



## **First polarization study of the M87 jet and active galactic nuclei at submillimeter wavelengths with ALMA**

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# First polarization study of the M87 jet and active galactic nuclei at submillimeter wavelengths with ALMA

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## ABSTRACT

**Aims.** We investigated the polarization and Faraday properties of Messier 87 (M87) and seven other radio-loud active galactic nuclei (AGNs) at  $\lambda 0.87$  mm (345 GHz) using the Atacama Large Millimeter/submillimeter Array (ALMA). Our goal was to characterize the linear polarization (LP) fractions, measure Faraday rotation measures (RMs), and examine the magnetic field structures in the emission regions of these AGNs.

**Methods.** We conducted full-polarization observations as part of the ALMA Band 7 very long baseline interferometry (VLBI) commissioning during the April 2021 Event Horizon Telescope (EHT) campaign. We analyzed the LP fractions and RMs to assess the nature of Faraday screens and magnetic fields in the submillimeter emission regions.

**Results.** We find LP fractions between 1% and 17% and RMs exceeding  $10^5$  rad  $m^{-2}$ , which are 1–2 orders of magnitude higher than typically observed at longer wavelengths ( $\lambda > 3$  mm). This suggests denser Faraday screens or stronger magnetic fields. Additionally, we present the first submillimeter polarized images of the M87 jet and the observed AGNs, revealing RM gradients and sign reversals in the M87 jet indicative of a kiloparsec-scale helical magnetic field structure.

**Conclusions.** Our results provide essential constraints for calibrating, analyzing, and interpreting VLBI data from the EHT at 345 GHz, representing a critical step toward submillimeter VLBI imaging.

**Key words.** black hole physics – magnetic fields – polarization – radiation mechanisms: non-thermal – instrumentation: interferometers – instrumentation: polarimeters

## 1. Introduction

The development of a phased-array capability for the Atacama Large Millimeter/submillimeter Array (ALMA) has revolutionized very long baseline interferometry (VLBI) at millimeter wavelengths (Matthews et al. 2018; Goddi et al. 2019a). By significantly enhancing the sensitivity of VLBI baselines, particularly at 230 GHz ( $\lambda \approx 1.3$  mm), the inclusion of phased-ALMA arrays into existing VLBI arrays made possible the first horizon-scale images of supermassive black holes, including M87\* in the Messier 87 Galaxy (Event Horizon Telescope Collaboration 2019a,b,c,d,e,f, 2021a,b, 2023, 2024a) and Sgr A\* at the heart of the Milky Way (Event Horizon Telescope Collaboration 2022a,b,c,d,e,f, 2024b).

In recent developments, ALMA’s phased array capabilities were extended to the submillimeter (ALMA Band 7;  $\nu \approx 345$  GHz;  $\lambda \approx 0.87$  mm), enabling a performance comparable to that of Band 6 ( $\nu \approx 230$  GHz) under suitable observing conditions (Crew et al. 2023). The first test of ALMA’s Band 7 phasing capability occurred in October 2018 during a global

VLBI campaign, marking the detection of the first VLBI fringes in the submillimeter regime between ALMA and three other Event Horizon Telescope (EHT) sites (Raymond et al. 2024). Following this success, a full end-to-end test of the Band 7 VLBI capability was conducted at the end of the ALMA-EHT science campaign in April 2021, with the aim of obtaining the first 345 GHz VLBI fringes toward the EHT key target M87\* and selected radio-loud active galactic nuclei (AGNs), as well as assessing the feasibility of VLBI imaging in the submillimeter regime.

Observations at  $\nu \approx 345$  GHz offer a 50% improvement in angular resolution over 230 GHz and are expected to enhance  $uv$  coverage through the combination with data from lower frequencies. Such a multifrequency synthesis will enable higher-fidelity imaging while minimizing interstellar scattering effects, which is particularly critical for imaging Sgr A\* (Event Horizon Telescope Collaboration 2022c).

To turn the ALMA array into a coherently phased aperture for millimeter VLBI and integrate it with other VLBI stations, the connected-element interferometric visibilities

must first be calibrated (Goddi et al. 2019b). This process inherently generates a full-polarization interferometric dataset as a by-product of VLBI observations with ALMA. For VLBI experiments, these interferometric datasets serve multiple purposes: they provide source-integrated parameters to refine and validate VLBI calibration and imaging workflows (Event Horizon Telescope Collaboration 2019b,c, 2021a, 2022b,c, 2024b) and provide critical observational constraints for theoretical models and physical interpretations (Event Horizon Telescope Collaboration 2021b, 2022d,e,f). However, beyond their role in VLBI, these datasets have significant standalone scientific value, enabling the derivation of millimeter emission, polarization, and Faraday properties of VLBI targets on arcsecond scales. At  $\lambda 0.87$  mm, synchrotron emission originates from a more optically thin region closer to the supermassive black hole when compared to  $\lambda > 1$  mm, providing a unique window into the jet base and accretion flow.

In this paper we present the analysis of connected-element interferometric data from the ALMA Band 7 VLBI test conducted in April 2021<sup>1</sup>. Results from the full 345 GHz VLBI campaign, including data from other stations, will be discussed in a follow-up paper (EHTC et al., in prep.; hereafter Paper II).

Section 2 details the observational setup (Sect. 2.1), observed targets (Sect. 2.2), calibration procedures (Sect. 2.3), and full-polarization image deconvolution (Sect. 2.4). Section 3 describes the data analysis, including the extraction of Stokes parameters, polarimetric and Faraday property estimation (Sect. 3.1), and polarized image production (Sect. 3.2). Section 4 presents the polarimetric properties (Sect. 4.1) and spectral indices (Sect. 4.2) of the observed AGN sources, with a focus on the submillimeter polarized emission in the M87 jet (Sect. 4.3). Section 5 summarizes our conclusions.

## 2. Observations, data calibration, and imaging

ALMA Band 7 observations were performed on April 19, 2021, at the end of the EHT science campaign conducted in Band 6. ALMA was operated as a phased array, and joined a global network of VLBI stations operating at this frequency for an end-to-end submillimeter VLBI commissioning observation.

Weather conditions at ALMA were excellent: typical opacity values were  $\tau_{225} \sim 0.05$ , while system temperatures ( $T_{\text{sys}}$ ) ranged from 95 to 276 K. The measured phasing efficiencies during the test varied between 81% and 97%, reflecting strong system performance under favorable conditions (Crew et al. 2023).

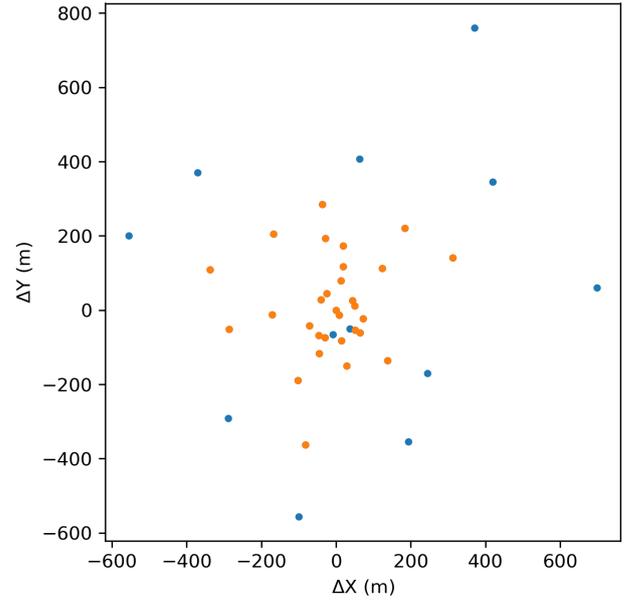
### 2.1. Observational setup

The observations at Band 7 spanned a continuous session of nearly 5 hours (from 01:16 to 06:08 UTC) and utilized 42 antennas. Of these, 31 were configured within a 400-meter radius around the reference antenna to form the phased array (baselines to 0.8 km). Eleven additional antennas (baselines to 1.3 km) completed the array and were used as un-phased comparison antennas for the determination of phasing efficiency estimates. The antenna locations on April 19 are plotted in Fig. 1.

The setup incorporated four spectral windows (SPWs) in dual linear polarization (LP); two in the lower sideband and two in the upper sideband, with central frequencies of 335.600, 337.5411<sup>2</sup>, 347.600, and 349.600 GHz. Each SPW offered an effective bandwidth of 1875 MHz. The ALMA interferomet-

<sup>1</sup> These ALMA commissioning data are made public with this publication.

<sup>2</sup> The Local Oscillator setup at ALMA does not allow evenly spaced basebands in Band 7, unlike in VLBI's Band 6.



**Fig. 1.** ALMA antenna locations for the phased array (orange points) and the un-phased comparison antennas (blue points) during the Band 7 observations on April 19, 2021. Positions are relative to the array reference antenna (Goddi et al. 2019b) and are plotted with positive values of X toward local east and positive values of Y toward local north.

ric data were averaged by a factor of 8 in frequency, resulting in 240 spectral channels per SPW with a channel spacing of 7.8125 MHz.

### 2.2. Observed targets

As these observations were designed as a comprehensive end-to-end commissioning test, the well-known quasar 3C279 was selected as the primary target. Given the novelty of VLBI at 345 GHz, there is currently limited guidance on suitable calibrators for this frequency. To address this, a selection of promising candidates was included for calibration and evaluation. This candidate sample comprises bright quasars and other compact extragalactic sources that had previously produced VLBI fringes in Band 6 observations. Notably, the sample includes M87\*, a planned future science target for 345 GHz VLBI observations. It should be noted that the observations include scans on three targets during gaps in the VLBI schedule. These scans were observed by ALMA for calibration purposes. In total, eight targets were observed in VLBI mode, while one target was observed exclusively by ALMA. A detailed list of the observed sources and their adopted calibration roles is provided in Table 1 (see also Sect. 2.3).

### 2.3. Data calibration

The ALMA data were calibrated using the Common Astronomy Software Applications (CASA) package, following the specialized “Quality Assurance Level 2” (QA2) procedures described in Goddi et al. (2019b, see also Crew et al. 2023). As detailed in those studies, the VLBI and non-VLBI scans are calibrated independently, resulting in separate calibration solutions.

3C273 was adopted as an absolute flux-density calibrator, a source routinely observed by ALMA as part of the flux-density

**Table 1.** VLBI sources observed on April 19, 2021, in Band 7.

Source		Coordinates		Calibration Intent <sup>a</sup>	T <sub>on source</sub> (min)
Common Name	J2000 Name	RA (J2000)	Dec (J2000)		
3C 279	J1256–0547	12:56:11.167	–05:47:21.52	Target 1	34.4
3C 273	J1229+0203	12:29:06.700	+02:03:08.60	Flux, Bandpass	25.8
M87	NGC 4486	12:30:49.412	+12:23:28.27	Target 2	16.9
4C 01.28	J1058+0133	10:58:29.604	+01:33:58.83	Polarization	10.5
S4 1144+40	J1146+3958	11:46:58.298	+39:58:34.28	...	7.5
PKS 1510–089	J1512–0905	15:12:50.533	–09:05:59.83	...	3.8
PKS 1335–127	J1337–1257	13:37:39.783	–12:57:24.70	...	3.8
non-VLBI scans <sup>b</sup>					
4C 01.28	J1058+0133	10:58:29.604	+01:33:58.83	Polarization	40.3
PKS 1243–072	J1246–0730	12:46:04.232	–07:30:46.575	Phase	24.7
3C 273	J1229+0203	12:56:11.167	–05:47:21.52	Flux, Bandpass	10.1

**Notes.** <sup>(a)</sup> In addition to the specified intent, amplitude and phase gains were estimated on each observed source. <sup>(b)</sup> These sources were observed by ALMA for calibration purposes during gaps in the VLBI schedules.

monitoring program with the ALMA Compact Array (ACA; see Appendix C). From the ALMA database, the flux density at 342 GHz was derived as  $S_{342\text{GHz}} = 4.164\text{Jy}$  with a spectral index  $\alpha = -0.79$ .

Accurate calibration of ALMA VLBI science data requires a full polarization calibration of interferometric visibilities. ALMA’s linearly polarized feeds simultaneously receive both orthogonal polarizations (X and Y), with all antennas in the phased array aligned to the reference antenna. Since the reference antenna’s phase is set to zero in both polarizations, a residual phase bandpass remains in the cross-hands of all baselines. Correcting this residual XY phase is essential for accurately combining cross- and parallel-hands to extract Stokes parameters. Thus, the primary calibration requirement for VLBI correlation is determining the X-Y phase difference and delay at the reference antenna (see Sect. 5 of Goddi et al. 2019b for details). This step enables the conversion of ALMA’s linearly polarized data into a circular polarization basis, ensuring consistency with other VLBI stations (Martí-Vidal et al. 2016; Goddi et al. 2019b). J1058+0133 (4C 01.28) was chosen as the polarization calibrator due to its adequate parallactic angle coverage and the presence of a compact, strongly polarized core (at the ~8% level). This selection enabled the simultaneous determination of the source polarization model and an estimate of the XY cross-phase at the reference antenna. The Stokes parameters derived from the polarization model were determined as  $IQUV = [1.50, 0.033, -0.126, 0.0]\text{Jy}$ .

#### 2.4. Full-Stokes imaging

All targets observed in Band 7 were imaged using the CASA task `tclean` in all Stokes parameters:  $I$ ,  $Q$ ,  $U$ , and  $V$ . A Briggs weighting scheme (Briggs 1995) was adopted with a robust parameter of 0.5 and a cleaning gain of 0.1.

A first cleaning step (100 iterations across all Stokes parameters) was performed within the inner  $4''$ . If significant emission ( $>7\sigma$ ) remained in the residual maps (e.g., for M87), an automatic script updated the cleaning mask, and a second, deeper cleaning (2000 iterations across all Stokes parameters) was conducted down to  $2\sigma$ . Both cleaning steps were run with the parameter `interactive=False`. A final interactive cleaning step (`interactive=True`) was performed to manually adjust the mask, capturing any real emission missed by the automatic

process and cleaning deeper sources with complex structures or high-residual signals.

The array configuration during phased-array observations provided synthesized beams ranging from  $[0':36-0':50] \times [0':29-0':34]$ , depending on the target. We produced  $576 \times 576$  pixel maps with a pixel size of  $0''.06$ , resulting in a field of view of  $34''.6 \times 34''.6$ , which comfortably covered the primary beam of ALMA Band 7 ( $18''$  at 350 GHz).

Maps were produced for individual SPWs and by combining SPWs in each sideband (SPW=0, 1 and SPW=2, 3) using `deconvolver='hogbom'` with `nterms=1`. Additionally, maps combining all four SPWs were created using `deconvolver='mtmfs'` with `nterms=2`, achieving better sensitivity and producing higher-quality images<sup>3</sup>. Thus, the combined SPW images were used for the imaging analysis presented in this paper, except in cases where a per-SPW analysis was necessary, such as for spectral index or Faraday rotation studies.

### 3. Data analysis

The data, calibrated and imaged as described in Sect. 2.1, were analyzed following the procedures outlined in Goddi et al. (2021). The analysis consisted of two main components:

- Extract the Stokes parameters in the compact cores of the observed targets and estimate their polarimetric and Faraday properties (Sect. 3.1).
- Produce polarized images for each target and determine the spatial distribution of the polarimetric quantities on arcsecond (kiloparsec) scales (Sect. 3.2).

#### 3.1. Polarization analysis of point sources

To extract flux values for Stokes  $I$ ,  $Q$ ,  $U$ , and  $V$ , we employed two alternative methods, one that utilizes the visibility data and the other full-Stokes images.

<sup>3</sup> The `deconvolver='mtmfs'` configuration outperformed `deconvolver='hogbom'` when combining all four SPWs, yielding 30–40% better sensitivity on average, as expected for  $\text{RMS} \propto 1/\sqrt{\Delta\nu}$ . However, `deconvolver='hogbom'` performed poorly when combining all four SPWs, especially for sources with steep spectral indices, resulting in up to 50% worse RMS compared to `deconvolver='mtmfs'`.

**Table 2.** Frequency-averaged polarization properties of AGN targets (at a representative frequency of 343 GHz).

Source	Day (2021)	$I^a$ (Jy)	Spectral Index ( $\alpha$ )	LP (%)	$\chi$ (deg)	$\chi_0$ (deg)	RM ( $10^5 \text{ rad m}^{-2}$ )
3C273	Apr 19	$4.17 \pm 0.42$	$-0.781 \pm 0.008$	$1.67 \pm 0.03$	$-64.19 \pm 0.52$	$-38.9 \pm 7.3$	$-5.8 \pm 1.7$
3C279	Apr 19	$8.73 \pm 0.87$	$-0.48 \pm 0.04$	$10.63 \pm 0.03$	$-46.711 \pm 0.081$	$-41.2 \pm 1.1$	$-1.26 \pm 0.26$
4C01.28	Apr 19	$1.50 \pm 0.15$	$-0.42 \pm 0.03$	$8.67 \pm 0.03$	$-37.72 \pm 0.11$	$-37.4 \pm 1.5$	$-0.08 \pm 0.35$
J1146+3958	Apr 19	$0.396 \pm 0.040$	$-0.4 \pm 0.2$	$2.66 \pm 0.07$	$87.27 \pm 0.77$	$107 \pm 11$	$-4.4 \pm 2.5$
J1337–1257	Apr 19	$3.13 \pm 0.31$	$-0.38 \pm 0.05$	$16.68 \pm 0.03$	$-7.310 \pm 0.055$	$-8.15 \pm 0.77$	$0.19 \pm 0.18$
J1512–0905	Apr 19	$1.74 \pm 0.17$	$-0.42 \pm 0.04$	$1.07 \pm 0.04$	$-65.72 \pm 0.95$	$-96 \pm 13$	$6.9 \pm 3.1$
M87	Apr 19	$0.970 \pm 0.097$	$-1.24 \pm 0.02$	$2.80 \pm 0.03$	$12.60 \pm 0.35$	$8.5 \pm 5.0$	$0.9 \pm 1.1$
non-VLBI scans							
3C273 <sup>b</sup>	Apr 19	$4.17 \pm 0.42$	$-0.68 \pm 0.01$	...	...	...	...
4C01.28	Apr 19	$1.50 \pm 0.15$	$-0.36 \pm 0.02$	$8.66 \pm 0.03$	$-37.68 \pm 0.11$	$-38.0 \pm 1.6$	$0.07 \pm 0.37$
J1246–0730	Apr 19	$0.280 \pm 0.028$	$-0.56 \pm 0.04$	$0.6 \pm 0.1$	$-17.0 \pm 6.6$	$-119 \pm 95$	$23 \pm 21$

**Notes.** <sup>(a)</sup> Errors in Stokes I include thermal and 10% absolute flux systematic uncertainty summed in quadrature. <sup>(b)</sup> The short observation on 3C273 did not result in sufficiently good data to enable a meaningful polarization analysis.

In the  $uv$ -plane analysis, we used the external CASA library UVMULTIFIT (Martí-Vidal et al. 2014). To optimize processing time, we first averaged all 240 frequency channels to produce eight-channel, four-SPW visibility  $uv$  files. Assuming the emission is dominated by a central point source at the phase center, we fit a delta function to the visibilities, extracting Stokes  $I$ ,  $Q$ ,  $U$ , and  $V$  parameters for each SPW individually.

For the image-based approach, we summed the central  $5 \times 5$  pixels of the CLEAN model component map. Summing only these central pixels isolates core emission from the surroundings, which is particularly important for sources with extended structure.

Goddi et al. (2021) performed a statistical comparison of flux-extraction methods in the  $uv$  and image planes for a sample of VLBI targets observed in Bands 3 and 6. They found that the median absolute deviation of the Stokes parameters between methods is  $<1\%$  for both point and extended sources. This agreement holds for the current Band 7 observations as well.

Using the measured Stokes parameters, we determined polarization properties for all targets, including the fractional LP ( $LP = \sqrt{Q^2 + U^2}/I$ ), the electric vector position angle (EVPA;  $2\chi = \arctan(U/Q)$ ), and its variation with frequency (Faraday rotation; see Sect. 3.1.2). Uncertainties in LP include the thermal errors of Stokes  $Q$  and  $U$  and a  $1\sigma$  systematic error (added in quadrature) associated with Stokes  $I$  leakage into Stokes  $Q$  and  $U$  (0.03% of Stokes  $I$ ). This analysis results in LP uncertainties  $<0.1\%$ , consistent with previous studies (Nagai et al. 2016; Bower et al. 2018; Goddi et al. 2021).

The polarization quantities, averaged across the four SPWs, are reported in Table 2. Table B.1 provides the polarimetric quantities for each SPW. We note that while the polarization parameters of 4C 01.28 from non-VLBI scans closely match those from VLBI scans (indicating consistent independent polarization calibrations), the short observation of 3C 273 did not yield data of sufficient quality for meaningful polarization analysis. Consequently, non-VLBI scans of 3C 273 were excluded, and the polarization analysis is based solely on VLBI scans.

### 3.1.1. Comparison with the AMAPOLA survey

For the purposes of absolute flux calibration, ALMA regularly monitors the flux density of bright sources (mainly blazars or quasi-stellar objects) distributed across the entire right ascension range (“the Grid”). These observations are conducted together

with solar system objects as part of the Grid Survey (GS) program, which operates on a cadence of approximately 10 days. The observations are executed with the Atacama Compact Array (ACA) in Bands 3, 6, and 7. Since the full-polarization mode is employed, it is possible to extract polarimetric information from the GS sources. This polarimetric analysis is performed using AMAPOLA<sup>4</sup>, a set of CASA-compatible Python scripts designed to reduce the full-Stokes polarimetry of GS observations. While the AMAPOLA values are primarily used for observation planning and the ACA and ALMA-VLBI arrays cover different  $uv$  ranges, comparing our data with the AMAPOLA database helps identify any systematic effects or clear inconsistencies due to variability within a week-long time frame.

The GS includes multiple measurements in Band 3 and Band 7 from April 2021 for all our targets, except M87. Our analysis shows that the polarimetric measurements are generally consistent with the historical trends reported for the Grid sources. Additional details, including comparison plots, are provided in Appendix C.

### 3.1.2. Rotation measure

Since we measured the EVPA at four distinct frequencies (one for each SPW) spanning a 16 GHz frequency range (334.6–350.6 GHz), we could estimate the Faraday rotation measure (RM) in the 0.87 mm band. Assuming that the Faraday rotation arises from a single, external, homogeneous Faraday screen (i.e., the rotation occurs outside the plasma responsible for the polarized emission), a linear dependence between the EVPA and the wavelength squared is expected.

We modeled this relationship by fitting the RM and the mean wavelength EVPA ( $\bar{\chi}$ ) using the standard linear relation:

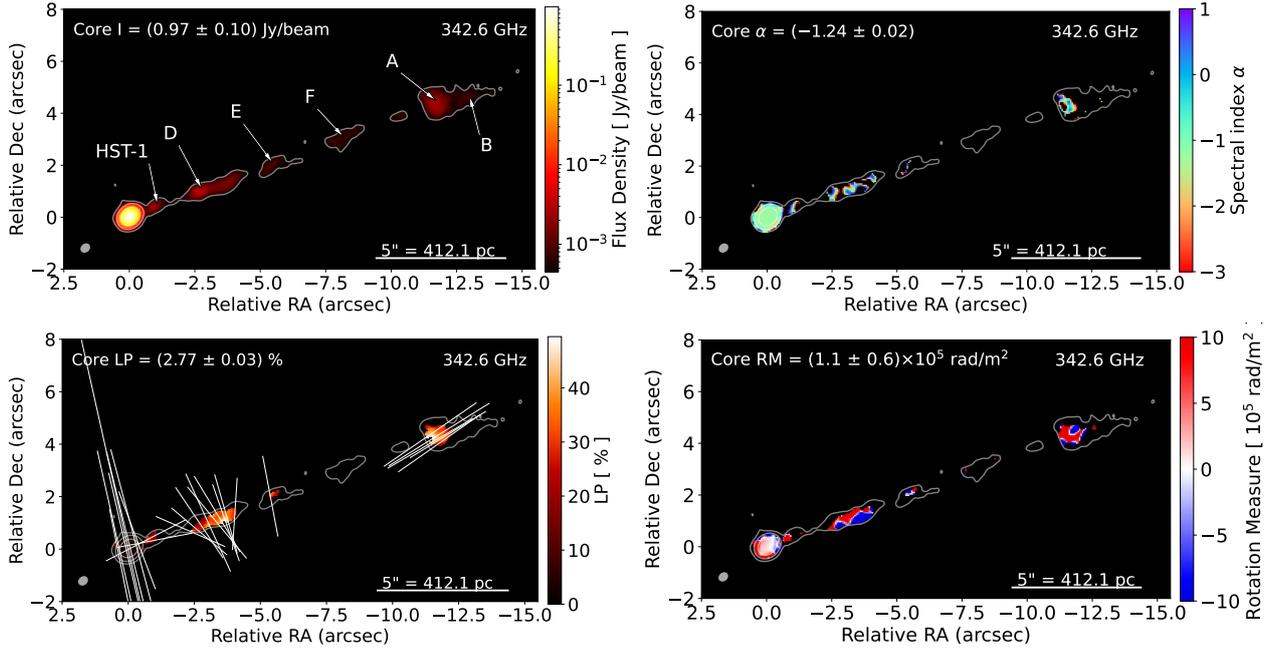
$$\chi = \bar{\chi} + \text{RM}(\lambda^2 - \bar{\lambda}^2), \quad (1)$$

where  $\chi$  is the observed EVPA at wavelength  $\lambda$ , and  $\bar{\chi}$  is the EVPA at the mean wavelength  $\bar{\lambda}$  (corresponding to the band center). Additionally, the EVPA extrapolated to zero wavelength (assuming the  $\lambda^2$  dependence holds) is given by

$$\chi_0 = \bar{\chi} - \text{RM}\bar{\lambda}^2. \quad (2)$$

The RM fitting is performed using a weighted least-squares method applied to  $\chi$  as a function of  $\lambda^2$ . The values of  $\bar{\chi}$ ,  $\chi_0$ ,

<sup>4</sup> <http://www.alma.cl/~skameno/AMAPOLA/>



**Fig. 2.** Polarization images of M87 at  $\lambda 0.87$  mm observed on April 19, 2021. The raster images in each panel cover an area of  $\approx 1.5 \times 0.8$  kpc and display the following: total intensity, spectral index, fractional LP, and Faraday RM (from the top left to bottom right). White vectors overlaid in the LP panel (bottom left) represent the orientation of the EVPAs, with vector lengths linearly proportional to the polarized intensity. In each panel, the white contour corresponds to the  $4\sigma_I$  level, where  $\sigma_I = 0.11$  mJy/beam is the RMS noise in the Stokes  $I$  map. The total intensity brightness is plotted using a logarithmic scale starting at the  $3\sigma$  level. For the spectral index map, we applied a threshold of  $5 \times \sigma$  in Stokes  $I$ . For the LP fraction and RM maps, thresholds are defined as  $3 \times \sigma_I$  for Stokes  $I$  and  $2 \times \sigma_{I_p}$  for the polarized flux density (here the  $\sigma_{I_p} = 0.08$  mJy/beam includes the thermal noise and the systematic error from Stokes  $I$  leakage into Stokes  $Q$  and  $U$ , combined in quadrature). The total intensity, spectral index, LP fraction, and RM values at the peak of the compact core are annotated in the upper-left corner of each panel. EVPAs are sampled every six pixels for clarity. The synthesized beam, represented as an ellipse in the lower-left corner of each panel, measures  $0''.40 \times 0''.32$  at a position angle of  $-48^\circ$ . Note that no primary beam correction is applied to these maps.

and the fitted RM are reported in the sixth, seventh, and eighth columns of Table 2. We determine EVPA uncertainties ranging from  $0.06^\circ$  to  $1^\circ$  (excluding J1246-0730, where  $LP < 1\%$ ), corresponding to RM propagated errors between 0.08 and  $3 \times 10^5$  rad  $m^{-2}$ . Despite these relatively high uncertainties and the limited frequency coverage at these high frequencies, we achieve  $3\sigma$  RM detections in 3C273 and 3C279. Details on the calculation of EVPA and RM uncertainties, as well as an evaluation of the robustness of RM fits derived from relatively narrow frequency coverage (16–18 GHz) at submillimeter wavelengths, are provided in Goddi et al. (2021).

### 3.2. Polarization images

We used the full-Stokes images (produced as described in Sect. 2.4) to determine the spatial distribution of the polarimetric quantities on arcsecond scales. Specifically, we executed custom Python scripts in CASA that take as input the Stokes  $I$ ,  $Q$ , and  $U$  images and output images of the linear polarized intensity, the fractional LP, the EVPA, and the Faraday RM. All these polarization quantities were calculated as described in Sects. 3.1 and 3.1.2 on a pixel-by-pixel basis after convolving all SPW images to the same synthesized beam. In generating the final images, we applied a threshold defined as  $3 \times \sigma$  (rms noise level) for Stokes  $I$  and  $2 \times \sigma$  for the polarized flux density<sup>5</sup>.

<sup>5</sup> The  $\sigma$  error on the polarized flux includes the thermal error (residual image rms noise) and the systematic error (added in quadrature) associated with Stokes  $I$  leakage onto Stokes  $Q$  and  $U$  (0.03% of Stokes  $I$ ).

Representative images are presented in Fig. 2, showcasing M87, and Fig. 3, featuring 3C 279, 3C 273, and 4C 01.28<sup>6</sup>. The raster image in each panel displays the total intensity, spectral index (for M87), LP fraction, and RM. White vectors overlaid on these images represent the orientation of the EVPAs, with their lengths linearly proportional to the polarized flux. It should be noted that these EVPAs are not corrected for Faraday rotation and that the magnetic field vectors should be rotated by  $90^\circ$ .

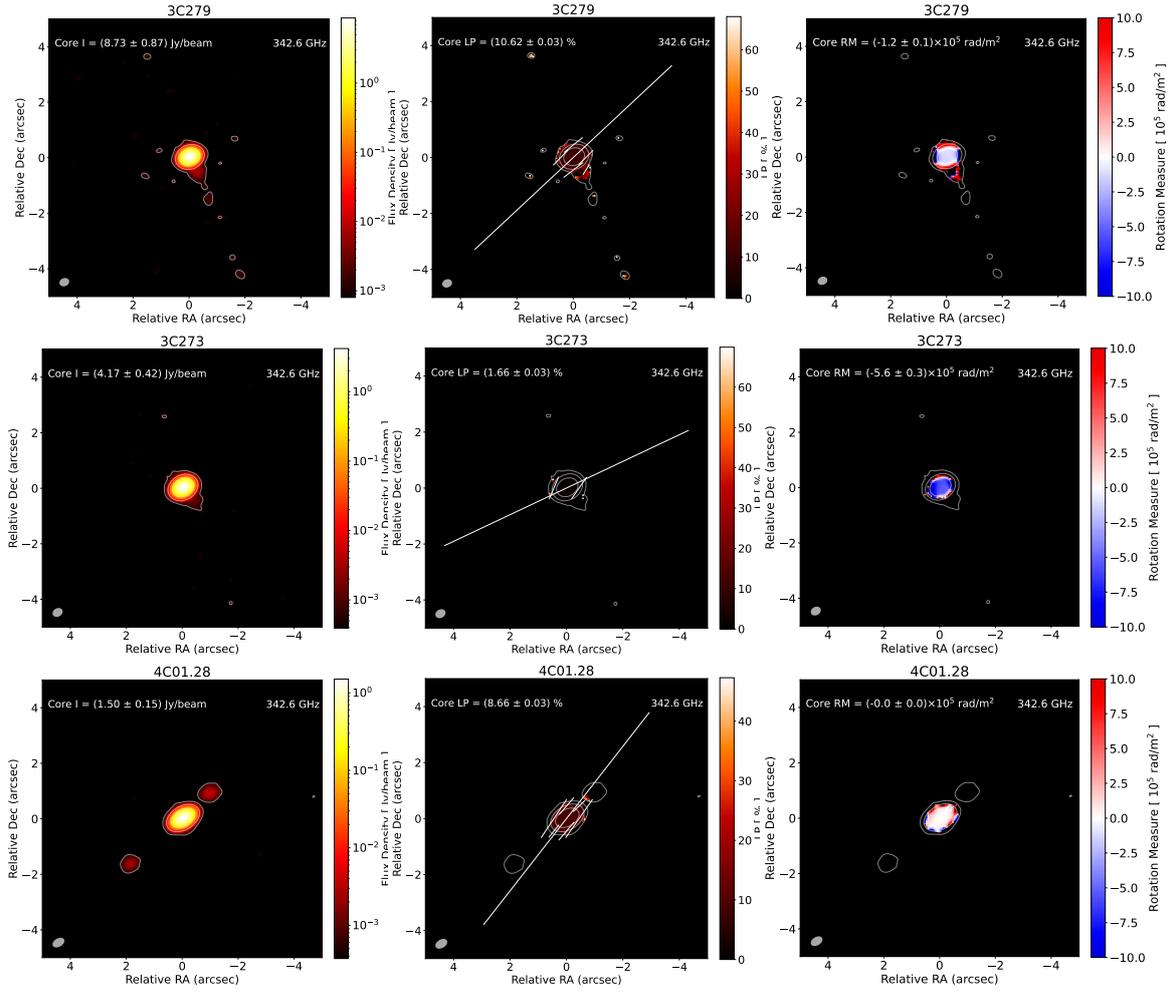
## 4. Results and discussion

We derived polarimetric properties and produced polarized images of eight AGNs observed in full-polarization mode with ALMA in the 0.87 mm band for the first time. In Sect. 4.1 we analyze the polarization characteristics of the AGN; in Sect. 4.2 we discuss their spectral indices; and in Sect. 4.3 we focus on the polarized submillimeter emission from the kiloparsec-scale jet in M87.

### 4.1. AGN polarization properties

The polarimetric quantities for our AGN targets, derived as described in Sect. 3.1, are summarized in Table 2. The table also includes the measured LP fractions, EVPAs, and RMs, which can provide key insights into the magnetic field structures and plasma conditions within the AGN jets. The LP fractions in the

<sup>6</sup> Images for other targets are not included, as they remain unresolved on arcsecond scales and do not provide additional information beyond what is already summarized in Table 2.



**Fig. 3.** Polarization images of selected AGNs observed with ALMA at 0.87 mm on April 19, 2021 (see Fig. 2 for a description of the plotted quantities). The synthesized beams (represented as an ellipse in the lower-left corner of each panel) have the following sizes (and position angles):  $0'':36 \times 0'':29$  ( $-67.5^\circ$ ) for 3C279,  $0'':38 \times 0'':30$  ( $-60.8^\circ$ ) for 3C273, and  $0'':46 \times 0'':30$  ( $-58.7^\circ$ ) for 4C01.28. Note that the EVPAs are not Faraday-corrected and that the magnetic field vectors should be rotated by  $90^\circ$ , ignoring Lorentz transformation and light aberration.

AGN central cores span a wide range, from  $\lesssim 1\%$  for weakly polarized targets (e.g., PKS1243-072 and PKS1510-089) to 10–17% for strongly polarized sources like 3C279 and PKS1335-127, consistent with previous measurements (Appendix C).

This is the first time Faraday RMs have been measured in the submillimeter. In 3C273, we measure  $RM = (-5.8 \pm 1.6) \times 10^5 \text{ rad m}^{-2}$  at 0.87 mm, which is consistent with previous ALMA observations at 1.3 mm (December 2016:  $RM = (5.0 \pm 0.3) \times 10^5 \text{ rad m}^{-2}$ ; April 2017:  $RM = (2.5 \pm 0.3) \times 10^5 \text{ rad m}^{-2}$ ; Hovatta et al. 2019; Goddi et al. 2021) and higher than 3 mm observations ( $RM = (-0.60 \pm 0.14) \times 10^4 \text{ rad m}^{-2}$ ; Goddi et al. 2021). Interestingly, the sign of the RM at 0.87 mm differs from previous 1.3 mm observations, which we interpret as a result of time variability rather than a frequency-dependent effect (Carlos et al., in prep.).

In 3C279, earlier measurements showed RM values ranging from 1800 to 2700  $\text{rad m}^{-2}$  at 3.5 mm (e.g., Lee et al. 2015; Goddi et al. 2021) and an upper limit of 5000  $\text{rad m}^{-2}$  at 1.3 mm (Goddi et al. 2021). Our new observations suggest a significant increase in RM at shorter wavelengths, supporting the idea that higher frequencies probe the innermost regions with stronger magnetic fields and denser plasma. Given the time variability of these sources, simultaneous multi-band observa-

tions are required to confirm this trend and establish the dependence of RM on wavelength. Some targets, including M87, 3C279, 3C273, and 4C 01.28, were observed at lower frequencies (Bands 3 and 6;  $\nu \approx 86 \text{ GHz}$  and  $230 \text{ GHz}$ , respectively) during the same week. A comparative analysis of polarization properties across ALMA bands is planned for a future study (Carlos et al., in prep.).

Our findings of RM exceeding  $10^5 \text{ rad m}^{-2}$  in AGN cores at  $\lambda \sim 0.87 \text{ mm}$  align with previous studies at millimeter wavelengths (e.g., Plambeck et al. 2014; Martí-Vidal et al. 2015; Hovatta et al. 2019; Goddi et al. 2021). These high RM values, which are 1–2 orders of magnitude greater than those typically reported for AGNs at  $\lambda > 3 \text{ mm}$  (e.g., Gabuzda et al. 2017; Peng et al. 2024), point to a denser Faraday screen or stronger magnetic fields in the submillimeter emission region.

#### 4.2. AGN spectral indices

In addition to polarization parameters, we derived the total intensity spectral index  $\alpha$  for all sources, where  $\alpha$  is defined such that  $I(\nu) \propto \nu^\alpha$ . The spectral index was computed “in-band” using a weighted least-squares fit across the four flux-density measurements in each SPW. The AGN cores systematically

show negative spectral indices in the range  $\alpha = -1.3$  to  $-0.4$ , consistent with previous findings at millimeter wavelengths (e.g., Agudo et al. 2018; Goddi et al. 2021). This contrasts with the flat spectral indices ( $\alpha \approx 0$ ) typically observed at longer centimeter wavelengths, further supporting the idea that AGN cores become progressively more optically thin at shorter wavelengths.

For M87, the compact core exhibits a spectral index of  $\alpha = -1.25$ , consistent with previous measurements at 1.3 mm (Goddi et al. 2021) but contrasting with the flatter spectra observed at 3 mm (e.g., Doi et al. 2013) and at centimeter wavelengths (e.g., Kravchenko et al. 2020). This steep spectral index suggests a spectral break between 3 mm and 1.3 mm, transitioning to a consistent power law from 1.3 mm to 0.87 mm. Such a break is likely due to the inclusion of contributions from both the compact core and the inner jet within the ALMA beam. While the compact (VLBI) core typically displays a flat spectrum, the jet component has a steeper spectral index, which dominates at higher frequencies due to decreased opacity at the jet base, as predicted by the standard jet model (Blandford & Königl 1979).

#### 4.3. Polarization properties of the M87 jet at 345 GHz

High-resolution polarization imaging of the relativistic jet in M87 at millimeter wavelengths has been achieved through ALMA observations at  $\lambda 3$  mm with a resolution  $\sim 2.5''$  (Peng et al. 2024) and  $\lambda 1.3$  mm with a resolution  $\sim 1''$  (Goddi et al. 2021). These studies revealed the narrow, straight kiloparsec-scale jet extending over  $\sim 25''$  from the nucleus, including several prominent knots (D, F, A, B, C) previously identified in optical and radio images. At  $\lambda 3$  mm, the jet-inflated radio lobes are also visible, with features imaged in greater detail at lower frequencies (e.g., the NRAO 20 cm Very Large Array image<sup>7</sup>).

Our  $\lambda 0.87$  mm ALMA observations show a similar structure to the  $\lambda 1.3$  mm image (e.g., see Fig. 2 in Goddi et al. 2021) but with improved angular resolution ( $\sim 0.3''$ ), allowing us to resolve HST-1 from the core. HST-1 is a bright, knot-like feature discovered with the *Hubble* Space Telescope and located approximately 0.85 arcseconds (about 60–70 parsecs) downstream from the central black hole (Biretta et al. 1999). HST-1 has exhibited remarkable properties over the years, including rapid variability, superluminal motion, and significant flaring activity (Cheung et al. 2007; Giroletti et al. 2012), making it a critical site for understanding particle acceleration, jet collimation, and magnetic field dynamics in the M87 jet and in AGNs in general.

The radio core dominates the Stokes  $I$  emission with a peak brightness of  $\sim 1$  Jy beam<sup>-1</sup> (Fig. 2, top left). In contrast, the jet knots become more prominent in the linearly polarized intensity, relative to the total intensity, as the fractional polarization increases from the core outward. The LP image (Fig. 2, bottom left) shows the lowest fractional polarization ( $\lesssim 3\%$ ) at the core, which rises to  $\sim 20\%$  toward HST-1 and peaks at  $\sim 55\%$ ,  $42\%$ , and  $40\%$  in between knots D, E, and A, respectively. The high degree of polarization observed in the knots is indicative of a well-ordered magnetic field structure within these regions, likely resulting from shock compression or shear flows in the jet plasma (e.g., Laing 1980).

The EVPA distribution observed at  $\lambda 0.87$  mm closely matches prior results at  $\lambda 1.3$  mm (Goddi et al. 2021) and  $\lambda 3$  mm (Peng et al. 2024) and agrees with centimeter-wave polarization measurements from the Very Large Array (VLA; Algaba et al. 2016; Pasetto et al. 2021). This consistency across multiple

epochs suggests a stable magnetic field configuration. The EVPA distribution is generally perpendicular to the jet axis, except in the regions HST-1 and Knot A. Ignoring Lorentz transformation and light aberration, rotating the EVPA by  $90^\circ$  (without Faraday correction) indicates that the magnetic field is mostly parallel to the jet axis, except in HST-1 and Knot A, where it becomes nearly perpendicular. These deviations are likely caused by recollimation or standing shocks, which can alter the helicity of the magnetic field locally due to variations in the radial profiles of the poloidal and toroidal components (Mizuno et al. 2015).

The RM, derived from EVPA measurements between 334.6 GHz and 350.6 GHz (Fig. 2, bottom right), reveals both gradients and sign reversals along the jet. Near the core, the RM exhibits an east-to-west gradient, ranging from  $(2.5 \pm 3.9) \times 10^5$  rad m<sup>-2</sup> at  $0.3''$  east to  $(-1.2 \pm 0.5) \times 10^5$  rad m<sup>-2</sup> at  $0.3''$  west. Downstream in HST-1, the RM reaches  $(44 \pm 18) \times 10^5$  rad m<sup>-2</sup> at  $0.96''$  from the core, decreasing with distance. Knot D exhibits RM values ranging from  $(47 \pm 23) \times 10^5$  rad m<sup>-2</sup> to  $(-29 \pm 11) \times 10^5$  rad m<sup>-2</sup>, while knot A displays an RM of  $(20.3 \pm 4.1) \times 10^5$  rad m<sup>-2</sup> that varies significantly across the region.

The observed reversals in RM sign along the jet indicate changes in the line-of-sight magnetic field direction, while the RM gradients across the jet width reveal oppositely directed line-of-sight magnetic fields at the jet edges. These results are consistent with previous lower-frequency studies, which reported similar RM gradients observed with ALMA at 3 mm (Peng et al. 2024) and the VLA at 1.7–7.5 cm (Pasetto et al. 2021). Such RM gradients and sign reversals are evidence of a helical magnetic field threading the jet, potentially persisting up to kiloparsec scales (Pasetto et al. 2021). This configuration aligns with theoretical predictions of dynamically significant poloidal magnetic fields being twisted into a helix by the rotation of the black hole-accretion disk system (Tchekhovskoy et al. 2011), and possibly with the “cosmic battery” model (e.g., Myserlis & Contopoulos 2021; Contopoulos et al. 2022). Independent support for this interpretation is provided by EHT observations of polarized emission near the M87 black hole (Event Horizon Telescope Collaboration 2021b), further corroborating the presence of a helical magnetic field structure.

A helical magnetic field could potentially also explain another feature observed in the M87 core, its high RM variability (Goddi et al. 2021). Several scenarios may account for this variability, including turbulence in the accretion flow causing internal Faraday rotation, a dynamically changing external Faraday screen, or a rapidly varying source at horizon scales with a static external screen. Alternatively, a helical magnetic field could introduce RM variability through beam-averaging effects. Variations in beam size across different observations may sample regions with oppositely directed magnetic fields along the line of sight, leading to distinct RM measurements.

Our submillimeter observations, however, are limited by angular resolution, sensitivity, and frequency coverage, preventing a definitive differentiation between internal and external Faraday rotation and a precise characterization of the helical magnetic field in the M87 jet. While spatial RM variations along the jet axis are evident, the current imaging sensitivity is insufficient to continuously recover the polarized emission structure on kiloparsec scales. This limitation prevents us from confirming whether consistent RM and LP transverse gradients persist along the full extent of the jet. Moreover, the limited resolution impedes the separation of emissions originating from the jet’s edges and central axis, leaving open the possibility that the observed RM gradients are due to external material rather than the jet’s intrinsic magnetic field. Simultaneous observations with

<sup>7</sup> <https://www.nrao.edu/archives/items/show/33382>

the EHT (which will be discussed in Paper II) are expected to shed light on the properties of the Faraday rotation medium in the core region and provide insights into the magnetic field structure at the base of the jet.

The observed time variability in RM (Goddi et al. 2021) adds another layer of complexity, as it precludes using nonsimultaneous datasets to reliably confirm RM-wavelength dependences. Addressing these challenges requires simultaneous, beam-matched, and multifrequency ALMA observations. To this end, we plan to analyze a comprehensive dataset of ALMA observations, covering multiple frequency bands (Bands 3, 6, and 7) and spanning several years, with data obtained on multiple days. This systematic approach will enable a detailed investigation of time- and frequency-dependent effects and thus a more robust determination of the RM's frequency dependence. Such an analysis will provide critical insights into whether the Faraday rotation is internal or external and refine our understanding of the helical magnetic field structure in the M87 jet. The findings from this extended analysis will be presented in an upcoming publication.

## 5. Summary and conclusions

We have presented the first submillimeter full-polarization study of radio-loud AGNs with ALMA, analyzing their polarization and Faraday properties. We find LP fractions ranging from 1% to 17% and RMs exceeding  $10^5 \text{ rad m}^{-2}$ , consistent with earlier studies at millimeter wavelengths (e.g., Plambeck et al. 2014; Martí-Vidal et al. 2015; Hovatta et al. 2019; Goddi et al. 2021). These RM values are 1–2 orders of magnitude higher than those observed in AGNs at  $\lambda > 3 \text{ mm}$  (e.g., Gabuzda et al. 2017; Peng et al. 2024), indicating a denser Faraday screen or stronger magnetic fields in the submillimeter emission regions.

We produced the highest-frequency polarized images ever of these AGNs, which included M87 and its jet. For M87, we observe RM gradients and sign reversals both along and across the jet axis, potentially reflecting reversals in the magnetic field direction relative to the line of sight. If confirmed, this would support a helical magnetic field configuration on kiloparsec scales, as suggested by lower-frequency VLA studies (e.g., Pasetto et al. 2021). In future work we will analyze multifrequency, multi-epoch ALMA data and explore the time- and frequency-dependent properties of Faraday rotation to better constrain the magnetic field structure in M87.

The ALMA data were obtained in the 0.87 mm band during VLBI commissioning tests conducted in collaboration with the EHT Collaboration. These data provided essential calibration and interpretation for simultaneous VLBI observations with the EHT. Notably, phased ALMA observations on April 19, 2021, enabled the first detection of VLBI fringes at 345 GHz for M87 and selected AGNs (Matthews & Crew 2024, Paper II). This milestone represents a key step in evaluating VLBI imaging feasibility in the submillimeter band.

Very long baseline interferometry observations at this frequency offer a 50% improvement in resolution compared to 230 GHz and substantially reduce interstellar scattering, which is especially relevant for imaging Sgr A\*. The combination of 230 GHz and 345 GHz data enhances  $uv$  coverage, enabling high-fidelity imaging with multifrequency synthesis. These advancements will open new pathways for studying black hole shadows in M87 and Sgr A\*, as well as accretion and jet formation in nearby AGNs. Reduced opacity at 345 GHz allows jet-launching regions closer to the black hole to be probed, offering critical insights into jet formation, collimation, acceleration, and the phenomenon of limb-brightening in inner jets, and thus deepening our understanding of AGN physics.

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## Appendix B: Stokes parameters per ALMA frequency band (SPW)

Table 2 in the main text reports the polarization quantities averaged across the four SPWs. Here we report the polarimetric quantities (Stokes  $IQUV$ , LP, and EVPA) for each SPW in Table B.1. The quoted uncertainties include the thermal error, the  $1\sigma$  systematic error associated with Stokes  $I$  leakage into Stokes  $Q$ ,  $U$  (0.03% of Stokes  $I$ ), and Stokes  $V$  (0.6% of Stokes  $I$ ), as recommended by the ALMA observatory (see ALMA Technical Handbook; Remijan et al. 2019). The total uncertainties are computed by combining these contributions in quadrature. The uncertainty in LP is primarily dominated by systematic errors, except for the weakest sources, where thermal noise becomes significant. Unlike Goddi et al. (2021), we did not apply an LP bias correction because our sample does not include low-polarization sources with  $LP \ll 1\%$ . We note that Stokes  $V$  is assumed to be zero during polarization calibration (see Goddi et al. 2019b), and no additional calibration step was applied to reliably constrain circular polarization (CP) in this experiment. Consequently, we do not claim CP detections for the observed sources. For an assessment of the reliability of Stokes  $V$  detections in interferometric observations using linearly polarized feeds, we refer the reader to Appendix G in Goddi et al. (2021).

## Appendix C: Comparison of Stokes parameters with the AMAPOLA polarimetric Grid Survey

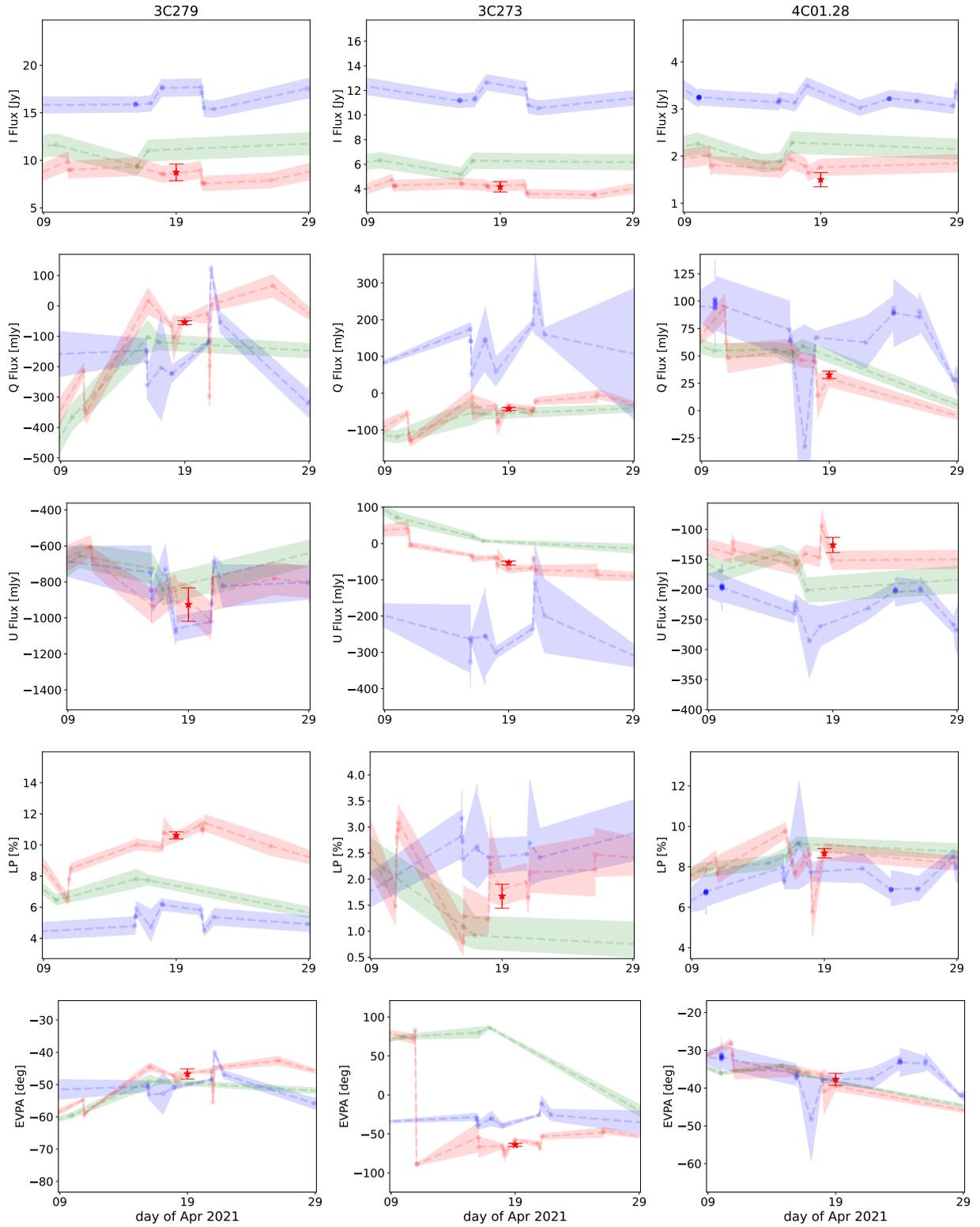
Figures C.1 and C.2 illustrate the polarimetric parameters reported in Table B.1 (green data points and error bars), specifically Stokes  $I$ ,  $Q$ ,  $U$ , LP, and EVPA. The shaded  $\pm 1\sigma$  regions represent the time-variance of the same parameters as measured by AMAPOLA over a 20-day period surrounding the VLBI observations (April 9 to 29, 2021). Inflections in the shaded trend lines indicate the times of GS observations. Blue corresponds to Band 3 measurements, while green and red corresponds to Bands 6 and 7, respectively. The figures show that most Band 7 measurements derived from our current study fall within the red-shaded regions, demonstrating general consistency with the AMAPOLA trends. The observed discrepancies may align with inter-GS cadence variability or differential time-variability between frequency bands, as also seen in some AMAPOLA monitoring cases.

We conclude that, despite differences in array configurations and data reduction methods, the ALMA-VLBI results are consistent with AMAPOLA. This consistency validates the accuracy of flux density and polarization calibration in VLBI mode for Band 7, extending the reliability previously demonstrated in Bands 3 and 6 (Goddi et al. 2019b, 2021).

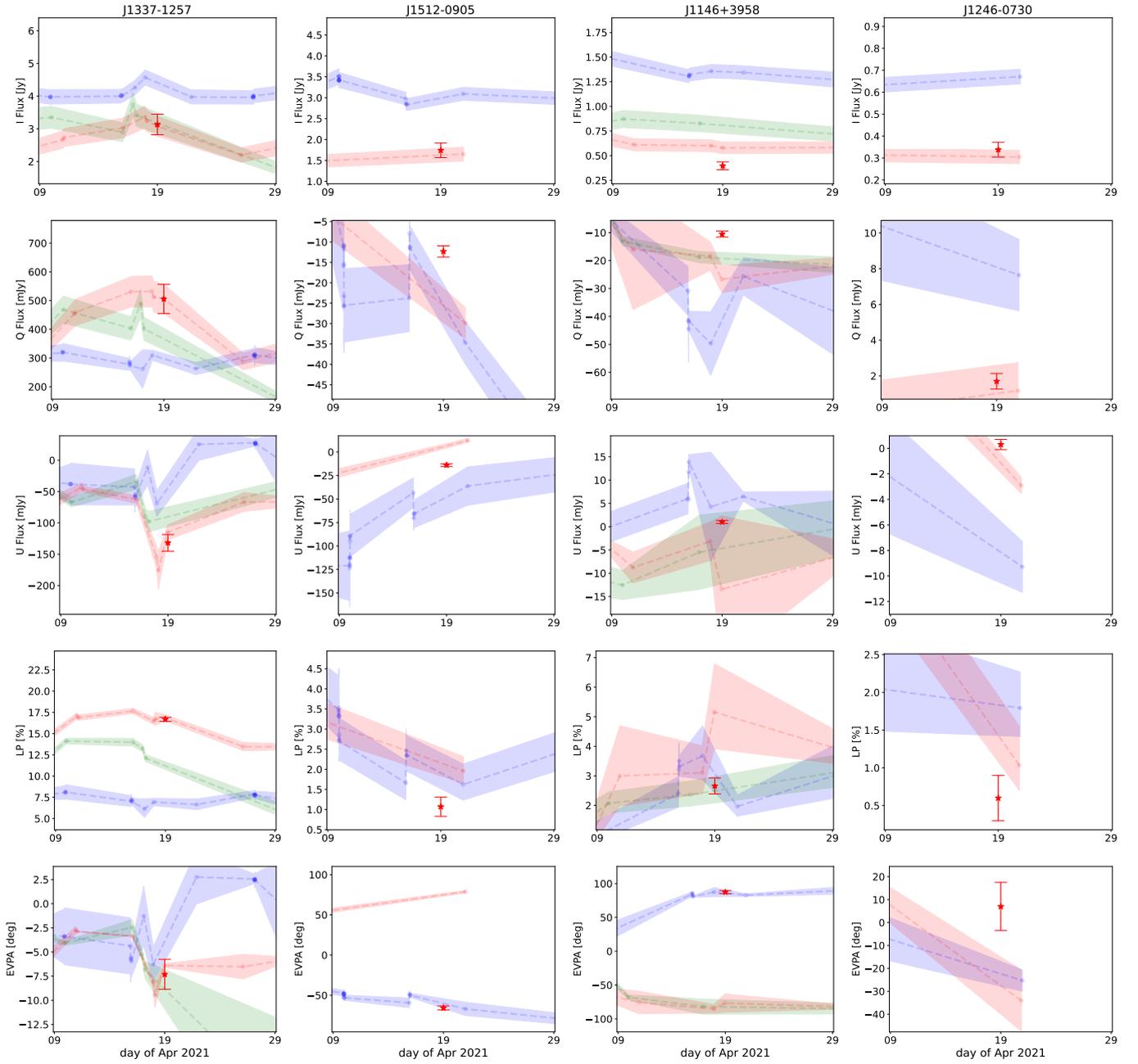
**Table B.1.** Polarization parameters of targeted AGN sources per frequency band (SPW).

Frequency (GHz)	I (Jy)	Q (mJy)	U (mJy)	V (mJy)	LP (%)	EVPA (deg)
VLBI scans						
3C279						
335.60	$8.8086 \pm 0.0003$	$-63.2 \pm 2.6$	$-923.0 \pm 2.6$	$-52 \pm 53$	$10.50 \pm 0.03$	$-46.960 \pm 0.082$
337.54	$8.7961 \pm 0.0003$	$-59.9 \pm 2.6$	$-927.2 \pm 2.6$	$-49 \pm 53$	$10.56 \pm 0.03$	$-46.849 \pm 0.081$
347.60	$8.6754 \pm 0.0003$	$-50.1 \pm 2.6$	$-929.0 \pm 2.6$	$-43 \pm 52$	$10.72 \pm 0.03$	$-46.542 \pm 0.080$
349.60	$8.6327 \pm 0.0003$	$-48.2 \pm 2.6$	$-923.8 \pm 2.6$	$-53 \pm 52$	$10.72 \pm 0.03$	$-46.492 \pm 0.080$
3C273						
335.60	$4.2364 \pm 0.0001$	$-45.4 \pm 1.3$	$-53.1 \pm 1.3$	$-44 \pm 25$	$1.65 \pm 0.03$	$-65.28 \pm 0.52$
337.54	$4.2165 \pm 0.0001$	$-45.1 \pm 1.3$	$-54.2 \pm 1.3$	$-41 \pm 25$	$1.67 \pm 0.03$	$-64.88 \pm 0.52$
347.60	$4.1226 \pm 0.0001$	$-41.1 \pm 1.2$	$-55.0 \pm 1.2$	$-36 \pm 25$	$1.67 \pm 0.03$	$-63.39 \pm 0.52$
349.60	$4.1019 \pm 0.0002$	$-40.9 \pm 1.2$	$-55.4 \pm 1.2$	$-43 \pm 25$	$1.68 \pm 0.03$	$-63.22 \pm 0.52$
M87						
335.60	$0.9944 \pm 0.0002$	$25.22 \pm 0.34$	$12.04 \pm 0.34$	$-14.9 \pm 6.0$	$2.81 \pm 0.03$	$12.76 \pm 0.35$
337.54	$0.9881 \pm 0.0002$	$25.33 \pm 0.33$	$12.03 \pm 0.33$	$-14.3 \pm 5.9$	$2.84 \pm 0.03$	$12.70 \pm 0.34$
347.60	$0.9526 \pm 0.0002$	$24.00 \pm 0.32$	$11.37 \pm 0.33$	$-12.2 \pm 5.7$	$2.79 \pm 0.03$	$12.67 \pm 0.35$
349.60	$0.9451 \pm 0.0002$	$23.66 \pm 0.33$	$10.79 \pm 0.33$	$-13.6 \pm 5.7$	$2.75 \pm 0.04$	$12.26 \pm 0.37$
4C01.28						
335.60	$1.5159 \pm 0.0002$	$32.89 \pm 0.50$	$-126.68 \pm 0.50$	$1.0 \pm 9.1$	$8.63 \pm 0.03$	$-37.72 \pm 0.11$
337.54	$1.5102 \pm 0.0002$	$32.80 \pm 0.49$	$-126.78 \pm 0.49$	$1.0 \pm 9.1$	$8.67 \pm 0.03$	$-37.75 \pm 0.11$
347.60	$1.4915 \pm 0.0002$	$32.61 \pm 0.49$	$-125.31 \pm 0.49$	$1.1 \pm 8.9$	$8.68 \pm 0.03$	$-37.71 \pm 0.11$
349.60	$1.4906 \pm 0.0002$	$32.67 \pm 0.50$	$-125.56 \pm 0.50$	$1.3 \pm 8.9$	$8.70 \pm 0.03$	$-37.71 \pm 0.11$
J1146+3958						
335.60	$0.4007 \pm 0.0003$	$-10.78 \pm 0.29$	$1.24 \pm 0.29$	$-7.0 \pm 2.4$	$2.71 \pm 0.07$	$86.73 \pm 0.76$
337.54	$0.3975 \pm 0.0002$	$-10.65 \pm 0.27$	$1.34 \pm 0.27$	$-6.2 \pm 2.4$	$2.70 \pm 0.07$	$86.42 \pm 0.71$
347.60	$0.3921 \pm 0.0003$	$-10.49 \pm 0.28$	$0.79 \pm 0.27$	$-5.8 \pm 2.4$	$2.68 \pm 0.07$	$87.85 \pm 0.74$
349.60	$0.3953 \pm 0.0003$	$-10.03 \pm 0.30$	$0.67 \pm 0.30$	$-7.0 \pm 2.4$	$2.54 \pm 0.08$	$88.08 \pm 0.86$
J1512-0905						
335.60	$1.7563 \pm 0.0004$	$-12.30 \pm 0.63$	$-14.67 \pm 0.62$	$-18 \pm 10$	$1.09 \pm 0.04$	$-64.99 \pm 0.94$
337.54	$1.7541 \pm 0.0003$	$-11.27 \pm 0.61$	$-14.16 \pm 0.61$	$-16 \pm 10$	$1.03 \pm 0.03$	$-64.25 \pm 0.97$
347.60	$1.7340 \pm 0.0004$	$-12.91 \pm 0.61$	$-14.21 \pm 0.61$	$-14 \pm 10$	$1.11 \pm 0.04$	$-66.14 \pm 0.92$
349.60	$1.7254 \pm 0.0004$	$-12.90 \pm 0.63$	$-12.91 \pm 0.63$	$-18 \pm 10$	$1.06 \pm 0.04$	$-67.48 \pm 0.99$
J1337-1257						
335.60	$3.1590 \pm 0.0004$	$505.4 \pm 1.0$	$-132.5 \pm 1.0$	$-105 \pm 19$	$16.54 \pm 0.03$	$-7.344 \pm 0.055$
337.54	$3.1469 \pm 0.0004$	$504.46 \pm 0.99$	$-129.79 \pm 0.99$	$-99 \pm 19$	$16.55 \pm 0.03$	$-7.214 \pm 0.055$
347.60	$3.1199 \pm 0.0004$	$508.98 \pm 0.99$	$-132.01 \pm 0.99$	$-88 \pm 19$	$16.85 \pm 0.03$	$-7.270 \pm 0.054$
349.60	$3.1042 \pm 0.0004$	$503.87 \pm 1.00$	$-133.33 \pm 1.00$	$-104 \pm 19$	$16.79 \pm 0.03$	$-7.411 \pm 0.055$
non-VLBI scans <sup>a</sup>						
J1246-0730						
335.60	$0.2826 \pm 0.0003$	$1.50 \pm 0.36$	$-0.49 \pm 0.36$	$0.7 \pm 1.7$	$0.6 \pm 0.1$	$-9.1 \pm 6.5$
337.54	$0.2822 \pm 0.0003$	$1.35 \pm 0.35$	$-1.03 \pm 0.35$	$0.5 \pm 1.7$	$0.6 \pm 0.1$	$-18.6 \pm 6.0$
347.60	$0.2776 \pm 0.0004$	$1.16 \pm 0.36$	$-1.25 \pm 0.36$	$0.0 \pm 1.7$	$0.6 \pm 0.1$	$-23.5 \pm 6.1$
349.60	$0.2761 \pm 0.0004$	$1.17 \pm 0.38$	$-0.76 \pm 0.38$	$0.4 \pm 1.7$	$0.5 \pm 0.1$	$-16.6 \pm 7.9$
4C01.28						
335.60	$1.5126 \pm 0.0002$	$33.13 \pm 0.52$	$-126.08 \pm 0.52$	$-0.0 \pm 9.1$	$8.62 \pm 0.03$	$-37.64 \pm 0.11$
337.54	$1.5079 \pm 0.0002$	$32.94 \pm 0.52$	$-126.40 \pm 0.52$	$-0.0 \pm 9.1$	$8.66 \pm 0.03$	$-37.70 \pm 0.11$
347.60	$1.4920 \pm 0.0003$	$32.82 \pm 0.51$	$-125.19 \pm 0.51$	$0.1 \pm 9.0$	$8.67 \pm 0.03$	$-37.66 \pm 0.11$
349.60	$1.4906 \pm 0.0003$	$32.66 \pm 0.52$	$-125.53 \pm 0.52$	$0.1 \pm 8.9$	$8.70 \pm 0.03$	$-37.71 \pm 0.12$

<sup>(a)</sup> Non-VLBI scans on 3C273 were not included in this polarization analysis.



**Fig. C.1.** Comparison of the polarimetric results obtained for all sources observed in VLBI mode on April 19, 2021, with those retrieved from the AMAPOLA polarimetric analysis of GS data. Each row shows a specific parameter (from top to bottom: Stokes  $I$ ,  $Q$ ,  $U$ , LP, and EVPA), while each column corresponds to a different source (labels in the top panel; see also Fig. C.2 for more sources). The measurements from the ALMA-VLBI observations are indicated as red stars with associated error bars. The shaded regions correspond to AMAPOLA's  $\pm 1\sigma$  uncertainties for Band 3 (97.5 GHz; blue shading), Band 6 (221.1 GHz; green shading), and Band 7 (343.4 GHz; red shading). These uncertainties are derived from the ACA GS data, while their temporal evolution (lines) is interpolated between individual GS measurements. This figure and Fig. C.2 highlight that, for most cases, the ALMA-VLBI measurements agree well with the AMAPOLA trends, helping confirm the reliability of the polarization calibration across bands.



**Fig. C.2.** Same as Fig. C.1 but for different sources. The apparent discrepancy between the Band 7 GS prediction and our measurement of the EVPA for J1512-0905 and J1146+3958 is attributed to the  $\pm\pi$  ambiguity in estimating the polarization direction (e.g., Taylor et al. 2009).