- Behrard, A., Dai, F., & Howard, A. W. 2022, *AJ*, 163, 160
- Benz, W., Broeg, C., Fortier, A., et al. 2021, *Exp. Astron.*, 51, 109
- Berger, T. A., Huber, D., Gaidos, E., van Saders, J. L., & Weiss, L. M. 2020, *AJ*, 160, 108
- Blackwell, D. E., & Shallis, M. J. 1977, *MNRAS*, 180, 177
- Bodenheimer, P., Stevenson, D. J., Lissauer, J. J., & D'Angelo, G. 2018, *ApJ*, 868, 138
- Bonfanti, A., Ortolani, S., Piotto, G., & Nascimbeni, V. 2015, *A&A*, 575, A18
- Bonfanti, A., Ortolani, S., & Nascimbeni, V. 2016, *A&A*, 585, A5
- Bonfanti, A., Delrez, L., Hooton, M. J., et al. 2021a, *A&A*, 646, A157
- Bonfanti, A., Fossati, L., Kubyskhina, D., & Cubillos, P. E. 2021b, *A&A*, 656, A157
- Bonomo, A. S., Zeng, L., Damasso, M., et al. 2019, *Nat. Astron.*, 3, 416
- Bonomo, A. S., Dumusque, X., Massa, A., et al. 2023, *A&A*, 677, A33
- Borucki, W. J., & Summers, A. L. 1984, *Icarus*, 58, 121
- Bourrier, V., Lovis, C., Cretignier, M., et al. 2021, *A&A*, 654, A152
- Brandeker, A., Heng, K., Lendl, M., et al. 2022, *A&A*, 659, A4
- Brown, T. M., Baliber, N., Bianco, F. B., et al. 2013, *PASP*, 125, 1031
- Buchner, J., Georgakakis, A., Nandra, K., et al. 2014, *A&A*, 564, A125
- Casagrande, L., Schönrich, R., Asplund, M., et al. 2011, *A&A*, 530, A138
- Castelli, F., & Kurucz, R. L. 2003, in *IAU Symposium*, 210, A20
- Chakraborty, H., Lendl, M., Akhavan, B., Petit dit de la Roche, D. J. M., & Deline, A. 2023, *A&A*, accepted, [arXiv:2311.16864]
- Choi, J., Dotter, A., Conroy, C., et al. 2016, *ApJ*, 823, 102
- Cochran, W. D., Fabrycky, D. C., Torres, G., et al. 2011, *ApJS*, 197, 7
- Collins, K. 2019, *AAS Meet. Abstr.*, 233, 140.05
- Collins, K. A., Kielkopf, J. F., Stassun, K. G., & Hessman, F. V. 2017, *AJ*, 153, 77
- Cubillos, P., Erkaev, N. V., Juvan, I., et al. 2017, *MNRAS*, 466, 1868
- Cummings, J. D., Deliyannis, C. P., Maderak, R. M., & Steinhauer, A. 2017, *AJ*, 153, 128
- Delgado Mena, E., Israelian, G., González Hernández, J. I., et al. 2014, *A&A*, 562, A92
- Deline, A., Hooton, M. J., Lendl, M., et al. 2022, *A&A*, 659, A74
- Delrez, L., Ehrenreich, D., Alibert, Y., et al. 2021, *Nat. Astron.*, 5, 775
- Dorn, C. & Lichtenberg, T. 2021, *ApJ*, 922, L4
- Dorn, C., Khan, A., Heng, K., et al. 2015, *A&A*, 577, A83
- Dorn, C., Venturini, J., Khan, A., et al. 2017, *A&A*, 597, A37
- Doyle, L. R., Wilcox, T. J., & Loree, J. J. 1984, *ApJ*, 287, 307
- Dragomir, D., Teske, J., Günther, M. N., et al. 2019, *ApJ*, 875, L7
- Dransfield, G., Mékarnia, D., Triaud, A. H. M. J., et al. 2022, *SPIE Conf. Ser.*, 12186, 121861F
- Espinoza, N., Kossakowski, D., & Brahm, R. 2019, *MNRAS*, 490, 2262
- Fausnaugh, M. M., Burke, C. J., Ricker, G. R., & Vanderspek, R. 2020, *RNAAS*, 4, 251
- Feroz, F., Hobson, M. P., & Bridges, M. 2009, *MNRAS*, 398, 1601
- Foreman-Mackey, D., Hogg, D. W., Lang, D., & Goodman, J. 2013, *PASP*, 125, 306
- Foreman-Mackey, D., Agol, E., Ambikasaran, S., & Angus, R. 2017, *AJ*, 154, 220
- Fossati, L., Erkaev, N. V., Lammer, H., et al. 2017, *A&A*, 598, A90
- Fulton, B. J., Petigura, E. A., Howard, A. W., et al. 2017, *AJ*, 154, 109
- Fulton, B. J., Petigura, E. A., Blunt, S., & Sinukoff, E. 2018, *PASP*, 130, 044504
- Gaia Collaboration (Brown, A. G. A., et al.) 2021, *A&A*, 649, A1
- Gaia Collaboration (Vallenari, A., et al.) 2023, *A&A*, 674, A1
- Garai, Z., Osborn, H. P., Gandolfi, D., et al. 2023, *A&A*, 674, A44
- Gardner, J. P., Mather, J. C., Abbott, R., et al. 2023, *PASP*, 135, 068001
- Ginzburg, S., Schlichting, H. E., & Sari, R. 2018, *MNRAS*, 476, 759
- Gorti, U., Liseau, R., Sándor, Z., & Clarke, C. 2016, *Space Sci. Rev.*, 205, 125
- Guerrero, N. M., Seager, S., Huang, C. X., et al. 2021, *ApJS*, 254, 39
- Guillot, T., Abe, L., Agabi, A., et al. 2015, *Astron. Nachr.*, 336, 638
- Gupta, A., & Schlichting, H. E. 2019, *MNRAS*, 487, 24
- Hadden, S., & Lithwick, Y. 2017, *AJ*, 154, 5
- Hakim, K., Rivoldini, A., Van Hoolst, T., et al. 2018, in *EGU General Assembly Conference Abstracts*, 15447
- Helled, R., & Stevenson, D. 2017, *ApJ*, 840, L4
- Henry, T. J., Soderblom, D. R., Donahue, R. A., & Baliunas, S. L. 1996, *AJ*, 111, 439
- Høg, E., Fabricius, C., Makarov, V. V., et al. 2000, *A&A*, 355, L27
- Hoyer, S., Guterman, P., Demangeon, O., et al. 2020, *A&A*, 635, A24
- Huang, C. X., Vanderburg, A., Pál, A., et al. 2020, *RNAAS*, 4, 206
- Ikoma, M., & Hori, Y. 2012, *ApJ*, 753, 66
- Inamdar, N. K., & Schlichting, H. E. 2015, *MNRAS*, 448, 1751
- Jenkins, J. M. 2002, *ApJ*, 575, 493
- Jenkins, J. M., Chandrasekaran, H., McCauliff, S. D., et al. 2010, *SPIE Conf. Ser.*, 7740, 77400D
- Jenkins, J. M., Twicken, J. D., McCauliff, S., et al. 2016, *Proc. SPIE*, 9913, 99133E
- Jenkins, J. M., Tenenbaum, P., Seader, S., et al. 2020, *Kepler Data Processing Handbook: Transiting Planet Search*, Kepler Science Document KSCI-19081-003
- Jensen, E. 2013, *Astrophysics Source Code Library* [record ascl:1306.007]
- Jin, S., & Mordasini, C. 2018, *ApJ*, 853, 163
- Johnstone, C. P., Güdel, M., Brott, L., & Lüftinger, T. 2015a, *A&A*, 577, A28
- Johnstone, C. P., Güdel, M., Stökl, A., et al. 2015b, *ApJ*, 815, L12
- Kempton, E. M. R., Bean, J. L., Louie, D. R., et al. 2018, *PASP*, 130, 114401
- Kimura, S. S., Kunitomo, M., & Takahashi, S. Z. 2016, *MNRAS*, 461, 2257
- Kreidberg, L. 2015, *PASP*, 127, 1161
- Krenn, A. F., Lendl, M., Patel, J. A., et al. 2023, *A&A*, 672, A24
- Kubyskhina, D., Fossati, L., Erkaev, N. V., et al. 2018, *A&A*, 619, A151
- Kubyskhina, D., Cubillos, P. E., Fossati, L., et al. 2019a, *ApJ*, 879, 26
- Kubyskhina, D., Fossati, L., Mustill, A. J., et al. 2019b, *A&A*, 632, A65
- Kurucz, R. 1993, *SYNTHESIS Spectrum Synthesis Programs and Line Data*, Kurucz CD-ROM (Cambridge), 18
- Lee, E. J., & Chiang, E. 2015, *ApJ*, 811, 41
- Lee, E. J., Karalis, A., & Thorngren, D. P. 2022, *ApJ*, 941, 186
- Leleu, A., Alibert, Y., Hara, N. C., et al. 2021, *A&A*, 649, A26
- Lendl, M., Bouchy, F., Gill, S., et al. 2020, *MNRAS*, 492, 1761
- Li, J., Tenenbaum, P., Twicken, J. D., et al. 2019, *PASP*, 131, 024506
- Lindgren, L., Bastian, U., Biermann, M., et al. 2021, *A&A*, 649, A4
- Lopez, E. D., & Fortney, J. J. 2013, *ApJ*, 776, 2
- Lozovsky, M., Helled, R., Dorn, C., & Venturini, J. 2018, *ApJ*, 866, 49
- Lucy, L. B., & Sweeney, M. A. 1971, *AJ*, 76, 544
- Luque, R., Nowak, G., Pallé, E., et al. 2019, *A&A*, 623, A114
- Maciel, W. J., Rodrigues, T. S., & Costa, R. D. D. 2011, *Rev. Mex. A&A*, 47, 401
- Marcy, G. W., Isaacson, H., Howard, A. W., et al. 2014, *ApJS*, 210, 20
- Marigo, P., Girardi, L., Bressan, A., et al. 2017, *ApJ*, 835, 77
- Masuda, K., & Winn, J. N. 2020, *AJ*, 159, 81
- Maxted, P. F. L., Ehrenreich, D., Wilson, T. G., et al. 2022, *MNRAS*, 514, 77
- McCully, C., Volgenau, N. H., Harbeck, D.-R., et al. 2018, *SPIE Conf. Ser.*, 10707, 107070K
- McLaughlin, D. B. 1924, *ApJ*, 60, 22
- Mékarnia, D., Guillot, T., Rivet, J. P., et al. 2016, *MNRAS*, 463, 45
- Millholland, S., Wang, S., & Laughlin, G. 2017, *ApJ*, 849, L33
- Mishra, L., Alibert, Y., Udry, S., & Mordasini, C. 2023, *A&A*, 670, A68
- Mol Lous, M., Helled, R., & Mordasini, C. 2022, *Nat. Astron.*, 6, 819
- Mordasini, C. 2020, *A&A*, 638, A52
- Morris, R. L., Twicken, J. D., Smith, J. C., et al. 2020, *Kepler Data Processing Handbook: Photometric Analysis*, Kepler Science Document KSCI-19081-003
- Naponiello, L., Mancini, L., Sozzetti, A., et al. 2023, *Nature*, 622, 255
- Nordström, B., Mayor, M., Andersen, J., et al. 2004, *A&A*, 418, 989
- Osborn, H. P. 2022, *Astrophysics Source Code Library* [record ascl:2204.020]
- Osborn, H. P., Nowak, G., Hébrard, G., et al. 2023, *MNRAS*, 523, 3069
- Otegi, J. F., Bouchy, F., & Helled, R. 2020, *A&A*, 634, A43
- Otegi, J. F., Helled, R., & Bouchy, F. 2022, *A&A*, 658, A107
- Owen, J. E., & Wu, Y. 2017, *ApJ*, 847, 29
- Pepe, F., Cristiani, S., Rebolo, R., et al. 2021, *A&A*, 645, A96
- Petigura, E. A., Rogers, J. G., Isaacson, H., et al. 2022, *AJ*, 163, 179
- Pollacco, D. L., Skillen, I., Collier Cameron, A., et al. 2006, *PASP*, 118, 1407
- Ricker, G. R., Winn, J. N., Vanderspek, R., et al. 2015, *J. Astron. Telescopes Instrum. Syst.*, 1, 014003
- Rogers, J. G., & Owen, J. E. 2021, *MNRAS*, 503, 1526
- Rossiter, R. A. 1924, *ApJ*, 60, 15
- Santos, N. C., Sousa, S. G., Mortier, A., et al. 2013, *A&A*, 556, A150
- Schanche, N., Hébrard, G., Collier Cameron, A., et al. 2020, *MNRAS*, 499, 428
- Schmider, F.-X., Abe, L., Agabi, A., et al. 2022, *SPIE Conf. Ser.*, 12182, 121822O
- Sestito, P., & Randich, S. 2005, *A&A*, 442, 615
- Skrutskie, M. F., Cutri, R. M., Stiening, R., et al. 2006, *AJ*, 131, 1163
- Smith, J. C., Stumpe, M. C., Van Cleve, J. E., et al. 2012, *PASP*, 124, 1000
- Snedden, C. A. 1973, PhD thesis, The University of Texas at Austin, USA
- Sousa, S. G. 2014, in *Determination of Atmospheric Parameters of B*, 297
- Sousa, S. G., Santos, N. C., Israelian, G., Mayor, M., & Monteiro, M. J. P. F. G. 2007, *A&A*, 469, 783
- Sousa, S. G., Santos, N. C., Mayor, M., et al. 2008, *A&A*, 487, 373

- Sousa, S. G., Santos, N. C., Adibekyan, V., Delgado-Mena, E., & Israelian, G. 2015, *A&A*, **577**, A67
- Sousa, S. G., Adibekyan, V., Delgado-Mena, E., et al. 2021, *A&A*, **656**, A53
- Speagle, J. S. 2020, *MNRAS*, **493**, 3132
- Stassun, K. G., Oelkers, R. J., Paegert, M., et al. 2019, *AJ*, **158**, 138
- Stumpe, M. C., Smith, J. C., Van Cleve, J. E., et al. 2012, *PASP*, **124**, 985
- Stumpe, M. C., Smith, J. C., Catanzarite, J. H., et al. 2014, *PASP*, **126**, 100
- Sun, L., Ioannidis, P., Gu, S., et al. 2019, *A&A*, **624**, A15
- Sun, Q., Deliyannis, C. P., Steinhauer, A., Anthony-Twarog, B. J., & Twarog, B. A. 2023, *ApJ*, **952**, 71
- Tokovinin, A. 2018, *PASP*, **130**, 035002
- Tsantaki, M., Sousa, S. G., Adibekyan, V. Z., et al. 2013, *A&A*, **555**, A150
- Tu, L., Johnstone, C. P., Güdel, M., & Lammer, H. 2015, *A&A*, **577**, A3
- Twicken, J. D., Clarke, B. D., Bryson, S. T., et al. 2010, *Proc. SPIE*, **7740**, 774023
- Twicken, J. D., Catanzarite, J. H., Clarke, B. D., et al. 2018, *PASP*, **130**, 064502
- Ulmer-Moll, S., Osborn, H. P., Tuson, A., et al. 2023, *A&A*, **674**, A43
- Vanderburg, A., & Johnson, J. A. 2014, *PASP*, **126**, 948
- Venturini, J., & Helled, R. 2017, *ApJ*, **848**, 95
- Venturini, J., Guilera, O. M., Haldemann, J., Ronco, M. P., & Mordasini, C. 2020, *A&A*, **643**, L1
- Weiss, L. M., Marcy, G. W., Petigura, E. A., et al. 2018, *AJ*, **155**, 48
- Wielen, R. 1977, *A&A*, **60**, 263
- Winn, J. N. 2010, in *Exoplanets*, ed. S. Seager, 55
- Winn, J. N., Holman, M. J., Henry, G. W., et al. 2007, *AJ*, **133**, 1828
- Wright, E. L., Eisenhardt, P. R. M., Mainzer, A. K., et al. 2010, *AJ*, **140**, 1868
- Zechmeister, M., & Kürster, M. 2009, *A&A*, **496**, 577
- Zeng, L., Sasselov, D. D., & Jacobsen, S. B. 2016, *ApJ*, **819**, 127
- Ziegler, C., Tokovinin, A., Briceno, C., et al. 2020, *VizieR Online Data Catalog: J/AJ/159/19*
- <sup>1</sup> Observatoire de Genève, Université de Genève, Chemin Pegasi 51, 1290 Versoix, Switzerland  
e-mail: angeliki.psaridi@unige.ch
- <sup>2</sup> Center for Space and Habitability, University of Bern, Gesellschaftsstrasse 6, 3012 Bern, Switzerland
- <sup>3</sup> Department of Physics and Kavli Institute for Astrophysics and Space Research, Massachusetts Institute of Technology, Cambridge, MA 02139, USA
- <sup>4</sup> Weltraumforschung und Planetologie, Physikalisches Institut, University of Bern, Gesellschaftsstrasse 6, 3012 Bern, Switzerland
- <sup>5</sup> Instituto de Astrofísica e Ciências do Espaço, Universidade do Porto, CAUP, Rua das Estrelas, 4150-762 Porto, Portugal
- <sup>6</sup> Departamento de Física e Astronomia, Faculdade de Ciências, Universidade do Porto, Rua do Campo Alegre, 4169-007 Porto, Portugal
- <sup>7</sup> Space Research Institute, Austrian Academy of Sciences, Schmiedlstrasse 6, 8042 Graz, Austria
- <sup>8</sup> Center for Astrophysics | Harvard & Smithsonian, 60 Garden Street, Cambridge, MA 02138, USA
- <sup>9</sup> Université Côte d'Azur, Observatoire de la Côte d'Azur, CNRS, Laboratoire Lagrange, Bd de l'Observatoire, CS 34229, 06304 Nice cedex 4, France
- <sup>10</sup> Institute for Computational Science, University of Zurich, Winterthurerstr. 90, 8057 Zurich, Switzerland
- <sup>11</sup> Astrophysics Group, Keele University, Staffordshire ST5 5BG, UK
- <sup>12</sup> Department of Astronomy & Astrophysics, University of Chicago, Chicago, IL 60637, USA
- <sup>13</sup> NASA Ames Research Center, Moffett Field, CA 94035, USA
- <sup>14</sup> Research Institute for Advanced Computer Science, Universities Space Research Association, Washington, DC 20024, USA
- <sup>15</sup> Department of Physics and Astronomy, Union College, 807 Union St., Schenectady, NY 12308, USA
- <sup>16</sup> Department of Physics, University of Warwick, Gibbet Hill Road, Coventry CV4 7AL, UK
- <sup>17</sup> Department of Astrophysical Sciences, Princeton University, Princeton, NJ 08544, USA
- <sup>18</sup> Department of Physics, Engineering and Astronomy, Stephen F. Austin State University, 1936 North St, Nacogdoches, TX 75962, USA
- <sup>19</sup> Dipartimento di Fisica e Astronomia “Galileo Galilei”, Università degli Studi di Padova, Vicolo dell'Osservatorio 3, 35122 Padova, Italy
- <sup>20</sup> INAF Osservatorio Astronomico di Padova, Vicolo dell'Osservatorio 5, 35122 Padova, Italy
- <sup>21</sup> Department of Astronomy, Stockholm University, AlbaNova University Center, 10691 Stockholm, Sweden
- <sup>22</sup> Dipartimento di Fisica, Università degli Studi di Torino, via Pietro Giuria 1, 10125, Torino, Italy
- <sup>23</sup> Cavendish Laboratory, JJ Thomson Avenue, Cambridge CB3 0HE, UK
- <sup>24</sup> Institute of Astronomy, University of Cambridge, Madingley Road, Cambridge, CB3 0HA, UK
- <sup>25</sup> Département de Physique, Institut Trottier de Recherche sur les Exoplanètes, Université de Montréal, Montréal, Québec H3T 1J4, Canada
- <sup>26</sup> Instituto de Astrofísica de Canarias, Via Lactea s/n, 38200 La Laguna, Tenerife, Spain
- <sup>27</sup> Departamento de Astrofísica, Universidad de La Laguna, Astrofísico Francisco Sanchez s/n, 38206 La Laguna, Tenerife, Spain
- <sup>28</sup> Institut de Ciències de l'Espai (ICE, CSIC), Campus UAB, Can Magrans s/n, 08193 Bellaterra, Spain
- <sup>29</sup> Institut d'Estudis Espacials de Catalunya (IEEC), Gran Capità 2–4, 08034 Barcelona, Spain
- <sup>30</sup> Admatis, 5. Kandó Kálmán Street, 3534 Miskolc, Hungary
- <sup>31</sup> Depto. de Astrofísica, Centro de Astrobiología (CSIC-INTA), ESAC campus, 28692 Villanueva de la Cañada (Madrid), Spain
- <sup>32</sup> Université Grenoble Alpes, CNRS, IPAG, 38000 Grenoble, France
- <sup>33</sup> Caltech/IPAC-NASA Exoplanet Science Institute, 770 S. Wilson Avenue, Pasadena, CA 91106, USA
- <sup>34</sup> Centre for Exoplanet Science, SUPA School of Physics and Astronomy, University of St Andrews, North Haugh, St Andrews KY16 9SS, UK
- <sup>35</sup> Université de Paris Cité, Institut de physique du globe de Paris, CNRS, 1 rue Jussieu, 75005 Paris, France
- <sup>36</sup> Institute of Planetary Research, German Aerospace Center (DLR), Rutherfordstrasse 2, 12489 Berlin, Germany
- <sup>37</sup> European Southern Observatory, Av. Alonso de Cordova 3107, Casilla 19001, Santiago de Chile, Chile
- <sup>38</sup> Centre for Mathematical Sciences, Lund University, Box 118, 221 00 Lund, Sweden
- <sup>39</sup> Department of Physics and McDonnell Center for the Space Sciences, Washington University, St. Louis, MO 63130, USA
- <sup>40</sup> Aix-Marseille Univ, CNRS, CNES, LAM, 38 rue Frédéric Joliot-Curie, 13388 Marseille, France
- <sup>41</sup> Astrobiology Research Unit, Université de Liège, Allée du 6 Août 19C, 4000 Liège, Belgium
- <sup>42</sup> Space sciences, Technologies and Astrophysics Research (STAR) Institute, Université de Liège, Allée du 6 Août 19C, 4000 Liège, Belgium
- <sup>43</sup> ETH Zurich, Institute for Particle Physics and Astrophysics, Wolfgang-Pauli-Strasse 27, 8093 Zurich Switzerland
- <sup>44</sup> Institut d'astrophysique de Paris, UMR7095 CNRS, Université Pierre & Marie Curie, 98bis bd. Arago, 75014 Paris, France
- <sup>45</sup> Airbus Defence and Space SAU, C/ Aviocar 2, 28906 Getafe, Madrid, Spain
- <sup>46</sup> Leiden Observatory, University of Leiden, PO Box 9513, 2300 RA Leiden, The Netherlands
- <sup>47</sup> Department of Space, Earth and Environment, Chalmers University of Technology, Onsala Space Observatory, 439 92 Onsala, Sweden
- <sup>48</sup> Department of Astrophysics, University of Vienna, Türkenschanzstrasse 17, 1180 Vienna, Austria
- <sup>49</sup> European Space Agency (ESA), European Space Research and Technology Centre (ESTEC), Keplerlaan 1, 2201 AZ Noordwijk, The Netherlands
- <sup>50</sup> Institute for Theoretical Physics and Computational Physics, Graz University of Technology, Petersgasse 16, 8010 Graz, Austria
- <sup>51</sup> Konkoly Observatory, Research Centre for Astronomy and Earth Sciences, 1121 Budapest, Konkoly Thege Miklós út 15–17, Hungary

- <sup>52</sup> ELTE Eötvös Loránd University, Institute of Physics, Pázmány Péter sétány 1/A, 1117 Budapest, Hungary
- <sup>53</sup> IMCCE, UMR8028 CNRS, Observatoire de Paris, PSL Univ., Sorbonne Univ., 77 av. Denfert-Rochereau, 75014 Paris, France
- <sup>54</sup> Proto-Logic LLC, 1718 Euclid Street NW, Washington, DC 20009, USA
- <sup>55</sup> European Southern Observatory, Karl-Schwarzschild-Straße 2, 85748 Garching bei München, Germany
- <sup>56</sup> INAF Osservatorio Astrofisico di Catania, via S. Sofia 78, 95123 Catania, Italy
- <sup>57</sup> Zentrum für Astronomie und Astrophysik, Technische Universität Berlin, Hardenbergstr. 36, 10623 Berlin, Germany
- <sup>58</sup> Institut fuer Geologische Wissenschaften, Freie Universitaet Berlin, Maltheserstrasse 74–100, 12249 Berlin, Germany
- <sup>59</sup> Department of Earth, Atmospheric and Planetary Sciences, Massachusetts Institute of Technology, Cambridge, MA 02139, USA
- <sup>60</sup> Department of Aeronautics and Astronautics, MIT, 77 Massachusetts Avenue, Cambridge, MA 02139, USA
- <sup>61</sup> ELTE Eötvös Loránd University, Gothard Astrophysical Observatory, 9700 Szombathely, Szent Imre h. u. 112, Hungary
- <sup>62</sup> HUN-REN–ELTE Exoplanet Research Group, Szent Imre h. u. 112, Szombathely, 9700, Hungary
- <sup>63</sup> NASA Goddard Space Flight Center, 8800 Greenbelt Road, Greenbelt, MD 20771, USA
- <sup>64</sup> Centre for Exoplanets and Habitability, University of Warwick, Gibbet Hill Road, Coventry CV4 7AL, UK

Appendix A: TESS light curves

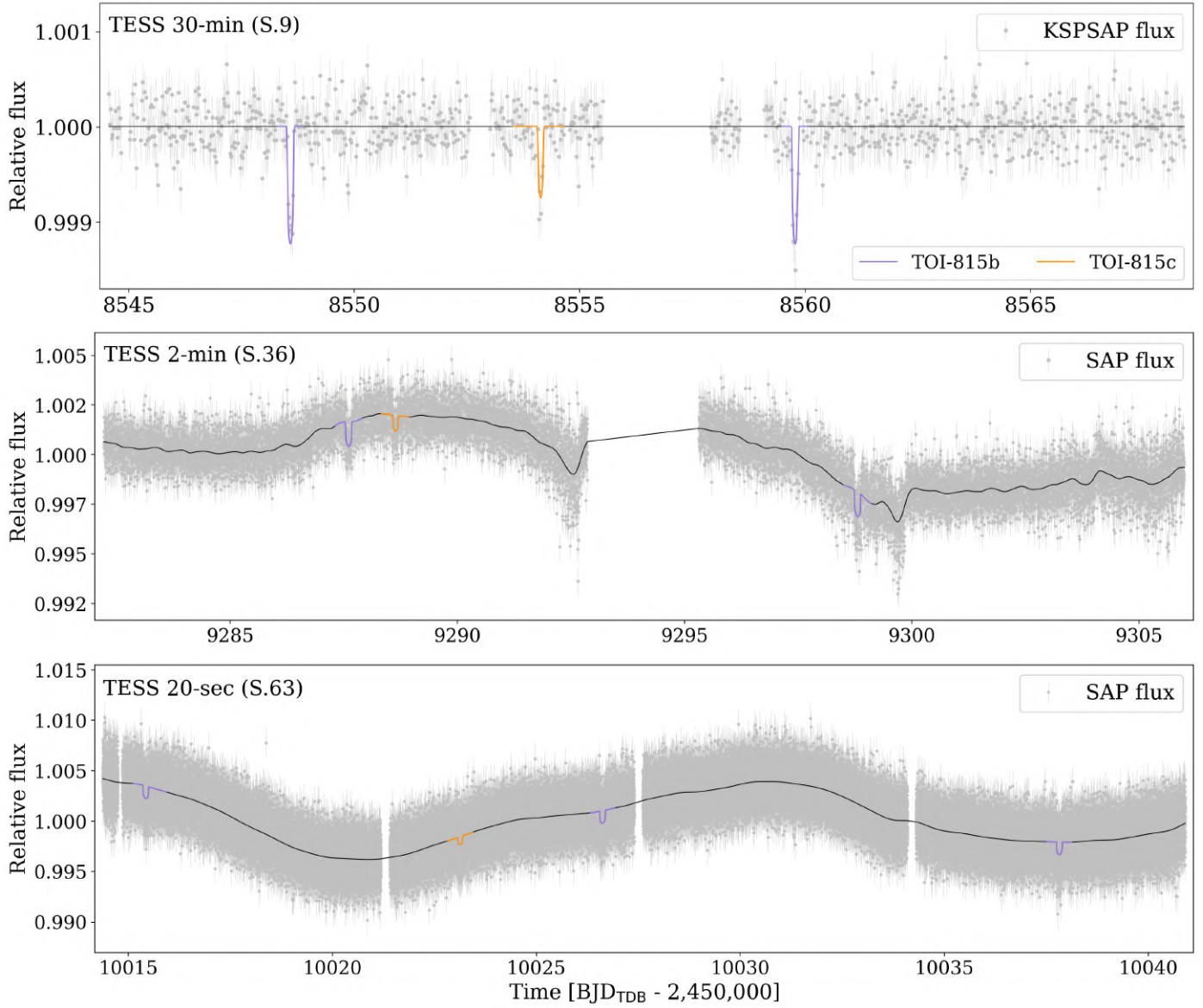


Fig. A.1: TESS SAP flux light curves with the best-fit juliet model shown with the black line (see Sect. 4 for details on the modeling). The transits of TOI-815b and TOI-815c are indicated with purple and orange, respectively.

## Appendix B: RVs with ESPRESSO

Table B.1: ESPRESSO RV observations for TOI-815, along with activity indicator measurements, i.e. FWHM, bisector, and  $\log R'_{\text{HK}}$ .

Time [BJD-2 450 000]	RV [m s <sup>-1</sup> ]	$\sigma_{\text{RV}}$ [m s <sup>-1</sup> ]	FWHM <sup>(2)</sup> [m s <sup>-1</sup> ]	BIS <sup>(2)</sup>	$\log R'_{\text{HK}}$ <sup>(3)</sup>	$t_{\text{exp}}$ [s]
9308.60433	-2973.3950	0.6145	6474.143	-33.038	-4.7084	600
9310.60679	-2970.1704	0.7233	6470.941	-33.052	-4.7338	600
9313.73506 <sup>(1)</sup>	-2962.6373	0.5712	6492.966	-30.291	-4.6757	600
9316.52823	-2967.0427	0.7459	6494.718	-23.625	-4.7107	600
9318.53033	-2968.9726	0.5863	6478.690	-29.060	-4.6762	600
9324.54484	-2979.7547	0.8117	6513.799	-28.181	-4.6859	600
9327.50777 <sup>(1)</sup>	-2965.7154	0.4397	6482.962	-32.938	-4.6229	600
9329.52592 <sup>(1)</sup>	-2971.0105	0.9842	6543.347	-29.283	-4.7529	600
9332.57219	-2979.7266	0.6474	6491.382	-26.738	-4.6824	600
9336.49956 <sup>(1)</sup>	-2982.4246	1.3115	6515.936	-40.000	-4.9931	600
9339.48187	-2973.6498	0.5988	6468.715	-33.267	-4.7189	600
9341.51194 <sup>(1)</sup>	-2966.5835	0.9108	6513.090	-35.131	-4.7516	600
9346.54064	-2972.2321	0.5959	6482.420	-24.888	-4.6930	600
9373.49668	-2973.5189	0.5467	6462.905	-33.991	-4.6367	600
9375.51475	-2971.6794	0.6513	6471.909	-32.600	-4.6265	600
9406.50596	-2975.2108	0.7212	6492.234	-32.504	-4.6149	600
9410.48054	-2971.8872	0.6759	6499.208	-30.202	-4.6978	600
9415.50797	-2971.5355	0.8107	6491.252	-26.700	-4.6620	600
9673.54348	-2977.4158	0.5684	6488.390	-34.257	-4.6894	700
9675.51515	-2973.7947	0.9406	6487.917	-25.694	-4.6418	700
9677.65231	-2972.6960	0.8151	6492.405	-31.643	-4.6580	700
9679.55323 <sup>(1)</sup>	-2973.4334	0.5445	6497.504	-32.766	-4.6080	700
9683.53185	-2981.6090	0.4694	6497.328	-29.483	-4.6329	700
9685.55122	-2980.4065	0.7458	6496.394	-24.517	-4.6203	700
9687.52944	-2972.0070	0.4487	6495.493	-29.672	-4.6432	700
9695.52289	-2967.3200	0.9486	6512.740	-29.566	-4.7299	700
9697.58163	-2965.9727	0.5759	6511.966	-26.858	-4.6611	700
9784.49262	-2977.5647	0.6128	6525.410	-22.486	-4.6248	700
9904.83330	-2969.5919	0.4342	6487.460	-35.249	-4.6440	700
9905.77809	-2970.1154	0.6533	6491.720	-34.630	-4.6797	700
9906.80910	-2970.1976	0.4551	6492.162	-33.028	-4.6399	700
9907.79772	-2969.7455	1.3028	6534.337	-38.113	-4.8075	700
9910.75047	-2966.0686	0.7246	6503.361	-25.248	-4.6902	700
9917.72905	-2978.0604	0.6064	6481.170	-33.996	-4.6673	700

**Notes:**

<sup>(1)</sup> These data did not pass the ESPRESSO flux quality control and therefore were excluded from the analysis.

<sup>(2)</sup> The uncertainties of FWHM and BIS are twice the uncertainties of the RVs ( $2\sigma_{\text{RV}}$ ).

<sup>(3)</sup> The median uncertainty of  $\log R'_{\text{HK}}$ <sup>(3)</sup> is 0.03.

Appendix C: Spectral synthesis

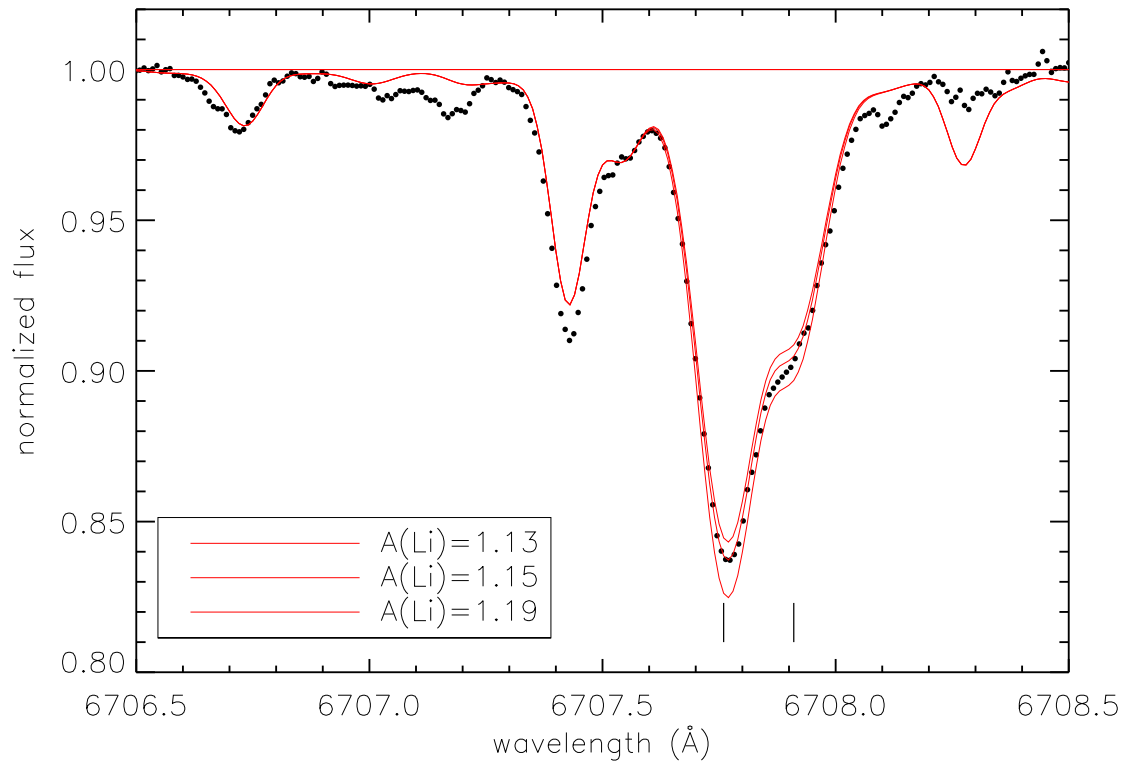


Fig. C.1: Spectral synthesis around the Li doublet. The vertical lines mark the position of the lines.

### Appendix D: Generalized Lomb-Scargle periodograms

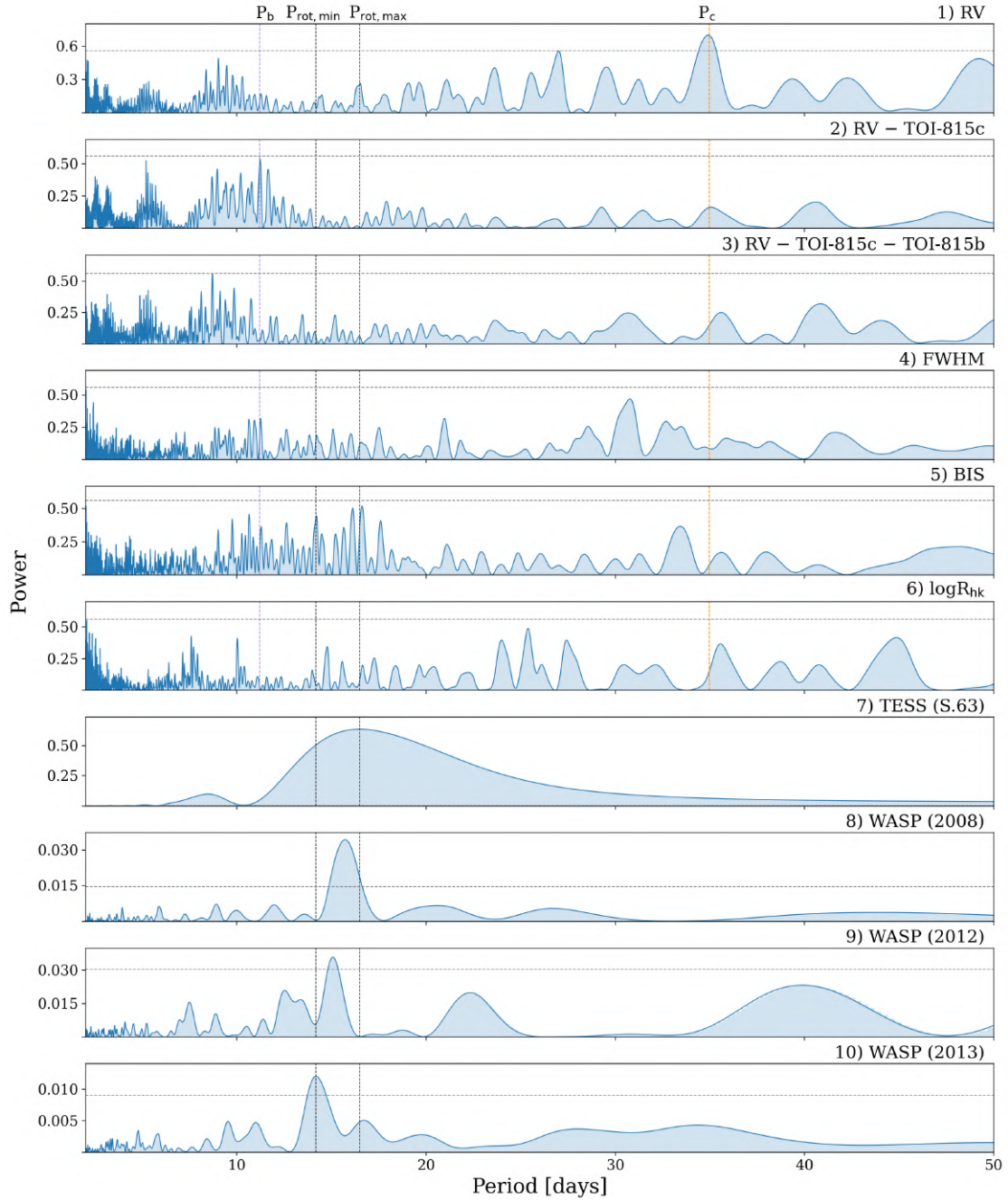


Fig. D.1: GLS periodogram of the TOI-815 spectroscopic and photometric time series. More specifically: (1) RV ESPRESSO measurements, (2) RV residuals, after the subtraction of TOI-815c signal, (2) RV residuals, after the subtraction of TOI-815b and c signals, (3)-(5) FWHM, bisector, and Ca II H & K line activity indicator ( $\log R'_{HK}$ ), (7) TESS Sector 63 photometric data, and (8)-(10) WASP photometric data (2008, 2012, and 2023). The two black vertical lines correspond to the minimum and maximum rotation periods ( $P_{rot,min} \sim 14.2$  days and  $P_{rot,max} \sim 16.5$  days). The purple and orange vertical lines correspond to the orbital periods of planets b ( $P_b \sim 11.2$  days) and c ( $P_c \sim 35$  days), respectively. The gray horizontal lines indicating the 1% FAP.

Appendix E: Priors

Table E.1: Maximum values and 68% confidence intervals of the posterior distributions from joint modeling with juliet.

Parameter	Prior <sup>(a)</sup>	Value <sup>(b)</sup>
<b>Jump stellar parameters</b>		
Stellar density, $\rho_*$ ( $\rho_\odot$ ) . . . . .	$N(2396, 139)$	2431 <sup>+122</sup> <sub>-125</sub>
<b>Jump parameters for TOI-815b</b>		
Orbital period, $P_{orb}$ (days) . . . . .	$\mathcal{U}(10, 12)$	11.197259 <sup>+0.000018</sup> <sub>-0.000017</sub>
Transit epoch, $T_0$ (BJD) . . . . .	$N(2459287.6, 0.1)$	2459287.6028 <sup>+0.0009</sup> <sub>-0.0009</sub>
Scaled planetary radius, $R_p/R_*$ . . . . .	$\mathcal{U}(0, 1)$	0.0349 <sup>+0.0004</sup> <sub>-0.0004</sub>
Impact parameter, $b$ . . . . .	$\mathcal{U}(0, 1)$	0.279 <sup>+0.058</sup> <sub>-0.073</sub>
Eccentricity, $e$ . . . . .	$\mathcal{F}(0)$	0 (adopted, $3\sigma < 0.20$ )
Argument of periastron, $\omega_*$ (deg) . . . . .	$\mathcal{F}(90)$	-
RV semi-amplitude, $K$ ( $\text{m s}^{-1}$ ) . . . . .	$\mathcal{U}(0, 100)$	2.6 <sup>+0.5</sup> <sub>-0.5</sub>
<b>Jump parameters for TOI-815c</b>		
Orbital period, $P_{orb}$ (days) . . . . .	$\mathcal{U}(34, 36)$	34.976145 <sup>+0.000099</sup> <sub>-0.000097</sub>
Transit epoch, $T_0$ (BJD) . . . . .	$N(2459288.6, 0.1)$	2459288.6331 <sup>+0.0023</sup> <sub>-0.0024</sub>
Scaled planetary radius, $R_p/R_*$ . . . . .	$\mathcal{U}(0, 1)$	0.0312 <sup>+0.0011</sup> <sub>-0.0010</sub>
Impact parameter, $b$ . . . . .	$\mathcal{U}(0, 1)$	0.839 <sup>+0.015</sup> <sub>-0.016</sub>
Eccentricity, $e$ . . . . .	$\mathcal{F}(0)$	0 (adopted, $3\sigma < 0.22$ )
Argument of periastron, $\omega_*$ (deg) . . . . .	$\mathcal{F}(90)$	-
RV semi-amplitude, $K$ ( $\text{m s}^{-1}$ ) . . . . .	$\mathcal{U}(0, 100)$	5.5 <sup>+0.5</sup> <sub>-0.5</sub>
<b>Instrumental photometric parameters</b>		
Offset relative flux, $M_{TESS9}$ ( $10^{-6}$ ) . . . . .	$N(0, 0.1)$	-8.6 <sup>+8.0</sup> <sub>-8.2</sub>
Jitter, $\sigma_{w,TESS9}$ (ppm) . . . . .	$\log\mathcal{U}(0.01, 1000)$	0.5 <sup>+6.5</sup> <sub>-0.5</sub>
Offset relative flux, $M_{TESS36}$ ( $10^{-6}$ ) . . . . .	$N(0, 0.1)$	-72 <sup>+411</sup> <sub>-417</sub>
Jitter, $\sigma_{w,TESS36}$ (ppm) . . . . .	$\log\mathcal{U}(0.01, 1000)$	466 <sup>+8</sup> <sub>-8</sub>
Offset relative flux, $M_{TESS63}$ ( $10^{-6}$ ) . . . . .	$N(0, 0.1)$	-3594 <sup>+3821</sup> <sub>-6137</sub>
Jitter, $\sigma_{w,TESS63}$ (ppm) . . . . .	$\log\mathcal{U}(0.01, 1000)$	457 <sup>+11</sup> <sub>-11</sub>
Offset relative flux, $M_{CHEOPS1}$ ( $10^{-6}$ ) . . . . .	$N(0, 0.1)$	-197 <sup>+46</sup> <sub>-44</sub>
Jitter, $\sigma_{w,CHEOPS1}$ (ppm) . . . . .	$\log\mathcal{U}(0.01, 1000)$	299 <sup>+23</sup> <sub>-22</sub>
Offset relative flux, $M_{CHEOPS2}$ ( $10^{-6}$ ) . . . . .	$N(0, 0.1)$	-474 <sup>+48</sup> <sub>-48</sub>
Jitter, $\sigma_{w,CHEOPS2}$ (ppm) . . . . .	$\log\mathcal{U}(0.01, 1000)$	404 <sup>+24</sup> <sub>-23</sub>
Offset relative flux, $M_{CHEOPS3}$ ( $10^{-6}$ ) . . . . .	$N(0, 0.1)$	-327 <sup>+34</sup> <sub>-35</sub>
Jitter, $\sigma_{w,CHEOPS3}$ (ppm) . . . . .	$\log\mathcal{U}(0.01, 1000)$	293 <sup>+19</sup> <sub>-19</sub>
Offset relative flux, $M_{LCO}$ ( $10^{-6}$ ) . . . . .	$N(0, 0.1)$	1089 <sup>+88</sup> <sub>-88</sub>
Jitter, $\sigma_{w,LCO}$ (ppm) . . . . .	$\log\mathcal{U}(0.01, 5000)$	986 <sup>+72</sup> <sub>-68</sub>
Offset relative flux, $M_{ASTEP}$ ( $10^{-6}$ ) . . . . .	$N(0, 0.1)$	761 <sup>+133</sup> <sub>-131</sub>
Jitter, $\sigma_{w,ASTEP}$ (ppm) . . . . .	$\log\mathcal{U}(0.01, 5000)$	3369 <sup>+89</sup> <sub>-85</sub>
<b>GP parameters</b>		
$\sigma_{GP,TESS9}$ (relative flux) . . . . .	$\mathcal{F}(0)$	-
$\rho_{GP,TESS9}$ (days) . . . . .	$\mathcal{F}(0)$	-
$\sigma_{GP,TESS36}$ ( $10^{-4}$ relative flux) . . . . .	$\log\mathcal{U}(10^{-6}, 100)$	12 <sup>+2</sup> <sub>-3</sub>
$\rho_{GP,TESS36}$ (days) . . . . .	$\log\mathcal{U}(0.001, 100)$	0.94 <sup>+0.13</sup> <sub>-0.11</sub>
$\sigma_{GP,TESS63}$ ( $10^{-4}$ relative flux) . . . . .	$\log\mathcal{U}(10^{-6}, 100)$	57 <sup>+44</sup> <sub>-22</sub>
$\rho_{GP,TESS63}$ (days) . . . . .	$\log\mathcal{U}(0.001, 100)$	9.11 <sup>+4.37</sup> <sub>-2.77</sub>
<b>Instrumental RV parameters</b>		
Jitter, $\sigma_{w,ESPRESSO}$ ( $\text{m s}^{-1}$ ) . . . . .	$\log\mathcal{U}(0.001, 100)$	1.88 <sup>+0.29</sup> <sub>-0.25</sub>
Systemic RV, $\mu_{ESPRESSO}$ ( $\text{m s}^{-1}$ ) . . . . .	$\mathcal{U}(-3500, -1000)$	-2973.61 <sup>+0.36</sup> <sub>-0.36</sub>
<b>Limb darkening parameters</b>		
$q_{1,ESPRESSO}$ . . . . .	$N(0.408, 0.017)$	0.407 <sup>+0.015</sup> <sub>-0.015</sub>
$q_{2,ESPRESSO}$ . . . . .	$N(0.349, 0.022)$	0.345 <sup>+0.020</sup> <sub>-0.019</sub>
$q_{1,CHEOPS}$ . . . . .	$N(0.519, 0.012)$	0.518 <sup>+0.011</sup> <sub>-0.011</sub>
$q_{2,CHEOPS}$ . . . . .	$N(0.408, 0.012)$	0.406 <sup>+0.011</sup> <sub>-0.011</sub>
$q_{1,LCO}$ . . . . .	$N(0.324, 0.008)$	0.324 <sup>+0.007</sup> <sub>-0.007</sub>
$q_{2,LCO}$ . . . . .	$N(0.314, 0.014)$	0.314 <sup>+0.012</sup> <sub>-0.012</sub>
$q_{1,ASTEP}$ . . . . .	$N(0.492, 0.016)$	0.492 <sup>+0.014</sup> <sub>-0.014</sub>
$q_{2,ASTEP}$ . . . . .	$N(0.376, 0.017)$	0.378 <sup>+0.015</sup> <sub>-0.016</sub>

**Notes:** <sup>(a)</sup>For the priors,  $N(\mu, \sigma^2)$  indicates a normal distribution with mean  $\mu$  and variance  $\sigma^2$ ,  $\mathcal{U}(a, b)$  a uniform distribution between  $a$  and  $b$ ,  $\log\mathcal{U}(a, b)$  a log-uniform distribution between  $a$  and  $b$  and  $\mathcal{F}(a)$  a parameter fixed to value  $a$ . <sup>(b)</sup>The posterior estimate indicates the median value and then error bars for the 68 % credibility intervals.



## Appendix F: CHEOPS alias hunting

Table F.1: CHEOPS photometric observations of TOI-815.

Visit	Expected planet	Start time [UTC]	Duration [h]	Eff. [%]	RMS ppm	Alias [d]	Detection	File key	Detrending <sup>(d)</sup>
1	c	2022-03-06 00:36	8.34	73	422	22.26	no	CH_PR110048_TG018001_V0300	(e)
2	c	2022-03-06 16:26	8.37	60	461	20.99	no	CH_PR110048_TG018901_V0300	(e)
3	c	2022-03-17 04:35	9.84	61	417	22.95	no	CH_PR110048_TG018101_V0300	(e)
4 <sup>(a)</sup>	b	2022-03-18 08:36	11.64	66	514	11.20	yes	CH_PR100031_TG050801_V0300	t, x, y, y <sup>2</sup> , sin( $\phi$ ), cos( $\phi$ ), cos(2 $\phi$ )
5	c	2022-04-13 01:20	8.47	60	429	17.92	no	CH_PR110048_TG025201_V0300	(e)
6	c	2022-04-30 06:47	8.34	51	497	29.38	no	CH_PR110048_TG027601_V0300	t (clear flux trend in data)
7	c	2023-02-15 11:09	8.29	69	520	31.93	no	CH_PR110048_TG036501_V0300	(e)
8	c	2023-02-22 01:40	9.04	68	447	25.33	no	CH_PR110048_TG036701_V0300	(e)
9	c	2023-02-27 14:48	8.22	64	580	19.85	no	CH_PR110048_TG036901_V0300	(e)
10 <sup>(b)</sup>	b	2023-02-28 14:16	8.39	62	460	18.83	yes	CH_PR110048_TG037001_V0300	t <sup>2</sup> , x, y, y <sup>2</sup> , bg, sin( $\phi$ ), sin(2 $\phi$ ), cos(2 $\phi$ )
11 <sup>(c)</sup>	c	2023-04-23 08:31	10.09	58	462	34.98	yes	CH_PR110048_TG042201_V0300	t, t <sup>2</sup> , y, y <sup>2</sup> , bg, cos( $\phi$ )

**Notes:**

<sup>(a)</sup> Figure 7, left.

<sup>(b)</sup> Figure 7, center.

<sup>(c)</sup> Figure 7, right.

<sup>(d)</sup> The correlation terms were determined using `pycheops`.

<sup>(e)</sup> Detrended with: x, y, bg, cos( $\phi$ ), cos(2 $\phi$ ), cos(3 $\phi$ ),  $\Delta T$ , sin( $\phi$ ), sin(2 $\phi$ ), sin(3 $\phi$ ), bg<sup>2</sup>, x<sup>2</sup>, y<sup>2</sup>

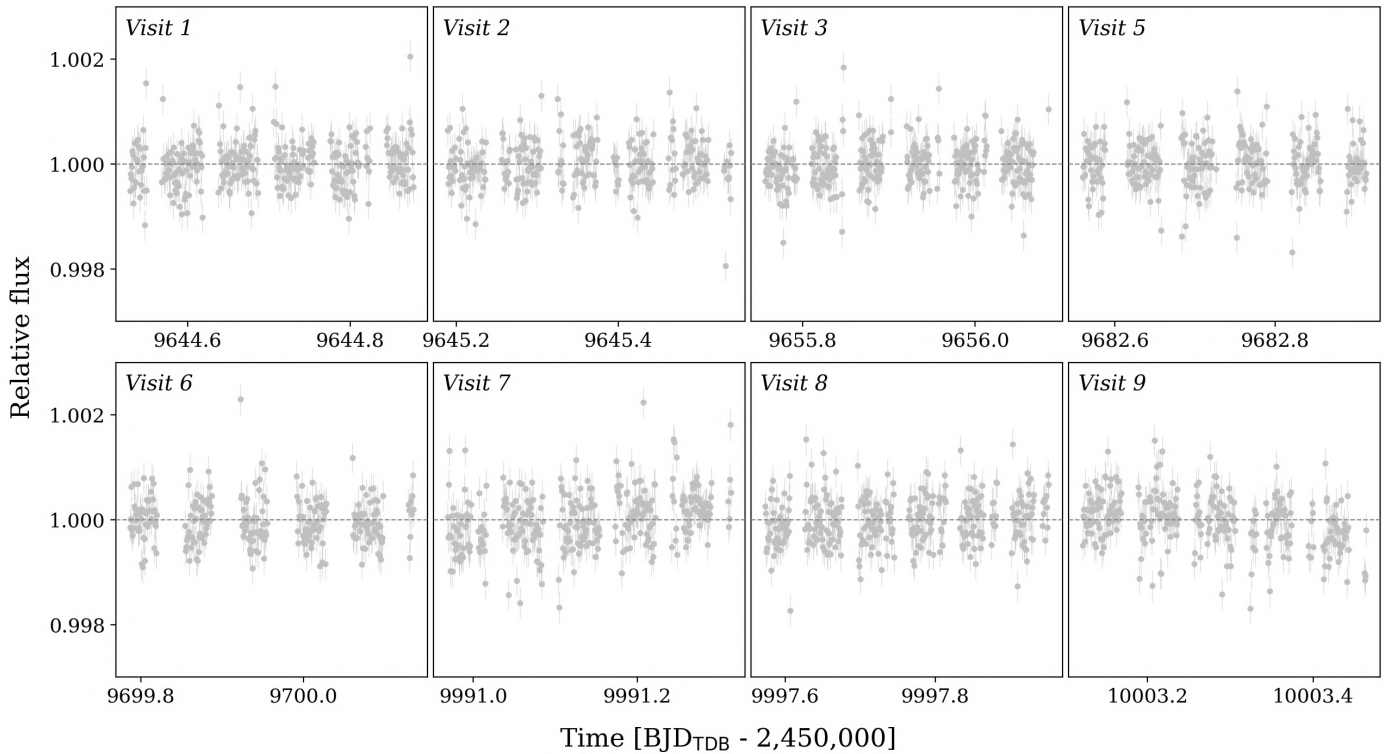


Fig. F.1: CHEOPS unsuccessful transit observations that ruled out the most likely period aliases of TOI-815c. The data have been detrended using the detrending terms presented in Table F.1.