



A Two-Component Lateral Distribution Function for the Reconstruction of Air-Shower Events in IceTop

Downloaded from: <https://research.chalmers.se>, 2025-06-01 07:52 UTC

Citation for the original published paper (version of record):

Soldin, D., Abbasi, R., Ackermann, M. et al (2024). A Two-Component Lateral Distribution Function for the Reconstruction of Air-Shower Events in IceTop. *Proceedings of Science*, 444. <http://dx.doi.org/10.22323/1.444.0357>

N.B. When citing this work, cite the original published paper.

A Two-Component Lateral Distribution Function for the Reconstruction of Air-Shower Events in IceTop

The IceCube Collaboration

(a complete list of authors can be found at the end of the proceedings)

E-mail: mark.weyrauch@kit.edu, dennis.soldin@kit.edu

The surface component of the IceCube Neutrino Observatory, IceTop, consists of an array of ice-Cherenkov tanks measuring the electromagnetic signal as well as low-energy (\sim GeV) muons from cosmic-ray air showers. In addition, accompanying high-energy (above a few 100 GeV) muons can be observed in coincidence in the deep in-ice detector. A combined measurement of the low- and high-energy muon content is of particular interest for tests of hadronic interaction models as well as for cosmic-ray mass discrimination. However, since IceTop does not feature dedicated muon detectors, an estimation of the low-energy muon component of individual air showers is challenging.

In this work, a two-component lateral distribution function (LDF), using separate descriptions for the electromagnetic and muon lateral distributions of the detector signals, is introduced as a new approach for the estimation of low-energy muons in air showers on an event-by-event basis. The principle of the air-shower reconstruction using the two-component LDF, as well as its reconstruction performance with respect to primary energy and number of low-energy muons will be discussed.

Corresponding authors: Mark Weyrauch^{1*}, Dennis Soldin^{2,3}

¹*Karlsruhe Institute of Technology, Institute for Astroparticle Physics, Karlsruhe, Germany*

²*Karlsruhe Institute of Technology, Institute of Experimental Particle Physics, Karlsruhe, Germany*

³*Department of Physics and Astronomy, University of Utah, Salt Lake City, UT 84112, USA*

* Presenter

The 38th International Cosmic Ray Conference (ICRC2023)
26 July – 3 August, 2023
Nagoya, Japan



1. Introduction

Cosmic rays enter the Earth's atmosphere where they produce extensive air showers (EASs) which can be measured with large detector arrays at the ground. Although cosmic rays have been measured over several orders of magnitude in energy, large uncertainties remain, for example in measurements of their mass composition [1]. The main challenge lies in the understanding of muon production in EAS. Various experiments performed measurements of muons in EASs over the last decades, where some reported discrepancies in the number of muons in simulated and observed air showers, while others reported no discrepancies [2–5]. IceTop has measured the GeV muon density in EASs at cosmic ray energies between 2.5 PeV and 120 PeV where no significant discrepancies have been observed [6]. However, this statistical analysis does not provide event-by-event information of the muon content in EASs which is important in order to correlate measurements of low-energy (\sim GeV) muons in IceTop and high-energy (above a few 100 GeV) muons in the deep-ice detector of the IceCube Neutrino Observatory [7–9]. This coincident measurement will provide strong constraints on hadronic interaction models [10, 11].

In this work, we will present a novel approach to reconstruct the number of muons in air showers measured with IceTop on an event-by-event basis. This approach uses a two-component lateral distribution function (LDF) to fit the EAS signals measured in IceTop in order to account for the electromagnetic and muonic signal contributions separately.

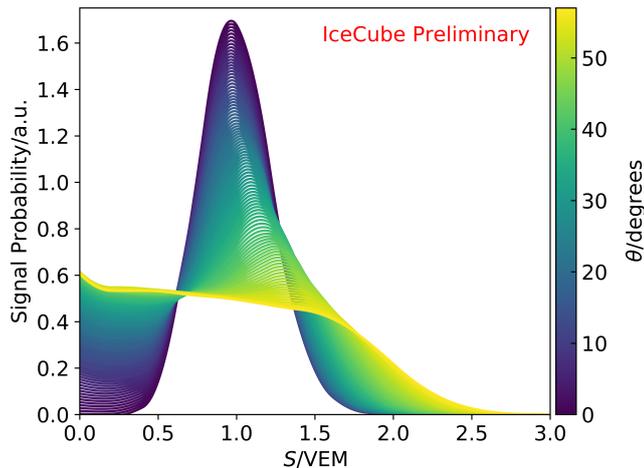
2. IceTop

IceTop [12] is the surface detector of the IceCube Neutrino Observatory (IceCube) [13] and it is located at the geographic South Pole at about 2.8 km above sea level, corresponding to an atmospheric depth of about 690 g/cm^2 . IceTop comprises 81 detector stations, each consisting of two cylindrical Cherenkov tanks, which are separated by approximately 10 m, deployed in a triangular grid with a spacing of 125 m. Each tank is filled with clear ice and houses two digital optical modules (DOMs) which measure the Cherenkov light generated by charged particles traversing the ice. An infill area in the center of the detector has a denser spacing of $< 50 \text{ m}$, which improves the sensitivity of cosmic ray measurements at low energies.

A local trigger occurs when the signal in a tank exceeds a predefined discriminator threshold, as described in Ref. [12]. Further, a hard local coincidence (HLC) is defined where both tanks in a station have a local trigger within a time window of $1 \mu\text{s}$. If there is a local trigger in only one tank, it is called a soft local coincidence (SLC). Following calibration, which accounts for the individual tank responses, all IceTop signals are expressed in units of vertical equivalent muons (VEM), which is the average signal produced by a muon traversing the tank vertically. For further analysis, various basic event cleanings are applied and only events that pass a cosmic ray filter are kept (see Ref. [12] for details). The subsequent EAS reconstruction is generally based on a fit to the lateral distribution of HLC hits (only) which provides the core position and direction of the air shower, as well as an estimate of its energy. In the following, we will describe an approach to extend this technique by including SLCs in order to reconstruct the lateral distribution of muons, which typically produce the SLC hits recorded by IceTop.

3. Signal Model

In order to perform a combined fit of two LDFs a detailed understanding of the corresponding signal models is required. The signal response for the electromagnetic part of the shower is well-known from the standard IceTop reconstruction described in Ref. [12]. For the muonic part, the response of the IceTop tank is modeled via dedicated response simulations. Muons with different multiplicities and zenith angles are injected into the tank, followed by the simulation of the detector response



using Geant4 [14] and a calibration to units of VEM. Generally, the signal distributions of muons mainly depend on the track length within the detector, i.e. are determined by the tank geometry. The tank response for single muons as function of the zenith angle is shown in Fig.1. The peak position for a vertical muon is by definition located at a charge deposit of 1 VEM. With higher inclination the maximum track length increases, causing a shift in the peak position proportional to $1/\cos \theta$. Simultaneously, the influence of muons traversing only the edge of the detector manifests in an increased contribution at lower charge deposits. Hence, the width of the signal distribution, i.e. the signal fluctuation increases as a function of the zenith angle. The signal distributions $p_{\mu,\text{sig}}(S|\theta, n)$ for multiplicities of up to $n = 15$ muons as function of muon inclination are saved as spline fits and can be retrieved for any given combination of observed signal strength and zenith angle. It is important to mention that the latter is given by the reconstructed primary zenith angle, assuming the muon directions follow the direction of the primary cosmic-ray. In order to obtain the total muon signal probability density function (PDF) given an expectation value $\langle N_{\mu} \rangle$ for the average number of muons, the signal distributions for an integer number of muons have to be weighted by the according Poisson probability

$$p_{\mu}(S|\theta, \langle N_{\mu} \rangle) = \sum_n \frac{\langle N_{\mu} \rangle^n}{n!} e^{-\langle N_{\mu} \rangle} p_{\mu,\text{sig}}(S|\theta, n), \quad (1)$$

where $\langle N_{\mu} \rangle$ is determined by the muon signal expectation value $\langle S_{\mu} \rangle$ given by the muon LDF multiplied with the effective IceTop tank area. For muon multiplicities above $n = 15$ the muon PDF is approximated by a Gaussian distribution. The total PDF for an observed tank signal S is determined by the convolution of electromagnetic (Gaussian) PDF

$$p_{\text{em}}(S|\theta, \langle S_{\text{em}} \rangle) = \frac{1}{\sigma_{\text{em}, \log_{10} S} \sqrt{2\pi}} \exp\left(-\frac{(\log_{10} S - \log_{10} \langle S_{\text{em}} \rangle)^2}{2\sigma_{\text{em}, \log_{10} S}^2}\right) \quad (2)$$

and the muon signal PDF. The electromagnetic signal is affected by the snow accumulation on top of the IceTop tanks. Thus, the resulting signal attenuation has to be taken into account in the reconstruction procedure.

4. Implementation

For the combined reconstruction of the electromagnetic and muonic part, respectively, two LDFs are fit to an observed charge distribution. The LDF used for the electromagnetic contribution is the Double Logarithmic Parabola (DLP) function [12]

$$S_{\text{em}} = S_{\text{em},125} \left(\frac{r}{r_{\text{em}}} \right)^{-\beta_{\text{em}}(S_{\text{em},125}) - \kappa(S_{\text{em},125}) \log_{10}(r/r_{\text{em}})}, \quad r_{\text{em}} = 125 \text{ m}. \quad (3)$$

The parameters β_{em} and κ describe the slope and the curvature of the function, respectively. $S_{\text{em},125}$ is given as the signal strength at a distance of 125 m perpendicular to the shower axis and is used as a proxy for the primary energy. This reference distance is found to minimize the primary mass dependence while providing a good reconstruction resolution [15]. For the parametrization of κ as a function of $S_{\text{em},125}$, a separate simulation set is used in which all muons are omitted for the derivation of the detector response. The DLP function is then fit to the average lateral distributions in bins of $S_{\text{em},125}$, resulting in both a decrease in the slope β_{em} and an increase in the curvature κ (i.e., steeper tail of the distribution) for an accurate description of the observed lateral shape of purely electromagnetic showers. While the parametrization of κ is fixed within the whole reconstruction procedure, the LDF slope is treated as free parameter with its parametrization taken as an initial choice. The LDF used to model the muon contribution is based on the Greisen function [16]:

$$S_{\mu} = S_{\mu,600} \left(\frac{r}{r_{\mu}} \right)^{-\beta_{\mu}(S_{\text{em},125})} \left(\frac{r + 320 \text{ m}}{r_{\mu} + 320 \text{ m}} \right)^{-\gamma(S_{\text{em},125})}, \quad r_{\mu} = 600 \text{ m}, \quad (4)$$

containing two slope parameters β_{μ} and γ . Both are parametrized as a function of $S_{\text{em},125}$, accounting for the steepening of the muon density profile with increasing primary energy. $S_{\mu,600}$ is used as a proxy for the low-energy muon number and is derived at a distance of 600 m in order to reduce the electromagnetic background while minimizing the influence of large fluctuations far from the shower axis. The distance of 320 m is the Greisen radius [16]. In order to constrain the parameters to the expected physical range and at the same time increase the stability of the reconstruction procedure, all parameters are fit while applying a penalty term

$$p_{\text{constr}} = 1/2 \sum_p \lambda_p (p - p_{\text{par}})^2, \quad p \in \{\log_{10} S_{\text{em},125}, \beta_{\text{em}}, \beta_{\mu}, \gamma\}, \quad \lambda_p = \sigma_p^{-2} \quad (5)$$

to the final likelihood (LLH) function (Eq. (9)). In the case of $S_{\text{em},125}$, a constant penalty is applied, corresponding to the observed maximum spread of 4% when omitting all muons from a standard IceTop reconstruction [12]. The effective strength λ_p of the constraints for the slope parameters is determined by the inverse of the observed spread σ_p^{-1} in the corresponding parameter distributions after application of the reconstruction procedure. This spread reduces as a function of $S_{\text{em},125}$ as the reconstruction becomes more stable with increasing multiplicity. Thus, the corresponding parametrization of λ_p results in an increased strength of the penalty with increasing $S_{\text{em},125}$. In order to minimize a possible bias introduced by a given constraint, its parametrization is performed iteratively. With each iteration p_{par} and λ_p are updated with the mean and spread of the parameter distributions, respectively. The reconstruction procedure itself is based on a negative log-likelihood minimization and is incorporated into a new reconstruction framework, designed for the reconstruction of multiple LDFs and/or detectors [17]. For the two-component fit, the combined PDF is utilized in the convolution regime to determine the likelihood value for a given tank signal:

$$p(S|\theta, \langle S_{em} \rangle, \langle N_{\mu} \rangle) = p_{trg} \begin{cases} p_{em}((S - S_{\mu})/c_{snow}|\theta, \langle S_{em} \rangle), & \sigma_{\mu} < \sigma_{em}/10 \\ p_{\mu}(S - S_{em}c_{snow}|\theta, \langle N_{\mu} \rangle), & \sigma_{em} < \sigma_{\mu}/10 \\ \int_0^S p_{em}(S'_{em}/c_{snow}|\theta, \langle S_{em} \rangle) p_{\mu}(S - S'_{em}|\theta, \langle N_{\mu} \rangle) dS'_{em}, & \text{else} . \end{cases} \quad (6)$$

To account for the snow accumulation, an exponential attenuation factor c_{snow} [18] is applied to the electromagnetic signal. If either the electromagnetic or muon signal uncertainty (determined based on the corresponding signal expectation) is small in comparison, the according signal PDF is approximated as a delta function. The resulting signal PDF is multiplied with the discriminator trigger p_{trg} [6] to account for the signal detection efficiency. In addition to the resulting signal likelihood, saturated as well as silent (i.e. non-triggered) detectors are taken into account. While the former (ll_{sat}) can be included as Gauss CDF [12], the likelihood of silent detectors (ll_{sil}) is based on the no-hit probability p_{nohit} derived from the convolution of the electromagnetic and muon contribution:

$$p_{nohit} = p_{trg} \begin{cases} \text{GaussCDF}((S_{thr} - S_{\mu})/c_{snow}), & \sigma_{\mu} < \sigma_{em}/10 \\ \sum_{n=0}^3 \frac{\langle N_{\mu} \rangle^n}{n!} e^{-\langle N_{\mu} \rangle} \int_{S_{\mu}=0}^{S_{thr}-S_{em}c_{snow}} p_{\mu, sig}, & \sigma_{em} < \sigma_{\mu}/10 \\ \sum_{n=0}^3 \frac{\langle N_{\mu} \rangle^n}{n!} e^{-\langle N_{\mu} \rangle} \int_{S_{\mu}=0}^{S_{thr}} \int_{S_{em}=0}^{(S_{thr}-S_{\mu})/c_{snow}} p_{\mu, sig} p_{em}, & \text{else} . \end{cases} \quad (7)$$

The convolution is calculated within similar integration regions as in Eq. (6) and describes the probability of an observed tank signal to be below the threshold given a particular signal expectation from both the electromagnetic and muon contribution (as determined by the corresponding LDF). If both tanks in one station are not triggered (silent HLC), the likelihood is derived based on the squared probability p_{nohit}^2 :

$$ll_{sil} = \log(1 - p_{nohit}^n), \quad n = 2/1 \text{ if silent HLC/SLC} , \quad (8)$$

which corresponds to the not-hit probability of one station. The global likelihood is derived by summation of all parts including the parameter constraints:

$$llh = ll_{sig} + ll_{sat} + ll_{sil} - p_{constr} . \quad (9)$$

The reconstruction procedure is realized as 6-step log-likelihood minimization. Starting out with a standard 3-step IceTop reconstruction as described in [17] as a basis for the two-component fit. In the last three steps, the two-component fit is performed while fixing the geometrical reconstruction derived from the previous steps. Since the standard reconstruction only includes HLC hits, the underlying distribution as well as the resulting DLP fit are dominated by the electromagnetic contribution. Thus, the derived S_{125} [19] parameter can serve as an anchor point for the electromagnetic function and determines the initial values of $S_{em,125}$ and all LDF slope parameters. Initially, $S_{em,125}$ is fixed and only the low-energy muon number estimator $S_{\mu,600}$ along with the muon and electromagnetic slope parameters are adjusted while applying the constraints shown in Eq. (5) in order to find the proper parameter space. In the last two steps, all LDF parameters except for κ are fit successively. An example fit obtained after the last reconstruction step is illustrated in Fig. 5, showing the electromagnetic and muon LDF as well as the resulting total LDF $S_{tot}(r) = S_{em}(r) + S_{\mu}(r)$.

5. Reconstruction Performance

The two-component reconstruction is applied to a simulation set produced with the Sibyll 2.1 [20] hadronic interaction model, containing air showers of four different primaries (proton, helium, oxygen and iron) with energies of 100 TeV – 100 PeV. For the quantification of the reconstruction analysis, the application of different quality cuts is required. Firstly, a set of standard filters for the IceTop reconstruction discussed in [19] is applied. Most importantly, this includes a containment cut with a simple geometrical requirement for the reconstructed shower core position to be within the IceTop array. Secondly, three additional cuts specific to the

analysis presented in this work are applied. For a preliminary quantification of the reconstruction performance with the two component LDF only nearly vertical showers $\theta_{\text{reco}} < 18^\circ$ are taken into account. To ensure a sufficient fit quality, only reconstructed fits resulting in a reduced χ^2 of below 20 are included. This cut predominantly affects low energy events, as the fit stability increases with the event multiplicity. In the reconstruction only SLCs of ≥ 0.7 VEM are included in order to increase the probability of hits far from the shower axis to be of muonic origin [6]. To ensure a significant muonic contribution to the measured charge distribution, a final cut on the SLC multiplicity (≥ 3 SLCs) is applied. Above ~ 20 PeV these cuts become practically independent of the primary mass. The resulting distribution for $S_{\text{em},125}$ and $S_{\mu,600}$ is shown in Fig. 3, where $N_{\mu,\text{true}}$ includes all muons above a threshold of 210 MeV. Both distributions show a linear dependence to the corresponding true quantity, with a narrowing toward higher energies. The width of the distributions reflects the reconstruction resolution of the primary energy and low-energy muon number, respectively. In addition, the average of the distributions is shown for each simulated primary type separately. While $S_{\text{em},125}$ as a function of the true energy shows only a small primary mass dependence, $S_{\mu,600}$ manifests a significant splitting between light and heavy primaries for the same $\log_{10} N_{\mu,\text{true}}$. Although the mass dependence decreases with increasing muon number, it can result in a significant systematic shift for analyses utilizing $S_{\mu,600}$, and is therefore subject to further study. Fig. 4 shows the reconstruction performance, i.e., the bias and resolution for both estimators. The primary energy is reconstructed with a minimal dependence (few percent level) on the cosmic-ray mass. Above 10 PeV the energy resolution is below 10% and improves toward higher energies. Around 100 PeV saturation effects start to become important, and slightly worsen the reconstruction resolution for primary protons. The mass dependent bias of the low-energy muon number reconstruction is a direct reflection of the mass dependence of $S_{\mu,600}$. Towards higher $\log_{10} N_{\mu,\text{true}}$ (i.e., higher primary energy), the bias reduces from 20% to around 10%. The corresponding reconstruction resolution improves as a function of $\log_{10} N_{\mu,\text{true}}$ to below 20%.

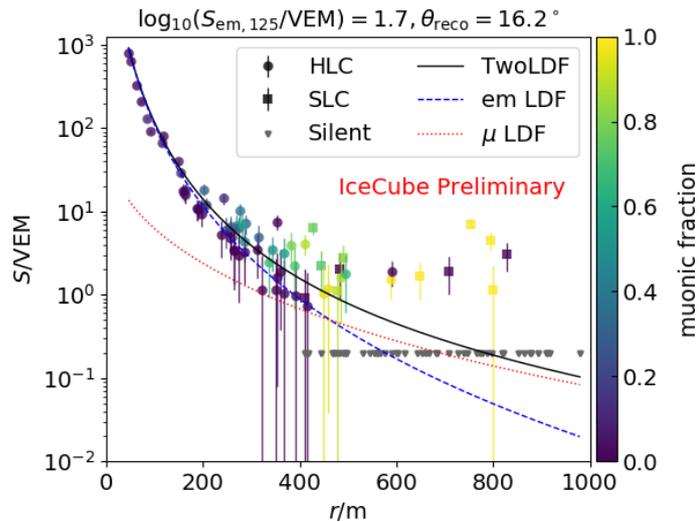


Figure 2: Two-component fit to the signal distribution of a simulated air-shower event with true primary energy of $\log_{10} E_{\text{true}} = 16.8$ eV. Silent detectors are visualized at a fixed value of 0.2 VEM. The color scale indicates the fraction of the measured signal produced by muons as derived from CORSIKA.

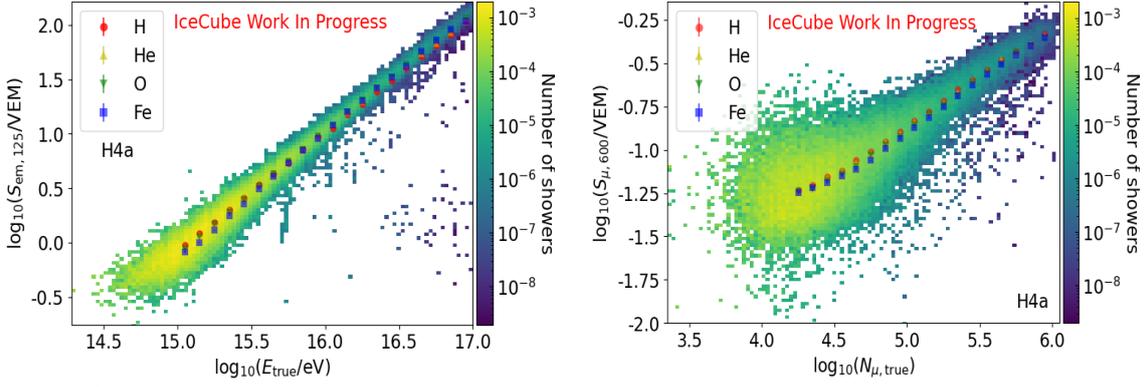


Figure 3: Distributions of the energy estimator $S_{em,125}$ (left) and the low-energy muon number estimator $S_{\mu,600}$ (right) for $\theta_{reco} < 18^\circ$ as a function of true primary energy and muon number, respectively, along with the mean values for each primary mass.

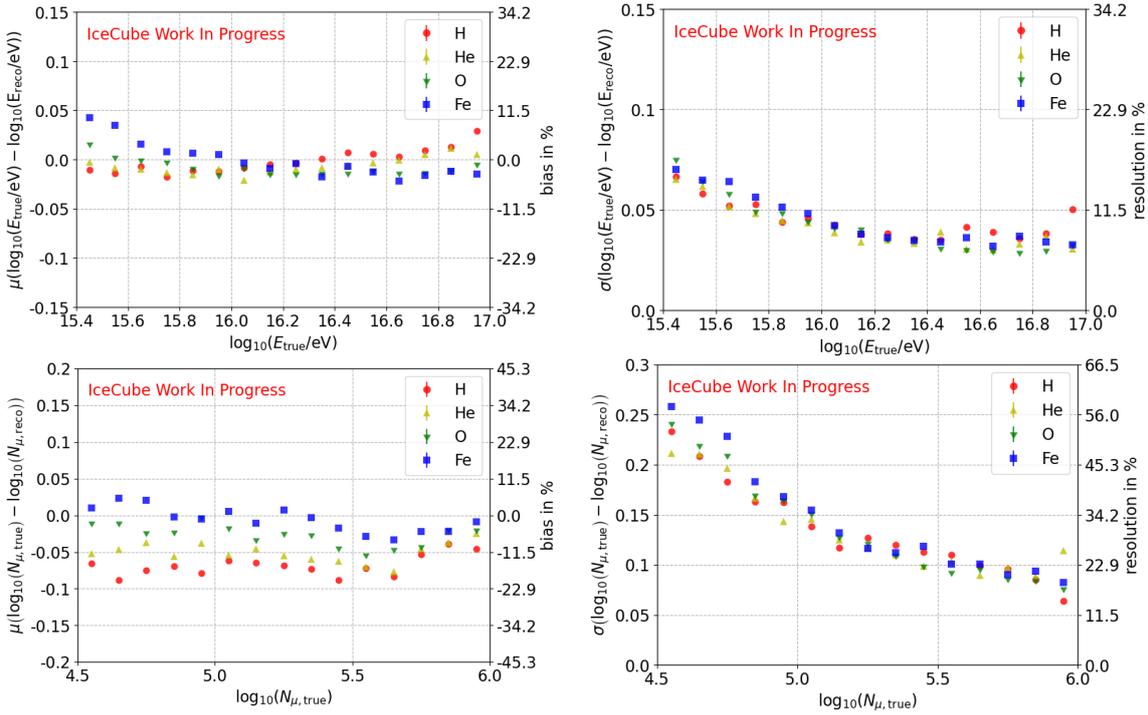


Figure 4: Reconstruction bias (left column) and resolution (right column) of the energy estimator $S_{em,125}$ (left) and the low-energy muon number estimator $S_{\mu,600}$ (right) as a function of true primary energy and muon number, respectively.

6. Conclusion

In this work a new method for a combined reconstruction of the primary energy and the low-energy muon number of cosmic-ray air showers on an event-by-event basis is presented. For this purpose, a separate description for the electromagnetic and muon LDFs is used in a combined likelihood approach in order to fit the signal footprint measured with IceTop. It is shown, that the parameters $S_{em,125}$ and $S_{\mu,600}$ of the electromagnetic and muonic LDF, respectively, can be used as a proxy for the primary energy and the low-energy muon number. While the former can be reconstructed with a minimal dependence on the mass of the cosmic ray, the low-energy muon number estimation can be further optimized in terms of reconstruction resolution as well as primary mass dependence. Several possible reasons for the mass dependence in $S_{\mu,600}$, such as a non-optimal choice of the reference distance as well as a persisting entanglement between the electromagnetic and muonic

part within the reconstruction procedure are currently under investigation. Further studies will include extended zenith and energy ranges as well as different levels of snow accumulation on top of the IceTop tanks. With increasing snow accumulation, the attenuation of the electromagnetic component becomes stronger, effectively increasing the potential of the IceTop tanks as muon detectors. This is of particular interest in the context of the IceTop enhancement [21] as well as the planned next generation detector IceCube-Gen2 [22]. An event-by-event based low-energy (\sim GeV) muon estimate is of high interest for tests of hadronic interaction models when combined with an estimation of the high-energy (above a few 100 GeV) muons of the same showers, allowing for a study of the corresponding correlations. Additionally, a separation of the electromagnetic and muonic shower component can be beneficial for mass composition analyses. Generally, any future study incorporating SLC information can potentially benefit from the application of the two-component LDF. In this context, the inclusion of IceTop uncontained events (shower core outside of IceTop) is of particular interest and will be part of future studies.

References

- [1] K.-H. Kampert and M. Unger *Astropart. Phys.* **35** (2012) 660–678.
- [2] J. Albrecht *et al.* *Astrophys. Space Sci.* **367** no. 3, (2022) 27.
- [3] **EAS-MSU, IceCube, KASCADE-Grande, NEVOD-DECOR, Pierre Auger, SUGAR, Telescope Array, Yakutsk EAS Array** Collaboration, H. P. Dembinski *et al.* *EPJ Web Conf.* **210** (2019) 02004.
- [4] **EAS-MSU, IceCube, KASCADE-Grande, NEVOD-DECOR, Pierre Auger, SUGAR, Telescope Array, Yakutsk EAS Array** Collaboration, D. Soldin *PoS ICRC2021* (2021) 349.
- [5] J. C. Arteaga-Velzquez *PoS ICRC2023* (these proceedings) 466.
- [6] **IceCube** Collaboration, R. Abbasi *et al.* *Phys. Rev. D* **106** no. 3, (2022) 032010.
- [7] **IceCube** Collaboration, S. Verpoest *et al.* *PoS ICRC2021* (2021) 357.
- [8] **IceCube** Collaboration, S. Verpoest *PoS ECRS* (2023) 074.
- [9] **IceCube** Collaboration, S. Verpoest *PoS ICRC2023* (these proceedings) 207.
- [10] A. Coleman *et al.* *Astropart. Phys.* **149** (2023) 102819.
- [11] D. Soldin *EPJ Web Conf.* **280** (2023) 04003.
- [12] **IceCube** Collaboration, R. Abbasi *et al.* *Nucl. Instrum. Meth. A* **700** (2013) 188–220.
- [13] **IceCube** Collaboration, M. G. Aartsen *et al.* *JINST* **12** no. 03, (2017) P03012.
- [14] **GEANT4** Collaboration, S. Agostinelli *et al.* *Nucl. Instrum. Meth. A* **506** (2003) 250–303.
- [15] K. G. Andeen, *First Measurements of Cosmic Ray Composition from 1-50 PeV using New Techniques on Coincident Data from the IceCube Neutrino Observatory*. PhD thesis, University of Wisconsin – Madison, 2011.
- [16] K. Greisen *Ann. Rev. Nucl. Part. Sci.* **10** (1960) 63–108.
- [17] **IceCube** Collaboration, A. Leszczyńska *et al.* *PoS ICRC2023* (these proceedings) 366.
- [18] **IceCube** Collaboration, K. Rawlins *PoS ICRC2023* (these proceedings) 337.
- [19] **IceCube** Collaboration, M. G. Aartsen *Phys. Rev. D* **88** (Aug, 2013) 042004.
- [20] E.-J. Ahn, R. Engel, T. K. Gaisser, P. Lipari, and T. Stanev *Phys. Rev. D* **80** (Nov, 2009) 094003.
- [21] **IceCube** Collaboration, A. Haungs *et al.* *PoS ICRC2021* (2021) 336.
- [22] **IceCube-Gen2** Collaboration, F. G. Schroeder *et al.* *PoS ICRC2021* (2021) 407.

Full Author List: IceCube Collaboration

R. Abbasi¹⁷, M. Ackermann⁶³, J. Adams¹⁸, S. K. Agarwalla^{40, 64}, J. A. Aguilar¹², M. Ahlers²², J.M. Alameddine²³, N. M. Amin⁴⁴, K. Andeen⁴², G. Anton²⁶, C. Argüelles¹⁴, Y. Ashida⁵³, S. Athanasiadou⁶³, S. N. Axani⁴⁴, X. Bai⁵⁰, A. Balagopal V.⁴⁰, M. Baricevic⁴⁰, S. W. Barwick³⁰, V. Basu⁴⁰, R. Bay⁸, J. J. Beatty^{20, 21}, J. Becker Tjus^{11, 65}, J. Beise⁶¹, C. Bellenghi²⁷, C. Benning¹, S. BenZvi⁵², D. Berley¹⁹, E. Bernardini⁴⁸, D. Z. Besson³⁶, E. Blaufuss¹⁹, S. Blot⁶³, F. Bontempo³¹, J. Y. Book¹⁴, C. Boscolo Meneguolo⁴⁸, S. Böser⁴¹, O. Botner⁶¹, J. Böttcher¹, E. Bourbeau²², J. Braun⁴⁰, B. Brinson⁶, J. Brostean-Kaiser⁶³, R. T. Burley², R. S. Busse⁴³, D. Butterfield⁴⁰, M. A. Campana⁴⁹, K. Carloni¹⁴, E. G. Carnie-Bronca², S. Chattopadhyay^{40, 64}, N. Chau¹², C. Chen⁶, Z. Chen⁵⁵, D. Chirkin⁴⁰, S. Choi⁵⁶, B. A. Clark¹⁹, L. Classen⁴³, A. Coleman⁶¹, G. H. Collin¹⁵, A. Connolly^{20, 21}, J. M. Conrad¹⁵, P. Coppin¹³, P. Correa¹³, D. F. Cowen^{59, 60}, P. Dave⁶, C. De Clercq¹³, J. J. DeLaunay⁵⁸, D. Delgado¹⁴, S. Deng¹, K. Deoskar⁵⁴, A. Desai⁴⁰, P. Desiati⁴⁰, K. D. de Vries¹³, G. de Wasseige³⁷, T. DeYoung²⁴, A. Diaz¹⁵, J. C. Díaz-Vélez⁴⁰, M. Dittmer⁴³, A. Domi²⁶, H. Dujmovic⁴⁰, M. A. DuVernois⁴⁰, T. Ehrhardt⁴¹, P. Eller²⁷, E. Ellinger⁶², S. El Mentawi¹, D. Elsässer²³, R. Engel^{31, 32}, H. Erpenbeck⁴⁰, J. Evans¹⁹, P. A. Evenson⁴⁴, K. L. Fan¹⁹, K. Fang⁴⁰, K. Farrag¹⁶, A. R. Fazely⁷, A. Fedynitch⁵⁷, N. Feigl¹⁰, S. Fiedlschuster²⁶, C. Finley⁵⁴, L. Fischer⁶³, D. Fox⁵⁹, A. Frankowski¹¹, A. Fritz⁴¹, P. Fürst¹, J. Gallagher³⁹, E. Ganster¹, A. Garcia¹⁴, L. Gerhardt⁹, A. Ghadimi⁵⁸, C. Glaser⁶¹, T. Glauch²⁷, T. Glüsenskamp^{26, 61}, N. Goehke³², J. G. Gonzalez⁴⁴, S. Goswami⁵⁸, D. Grant²⁴, S. J. Gray¹⁹, O. Gries¹, S. Griffin⁴⁰, S. Griswold⁵², K. M. Groth²², C. Günther¹, P. Gutjahr²³, C. Haack²⁶, A. Hallgren⁶¹, R. Halliday²⁴, L. Halve¹, F. Halzen⁴⁰, H. Hamdaoui⁵⁵, M. Ha Minh²⁷, S. K. Hanson⁴⁰, J. Hardin¹⁵, A. A. Harnisch²⁴, P. Hatch³³, A. Haungs³¹, K. Helbing⁶², J. Hellrung¹¹, F. Henningsen²⁷, L. Heuermann¹, N. Heyer⁶¹, S. Hickford⁶², A. Hidvegi⁵⁴, C. Hill¹⁶, G. C. Hill², K. D. Hoffman¹⁹, S. Hori⁴⁰, K. Hoshina^{40, 66}, W. Hou³¹, T. Huber³¹, K. Hultqvist⁵⁴, M. Hünnefeld²³, R. Hussain⁴⁰, K. Hyman²³, S. In⁵⁶, A. Ishihara¹⁶, M. Jacquart⁴⁰, O. Janik¹, M. Jansson⁵⁴, G. S. Japaridze⁵, M. Jeong⁵⁶, M. Jin¹⁴, B. J. P. Jones⁴, D. Kang³¹, W. Kang⁵⁶, X. Kang⁴⁹, A. Kappes⁴³, D. Kappesser⁴¹, L. Kardum²³, T. Karg⁶³, M. Karl²⁷, A. Karle⁴⁰, U. Katz²⁶, M. Kauer⁴⁰, J. L. Kelley⁴⁰, A. Khattee Zathul⁴⁰, A. Kheirandish^{34, 35}, J. Kiryuk⁵⁵, S. R. Klein^{8, 9}, A. Kochocki²⁴, R. Koirala⁴⁴, H. Kolanoski¹⁰, T. Kontrimas²⁷, L. Köpke⁴¹, C. Kopper²⁶, D. J. Koskinen²², P. Koundal³¹, M. Kovacevich⁴⁹, M. Kowalski^{10, 63}, T. Kozynets²², J. Krishnamoorthi^{40, 64}, K. Kruiswijk³⁷, E. Krupczak²⁴, A. Kumar⁶³, E. Kun¹¹, N. Kurahashi⁴⁹, N. Lad⁶³, C. Lagunas Gualda⁶³, M. Lamoureux³⁷, M. J. Larson¹⁹, S. Latseva¹, F. Lauber⁶², J. P. Lazar^{14, 40}, J. W. Lee⁵⁶, K. Leonard DeHoltan⁶⁰, A. Leszczyńska⁴⁴, M. Lincetto¹¹, Q. R. Liu⁴⁰, M. Liubarska²⁵, E. Lohfink⁴¹, C. Love⁴⁹, C. J. Lozano Mariscal⁴³, L. Lu⁴⁰, F. Lucarelli²⁸, W. Luszczak^{20, 21}, Y. Lyu^{8, 9}, J. Madsen⁴⁰, K. B. M. Mahn²⁴, Y. Makino⁴⁰, E. Manao²⁷, S. Mancina^{40, 48}, W. Marie Sainte⁴⁰, I. C. Mariş¹², S. Marka⁴⁶, Z. Marka⁴⁶, M. Marsee⁵⁸, I. Martinez-Soler¹⁴, R. Maruyama⁴⁵, F. Mayhew²⁴, T. McElroy²⁵, F. McNally³⁸, J. V. Mead²², K. Meagher⁴⁰, S. Mechal⁶³, A. Medina²¹, M. Meier¹⁶, Y. Merckx¹³, L. Merten¹¹, J. Micallef²⁴, J. Mitchell⁷, T. Montaruli²⁸, R. W. Moore²⁵, Y. Morii¹⁶, R. Morse⁴⁰, M. Moulai⁴⁰, T. Mukherjee³¹, R. Naab⁶³, R. Nagai¹⁶, M. Nakos⁴⁰, U. Naumann⁶², J. Necker⁶³, A. Negi⁴, M. Neumann⁴³, H. Niederhausen²⁴, M. U. Nisa²⁴, A. Noell¹, A. Novikov⁴⁴, S. C. Nowicki²⁴, A. Obertacke Pollmann¹⁶, V. O'Dell⁴⁰, M. Oehler³¹, B. Oeyen²⁹, A. Olivias¹⁹, R. Ørsøe²⁷, J. Osborn⁴⁰, E. O'Sullivan⁶¹, H. Pandya⁴⁴, N. Park³³, G. K. Parker⁴, E. N. Paudel⁴⁴, L. Paul^{42, 50}, C. Pérez de los Heros⁶¹, J. Peterson⁴⁰, S. Philippen¹, A. Pizzuto⁴⁰, M. Plum⁵⁰, A. Pontén⁶¹, Y. Popovych⁴¹, M. Prado Rodriguez⁴⁰, B. Pries²⁴, R. Procter-Murphy¹⁹, G. T. Przybylski⁹, C. Raab³⁷, J. Rack-Helleis⁴¹, K. Rawlins³, Z. Rechav⁴⁰, A. Rehman⁴⁴, P. Reichherzer¹¹, G. Renzi¹², E. Resconi²⁷, S. Reusch⁶³, W. Rhode²³, B. Riedel⁴⁰, A. Rifaie¹, E. J. Roberts², S. Robertson^{8, 9}, S. Rodan⁵⁶, G. Roellinghoff⁵⁶, M. Rongen²⁶, C. Rott^{53, 56}, T. Ruhe²³, L. Ruohan²⁷, D. Ryckbosch²⁹, I. Safa^{14, 40}, J. Saffer³², D. Salazar-Gallegos²⁴, P. Sampathkumar³¹, S. E. Sanchez Herrera²⁴, A. Sandrock⁶², M. Santander⁵⁸, S. Sarkar²⁵, S. Sarkar⁴⁷, J. Savelberg¹, P. Savina⁴⁰, M. Schaufel¹, H. Schieler³¹, S. Schindler²⁶, L. Schlickmann¹, B. Schlüter⁴³, F. Schlüter¹², N. Schmeisser⁶², T. Schmidt¹⁹, J. Schneider²⁶, F. G. Schröder^{31, 44}, L. Schumacher²⁶, G. Schwefer¹, S. Sclafani¹⁹, D. Seckel⁴⁴, M. Seikh³⁶, S. Seunarine⁵¹, R. Shah⁴⁹, A. Sharma⁶¹, S. Shefali³², N. Shimizu¹⁶, M. Silva⁴⁰, B. Skrzypek¹⁴, B. Smithers⁴, R. Snihur⁴⁰, J. Soedingrekso²³, A. Søgaard²², D. Soldin³², P. Soldin¹, G. Sommani¹¹, C. Spannfellner²⁷, G. M. Spiczak⁵¹, C. Spiering⁶³, M. Stamatikos²¹, T. Stanev⁴⁴, T. Stezelberger⁹, T. Stürwald⁶², T. Stuttard²², G. W. Sullivan¹⁹, I. Taboada⁶, S. Ter-Antonyan⁷, M. Thiesmeyer¹, W. G. Thompson¹⁴, J. Thwaites⁴⁰, S. Tilav⁴⁴, K. Tollefson²⁴, C. Tönnis⁵⁶, S. Toscano¹², D. Tosi⁴⁰, A. Trettin⁶³, C. F. Tung⁶, R. Turcotte³¹, J. P. Twagirayezu²⁴, B. Ty⁴⁰, M. A. Unland Elorrieta⁴³, A. K. Upadhyay^{40, 64}, K. Upshaw⁷, N. Valtonen-Mattila⁶¹, J. Vandenbroucke⁴⁰, N. van Eijndhoven¹³, D. Vannerom¹⁵, J. van Santen⁶³, J. Vara⁴³, J. Veitch-Michaelis⁴⁰, M. Venugopal³¹, M. Vereecken³⁷, S. Verpoest⁴⁴, D. Veske⁴⁶, A. Vijai¹⁹, C. Walck⁵⁴, C. Weaver²⁴, P. Weigel¹⁵, A. Weindl³¹, J. Weldert⁶⁰, C. Wendt⁴⁰, J. Werthebach²³, M. Weyrauch³¹, N. Whitehorn²⁴, C. H. Wiebusch¹, N. Willey²⁴, D. R. Williams⁵⁸, L. Witthaus²³, A. Wolf¹, M. Wolf²⁷, G. Wrede²⁶, X. W. Xu⁷, J. P. Yanez²⁵, E. Yildizci⁴⁰, S. Yoshida¹⁶, R. Young³⁶, F. Yu¹⁴, S. Yu²⁴, T. Yuan⁴⁰, Z. Zhang⁵⁵, P. Zheleznin¹⁴, M. Zimmerman⁴⁰

¹ III. Physikalisches Institut, RWTH Aachen University, D-52056 Aachen, Germany

² Department of Physics, University of Adelaide, Adelaide, 5005, Australia

³ Dept. of Physics and Astronomy, University of Alaska Anchorage, 3211 Providence Dr., Anchorage, AK 99508, USA

⁴ Dept. of Physics, University of Texas at Arlington, 502 Yates St., Science Hall Rm 108, Box 19059, Arlington, TX 76019, USA

⁵ CTSPS, Clark-Atlanta University, Atlanta, GA 30314, USA

⁶ School of Physics and Center for Relativistic Astrophysics, Georgia Institute of Technology, Atlanta, GA 30332, USA

⁷ Dept. of Physics, Southern University, Baton Rouge, LA 70813, USA

⁸ Dept. of Physics, University of California, Berkeley, CA 94720, USA

⁹ Lawrence Berkeley National Laboratory, Berkeley, CA 94720, USA

¹⁰ Institut für Physik, Humboldt-Universität zu Berlin, D-12489 Berlin, Germany

¹¹ Fakultät für Physik & Astronomie, Ruhr-Universität Bochum, D-44780 Bochum, Germany

¹² Université Libre de Bruxelles, Science Faculty CP230, B-1050 Brussels, Belgium

¹³ Vrije Universiteit Brussel (VUB), Dienst ELEM, B-1050 Brussels, Belgium

¹⁴ Department of Physics and Laboratory for Particle Physics and Cosmology, Harvard University, Cambridge, MA 02138, USA

- ¹⁵ Dept. of Physics, Massachusetts Institute of Technology, Cambridge, MA 02139, USA
¹⁶ Dept. of Physics and The International Center for Hadron Astrophysics, Chiba University, Chiba 263-8522, Japan
¹⁷ Department of Physics, Loyola University Chicago, Chicago, IL 60660, USA
¹⁸ Dept. of Physics and Astronomy, University of Canterbury, Private Bag 4800, Christchurch, New Zealand
¹⁹ Dept. of Physics, University of Maryland, College Park, MD 20742, USA
²⁰ Dept. of Astronomy, Ohio State University, Columbus, OH 43210, USA
²¹ Dept. of Physics and Center for Cosmology and Astro-Particle Physics, Ohio State University, Columbus, OH 43210, USA
²² Niels Bohr Institute, University of Copenhagen, DK-2100 Copenhagen, Denmark
²³ Dept. of Physics, TU Dortmund University, D-44221 Dortmund, Germany
²⁴ Dept. of Physics and Astronomy, Michigan State University, East Lansing, MI 48824, USA
²⁵ Dept. of Physics, University of Alberta, Edmonton, Alberta, Canada T6G 2E1
²⁶ Erlangen Centre for Astroparticle Physics, Friedrich-Alexander-Universität Erlangen-Nürnberg, D-91058 Erlangen, Germany
²⁷ Technical University of Munich, TUM School of Natural Sciences, Department of Physics, D-85748 Garching bei München, Germany
²⁸ Département de physique nucléaire et corpusculaire, Université de Genève, CH-1211 Genève, Switzerland
²⁹ Dept. of Physics and Astronomy, University of Gent, B-9000 Gent, Belgium
³⁰ Dept. of Physics and Astronomy, University of California, Irvine, CA 92697, USA
³¹ Karlsruhe Institute of Technology, Institute for Astroparticle Physics, D-76021 Karlsruhe, Germany
³² Karlsruhe Institute of Technology, Institute of Experimental Particle Physics, D-76021 Karlsruhe, Germany
³³ Dept. of Physics, Engineering Physics, and Astronomy, Queen's University, Kingston, ON K7L 3N6, Canada
³⁴ Department of Physics & Astronomy, University of Nevada, Las Vegas, NV, 89154, USA
³⁵ Nevada Center for Astrophysics, University of Nevada, Las Vegas, NV 89154, USA
³⁶ Dept. of Physics and Astronomy, University of Kansas, Lawrence, KS 66045, USA
³⁷ Centre for Cosmology, Particle Physics and Phenomenology - CP3, Université catholique de Louvain, Louvain-la-Neuve, Belgium
³⁸ Department of Physics, Mercer University, Macon, GA 31207-0001, USA
³⁹ Dept. of Astronomy, University of Wisconsin–Madison, Madison, WI 53706, USA
⁴⁰ Dept. of Physics and Wisconsin IceCube Particle Astrophysics Center, University of Wisconsin–Madison, Madison, WI 53706, USA
⁴¹ Institute of Physics, University of Mainz, Staudinger Weg 7, D-55099 Mainz, Germany
⁴² Department of Physics, Marquette University, Milwaukee, WI, 53201, USA
⁴³ Institut für Kernphysik, Westfälische Wilhelms-Universität Münster, D-48149 Münster, Germany
⁴⁴ Bartol Research Institute and Dept. of Physics and Astronomy, University of Delaware, Newark, DE 19716, USA
⁴⁵ Dept. of Physics, Yale University, New Haven, CT 06520, USA
⁴⁶ Columbia Astrophysics and Nevis Laboratories, Columbia University, New York, NY 10027, USA
⁴⁷ Dept. of Physics, University of Oxford, Parks Road, Oxford OX1 3PU, United Kingdom
⁴⁸ Dipartimento di Fisica e Astronomia Galileo Galilei, Università Degli Studi di Padova, 35122 Padova PD, Italy
⁴⁹ Dept. of Physics, Drexel University, 3141 Chestnut Street, Philadelphia, PA 19104, USA
⁵⁰ Physics Department, South Dakota School of Mines and Technology, Rapid City, SD 57701, USA
⁵¹ Dept. of Physics, University of Wisconsin, River Falls, WI 54022, USA
⁵² Dept. of Physics and Astronomy, University of Rochester, Rochester, NY 14627, USA
⁵³ Department of Physics and Astronomy, University of Utah, Salt Lake City, UT 84112, USA
⁵⁴ Oskar Klein Centre and Dept. of Physics, Stockholm University, SE-10691 Stockholm, Sweden
⁵⁵ Dept. of Physics and Astronomy, Stony Brook University, Stony Brook, NY 11794-3800, USA
⁵⁶ Dept. of Physics, Sungkyunkwan University, Suwon 16419, Korea
⁵⁷ Institute of Physics, Academia Sinica, Taipei, 11529, Taiwan
⁵⁸ Dept. of Physics and Astronomy, University of Alabama, Tuscaloosa, AL 35487, USA
⁵⁹ Dept. of Astronomy and Astrophysics, Pennsylvania State University, University Park, PA 16802, USA
⁶⁰ Dept. of Physics, Pennsylvania State University, University Park, PA 16802, USA
⁶¹ Dept. of Physics and Astronomy, Uppsala University, Box 516, S-75120 Uppsala, Sweden
⁶² Dept. of Physics, University of Wuppertal, D-42119 Wuppertal, Germany
⁶³ Deutsches Elektronen-Synchrotron DESY, Platanenallee 6, 15738 Zeuthen, Germany
⁶⁴ Institute of Physics, Sachivalaya Marg, Sainik School Post, Bhubaneswar 751005, India
⁶⁵ Department of Space, Earth and Environment, Chalmers University of Technology, 412 96 Gothenburg, Sweden
⁶⁶ Earthquake Research Institute, University of Tokyo, Bunkyo, Tokyo 113-0032, Japan

Acknowledgements

The authors gratefully acknowledge the support from the following agencies and institutions: USA – U.S. National Science Foundation-Office of Polar Programs, U.S. National Science Foundation-Physics Division, U.S. National Science Foundation-EPSCoR, Wisconsin Alumni Research Foundation, Center for High Throughput Computing (CHTC) at the University of Wisconsin–Madison, Open Science Grid (OSG), Advanced Cyberinfrastructure Coordination Ecosystem: Services & Support (ACCESS), Frontera computing project at the Texas Advanced Computing Center, U.S. Department of Energy-National Energy Research Scientific Computing

Center, Particle astrophysics research computing center at the University of Maryland, Institute for Cyber-Enabled Research at Michigan State University, and Astroparticle physics computational facility at Marquette University; Belgium – Funds for Scientific Research (FRS-FNRS and FWO), FWO Odysseus and Big Science programmes, and Belgian Federal Science Policy Office (Belspo); Germany – Bundesministerium für Bildung und Forschung (BMBF), Deutsche Forschungsgemeinschaft (DFG), Helmholtz Alliance for Astroparticle Physics (HAP), Initiative and Networking Fund of the Helmholtz Association, Deutsches Elektronen Synchrotron (DESY), and High Performance Computing cluster of the RWTH Aachen; Sweden – Swedish Research Council, Swedish Polar Research Secretariat, Swedish National Infrastructure for Computing (SNIC), and Knut and Alice Wallenberg Foundation; European Union – EGI Advanced Computing for research; Australia – Australian Research Council; Canada – Natural Sciences and Engineering Research Council of Canada, Calcul Québec, Compute Ontario, Canada Foundation for Innovation, WestGrid, and Compute Canada; Denmark – Villum Fonden, Carlsberg Foundation, and European Commission; New Zealand – Marsden Fund; Japan – Japan Society for Promotion of Science (JSPS) and Institute for Global Prominent Research (IGPR) of Chiba University; Korea – National Research Foundation of Korea (NRF); Switzerland – Swiss National Science Foundation (SNSF); United Kingdom – Department of Physics, University of Oxford.