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Census of the most obscured galaxy nuclei over cosmic time to be revealed by PRIMA

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ABSTRACT. Characterizing the growth of supermassive black holes (SMBHs) is critical to the evolution of galaxies; however, the majority of this activity is obscured, rendering traditional tracers of active SMBHs, such as in the restframe optical/ultraviolet, ineffective. The mid-infrared has been particularly successful in revealing obscured active galactic nucleus activity; however, much of this work is confined to the local universe due to the lack of a far-infrared (far-IR) telescope with the required sensitivity and wavelength coverage. We demonstrate the effectiveness of PRobe Far-Infrared Mission for Astrophysics (PRIMA), a concept 1.8-m far-IR observatory, to detect and characterize deeply obscured galaxy nuclei over cosmic time. With the PRIMAgger instrument covering 25 to 235 μm , we find that we can accurately detect obscured nuclei via the deep silicate absorption at restframe 9.8 μm between $z = 2-7$. In addition, the FIRESS spectrograph can produce $R \sim 100$ spectra of obscured nuclei out to $z \sim 7$, detecting polycyclic aromatic hydrocarbons, ices, and ionized and molecular gas. With the large number of deeply obscured nuclei that PRIMA can detect and characterize, such a mission is critical to understand the growth of SMBHs.

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1 Introduction

Active galactic nuclei (AGNs) are thought to play a key role in how galaxies evolve, where the interaction with the host galaxy can regulate star formation on a variety of scales.^{1,2} To understand this coevolution, one must understand how supermassive black holes (SMBHs) grow and affect their host galaxy; however, much of this activity is hidden behind vast quantities of dust.³ In the unified model, obscuration was thought to be a viewing angle effect where type 2 AGNs are viewed through a dust obscuring torus.⁴ However, this picture has evolved in recent times where obscuration may be an evolutionary stage, where rapid SMBH growth is hidden before the AGN eventually clears its dust.^{5,6} Evidence for a population of dust-enshrouded AGN can be found in the cosmic X-ray background,⁷ which postulates the presence of a large fraction of sources with column densities in excess of 10^{22} cm^{-2} . At the extreme end of this population with $N_{\text{H}} > 10^{25} \text{ cm}^{-2}$, a significant fraction of high-energy photons can be absorbed, making even X-rays ineffective.⁸

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As the obscuring dust is heated and reprocesses photons from the accretion disk, the restframe infrared is extremely bright in deeply obscured objects. With the advent of the JWST/MIRI,⁹ imaging has allowed the selection of AGN at high- z ,¹⁰ via the hot dust continuum, which appears in the near-infrared, finding significantly more objects than X-ray selected AGN. However, the wavelength coverage of JWST is not sufficient to observe the restframe mid-infrared at redshifts beyond cosmic noon ($z > 3$). In addition, we can only detect objects with an extremely hot dust component as this is required to produce a strong dust continuum in the near-infrared, and so, most obscured targets, which typically do not have such hot dust, are missed. To unveil deeply obscured nuclei sources at high- z ($z > 3$), access to the restframe ~ 5 to $30\ \mu\text{m}$ regime is required.

In the local universe, the most obscured nuclei exhibit column densities of $N_{\text{H}} > 10^{25}\ \text{cm}^{-2}$. These so called compact obscured nuclei (CONs) have been identified via dense gas tracers in the submillimeter, namely, HCN-vib emission.^{11–13} Unlike typical type 2 AGN, CONs show a significantly deeper silicate feature and a colder dust continuum^{14,15} in the mid-infrared. Moreover, these objects are more common than one may think considering their extreme nature, comprising $\sim 40\%$ of local ULIRGs,^{13,14} and considering the increase in luminous infrared galaxies at higher redshifts,¹⁶ these objects are likely numerous. Although detecting HCN-vib in a sample of local targets has successfully identified CONs,¹³ this tracer is extremely faint, making finding CONs at high- z a challenge.

The mid-infrared continuum of deeply obscured nuclei exhibits deep silicate absorption at $\sim 9.8\ \mu\text{m}$, which can be used to identify such targets.¹⁷ A new method to select deeply obscured nuclei based on the ratio of the equivalent widths (EWs) of polycyclic aromatic hydrocarbon (PAH) features within the silicate band¹⁵ has proven to be very successful as it is less affected by contamination by the host galaxy. The method relies on the contrast between the extinction of the PAHs (tracing the circumnuclear star formation) and the continuum (resulting from the obscured nucleus), successfully recovering the CONs identified in the submillimeter.^{14,15} In addition, ice absorption features are present in the spectra¹⁸ and can be used to measure the obscuring column density¹⁹ based on their optical depth.

Higher ionization potential lines such as [Ne V] and [O IV] have been proposed to measure the black hole accretion rate;²⁰ however, this method becomes ineffective for deeply obscured nuclei where these lines are typically undetected. In such deeply obscured objects, the dust absorbs the ionizing photons, preventing the gas from being ionized.^{21,22} This could explain the absence of high ionization lines in the spectra of deeply obscured nuclei,^{23,24} even with the sensitivity of JWST/MIRI in the local universe. The restframe mid-infrared continuum is therefore key to unveiling the most obscured nuclei across cosmic time, which requires a far-infrared telescope that is sensitive enough to detect these objects.

PRobe Far-Infrared Mission for Astrophysics^{26,27} (PRIMA²⁵) is a concept for a 1.8-m cryogenically cooled FIR telescope, which has recently been selected for phase A study by NASA. The telescope will contain two instruments, PRIMAgger and FIRESS. The latter is a spectrometer between 24 and $235\ \mu\text{m}$, whereas the former is an imager. With FIRESS, there are two modes: a high sensitivity low-resolution mode at $R \sim 100$ and a high-resolution FTS mode, which allows $R \sim 20000$ at $25\ \mu\text{m}$ to $R \sim 4400$ at $112\ \mu\text{m}$.

In this paper, we demonstrate PRIMA's ability to identify deeply obscured nuclei, which have a deep silicate feature indicative of column densities of $N_{\text{H}} \gtrsim 10^{25}\ \text{cm}^{-2}$ that have potentially remained undetected through recent/current X-ray or near-infrared missions. We show that the identification of this population is now possible through imaging and spectroscopy with PRIMA.

In Sec. 3.1.1, we demonstrate PRIMAgger's capability of finding candidates with both real data in the local universe and simulated observations. In Sec. 3.1.2, we discuss the use of FIRESS to study the nature and evolution of these sources. In this work, we assume Λ CDM cosmology with $H_0 = 70\ \text{km s}^{-1}\ \text{Mpc}^{-1}$, $\Omega_m = 0.27$, and $\Omega_\Lambda = 0.73$.

2 Methods

2.1 PRIMAgger

The PRIMAgger instrument contains a hyperspectral imager spanning two bands, PHI1 and PHI2, between 25 and $85\ \mu\text{m}$ with a spectral resolution of $R \sim 10$. In addition, there are four filters in the polarimeter imager between 96 and $235\ \mu\text{m}$. In this work, we represent each band of the

Table 1 PRIMAgger filter information and sensitivities. The sensitivity is given as the 5σ flux for a survey of 1500 h/deg^2 . The classical confusion limits and achieved sensitivities via Bayesian de-blending techniques are from Donnellan et al.²⁹

Filter	Central wavelength (μm)	Filter width (μm)	5σ sensitivity (μJy)	Confusion limit (μJy)	XID+, deep prior (μJy)
PHI1_1	25.0	2.5	75.1	20	82
PHI1_2	27.8	2.8	85.6	27	86
PHI1_3	30.9	3.1	96.0	37	93
PHI1_4	34.3	3.4	106.5	51	108
PHI1_5	38.1	3.8	116.9	71	116
PHI1_6	42.6	4.3	127.4	107	149
PHI2_1	47.4	4.7	140.4	161	95
PHI2_2	52.3	5.2	160.0	249	117
PHI2_3	58.1	5.8	179.6	401	138
PHI2_4	64.5	6.5	199.2	667	167
PHI2_5	71.7	7.2	218.8	1120	229
PHI2_6	79.7	8.0	238.4	1850	285
PPI1	96	23	89.8	4250	281
PPI2	126	15	117.6	12300	747
PPI3	172	20.3	160.8	28400	2650
PPI4	235	27.6	187.0	46000	7030

hyperspectral imager with six photometric filters^{28,29} following previous work, where each has a top hat profile with central wavelength and width given in Table 1. We show the 5σ sensitivities for a survey of 1500 h for an area of 1 square degree.

Considering PRIMA’s 1.8-m mirror and the wavelength coverage, confusion will naturally be a concern for source identification. We refer the reader to Donnellan et al.²⁹ for details on how best to mitigate this problem. In Table 1, we report the sensitivities corresponding to the classical confusion limit as well as those based on mitigating techniques, particularly the use of XID+ from Hurley et al.,³⁰ which uses a Bayesian framework to de-blend sources using some prior. We report the sensitivities using the deep catalog prior in Table 1. Although the sensitivity is reduced significantly for the longest wavelength filters ($\gtrsim 80 \mu\text{m}$) when considering confusion, the dust continuum is typically bright at these wavelengths (restframe $\gtrsim 12 \mu\text{m}$), and so, we find that accounting for confusion makes little difference to the total number of objects that can be detected.

At an $R \sim 10$, the resolution at 25 to $85 \mu\text{m}$ is sufficient to detect the continuum shape of extremely obscured objects, as shown in Fig. 1. In particular, the $9.8 \mu\text{m}$ silicate absorption will be measurable from $z \sim 2$ to 7, allowing for the selection of obscured nuclei using different combinations of colors¹⁵ from the filters shown in Table 1. We demonstrate this in Fig. 1, where the Spitzer IRS¹⁷ + Akari³¹ spectrum of NGC 4418 is shown at different redshifts to represent the typical spectrum of a deeply obscured nucleus. We also show the spectrum of NGC 7714 as a typical star-forming galaxy where the spectrum is dominated by PAHs. We use the spectrum from the GALSEDATLAS,³² which combines the Spitzer and Akari spectrum.

An important consideration for subsequent analysis is well-constrained photometric redshifts. The presence of the relatively broad PAH features and/or silicate absorption features in combination with the good filter coverage of PRIMAgger is sufficient to simultaneously determine a photometric redshift. A detailed analysis of the effectiveness of the photo- z measurements with PRIMA is beyond the scope of this work and deserves a dedicated analysis.

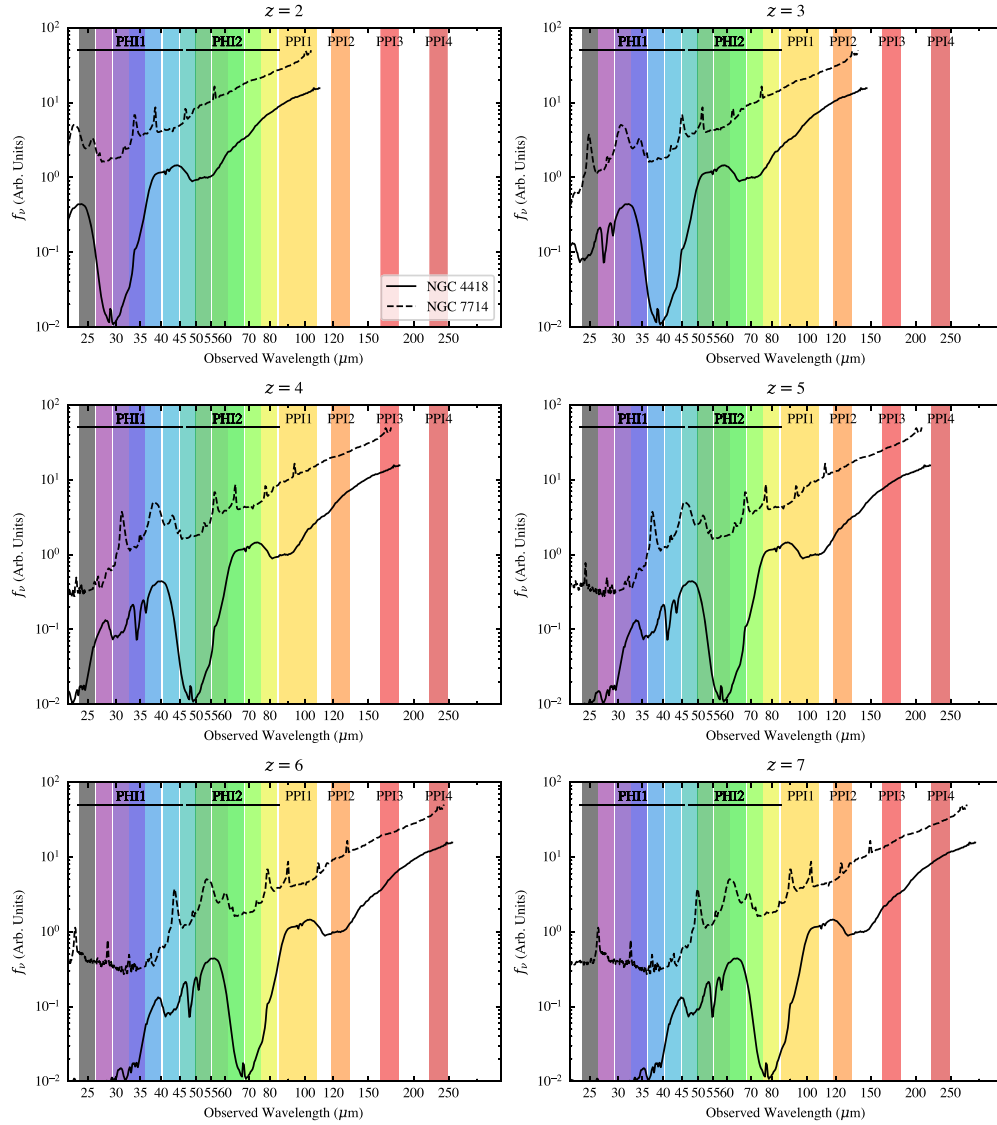


Fig. 1 Demonstration of PRIMA's filters at detecting various spectral features, such as the 9.8- μm silicate absorption, at different redshifts. The solid line shows the Spitzer IRS + Akari spectrum of NGC 4418, whereas the dashed line shows NGC 7714—a typical star-forming galaxy. The filters, shown as the shaded bands, are described in Table 1.

However, previous studies with AKARI³³ found that mid-IR-determined photometric redshifts are recovered to an accuracy within $\sim 10\%$. Considering the wider filters of AKARI compared with PRIMA, photometric redshifts can be accurately constrained to within $\sim 10\%$. Moreover, the uncertainty of this measurement is likely smaller than the width of the redshift bins we use in our color selection (see Sec. 3.1.1).

2.1.1 Galaxy samples

In this work, we demonstrate PRIMA's ability to identify and characterize deeply obscured nuclei at $2 < z < 7$. To do this, we use archival Spitzer spectra of local (U)LIRGs in the IDEOS database^{18,34} and follow the same methodology as Donnan et al.¹⁴

We use the *HERschel* Ultra luminous infrared galaxy Survey (HERUS) sample³⁵ to obtain a sample of 42 local ULIRGs. In addition, we use the Great Observatories All-sky LIRG Survey (GOALS)³⁶ survey, which contains 179 LIRGs and 22 ULIRGs. Out of the full sample, we use those that have Spitzer IRS data, and thus, we have 143 LIRGs in our sample. Finally, we use a pure star-forming sample,³⁷ which consists of 106 galaxies.

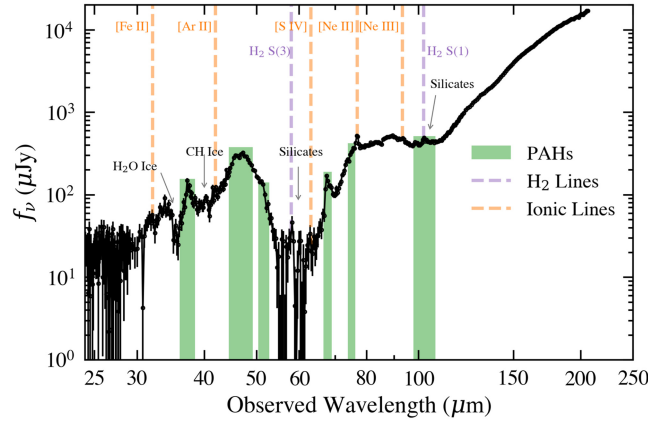


Fig. 2 Example spectra as observed by FIRESS of an obscured nucleus within a HIRG at $z = 5$ after 5 h of integration. We use the spectrum of Arp 220 as a template scaled to an infrared luminosity of $10^{13.15} L_{\odot}$ at $z = 5$. Key spectral features are labeled.

2.1.2 SPRITZ simulation

We use the Spectro-Photometric Realisations of IR-Selected Targets at all- z (SPRITZ)³⁸ simulation,³⁹ which provides predictions of the number of galaxies at different redshifts, following the work of Bisigello et al.²⁸ This simulation is empirically driven, with the population of galaxies being generated based on observed luminosity functions of different galaxy types, namely, star-forming (SF), AGN, composite, and passive. The proportion of each galaxy type for a given luminosity and redshift is selected from observed IR luminosity functions.⁴⁰ Each type has an associated SED template assigned depending on the properties of each object; however, we are cautious to not rely on the templates as they do not properly account for deeply obscured nuclei. For more details, see Bisigello et al.^{28,39}

2.2 FIRESS

The FIRESS instrument is a spectrograph, observing between 23 and 235 μm with a spectral resolution of $R \sim 100$. With FIRESS, a 5σ line flux of $3 \times 10^{-19} \text{ Wm}^{-2}$ can be obtained with 1 h of integration. This corresponds to a flux level of $\sim 100 \mu\text{Jy}$, which is sufficient to obtain quality spectra of a significant number of candidates detected in the PRIMAGER surveys. Increasing the exposure time to 5 h yields a depth of $\sim 45 \mu\text{Jy}$. To test the ability of FIRESS to measure the spectra of deeply obscured nuclei, we simulate a number of objects of varying luminosity at different redshifts. We use the spectrum of Arp 220, which hosts a deeply obscured nucleus¹³ as well as circumnuclear star formation.

To simulate the spectra, we scale the Spitzer spectra to match the photometry of objects in the SPRITZ simulation that are identified as deeply obscured nuclei (see Sec. 3.1.1). We show the resulting simulated spectra of an obscured nucleus at $z = 5$, in Fig. 2. We find that the continuum and emission/absorption features are present, allowing the recovery of PAH fluxes, ionic line fluxes such as [Ne II], and absorption from H_2O ice at restframe $\sim 6 \mu\text{m}$. The spectra is therefore of a sufficiently high quality to characterize the nature of deeply obscured nuclei candidates identified from the imaging surveys.

3 Results

3.1 Population of Deeply Obscured Nuclei at $2 < z < 7$

3.1.1 PRIMAGER

To identify deeply obscured nuclei in the PRIMAGER survey, we select filters from Table 1 that are sensitive to the silicate absorption at 9.8 μm at a given redshift. Although there is enough wavelength range to go beyond $z = 7$, the sensitivity limits detecting a reasonable number of targets, as discussed in Sec. 2.1.2. For each redshift slice, we select a combination of three filters. One of them covers the silicate absorption band,¹⁵ whereas the other two are selected to measure

the continuum outside the silicate band. In particular, we choose one to be at $\sim 13\text{-}\mu\text{m}$ restframe just outside the silicate absorption and the other to be at $\sim 30\text{ }\mu\text{m}$ restframe, which is sensitive to the relatively colder dust. This arrangement helps to discriminate against type 2 AGN, which may have a deep silicate feature but show a hotter far-infrared continuum.⁴¹ It is worth noting that other filter combinations can also be used, such as using the 3 to 5 μm slope, as was presented by García-Bernete et al. submitted.

In addition, due to the good filter coverage, PRIMA will be able to sample the SED well (see Sec. 3.2), and so, one can both determine a photometric redshift and classify the nature of the object purely via SED fitting without any prior color selection. In this work, we choose to show the color section as it clearly illustrates the strength of the silicate absorption for identifying deeply obscured nuclei and can be quantified, without relying on specific SED fitting codes.

Figure 3 shows the color measurements for the star-forming,³⁷ HERUS, and GOALS samples for different redshift slices. The red dots represent objects identified as hosting deeply obscured nuclei as per Donnan et al.¹⁴

Donnan et al.¹⁴ provided two criteria to select deeply obscured nuclei. The first was based on PAH EW ratios,¹⁵ whereas the second was based on modeling the spectra to infer the optical depth of the nuclear continuum. In this work, we use the second classification. For a comparison between the two, see Donnan et al.¹⁴

Using Donnan et al.'s¹⁴ selected deeply obscured nuclei as the ground truth, we create selection criteria for PRIMA using the pure star-forming sample.³⁷ To generate a color selection for deeply obscured nuclei, we first measure the mean color and standard deviation of the star-forming sample. We then create the deeply obscured nuclei color selection as being 1σ , 2σ , and 3σ greater than the star-forming mean. This is shown in Fig. 3.

We find that the color selection does a good job at identifying the most deeply obscured objects; however, the method may miss objects where the emission from the host galaxy is very strong. This is indicated by the size of the points plotted in Fig. 3, where the smaller the marker, the lower the nuclear fraction, as measured by Donnan et al.¹⁴ The nuclear fraction is defined as the fractional contribution to the mid-IR flux (between 5 and 14 μm) from the nuclear component using PAHDECOMP (<https://github.com/FergusDonnan/PAHDecomp>). Those objects with the highest nuclear fraction ($\gtrsim 60\%$) appear firmly within the selection criteria. We find that the 1σ criteria selects 80% of the total obscured nuclei, whereas the 2σ and 3σ criteria find 65% and 60%, respectively. Although the 1σ criteria find a majority of the obscured nuclei, up to $\sim 45\%$ of the targets selected are ULIRGs and LIRGs that do not host deeply obscured nuclei. This drops to zero with a strict color cutoff but would limit to only the most extreme objects being selected. The 2σ limit provides a reasonable selection as the number of false positives counteracts the number of objects missed, and so, if one is simply interested in the total fraction of deeply obscured nuclei for a given luminosity and/or redshift bin, the fraction of obscured nuclei is reasonably accurate.

The effect of contamination by the host galaxy due to the size of the PRIMA primary mirror and therefore the afforded spatial resolution will limit the completeness of obscured nuclei that PRIMA can detect. This issue will impact the detection of deeply obscured nuclei in the high- z universe; however, PRIMA will be able to detect a large number of deeply obscured candidates, more than any other current or future space/ground facility.

To test how many galaxies PRIMA can detect, we use the SPRITZ simulation results for a survey of 1500 h over 1 square degree. We follow Bisigello et al.²⁸ and consider a galaxy detected if at least 6 out of the 12 photometric bands have a flux greater than 5σ , where σ is the noise at a given wavelength. We plot the number of detected galaxies in Fig. 4, where we find a reasonable number of galaxies detected out to $z \sim 7$. In particular, we find ~ 9000 galaxies at $z = 1.5$ to 2.5 , which lowers to 53 at $z = 6.5$ to 7.5 . At the highest redshift bin, only hyper-luminous infrared galaxies (HLIRGs, $L_{\text{IR}} > 10^{13} L_{\odot}$) are bright enough to be detected.

Predicting the exact number of galaxies that host deeply obscured nuclei is challenging, as the fraction of HLIRGs, ULIRGs, and LIRGs that host such nuclei is not well constrained beyond the local universe. Indeed, PRIMA is well suited to determine the fraction of deeply obscured nuclei in the universe and therefore measure this fraction. We therefore provide a range of predictions depending on the assumed intrinsic fraction, to provide the reader with a general idea of how many objects PRIMA will be able to detect. We plot these predictions in Fig. 5 for three

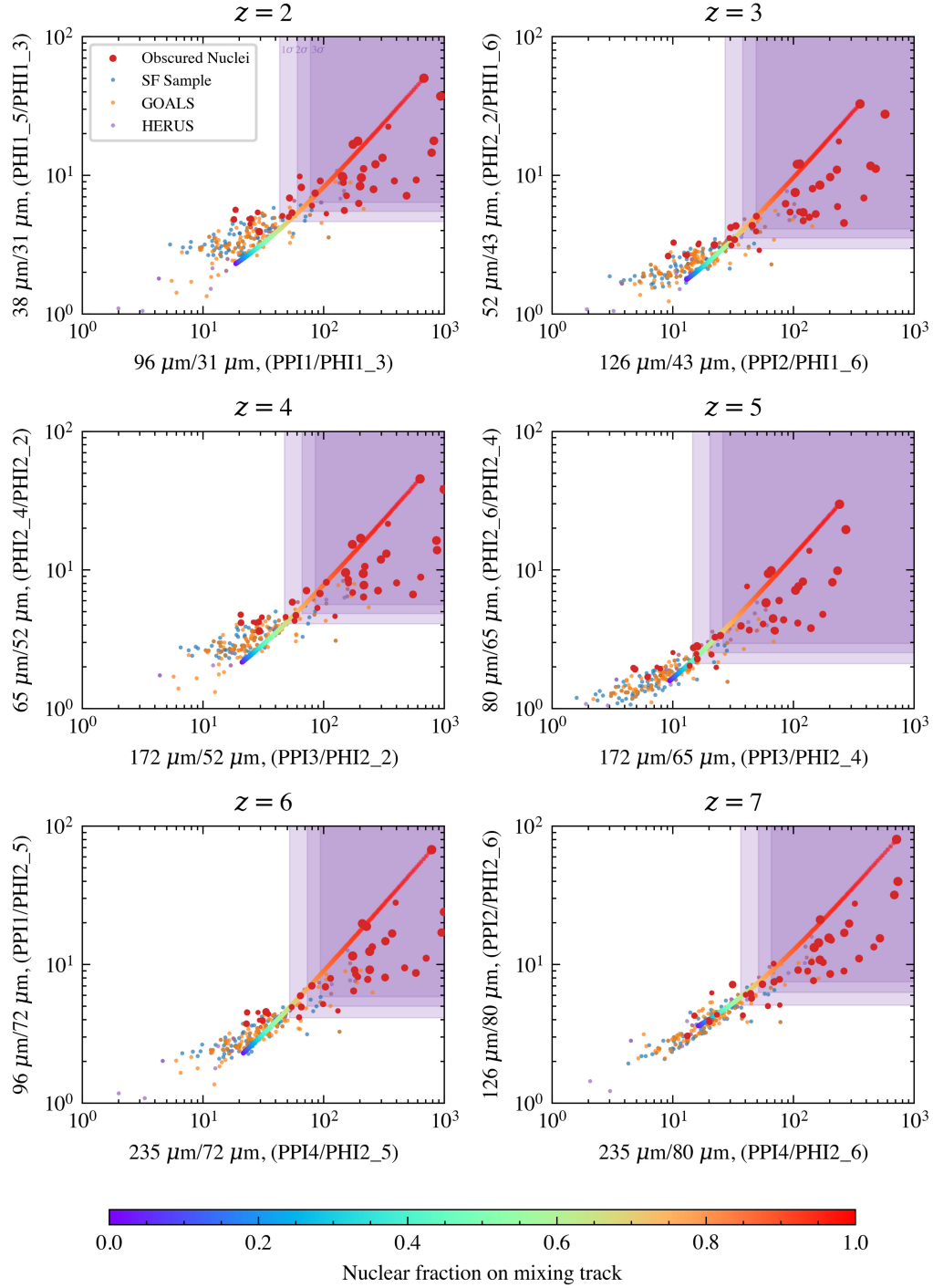


Fig. 3 Color selection for identifying deeply obscured nuclei with PRIMAGER. Each panel shows a different redshift slice with varying filter combinations. The blue dots show the pure star-forming sample,³⁷ the orange shows the GOALS³⁶ sample, whereas the purple points show the HERUS sample.³⁵ The red points show objects identified as hosting a deeply obscured nucleus in the GOALS and HERUS samples from Spitzer spectra,¹⁴ where the size of the points reflect the nuclear fraction as measured in Donnan et al.¹⁴ The purple box displays the selection criteria for deeply obscured nuclei. This is defined based on the mean and standard deviation, σ of the pure star-forming sample, where three boxes are shown as 1σ , 2σ , and 3σ from this mean. The color bar shows a track for a simulated spectrum consisting of a nuclear component, NGC 4418, and a star-forming component, NGC 7714. The color indicates an increasing nuclear fraction to the total SED.

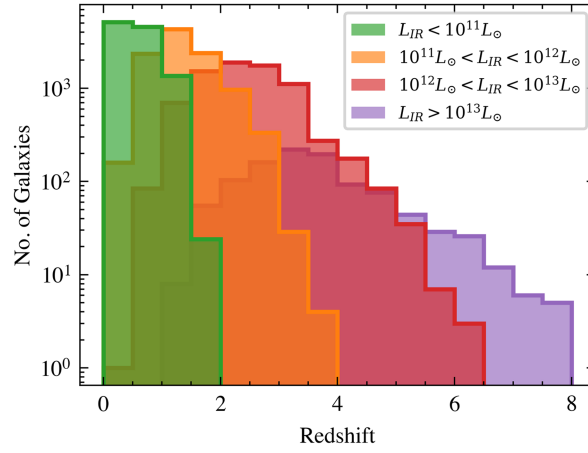


Fig. 4 Number of galaxies of varying infrared luminosity that can be detected by PRIMAGER using the output of the SPRITZ simulation. These values are for a 1500 h/deg², accounting for confusion using the values shown in Table 1.

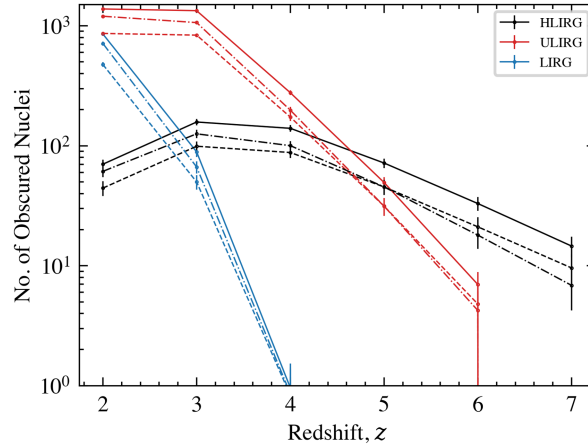


Fig. 5 Number of deeply obscured nuclei that PRIMA can detect as a function of wavelength and infrared luminosity. HLIRGs are shown in black, ULIRGs in red, and LIRGs in blue. The solid line shows an assumed fraction of sources consistent with local galaxies. The dashed line shows half the fraction of local galaxies. The dashed dot line shows the evolving fraction of obscured sources that decrease from the local value to zero at $z = 10$. This has the form, $0.4(1 - \frac{z}{10})$, starting at 0.4 at $z = 0$.

different assumptions. For each, we calculate an error on the number of obscured sources, based on a beta distribution,⁴² which scales with the size of the total sample.

Our first prediction is based on the assumption that 40% of ULIRGs and 20% of LIRGs host deeply obscured nuclei, in line with the local population. As there are no local HLIRGs, we also assume 40% for this population to be conservative, but it is likely higher considering the increasing fraction with infrared luminosity. This is shown as the solid lines in Fig. 5. We then assumed half of the local fraction, i.e., 20% of ULIRGs and HLIRGs, 10% of LIRGs. This is shown as the dashed lines in Fig. 5. We also consider an evolving obscured fraction, which is likely closer to reality. For simplicity, we assume a decreasing fraction from local values of 40% of HLIRGs, 40% of ULIRGs, and 20% of LIRGs to 0 at $z = 10$. This is shown as the dashed-dot line in Fig. 5.

We find that PRIMA will be able to detect ~ 100 to 1000 LIRGs and ULIRGs and ~ 100 HLIRGs hosting deeply obscured nuclei at cosmic noon. The number of LIRGs drops rapidly due to their relative faintness, reaching down to zero beyond $z \sim 4$. The number of ULIRGs hosting deeply obscured nuclei drops to ~ 10 at $z \sim 6$. Beyond that redshift, they are no longer detected. The number of HLIRGs remains $\lesssim 10$ out to $z \sim 7$. These numbers vary depending on the assumed intrinsic fraction of obscured sources; however, even when this fraction is as low as

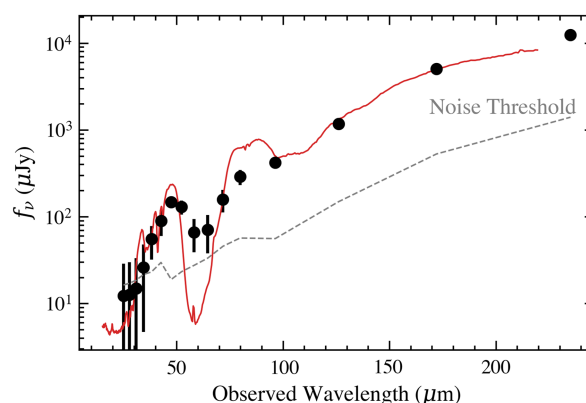


Fig. 6 Example of a deeply obscured nucleus at $z \sim 5$, selected from the SPRITZ simulation. This particular object has an infrared luminosity of $10^{13.38} L_{\odot}$. The figure shows the 9.8- μm silicate absorption at an observed wavelength of $\sim 60 \mu\text{m}$, where the absorption is sufficiently deep to hit the noise threshold, as shown by the gray dashed line. In comparison is the Spitzer + Akari spectrum of NGC 4418, which has an extremely deep silicate feature—note that this is not simulated but purely shown for reference.

$\sim 10\%$ of HLIRGs at $z \sim 7$, a non-zero number of sources will be detected. This exercise demonstrates that PRIMA will likely make considerable strides in constraining the number of deeply obscured nuclei up to $z \sim 7$.

It is also worth noting that if the silicate absorption feature is sufficiently deep, the flux in the bands covering this feature will drop below the noise threshold of the instrument, and thus, the measured color becomes a lower limit. We show an example SED of an object showing a very deep silicate feature at $z \sim 5$ in Fig. 6 in comparison to the Spitzer spectrum of NGC 4418, which has one of the deepest silicate features in the local universe. Although such objects will still be successfully classified as hosting deeply obscured nuclei, they will not appear in the top right corner of the color–color plot but instead appear as an upper limit within the box.

3.1.2 FIRESS

Photometrically selected candidates can be followed up to confirm their nature using FIRESS spectroscopy. As we demonstrated in Fig. 2, we can obtain quality spectra within a reasonable integration time. Using spectroscopic observations, we will be able to properly understand the nature of these targets. Only through spectroscopy we will be able to confirm the presence of a deeply obscured nucleus and study their properties through detailed modeling of the FIRESS spectra.

As demonstrated in Fig. 2, the spectra will include a number of important features. Apart from the silicate absorption features, absorption from ices and emission from PAHs are clear. In addition, many emission lines can be detected, depending on the nature of the target.

Ice features of H_2O and CH can be seen at ~ 36 and $40 \mu\text{m}$, respectively, for a galaxy at $z = 5$. These features have been shown to trace the obscuring gas column density well¹⁹ and are commonplace in obscured nuclei and even in some “typical” Seyfert AGN. Although the role of ices in the dusty structure in galaxy nuclei is not well understood, by the launch of PRIMA, work with JWST data will have advanced our understanding; however, this will be confined to $z < 3$. With PRIMA, we will be able to trace the feature over cosmic time to understand how ices form and evolve in the nuclei of galaxies.

With a representative sample of targets, we can search for signs of evolution in the nuclear structure via high ionization potential lines such as $[\text{Ne V}]$ ($14.32 \mu\text{m}$ restframe). Most obscured nuclei in the local universe show no emission from these lines, even with high sensitivity JWST/MIRI spectra in the local universe;^{23,24} however, from the large sample Spitzer selected objects in Donnan et al.,¹⁴ only one target showed evidence of $[\text{Ne V}]$, IRAS 19254-7245. With the large number of targets that PRIMA can identify, we will be able to see if such high-ionization lines are present at high- z . Their presence would indicate that these objects are fundamentally different

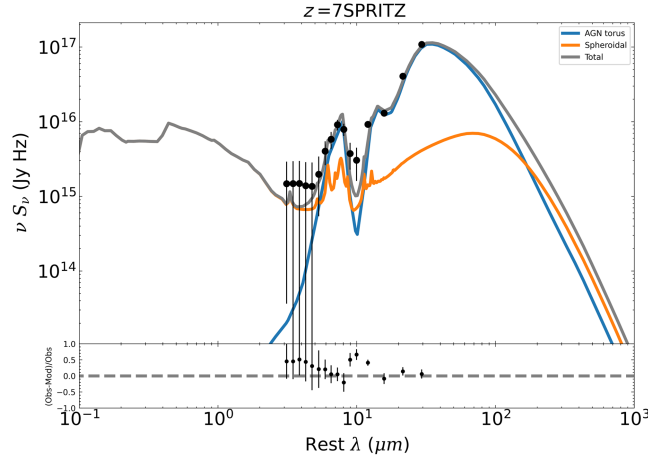


Fig. 7 Example of a fit to an H LIRG hosting an obscured nucleus identified from the SPRITZ simulation at $z = 7$ using the SMART SED fitting code.⁴³ This particular source has an infrared luminosity of $10^{13.78} L_{\odot}$.

compared with their local counterparts. PRIMA therefore has the ability to measure how these objects evolve over cosmic time.

3.2 Torus Modeling

At present, the majority of deeply obscured nuclei lack high-ionization potential lines and X-ray emission; therefore, the infrared SED is the only reliable diagnostic of the properties of the buried AGN. However, at high- z , there is a lack of understanding of AGN tori leading to torus templates often being blindly applied at high- z . As the majority of AGNs are obscured at $z > 1$,³ understanding the nature of the dusty obscuring structure is critical to understanding the growth of SMBHs and ultimately how galaxies evolve. With the wavelength coverage PRIMAgger and FIRESS, one can fit and refine the torus models to understand the nature of the most obscured nuclei.

As an example, we fit a candidate selected from the SPRITZ simulation at $z = 7$ and fit the photometry with the SMART SED fitting code.⁴³ This tool uses radiative transfer simulations for all the components of the SED, including a spheroidal host component and an AGN torus. We use the CYGNUS library of torus models,^{44,45} to fit this object.

The fit is shown in Fig. 7. As expected, the continuum is dominated by the torus, with the spheroidal host contributing to the PAH features.

Iterating and refining the torus models to fit the obscured nuclei that we will detect are critical for inferring the physical parameters of the AGN within. For example, the intrinsic luminosity and thus accretion rate strongly depend on the torus geometry,⁴⁴ and so, for determining the cosmic accretion rate, the density to map the growth of SMBHs across cosmic time requires such a detailed understanding of the torus.

4 Conclusions

We have demonstrated the power of PRIMA to detect and characterize the deeply obscured galaxy nuclei over cosmic time. The main strength of PRIMA is the large sample of galaxies that can be observed in the reframe mid-IR up to $z = 7$. We showed that with PRIMAgger, the good wavelength coverage and sensitivity allow the successful detection of deeply obscured nuclei between $z = 2$ to 7, where $\sim 65\%$ of obscured nuclei can be recovered. Those recovered have the deepest silicate feature, whereas the remaining $\sim 35\%$ are likely to suffer from contamination by the host galaxy. With different assumptions on the intrinsic fraction of sources hosting deeply obscured nuclei, we have shown that PRIMA can effectively constrain this number down to $\sim 10\%$ of H LIRGs at $z \sim 7$.

With FIRESS, high-quality spectra can be obtained, with 5 h of integration, which enables us to study/constrain the properties of deeply obscured nuclei. In addition, spectral features such as H₂O ice can be studied at high- z , opening a new window into how dust evolves in the early universe.

We finally showed that the spectral coverage of PRIMAgger is sufficient to constrain torus models, even out to $z = 7$, which will revolutionize our understanding of the evolution of dust and its role in AGN obscuration. With PRIMA, we will therefore be able to construct a complete census of the most obscured galaxy nuclei over cosmic time and therefore reveal the hidden growth of SMBHs.

Disclosures

The authors declare that there are no conflicts of interest.

Code and Data Availability

The Spitzer IRS spectra used in this work are available from the IDEOS database.³⁴

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