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ALMA Detection of Masers and Dasars in the Hydrogen Recombination Lines of the Planetary Nebula Mz3

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Abstract

The hydrogen recombination lines H30 α , H40 α , H42 α , H50 β , and H57 γ and the underlying bremsstrahlung continuum emission were detected with ALMA in the bipolar nebula Mz3. The source was not spatially resolved, but the velocity profile of the H30 α line shows clear indication of maser amplification, confirming previous reports of laser amplification in the far-infrared H recombination lines observed with Herschel Space Observatory. Comparison between the flux densities of the H50 β , H40 α , and H42 α lines show overcooling, or darkness amplification by stimulated absorption (dasar effect) at the LSR velocity of about -25 km s^{-1} , which constrains the density of the absorbing region to about 10^3 cm^{-3} . The H30 α line, on the other hand, presents maser lines at LSR velocities of -69 and -98 km s^{-1} , which indicates ionized gas with densities close to 10^7 cm^{-3} . Although the source of emission was not resolved, it was possible to find the central position of the images for each velocity interval, which resulted in a well defined position–velocity distribution.

Unified Astronomy Thesaurus concepts: Planetary nebulae (1249); Interstellar masers (846); Astrophysical masers (103)

1. Introduction

Menzel 3 (Mz3; Ant Nebula; PN G331.7–01.0), is seen as a bipolar nebula in optical light, with a high degree of axial symmetry. It exhibits pairs of multiple nested outflows, a bipolar structure of spokes, and a puzzling ring structure surrounding its waist in a tilted position with respect to the main nebular symmetry axis (“chakram”; M. A. Guerrero et al. 2004; M. Santander-García et al. 2004; N. Clyne et al. 2015).

The mid-infrared emission from its core is consistent with the presence of an edge-on circumstellar disk with its axis inclined by $\sim 74^\circ \pm 3^\circ$ in relation to the line of sight, as determined by O. Chesneau et al. (2007). The disk obscures the central stellar system from our view (M. Cohen et al. 1978; O. Chesneau et al. 2007).

The symmetrical bipolar structure may be considered an indirect evidence of a binary stellar system in its core. This object has been classified as a young planetary nebula (PN), a pre-PN, a symbiotic star, and a symbiotic Mira (M. Cohen et al. 1978; J. A. Lopez & J. Meaburn 1983; S. Schmeja & S. Kimeswenger 2001; I. Bains et al. 2004; M. A. Guerrero et al. 2004; S. Akras et al. 2019), but its true nature is still largely unknown. The parallax of the central star of Mz3 was measured by GAIA Space Observatory as $0.3658 \pm 0.0856 \text{ mas}$ (GAIA collaboration 2022) corresponding to

a distance of $2.73 + 0.84 / -0.52 \text{ kpc}$ (C. A. L. J. Bailer-Jones et al. 2021).

The parameters of the ionizing source have been inferred from the outflow emission. The ionization of the lobes can be explained by a source with effective temperature $T_{\text{eff}} = 30\text{--}40 \text{ kK}$ and luminosity of $L_* \sim 10^4 L_\odot$ (N. Smith 2003; S. R. Pottasch & R. Surendranath 2005). This value may be underestimated as the dense material in the core may be obscuring the ionizing source. While the lobes have densities around $\sim 4 \times 10^3 \text{ cm}^{-3}$, the density in the core exceeds 10^6 cm^{-3} (Y. Zhang & X.-W. Liu 2002; N. Smith 2003). J. H. Kastner et al. (2003) discovered diffuse X-ray emission within the lobes (collimated outflows, perhaps a jetlike structure) and a compact X-ray source in the core of Mz3. The central source can be an intrinsically hard X-ray source or deeply embedded in X-ray-absorbing gas. Those authors suggested that this emission configuration is consistent with an accretion disk. The far-infrared-to-submillimeter spectrum ($55\text{--}605 \mu\text{m}$) was obtained with the Herschel Space Observatory as part of the Herschel Planetary Nebula Survey (T. Ueta et al. 2014). In addition to the usual bright ionic forbidden lines typically seen in ionized nebulae in this range, its spectrum exhibits unusually bright hydrogen recombination lines (HRLs). I. Aleman et al. (2018) proposed that stimulated emission amplification (laser effect) produced in the dense ionized gas in the core of Mz3 could explain the high intensity of those lines.

The range of physical conditions where stimulated emission amplification can be significant is limited (e.g., V. S. Strelitskiy et al. 1996; C. Hengel & W. H. Kegel 2000). HRL laser emission



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can be produced in dense and compact ionized regions (densities higher than 10^6 cm^{-3}), making them an interesting tool for inferring physical conditions and kinematics where the forbidden lines typically used as diagnostics are suppressed by the significant collision excitation process due to the densities higher than the critical densities; they are so far a rare detection. Only a few objects show this emission: MWC 349 (B[e] SG star), first radio recombination line (RRL) maser detected by J. Martin-Pintado et al. (1989), MWC 922 (B[e] FS CMa star), η Carinae (luminous blue variable), Cepheus A HW2 (young stellar object), MonR2-IRS2 (young stellar object), G45.47+0.05 (young stellar object), and M 82 (starburst galaxy) (P. Cox et al. 1995; E. R. Seaquist et al. 1996; I. Jiménez-Serra et al. 2011, 2013; Z. Abraham et al. 2014; C. Sánchez Contreras et al. 2017; Y. Zhang et al. 2019; Z. Abraham et al. 2020).

We have carried out ALMA observations of HRLs in Mz3 in order to understand the properties of its central ionized region. We have discovered evidence of maser and, unexpectedly, dasar emission (darkness amplification by stimulated emission) from this region. This is the first time that dasar emission, theoretically predicted almost three decades ago by V. S. Strelitski et al. (1996), has been identified in any astronomical object. The simultaneous detection of maser and dasar emission is important for the accurate determination of the densities in both Galactic and extragalactic sources such as compact H II regions, the winds of Wolf–Rayet and Be stars, active galactic nuclei, and starburst galaxies (V. S. Strelitski et al. 1996). The observations are presented in Section 2, and the results are described in Section 3 and discussed in Section 4. Conclusions are summarized in Section 5.

2. Observations and Data Analysis

Menzel 3¹¹ was observed with the ALMA 12 m array, in bands 3 and 6, during Cycle 8 (project code 2021.1.00912.S).

The band 3 observations were performed on 2021 November 5–7, totaling 74 minutes on both blocks, using four 1 GHz spectral windows (SPWs) with 1920 channels, 488 kHz each. The central frequencies were 86.156, 87.741, 98.557, and 100.034 GHz. The source J1617-5848 was used as a bandpass and flux calibrator, while sources J1604-441 and J1603-4904 were used as complex gain calibrators. In this band, we detected the H40 α , H42 α , H50 β , and H57 γ lines and the underlying continuum.

The band 6 data were taken on 2022 August 27, using four 2 GHz SPWs, three with 128 channels, centered on 217.991, 218.773, and 230.902 GHz to detect the continuum emission and one 1920 MHz SPW with high-velocity resolution and 976 kHz channels centered at 230.957 GHz.

Band 6 data were obtained using the same bandpass and flux calibrator as Band 3 and J1603-4904 as a phase calibrator. We detected the H30 α line and the underlying continuum.

In both bands, we used the Hogbom algorithm in the *tclean* task of CASA,¹² Briggs weighting of 0, and a pixel size of 4 mas; the channel width resulted in 1.26 km s⁻¹ resolution for H30 α ; 1.48 km s⁻¹ for H40 α , H50 β and H57 γ ; and 1.71 km s⁻¹ for H42 α . The corresponding synthesized beam

¹¹ International Celestial Reference System (ICRS) coordinates: R.A. $16^{\text{h}}17^{\text{m}}13^{\text{s}}391$, decl. $-51^{\circ} 59' 10''.712$; proper motions $-3.549 \pm 0.100 \text{ mas yr}^{-1}$ in R.A. and $-3.914 \pm 0.082 \text{ mas yr}^{-1}$ in decl. (GAIA collaboration 2022).

¹² CASA: Common Astronomy Software Applications for radio astronomy (The CASA Team 2022).

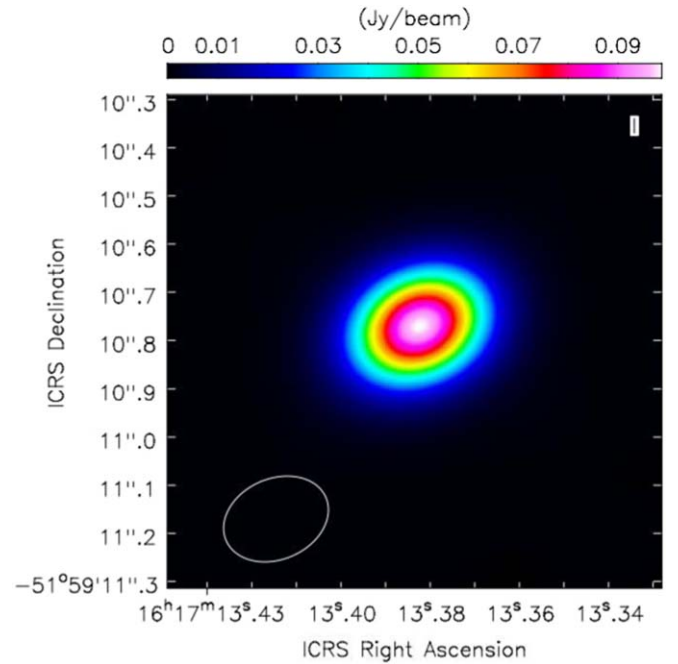


Figure 1. Continuum image of Mz3 obtained with ALMA at 93 GHz. The color wedge is presented at the top of the graph and the beam ellipse at the bottom left corner.

sizes are ($0''.396 \times 0''.343$, -79°), ($0''.252 \times 0''.189$, -61°), and ($0''.289 \times 0''.219$, -66°).

We also cleaned a cube with 12.6 km s^{-1} resolution (ten channels) for the SPWs that include the H30 α line, and 13.3 and 15.4 km s^{-1} (nine channels) for the H40 α and H42 α lines, respectively, to obtain relative astrometric positions of the centroid of the image for each velocity channel, with a better signal-to-noise ratio.

3. Results

Figure 1 presents the band 3 continuum image, with resolution ($0''.225 \times 0''.165$). This was the best resolution obtained in the observations, and we can see that the source was not resolved. The integrated flux density is $0.12 \pm 0.01 \text{ Jy}$ at 93 GHz and $0.17 \pm 0.03 \text{ Jy}$ at 224 GHz, resulting in a spectral index of $0.39^{+0.26}_{-0.31}$. This value is compatible with the emission of an ionized wind or of a compact H II region with the turnover of the bremsstrahlung spectrum above 90 GHz although there is still the possibility that the continuum spectrum is the combination of the spectra of more than one component, as discussed in Section 4.2.

Figure 2 presents the spectra of the detected lines (H30 α , H40 α , H42 α , H50 β , and H57 γ). The vertical lines show the frequencies of the corresponding H and He recombination lines. The corresponding frequencies in the LSR reference frame differ by a factor 7×10^{-6} of the rest frequency, corresponding to the velocity of about 2 km s^{-1} of the LSR relative to the baricenter. The presence of maser effect in the H30 α recombination line becomes clear when we compare its profile with those of the H40 α and H42 α lines. In addition, we must point out the difference in the shapes of the H50 β and H40 α lines, the former having a flatter profile.

A very weak feature can be seen at the high-frequency end of the H40 α and H42 α lines and could be part of the corresponding He lines; if the half-power width of the H and

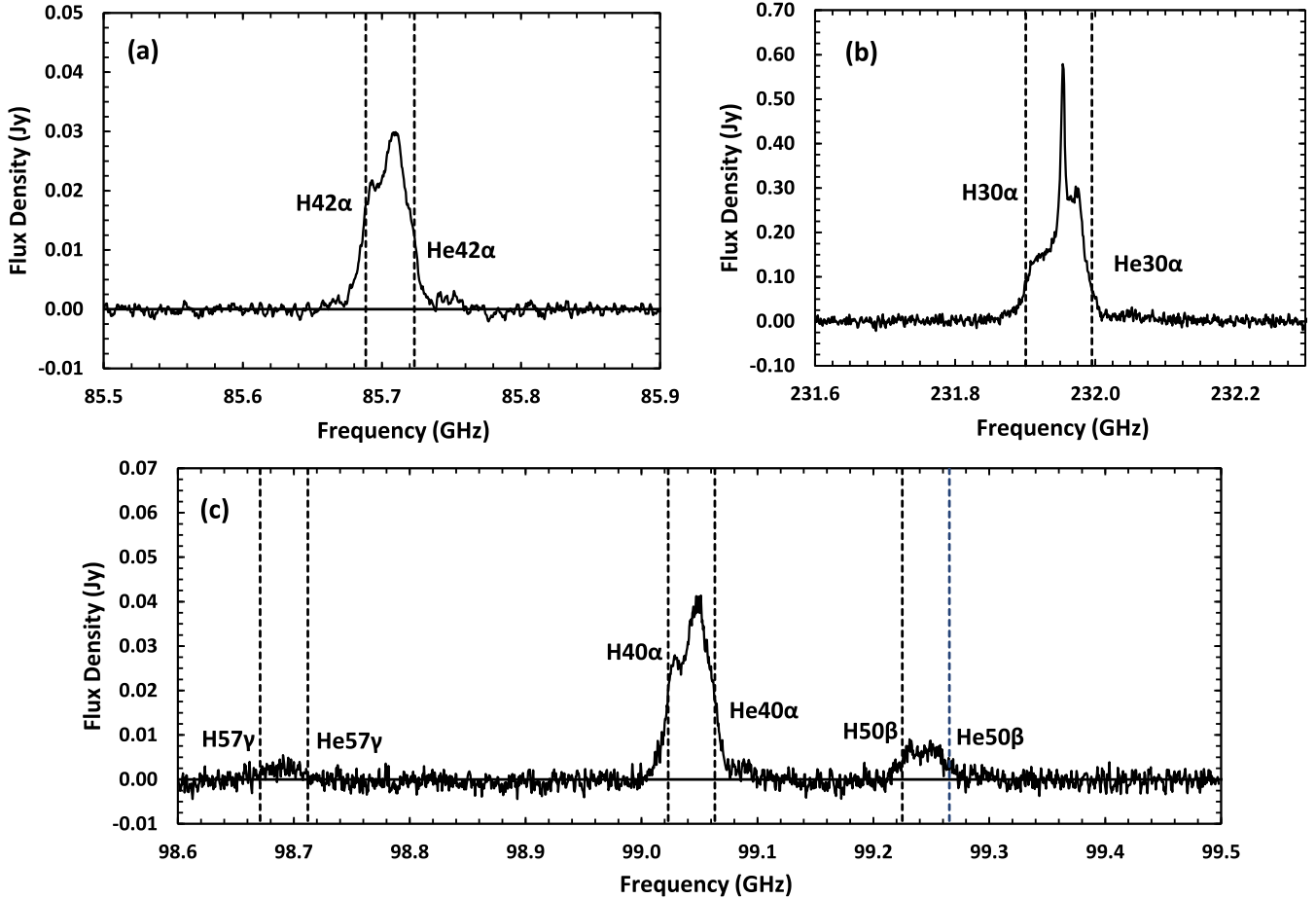


Figure 2. (a) H42 α , (b) H30 α , and (c) H40 α , H50 β , and H57 γ line profiles of Mz3. The vertical lines represent the rest frequencies of the respective H and He lines. Notice the factor of 10 difference in scale between the H30 α and the other two lines.

He recombination lines is the same, part of the He line would be blended with the corresponding H line. The helium lines are remarkably faint compared to the published helium abundance (one-fifth of H; S. R. Pottasch & R. Surendiranath 2005), probably reflecting the ionization condition of the region, but they are too weak to be analyzed in detail.

In order to obtain a quantitative evaluation of the magnitude of the maser effect in the different recombination lines, we use the fact that when the physical conditions are appropriate for maser amplification in the H30 α line, the H50 β line is formed in LTE (V. S. Strel'nitski et al. 1996).

To calculate quantitatively the amplification of the H $n\alpha$ lines, with $n = 30, 40,$ and 42 , we can use the observed H50 β line to estimate the profile that each of the H $n\alpha$ lines would have if they were emitted in LTE. Since the H50 β line is too noisy, we fitted three Gaussians to better describe its profile. Figure 3 shows the observed H50 β line profile in the LSR velocity scale, the Gaussians, the model profile (sum of the Gaussians), and the residuals of the fitting. The vertical line shows the systemic LSR velocity of the nebula, -53.4 km s^{-1} (-55 km s^{-1} heliocentric), as determined by M. P. Redman et al. (2000).

We used the model profile $S_{50\beta}^{\text{model}}(\nu)$ of the observed H50 β line to calculate the expected line profiles in LTE of the H $n\alpha$

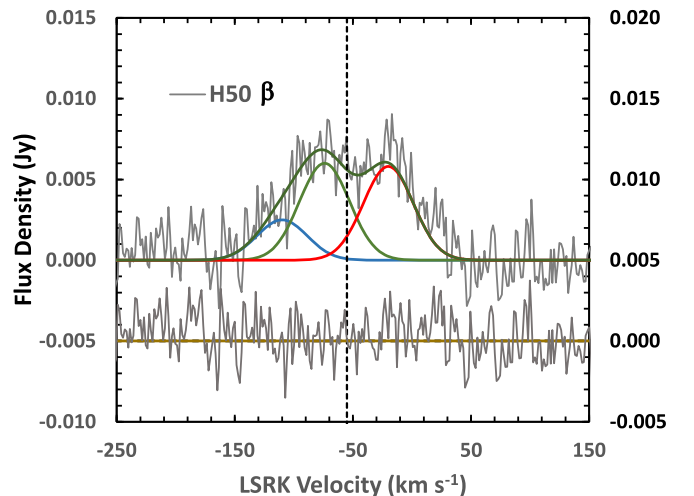


Figure 3. Observed H50 β line (gray line), together with a three-Gaussian fit (individual Gaussians are shown in blue, green, and red; the sum is shown in black). The flux density of the lines (top plot) is shown in the left vertical axis, and the residuals of the fitting (bottom plot) are shown in the right vertical axis. The vertical dashed line represents the systemic velocity of Mz3. The parameters of the Gaussians (amplitude in Jy; center velocity and half-power width in kilometers per second) are (0.0025, $-110, 20.8$), (0.0060, $-74, 20.8$), and (0.0058, $-20, 20.8$).

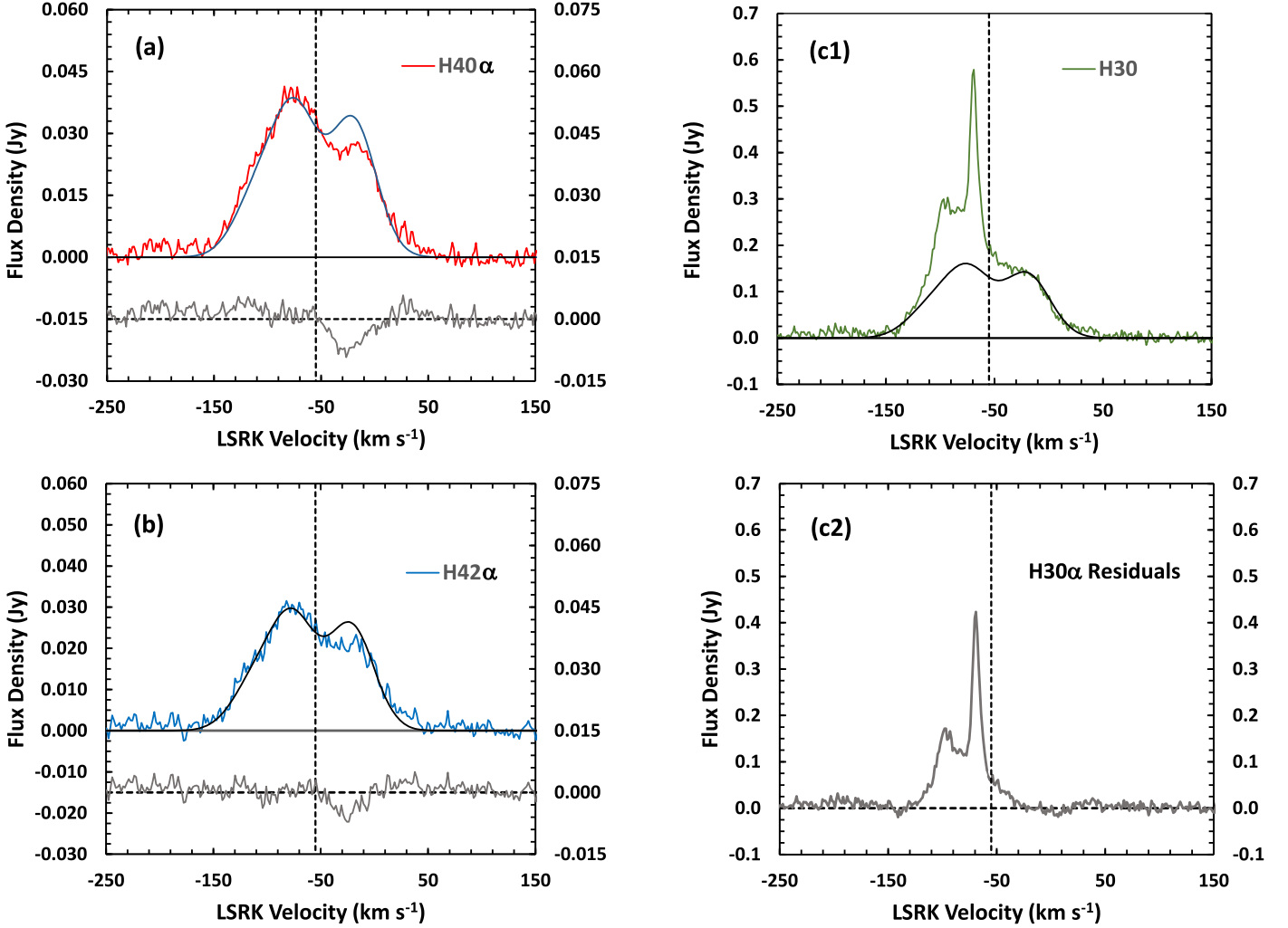


Figure 4. (a) Upper half of the graph shows the observed H40 α line of Mz3, superposed on a model line, constructed assuming that it was formed in LTE from the H50 β model profile presented in Figure 3. The flux density scale is presented in the left vertical axis. Lower half of the graph shows the difference between observed and model profiles, showing the absorption due to the dasar effect. The flux density scale is presented in the right vertical axis. (b) The same as (a) for the H42 α line; (c1) observed H30 α line, superposed on the LTE model line; and (c2) difference between observed and model H30 α line profiles, where the maser lines can be seen.

lines, $S_{n\alpha}^{\text{LTE}}(\nu)$, using the following expression:

$$S_{n\alpha}^{\text{LTE}}(\nu) = C(n\alpha, 50\beta) S_{50\beta}^{\text{model}}(\nu) \quad (1)$$

with

$$C(n\alpha, 50\beta) = \frac{f_{n+1,n}}{f_{52,50}} \left(\frac{52}{50}\right)^2 \left(\frac{n}{n+1}\right)^2 \left(\frac{\nu_{n\alpha}}{\nu_{50\beta}}\right)^2, \quad (2)$$

where $f_{m,n}$ is the oscillator strength of the transition between levels n and m , and ν_n is the transition frequency. The values of $C(n\alpha, 50\beta)$ are 23.8, 5.3, and 4.6 for the H30 α , H40 α , and H42 α lines, respectively (H. Goldwire 1968).

Figure 4 shows the difference between the H $n\alpha$ observed and model LTE profiles, obtained from Equation (1), and the respective residuals. The strong narrow lines in the H30 α residuals confirm the maser nature of the emission. However, the negative residuals seen in the H40 α and H42 α lines, centered at LSR velocity -25 km s^{-1} , were unexpected; they imply that the intensities of these lines are smaller than those of the lines in LTE. This condition, defined as overcooling, or

dasar effect, was introduced for the first time by V. S. Strelnit-ski et al. (1996) in the theory of HRLs; it appeared when the angular momentum l was taken into account in the calculation of b_n , the population of the levels of quantum number n relative to their LTE values (P. J. Storey & D. G. Hummer 1995). The dasar effect implies that the excitation temperature T_x of the line is lower than the kinetic temperature, while the maser effect results in negative excitation temperatures, which are defined as

$$T_x = \frac{h\nu_0/k}{\ln(n_1/n_2)}, \quad (3)$$

where ν_0 is the frequency of the line, n_i the population density per one degeneracy sublevel, and subindices 1 and 2 denote the lower and upper levels of the transition. The ratio n_1/n_2 can be written in terms of their LTE values as

$$\frac{n_1}{n_2} = \frac{b_1 n_1^*}{b_2 n_2^*}, \quad (4)$$

where * indicates LTE values (V. S. Strel'nitski et al. 1996). A possible model that explains the simultaneous existence of masers and dasars in the HRLs is presented in Section 4.3.

4. Discussion

Mz3 was not resolved by ALMA, neither in the continuum of 93 and 224 GHz nor in the HRLs H30 α , H40 α , H42 α , or H50 β . Interferometric observations of the 8.64 GHz continuum, with resolution ($1''.2 \times 0''.9$) obtained on 2002 December 13–15, with the Australia Telescope Compact Array (ATCA; I. Bains et al. 2004) also revealed a compact unresolved source with flux density of 0.20 mJy and spectral index between 4.8 and 8.6 GHz of 0.36, obtained from Gaussian fitting to the uniformly weighted radio maps. Assuming that the compact emission detected by ATCA and ALMA are produced by the same source, we obtain a spectral index between 8.6 and 93 GHz of 0.45, compatible with the value of $0.39^{+0.26}_{-0.31}$ for frequencies between 93 and 224 GHz, presented in Section 3.

From the flux density at 93 GHz and assuming the optical depth of the continuum as $\tau_C = 1$ and the electron temperature $T_e = 10^4$ K, we found the solid angle of the source as $\Omega = 7.5 \times 10^{-14}$ sr, which corresponds to an angular radius of $0''.03$ or a linear radius $R = 80$ au at a distance of 2.7 kpc. Assuming this value for the length of the source in the direction of the observer, we find an electron density $n_e = 10^7$ cm $^{-3}$, compatible with the existence of high excitation lines of heavy elements, as detected by Y. Zhang et al. (2019) and with what is expected from the existence of maser effects in the HRLs.

4.1. The Structure of the Emitting Region

The detection of maser and dasar effects in the HRLs of Mz3 puts constraints on the physical conditions of the emitting regions. The difference in the LSR velocities of the emission and absorption residuals implies that they can originate in different sources. In fact, according to V. S. Strel'nitski et al. (1996) masers in the H30 α line occur at densities of the order of 10^7 cm $^{-3}$, while dasars occur in the H40 α and H42 α lines at densities of the order of 10^3 cm $^{-3}$. The LSR velocities of the maser lines are -69 and -98 km s $^{-1}$, and -25 km s $^{-1}$ for the dasar, with the three of them close to the velocities used for the Gaussians in the H50 β line model. The velocities, relative to the systemic motion of the nebula, are -15.6 and -44.6 km s $^{-1}$ for the masers and $+28.4$ km s $^{-1}$ for the dasar.

To investigate if the lines come from different positions in the nebula, we resampled our cube in bins of ten channels for the H30 α line and nine channels for H40 α , resulting in a velocity resolution of 12.6 and 13.3 km s $^{-1}$, respectively. In each of those cubes, we fitted a two-dimensional Gaussian in the image plane of each bin and obtained the position of its centroid.¹³

Figure 5(a) presents the position of the centroids for each velocity channel, for lines H40 α and H30 α . The velocities at each position are indicated by different colors and represented in Figures 5(b) and (c) as straight lines superposed on the line profiles of the two recombination lines.

Figures 5(d) and (e) present the velocity versus R.A.(J2000) and decl.(J2000) of the centroids of the H30 α images; the vertical lines represent the position of the continuum source in

2022 August. The errors in R.A. and decl. obtained from the fitting are also shown in these figures; they are compatible with the thermal noise (A. R. Thompson et al. 2017), given by

$$\sigma_\theta = \frac{1}{2} \frac{\theta_{\text{res}}}{\mathfrak{R}_{\text{sn}}} \sim 1.2 \text{ mas}, \quad (5)$$

where θ_{res} is the resolution of the interferometer and \mathfrak{R}_{sn} is the signal-to-noise ratio, which in our case was $\mathfrak{R}_{\text{sn}} \simeq 100$ for each channel of both lines. Similar errors were obtained for the centroids of the H40 α images.

The calculated positions can be used only to compare the relative positions of the different velocity channel images. The uncertainty in their absolute positions must be evaluated taking into account the prescription given by ALMA:

$$\sigma_\theta^{\text{abs}} = \frac{\theta_{\text{res}}/9}{\min(\mathfrak{R}_{\text{sn}}, 20)} \sim 25 \text{ mas}. \quad (6)$$

Therefore, the absolute positions of the H30 α and H40 α lines shown in Figure 5(a) coincide within the precision of the ALMA observations.

Figures 5(d) and (e) also show the central position of the continuum source, its ICRS coordinates differ from those of Mz3 by 0.005 s in R.A. and $0''.004$ in decl., showing that the radio source is very close to the star.

In Figure 5(a), we can see an almost linear structure, extending for about $0''.02$, or 54 au, at a distance of 2.7 kpc, with a well-defined velocity gradient. The positions corresponding to the velocity bins of the strongest maser line are very close, indicating that the solid angle of the source is very small, as usual in maser sources. The position of the dasar source, with LSR velocity around -25 km s $^{-1}$ seems to be separated from that of the masers, but it is not possible to know if there is still a third source, which corresponds to the emission profile seen in the H50 β line. In Figure 5(f), we present the X-ray and optical images of the extended hourglass nebula (J. H. Kastner et al. 2003; M. A. Guerrero et al. 2004), indicating the direction perpendicular to this linear structure.

4.2. Comparison with Other Sources

There is a close resemblance between the HRL profiles of Mz3, presented in this work, and those of the B[e]-type stars MWC 349A and MWC 922. They all present two narrow lines in the H30 α spectrum, attributed to maser amplification, as well as a bipolar nebula surrounding the stars (J. Martín-Pintado et al. 1989; V. V. Gvaramadze & K. M. Menten 2012; C. Sánchez Contreras et al. 2017). RCW 349A was extensively studied with interferometric techniques; it was resolved with ALMA in the continuum at the frequency of 345 GHz and the position of the maser sources determined by the centroids of the Gaussian images along the profiles of the H26 α and H30 α RRLs (S. Prasad et al. 2023). The authors reported a linear structure, interpreted as an edge-on disk, rotating with Keplerian velocities around the central star, and a low-velocity wind, photoevaporated from the disk. A similar model was used to fit the ALMA observations of MWC 922 (C. Sánchez Contreras et al. 2017). The linear structure defined by the centroids of the velocity channel images of Mz3 are also compatible with rotation, but the low spatial resolution does not allow us to define the rotation model. Notice that both maser lines are blueshifted in Mz3, while one of them is redshifted in MWC 349A.

¹³ We used the CASA fitting routine *fit*.

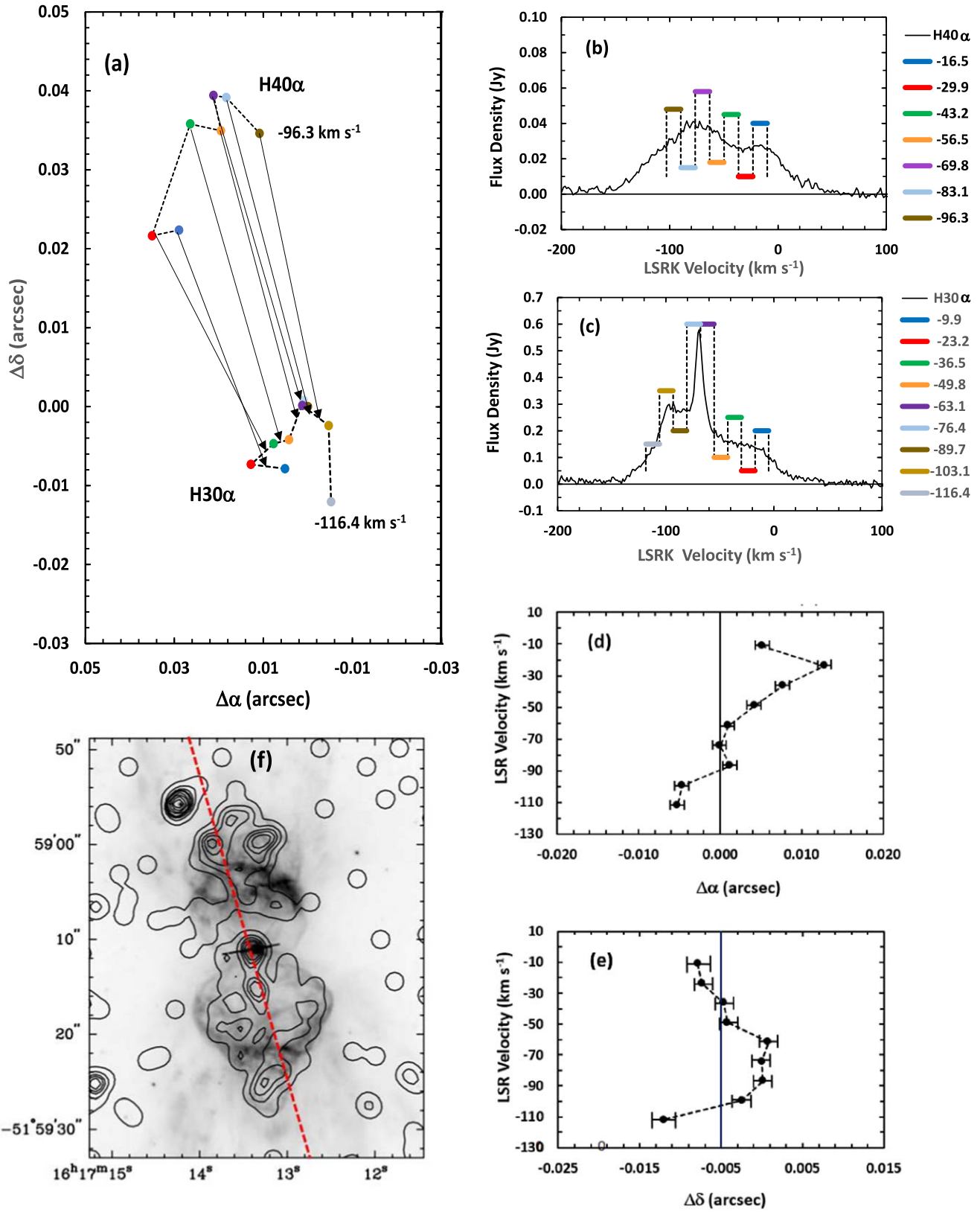


Figure 5. Panel (a): positions of the centroids of a double-Gaussian model for each of the velocity-binned images of the H40 α and H30 α lines. The axis represents offsets from R.A. $16^{\text{h}}17^{\text{m}}, 13^{\text{s}}381$, decl. $-51^{\circ}59'10''.802$. The different colors correspond to different velocities. The relative errors of the coordinates are always smaller than 1.5 mas and are shown in panels (d) and (e). The errors in the absolute astrometric coordinates are of the order of 20 mas. The black arrows join points of approximately the same velocity in both lines. Panels (b) and (c): profiles of the H40 α and H30 α lines, with the colored horizontal lines representing the velocity intervals used to determine the positions shown in (a). Panels (d) and (e): LSR velocities vs. R.A. and decl. for the H30 α line. Vertical lines represent the position of the centroid of the continuum image. Panel (f): optical and X-ray image of Mz3 (M. A. Guerrero et al. 2004), with a dashed line showing the direction perpendicular to the linear source structure.

However, the possible existence of several isolated sources in Mz3 cannot be ignored, an example being the strong HRL maser lines in η Carinae (Z. Abraham et al. 2020), which turned out to be the superposition of 16 spatially resolved sources with similar velocities, in a region of $0''.6$.

4.3. The Coexistence of Masers and Dasars

The simultaneous observation of masers and dasar absorption in the HRLs of Mz3 can be explained if we assume that the anomalous absorption is produced by an ionized low-density ($\sim 10^3 \text{ cm}^{-3}$) cloud, with temperature of about $7 \times 10^3 \text{ K}$, located between the maser and the observer, moving with a radial LSR velocity of -25 km s^{-1} (28 km s^{-1} relative to Mz3). The maser source can be either a dense rotating disk or the combination of several compact sources, with electron densities and temperatures of the order of 10^7 cm^{-3} and 10^4 K , respectively. The existence of the low-density clouds is supported by the extended radio continuum detected with ATCA (I. Bains et al. 2004), the physical conditions of the plasma in Mz3 derived from the optical spectrum (Y. Zhang et al. 2019), and the accretion disk postulated by J. H. Kastner et al. (2003) to explain the X-ray absorption in the core although the redshifted velocity of the dasar could point to an external origin for the cloud.

The observed brightness temperature T_B^{obs} , with the background emission subtracted, will be

$$\Delta T_B^{\text{obs}}(\nu) = [T_B^{\text{das}}(\nu) - T_B^{\text{mas}}(\nu)][1 - e^{-\tau(\nu)}], \quad (7)$$

where T_B^{das} and T_B^{mas} are the brightness temperatures of the dasar and background maser sources, respectively, and $\tau^{\text{das}} = \tau_L^{\text{das}} + \tau_C^{\text{das}}$ represents the optical depth of the dasar (line plus continuum).

Since $T_B^{\text{das}}(\nu) < T_B^{\text{mas}}(\nu)$, we will see the line in absorption, and assuming $\tau^{\text{das}} \ll 1$, Equation (7) becomes

$$\Delta T_B^{\text{obs}}(\nu) = [T_B^{\text{das}}(\nu) - T_B^{\text{mas}}(\nu)][\tau_L^{\text{das}}(\nu) + \tau_C^{\text{das}}(\nu)], \quad (8)$$

with τ_C^{das} and τ_L^{das} given by M. Brocklehurst & M. J. Seaton (1971):

$$\tau_C^{\text{das}} = 6.94 \times 10^{-8} \nu^{-2} \left(\frac{10^4}{T_e} \right)^{3/2} g_{ff} EM \quad (9)$$

$$\tau_L^{\text{das}} = 1.064 \times 10^{-12} b_n \beta_{n,n+m} \left(\frac{10^4}{T_e} \right)^{5/2} \frac{f_{n,n+m}(\nu \phi_\nu) EM}{\nu} \quad (10)$$

where g_{ff} is the Gaunt factor for free-free emission, $EM = n_e^2 L$ is the emission measure, ϕ_ν the line profile, and $\beta_{n,n+m}$ the amplification factor, given by

$$\beta_{n,n+m} = \frac{1 - b_{n+m}/b_n e^{-h\nu_n/kT_e}}{1 - e^{-h\nu_n/kT_e}}; \quad (11)$$

ϕ is the line profile, which we assume to be a Gaussian, with a half-power width $\Delta\nu$, corresponding to a velocity interval Δv :

$$\nu \phi_\nu = \left(\frac{4 \ln 2}{\pi} \right)^{1/2} \left(\frac{\nu}{\Delta\nu} \right) = \left(\frac{4 \ln 2}{\pi} \right)^{1/2} \left(\frac{c}{\Delta v} \right). \quad (12)$$

The values of $\tau_C^{\text{das}}(\nu) + \tau_L^{\text{das}}(\nu)$ at the centers of the H30 α , H40 α , H42 α , and H50 β lines are presented in Figure 6 for the electron temperature 7500 K, $\Delta v = 20 \text{ km s}^{-1}$, and electron densities 10^2 , 10^3 , and 10^4 cm^{-3} . They were calculated using

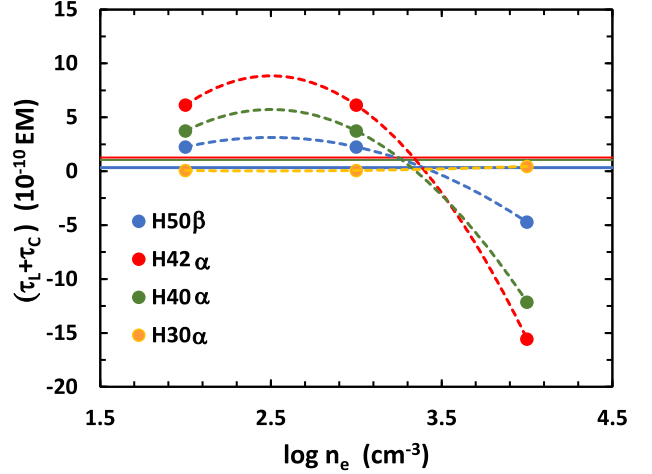


Figure 6. Optical depth of the H30 α , H40 α , H42 α , and H50 β lines for an electron temperature of 7500 K and different values of the electron density (circles). Dashed lines are quadratic fittings to the data. Horizontal lines represent the optical depth values assuming LTE.

the b_n coefficients published by P. J. Storey & D. G. Hummer (1995) and the Gaunt factors provided by P. A. M. van Hoof et al. (2014). The LTE values of the optical depths are shown as horizontal lines.

We can see that the values of τ^{das} differ from those in LTE by different factors, depending on the value of n , being higher for the H40 α and H42 α lines than for H30 α and H50 β , explaining the dasar effect.

Notice that the values of b_n do not include either the continuum field produced by the emitting plasma or the background photon field, and therefore the values of the optical depths are only approximations (F.-Y. Zhu et al. 2022).

5. Conclusions

The recombination lines H40 α , H42 α , H50 β , and H56 γ were detected with ALMA in Mz3 in 2021 November, while the H30 α line was detected in 2022 August. Comparison of the profiles and line intensities led to the following conclusions:

1. The H30 α line shows clear signs of maser amplification, with two strong lines at velocities of -16 and -46 km s^{-1} , relative to the systemic velocity of the nebula.
2. The density of the gas responsible for the maser emission must of the order of 10^7 cm^{-3} . At this density, the H50 β line is formed in LTE.
3. The line profiles of the H40 α and H42 α lines are different from that of the H50 β line, showing lower intensity than expected in LTE at velocities close to $+28 \text{ km s}^{-1}$ relative to the systemic velocity of the nebula, which we interpret as overcooling or dasar effect.
4. The coexistence of masers and dasars can be explained if we assume that the absorption occurs in a low-density cloud located in front of the maser source.
5. The density necessary for the occurrence of a dasar is about 10^3 cm^{-3} , implying the existence of at least two different emitting regions.
6. Although the source was not spatially resolved, it was possible to model the images of the velocity cube with

two-dimensional Gaussians and determine their centroid positions.

7. These positions are distributed along a line with a well-defined velocity gradient, being compatible with rotation.
8. The detection of dasars in HRLs, formed in high-density regions, shows the importance of including also other possible lower-density regions, with their corresponding velocities, in the models that calculate the line profiles.


The determination of the real structure of the recombination line region needs observation with higher spatial resolution, which can be achieved by ALMA at higher frequencies and more extended configurations.

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