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Evidence for Neutrino Emission from X-Ray-bright Active Galactic Nuclei with IceCube

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Abstract

Recently, IceCube reported neutrino emission from the Seyfert galaxy NGC 1068. Using 13.1 yr of IceCube data, we present a follow-up search for neutrino sources in the northern sky. NGC 1068 remains the most significant neutrino source among 110 preselected gamma-ray emitters while also being spatially compatible with the most significant location in the northern sky. Its energy spectrum is characterized by an unbroken power-law with spectral index $\gamma = 3.4 \pm 0.2$. Consistent with previous results, the observed neutrino flux exceeds its gamma-ray counterpart by at least 2 orders of magnitude. Motivated by this disparity and the high X-ray luminosity of the source, we selected 47 X-ray-bright Seyfert galaxies from the Swift/BAT spectroscopic survey that were not included in the list of gamma-ray emitters. When testing this collection for neutrino emission, we observe a 3.3σ excess from an ensemble of 11 sources, with NGC 1068 excluded from the sample. Our results strengthen the evidence that X-ray-bright cores of active galactic nuclei are neutrino emitters.

Unified Astronomy Thesaurus concepts: Neutrino astronomy (1100); X-ray active galactic nuclei (2035); High energy astrophysics (739); Seyfert galaxies (1447)

1. Introduction

For over a decade, the IceCube Neutrino Observatory has been consistently detecting a diffuse flux of high-energy cosmic neutrinos (IceCube Collaboration 2013). While a fraction of this flux has recently been linked to the Galactic plane (IceCube Collaboration 2023), the majority remains isotropic, pointing to an extragalactic origin. Neutrinos are expected to be produced in interactions of protons with ambient matter and/or radiation within their cosmic sources. These interactions generate charged and neutral pions, which subsequently decay to produce neutrinos and gamma rays. While gamma rays can also arise from purely leptonic processes, neutrinos are only produced in hadronic interactions and are expected to be accompanied by gamma-ray emission (e.g., F. Halzen & D. Hooper 2002 and references therein). In 2017, IceCube detected a high-energy neutrino with a high probability of being of astrophysical origin from the direction of the blazar TXS 0506+056. At the time of the neutrino arrival, the source exhibited enhanced gamma-ray activity. The chance probability of such an association was excluded at the 3σ level, making TXS 0506+056 the first candidate nonstellar astrophysical neutrino source (IceCube Collaboration 2018a) and supporting the theoretically anticipated correlation between neutrino and gamma-ray emissions (see M. Ahlers & F. Halzen 2018 and references therein).

Beyond the observation of TXS 0506+056, IceCube has reported 4.2σ evidence for TeV neutrino emission from the nearby active galactic nucleus (AGN) NGC 1068 (IceCube Collaboration 2022), notably without corresponding gamma-

ray emission at similar energies. In fact, the GeV gamma-ray emission detected by Fermi-LAT (A. A. Abdo et al. 2010) from NGC 1068 is likely dominated by star formation activity (B. Eichmann et al. 2022), and its photon flux is lower than the observed neutrino flux. Its neutrino emission is likely produced in the immediate vicinity of the supermassive black hole (SMBH) that powers the AGN, most plausibly within the AGN’s corona (Y. Inoue et al. 2019; K. Murase et al. 2020; K. Murase 2022; P. Padovani et al. 2024). The corona—a plasma of extremely hot electrons ($\sim 10^9$ K)—is responsible for the characteristic X-ray emission of AGN (e.g., P. Padovani et al. 2017 and references therein). In this environment, ultraviolet photons are Compton upscattered by hot electrons to keV energies, producing X-rays (E. P. T. Liang 1979). These coronal X-ray photons serve as effective targets for photomeson production when interacting with protons of energies around 100 TeV, leading to the generation of 1–10 TeV neutrinos as observed by IceCube (P. Padovani et al. 2024). Remarkably, NGC 1068 stands out as one of the X-ray-brightest AGN in the sky (A. Marinucci et al. 2016; C. Ricci et al. 2017).

Building on the evidence of neutrino emission from NGC 1068 and its interpretation as originating in the AGN corona, we present the results of dedicated neutrino searches aimed at further investigating this phenomenon. In Section 2, we describe the dataset, obtained by extending the one used in IceCube Collaboration (2022) with $\sim 50\%$ more statistics, resulting in a total of 13.1 yr of ν_μ -induced events from the northern sky. In Section 3, we introduce the analysis framework and the performed analyses, and in Section 4, we present the results. A survey of neutrino emission in the Northern Hemisphere reveals that the most significant excess remains spatially consistent with the position of NGC 1068. In addition, we examine a legacy list of 110 gamma-ray sources selected from the fourth Fermi Large Area Telescope (4FGL-DR2) catalog (S. Abdollahi et al. 2020), testing for neutrino emission from both individual objects and the list as a whole. Motivated by the gamma-ray-obscured and X-ray-bright nature of NGC 1068, we also compile a new, targeted sample of 47 AGN selected among the intrinsically brightest objects in the BAT AGN Spectroscopic Survey (BASS) catalog (M. J. Koss et al. 2022) and perform dedicated searches for

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neutrino emission from this population. The findings of the analyses are discussed in Section 5.

A previous search by the IceCube Collaboration (R. Abbasi et al. 2025a) investigated potential neutrino emission from a list of Seyfert galaxies, motivated by the neutrino excess from NGC 1068. That study examined 27 Seyfert galaxies selected as the intrinsically brightest in the 2–10 keV band from the BASS catalog but did not find a statistically significant neutrino signal. This work updates the candidate source selection, focusing on X-ray AGN in the same catalog that are especially bright in the 20–50 keV band. The harder X-ray emission is more robust against both line-of-sight absorption and spectral features associated with low-temperature AGN coronae. Despite substantial overlap with the previous source list, the newly adopted selection criteria and increased statistics provide evidence for neutrino emission from a population of X-ray-bright AGN in the northern sky.

2. Neutrino Dataset

The IceCube Neutrino Observatory is comprised of 86 strings, each instrumented with 60 digital optical modules (DOMs). The DOMs record the Cherenkov light produced by charged particles traveling through the Antarctic ice and are deployed at depths between 1450 and 2450 m (M. G. Aartsen et al. 2017). This full configuration, referred to as IC86, has been in continuous operation since 2011 May 13. Prior to that, from 2010 June 1, IceCube operated with 79 deployed strings (IC79), a near-complete geometry missing only seven strings. Of these, two belong to the DeepCore subarray (M. G. Aartsen et al. 2017), while the remaining five form an outer line on one edge of the array (see Appendix A). The IC79 dataset, comprising approximately 312 days of livetime, was reprocessed for this work using improved calibration and event filtering (R. Abbasi et al. 2021), ensuring consistency with the IC86 data. The total dataset used in this study corresponds to 13.1 yr of data, combining 0.9 yr from IC79 and 12.2 yr from IC86, collected up to 2023 November 28 and totaling approximately 1 million events.

In this work, we focus on muons produced in charged-current interactions of muon neutrinos. As they traverse the detector, these muons emit Cherenkov light, creating track-like signatures. *Tracks* are characterized by subdegree angular resolution at energies above ~ 1 TeV (see Appendix B), which makes this signature optimal for astrophysical source searches. We select events reconstructed with a decl. (δ) between -5° and 90° , encompassing the analyzed sky region ($-3^\circ < \delta < 81^\circ$). By focusing on the northern sky, the selection effectively suppresses atmospheric muon contamination, as such muons are absorbed when crossing the Earth before reaching the detector. Overall, this selection yields a neutrino purity of 99.8% (R. Abbasi et al. 2022). The selected ν_μ -induced sample is predominantly composed of muons produced by atmospheric neutrinos and a subdominant unresolved diffuse astrophysical component, which contributes $\sim 0.4\%$ to the total statistics of the sample, assuming the astrophysical flux measured by R. Abbasi et al. (2022). Both are treated as backgrounds to the search for pointlike neutrino emission. For each detected muon, we reconstruct the analysis observables: the muon track direction, \mathbf{d}_μ ; energy, E_μ ; and angular uncertainty, σ_μ . The 13 yr dataset used in this analysis was processed following the same procedures as in IceCube Collaboration (2022, 2025a). While including more IC 86 data

is straightforward, incorporating IC 79 data requires additional care due to the slightly different detector geometry. IC 79 data and Monte Carlo (MC) simulations were reprocessed with the same reconstruction algorithms as IC 86, and machine learning models (IceCube Collaboration 2022) trained on IC 86 simulations were successfully applied to IC 79 events, achieving consistent performance across both configurations (C. Bellenghi et al. 2023). In Figure 1, we show the agreement between the experimental data and the simulations for all relevant observables used in the analysis for both IC 79 and IC 86 configurations. Across all considered cases, the level of agreement is excellent, with maximal discrepancies of only a few percent in the high-statistics regions.

3. Analysis Method

In the search for neutrino point sources, we employ an unbinned maximum-likelihood method, along with likelihood ratio hypothesis testing, as described in J. Braun et al. (2008). The background-only hypothesis consists purely of contributions from the atmospheric and astrophysical diffuse emissions, while the signal hypothesis assumes an additional accumulation of astrophysical neutrinos clustered around a pointlike source. The likelihood function for the total number of events in the sample N is defined as

$$\mathcal{L}(n_s, \gamma, \mathbf{d}_{\text{src}}|\mathbf{x}) = \prod_{i=1}^N \left\{ \frac{n_s}{N} \cdot f_S(\mathbf{x}_i|\gamma, \mathbf{d}_{\text{src}}) + \left(1 - \frac{n_s}{N}\right) \cdot f_B(\mathbf{x}_i) \right\}, \quad (1)$$

where f_S and f_B are the signal and background probability density functions (pdfs), respectively. These pdfs are joint distributions that can be factorized into an energy and a spatial term, which together define the probability that a given event arises from a signal source or from background. An event \mathbf{x}_i is defined by the vector of observables ($E_{\mu,i}, \mathbf{d}_{\mu,i}, \sigma_{\mu,i}$) (defined in Section 2), \mathbf{d}_{src} denotes the direction of the analyzed source location in the sky, n_s is the mean number of signal events, and γ is the spectral index of an assumed unbroken power-law spectrum of $\Phi = \Phi_0 \cdot (E/E_0)^{-\gamma}$, with neutrino energy E and flux normalization Φ_0 proportional to n_s at pivot energy $E_0 = 1$ TeV. The free parameters of the likelihood are the mean number of signal events, n_s , and the spectral index, γ . The excellent agreement between data and simulations (see Figure 1) allows us to derive all observables' pdfs directly from the MC using the kernel density estimation (KDE) method (A. Poluektov 2015), as done in previous works (IceCube Collaboration 2022, 2025a). For this analysis, we use updated KDE pdfs for IC86—now based on ~ 2.5 times more simulated events, including muons from τ decays in ν_τ interactions—and construct them for IC 79 for the first time. The two samples are then combined in the analysis by multiplying their likelihood functions, weighted according to their respective detection efficiencies.

The test statistic (TS) adopted in this work is the negative logarithm of the likelihood ratio between the background and the signal hypotheses, with the signal likelihood maximized

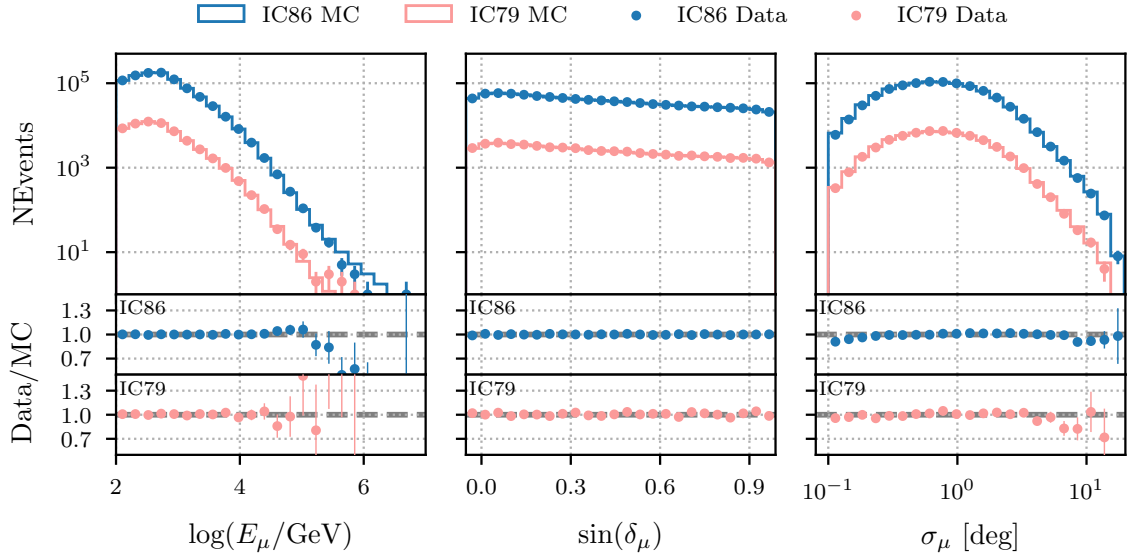


Figure 1. Agreement between the distributions of the analysis observables for experimental data (dots) and simulations (solid lines) for both IC 79 (pink) and IC 86 (blue). From left to right, we show the reconstructed muon energy, E_μ , which ranges from 100 GeV to a few PeV; the sine of the reconstructed decl., $\sin(\delta_\mu)$; and the reconstruction quality estimator on the muon track direction, σ_μ . In the lower panels, we display the ratios between the experimental data and the simulations for both detector configurations. The simulations show the sum of the atmospheric and diffuse astrophysical neutrino flux components. The atmospheric neutrino flux assumes the GST model (T. K. Gaisser et al. 2013) for the primary cosmic-ray spectrum and *Sybill2.3c* as the interaction model (F. Riehn et al. 2018). The diffuse astrophysical component assumes a single power law with spectral index $\gamma_{\text{astro}} = 2.37$ and normalization $\phi_{\text{astro}} = 1.44 \times 10^{-18} \text{ GeV}^{-1} \text{ cm}^{-2} \text{ s}^{-1} \text{ sr}^{-1}$ (R. Abbasi et al. 2022).

over the two source parameters, n_s and γ ,

$$\begin{aligned} \text{TS} &= -2 \log \left(\frac{\mathcal{L}(n_s = 0 | \mathbf{x})}{\mathcal{L}(\hat{n}_s, \hat{\gamma}, \mathbf{d}_{\text{src}} | \mathbf{x})} \right) \\ &= 2 \sum_i^N \log \left\{ \frac{\hat{n}_s}{N} \left(\frac{f_S(\mathbf{x}_i | \hat{\gamma}, \mathbf{d}_{\text{src}})}{f_B(\mathbf{x}_i)} - 1 \right) + 1 \right\}, \end{aligned} \quad (2)$$

where \hat{n}_s and $\hat{\gamma}$ are the parameter values that maximize the likelihood. The last equality in Equation (2) uses Equation (1) for the signal and background cases and highlights the dependence of the TS on the ratio of the signal and background pdfs, $f_S(\mathbf{x}_i | \hat{\gamma}, \mathbf{d}_{\text{src}}) / f_B(\mathbf{x}_i)$, also referred to as S/B . The TS is used to assess the significance of the signal hypothesis relative to the background hypothesis: higher TS values indicate that the observed data are less compatible with the background-only hypothesis. The likelihood optimization is performed, limiting the parameters in the ranges $[0, 1000]$ for n_s and $[0.6, 4.4]$ for γ , unless the source hypothesis assumes another spectral shape (e.g., a fixed spectral index or a spectrum different from the power law).

This approach is applied to all the results presented in Section 4. First, we report on an unbiased search for neutrino emission across the northern sky. This method consists of dividing the sky into a grid of points and testing each of them for astrophysical neutrino emission using the aforementioned maximum-likelihood ratio method. Next, we present results from a targeted search for sources within two predefined lists: one containing 110 gamma-ray emitters and the other comprising 47 X-ray-bright AGN. For each list, we perform two tests.

1. A *catalog search*, where we evaluate the significance of neutrino emission with respect to the background for each source individually.

2. A *binomial test*, to identify ensembles of sources that may show weak individual signals but yield a significant collective excess. For each of the N sources, we compute the local p -value and sort them in ascending order. For each rank $k = 1, \dots, N$, we then calculate the probability of observing k or more sources with p -values below the k th smallest value in the list in a background-only scenario. In mathematical terms, the binomial p -value is defined as

$$p_{\text{binom}} = \sum_{i=k}^N \binom{N}{i} p_k^i (1 - p_k)^{N-i}. \quad (3)$$

The outcome of the test consists of the number of sources k that provide the most significant collective excess above the background, i.e., the smallest binomial p -value among all the tested cases (ATLAS Collaboration 2020).

4. Results

4.1. Survey of the Northern Sky

We performed an unbiased search for neutrino emission across the northern sky by dividing it into a grid of pixels of $\approx 0.052 \text{ deg}^2$ using Healpix (K. M. Górski et al. 2005) with resolution parameter $\text{NSIDE} = 256$. At each pixel, we optimized the likelihood ratio and extracted the best-fit values \hat{n}_s , $\hat{\gamma}$ and the TS. The TS is then converted into a p -value by comparing it to the distribution of TS values from background-only (atmospheric and diffuse astrophysical neutrinos) simulations at the corresponding decl. (see Appendix C for details about the significance calculation). To refine the localization of the most significant excess, we identified the 20 most significant hot spots⁶⁹ and scanned these regions using a finer grid of pixels of $\approx 0.003 \text{ deg}^2$ obtained by increasing the NSIDE

⁶⁹ Defined as the lowest p -value pixel within a 1° radius.

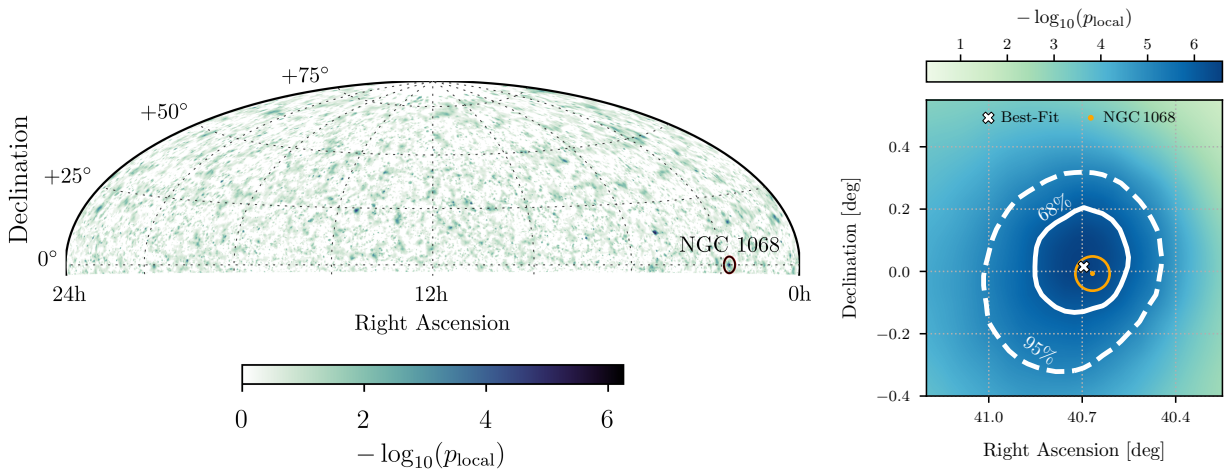


Figure 2. On the left, we show the p -value map of the northern sky, obtained under the hypothesis of a free spectral index. The map is shown in equatorial coordinates on a Hammer–Aitoff projection. The color bar represents the local significance of each pixel in the sky, and we highlight the location of the strongest emission, found to be spatially compatible with the Seyfert II galaxy NGC 1068. On the right, we show a zoom-in of the hottest spot. The white cross represents the best-fit location, the solid (dashed) line represents the 68% (95%) uncertainty contour of the excess, and the orange dot and circle represent, respectively, the source location and its optical size (G. Paturel et al. 2003).

resolution parameter to 2048. The lowest p -value from this refined search defines the final hottest spot in the sky. As in the previous analysis (IceCube Collaboration 2022), this search was performed under three spectral assumptions: with γ free and fixed at $\gamma = 2.0$ and $\gamma = 2.5$. We report here the most significant excess out of the three: the one obtained for the free spectral index search. The other two result in most significant excesses at different locations, as detailed in Appendix D.1. Figure 2 shows the p -value map under the free γ hypothesis, with the coordinates and best-fit parameters of the most significant excess listed in Table 1.

Consistent with IceCube Collaboration (2022), the analysis confirms that the most significant pixel lies within the optical extent of NGC 1068, at an angular offset of 0.04° from its optical center (see Figure 2, right panel). The excess reaches a local significance of 5.0σ , which corresponds to a global significance of 1.4σ after accounting for the look-elsewhere effect from scanning the entire observable northern sky and testing three spectral hypotheses. The location of the most significant spot remains fully compatible with NGC 1068, and the best-fit point now aligns even more closely with the source than in the previous result, which placed it 0.11° away (IceCube Collaboration 2022). However, the spatial likelihood contours have slightly widened due to the relatively large angular uncertainty of several newly detected low-energy events (with reconstructed energies mostly between 100 GeV and 3 TeV) associated with the source. The new events populate a part of the spectrum where the atmospheric background is more prominent and, together with statistical fluctuations, may also explain the reduction in global significance from 2.0σ (IceCube Collaboration 2022) to 1.4σ . Despite this, the results remain consistent with the expected evolution of a steady neutrino emitter scenario for NGC 1068.

4.2. Search from a List of Gamma-Ray Emitters

All-sky searches suffer from a large look-elsewhere effect due to the vast number of independent tests. To mitigate this penalty factor and improve sensitivity to weaker but persistent sources, we perform complementary searches on predefined lists of candidate emitters. Specifically, we run a catalog

search and a binomial test on the legacy list of 110 gamma-ray sources, as employed in IceCube Collaboration (2022). Using 13.1 yr of data, we confirm NGC 1068 as the most significant source in the list, now with a global significance of 4.0σ , after accounting for having tested 110 candidate sources. The best-fit spectrum has a flux normalization at 1 TeV of $\hat{\phi}_0 = 4.7_{-1.3}^{+1.1} \times 10^{-11} \text{ TeV}^{-1} \text{ cm}^{-2} \text{ s}^{-1}$ and a spectral index $\hat{\gamma} = 3.4 \pm 0.2$.⁷⁰

As previously noted, the global significance has slightly decreased compared to IceCube Collaboration (2022), primarily due to a shift toward lower neutrino energies, where the atmospheric background is more prominent. As a result, the best-fit spectral index has slightly softened, increasing from $\hat{\gamma} = 3.2$ to 3.4 (see Figure 3). However, the two spectral indices remain largely compatible within their uncertainties (see Figure 12 for a comparison of the best-fit spectrum measured in this work with the one previously reported in IceCube Collaboration 2022). The fitted mean number of signal events has increased from $\hat{n}_s = 79$ to 102 (both affected by a 1σ statistical uncertainty of $\sim 25\%$ as derived from the likelihood contours), consistent with a steady emission scenario.

To complement the single-source search, we also perform a binomial test on the same list, following the approach in IceCube Collaboration (2022). As in the all-sky analysis, the test was conducted for three spectral assumptions: free γ , $\gamma = 2.0$, and $\gamma = 2.5$. Here, we focus on the free- γ case, which yields the most significant result (see Appendix D.2 for the others). The most significant excess is found for 3 out of 110 sources: NGC 1068, PKS 1424+240, and TXS 0506+056, in ascending order of p -value (see left panel of Figure 4), consistent with previous results (IceCube Collaboration 2022). Their fit parameters and local p -values are listed in Table 1. The global significance of the excess is 3.0σ , slightly reduced compared to the 3.4σ reported in IceCube Collaboration (2022). This decrease arises because the time-integrated significance of TXS 0506+056 diminishes as additional data increase the background without new signal events, consistent with its reported time-variable behavior

⁷⁰ The 1σ statistical uncertainties on each flux parameter are derived from the one-dimensional profile likelihoods fixing the other parameter to its best-fit value.

Table 1
Summary of the Results

		R.A.	Decl.	\hat{n}_s	$\hat{\gamma}$	$-\log_{10} p_{\text{local}}$	$-\log_{10} p_{\text{global}}$
Northern sky survey	Hottest spot	40.69	0.02	102.6	3.4	6.6 (5.0 σ)	1.1 (1.4 σ)
110 Gamma-Ray Sources							
Most significant source	NGC 1068	40.67	-0.01	102.2	3.4	6.6 (5.0 σ)	4.5 (4.0 σ)
Binomial test	3 sources	4.8 (4.2 σ)	2.9 (3.0 σ)
Other sources in binomial excess	PKS 1424+240	216.76	23.80	96.3	3.6	4.1 (3.8 σ)	...
	TXS 0506+056	77.36	5.69	4.9	1.9	3.4 (3.3 σ)	...
47 X-Ray-bright AGN							
Most significant source	NGC 7469	345.82	8.87	5.5	1.9	4.1 (3.8 σ)	2.1 (2.4 σ)
Binomial test	11 sources	4.8 (4.2 σ)	3.3 (3.3 σ)
Other sources in binomial excess	NGC 4151 [†]	182.64	39.41	27.6	2.7	2.9 (3.1 σ)	...
	CGCG 420-015*	73.36	4.06	35.3	2.7	2.4 (2.7 σ)	...
	Cygnus A*	299.87	40.73	3.4	1.6	2.2 (2.5 σ)	...
	LEDA 166445	42.68	54.70	57.1	4.4	1.8 (2.1 σ)	...
	NGC 4992	197.27	11.63	27.3	2.9	1.6 (2.0 σ)	...
	NGC 1194*	45.95	-1.10	43.2	4.4	1.5 (1.8 σ)	...
	Mrk 1498	247.02	51.78	39.9	3.6	1.4 (1.7 σ)	...
	MCG +4-48-2*	307.15	25.73	36.7	3.2	1.4 (1.7 σ)	...
	NGC 3079	150.49	55.68	33.8	3.6	1.3 (1.7 σ)	...
	Mrk 417	162.38	22.96	4.4	2.0	1.3 (1.6 σ)	...

Note. We report the most significant results from all tests performed in this work: the northern sky survey, the catalog searches, and the associated binomial tests. The type of test is indicated in the leftmost column, with the corresponding results shown to the right. For each source, we report the equatorial coordinates (J2000 equinox) and the likelihood search results: number of signal events \hat{n}_s , spectral index $\hat{\gamma}$, local $-\log_{10} p_{\text{local}}$, and global $-\log_{10} p_{\text{global}}$ (after accounting for all performed trials) p -values with their corresponding significance in brackets. The sources highlighted with an asterisk were included in the R. Abbasi et al. (2025a) selection.

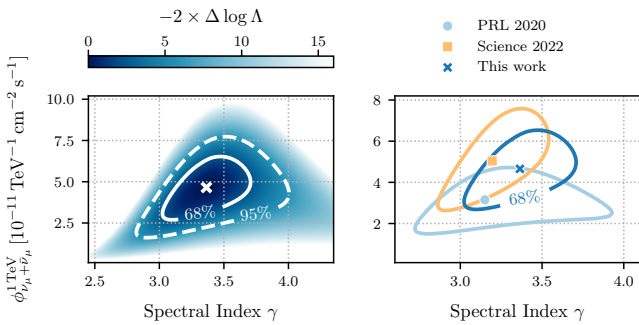


Figure 3. Left: profile likelihood scan for the flux parameters of NGC 1068. We highlight in white the best-fit result (cross) and the 68% and 95% uncertainty contours (solid and dashed lines), derived from Wilks’s theorem (S. S. Wilks 1938). The color scale indicates the log-likelihood ratio difference between any point in the parameter space and the overall best-fitting point. Right: best-fit and 68% contour comparison between this work (blue), IceCube Collaboration 2022 (yellow), and M. G. Aartsen et al. 2020 (light blue)—the first IceCube analysis that found NGC 1068 as the most significant source in a list, with a global significance of 2.9 σ . All contours include only statistical uncertainties.

(IceCube Collaboration 2018a, 2018b). Surprisingly, no additional gamma-ray sources from the list contribute significantly to the binomial excess, despite the increased exposure, highlighting the continued lack of broad correlation between gamma-ray brightness and neutrino emission in the current sample.

4.3. Search from a List of X-Ray-bright Active Galactic Nuclei

The emergence of NGC 1068 as a neutrino source motivated the dedicated follow-up presented here. As one of the closest and brightest Seyfert galaxies, this association suggested that other AGN with similar properties might also emit neutrinos. A

first attempt in this direction was made in R. Abbasi et al. (2025a), which reported a 2.7 σ binomial excess from a set of 27 X-ray-bright Seyfert galaxies selected from the BASS catalog.

Building on this, we constructed an updated list of nearby X-ray-bright AGN as candidate neutrino sources. We selected sources classified as Seyfert galaxies in the BASS catalog that exhibit intrinsic hard X-ray fluxes between 20 and 50 keV of at least 20% of that of NGC 1068. Although Seyfert galaxies are typically radio-quiet, the BASS classification is based solely on optical spectroscopy, and some nearby radio-loud galaxies may therefore appear in the sample. Since our focus is the correlation between neutrino production and X-ray brightness, such contamination is not a concern. The energy band of 20–50 keV was chosen to ensure robustness against line-of-sight absorption due to obscuring material in the circumnuclear environment. Absorption becomes negligible above ~ 10 keV for column densities up to $\log N_{\text{H}} \approx 23.5$ (M. J. Koss et al. 2022), making this range a more reliable proxy for intrinsic source power than the softer 2–10 keV band used in R. Abbasi et al. (2025a). Using intrinsic fluxes up to 50 keV also minimizes the potential bias against AGN with low-temperature coronae, whose spectra can cut off below 100 keV (A. C. Fabian et al. 2017), making them appear fainter in broader bands such as 14–195 keV. While the photon energies relevant for hadronic neutrino production are expected to lie in the softer ~ 1 –10 keV range, that band is highly susceptible to absorption. The 20–50 keV range therefore serves as a practical compromise: it is high enough to ensure robustness against absorption but not so high as to exclude plausible neutrino sources due to coronal spectral features. The resulting sample comprises 47 galaxies, excluding NGC 1068 itself, and

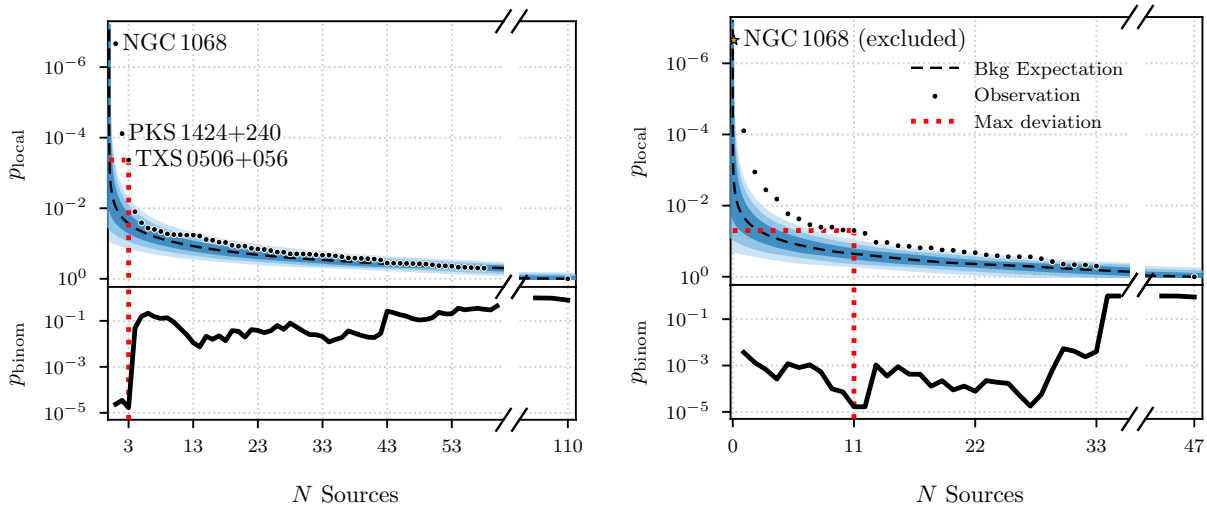


Figure 4. Visualization of the binomial test excess for the list of 110 gamma-ray emitters (left) and 47 X-ray-bright AGN (right). In the upper panels, we show the local p -values for each source, ordered from the lowest to the highest (black points), the background expectation (dashed black line), and its 1σ , 2σ , and 3σ Poissonian uncertainties (shaded blue bands). In the lower panels, we report the binomial probability for each subset of sources. The most significant excess is highlighted by dotted red lines. To preserve readability, the plot is truncated at the point where sources reach a p -value of 1.0, beyond which the binomial probability is also equal to 1.0.

includes both Seyfert I and II types, thereby spanning a range of obscuration levels and ensuring a representative set of potential neutrino-emitting environments.

The selected X-ray-bright, nonblazar AGN were tested under two spectral assumptions: a power law with free spectral index γ and source-specific spectra from the core–corona model of A. Kheirandish et al. (2021), as in R. Abbasi et al. (2025a). Here, we report only the results for the power-law case, which resulted in the highest statistical significance (see Appendix D for the core–corona model results).

The most significant individual source, excluding NGC 1068, is NGC 7469, with a local significance of 3.8σ and a global significance of 2.4σ , accounting for 47 tested sources and two spectral hypotheses. The likelihood fit for this source returned a very hard spectral index of $\hat{\gamma} = 1.9$, and the excess is fully dominated by two high-energy events, already identified by IceCube as likely astrophysical and issued as neutrino alerts: IC 220424A⁷¹ and IC 230416A⁷² (see Appendix D.3 for more details and G. Sommani et al. 2025 for an independent study of this neutrino association). Most interestingly, the binomial test of this list reveals a collective excess from 11 out of 47 sources, with a local significance of 4.2σ and a global significance of 3.3σ after accounting for having tested different p -value thresholds and two spectral hypotheses (see Appendix D for more details). Among the 11 contributing sources, each has a local p -value of $p_{\text{local}} \leq 6\%$, while MC simulations show that background-only realizations typically yield about five such sources. The contributing sources are listed in Table 1 and shown in the right panel of Figure 4. Including NGC 1068 in the list (yielding 48 sources) increases the excess to 12 sources with a local significance of 4.5σ but does not qualitatively alter the result.

5. Discussion

Thanks to IceCube’s continuous data-taking and remarkable long-term stability, we now have a view of the neutrino sky

with unprecedented sensitivity (see Appendix B). This study builds on that progress, presenting updated results from 13.1 yr of observations. While TXS 0506+056 remains the only case claimed by IceCube linking high-energy neutrinos to a gamma-ray source, our results provide new evidence supporting an association between neutrino emission and X-ray-bright AGN. Specifically, we find a population-level excess from a sample of 47 X-ray-bright, nonblazar AGN. Most of these sources do not exhibit high-energy gamma-ray emission. The few exceptions include NGC 1068, whose GeV emission is likely related to starburst activity (M. Ajello et al. 2020), and NGC 4151, possibly associated with 0.1–100 GeV gamma-ray emission at the 5.5σ level, perhaps connected to its ultrafast outflow (UFO; E. Peretti et al. 2025). The interpretation of NGC 4151 is further complicated by the presence of nearby blazars (A. Omeliukh et al. 2025; E. Peretti et al. 2025), but in either case—whether the gamma rays originate in the UFO or in the blazars—theoretical expectations indicate that the observed neutrino flux cannot be explained (P. Padovani et al. 2024; A. Omeliukh et al. 2025; E. Peretti et al. 2025), strengthening the case for neutrino generation in the X-ray corona, where TeV gamma rays are expected to be absorbed.

Among the contributing AGN, we find both Seyfert I (e.g., NGC 4151, NGC 7469) and Seyfert II (e.g., NGC 1068, CGCG 420–015) galaxies, suggesting that the level of nuclear obscuration does not significantly impact the likelihood of neutrino emission. In Figure 5, we show the best-fit spectra for the four most significant X-ray-bright AGN. Interestingly, the energy ranges of the contributing events vary across sources: NGC 7469’s excess is dominated by two >100 TeV events, while NGC 1068’s spectrum has shifted to lower energies, now constrained between 0.2 and 20.6 TeV at 95% C.L. (see Appendix D for more details on the energy range calculation), though still consistent with previous measurements.

The reduced significance of the NGC 1068 signal—despite the increased statistics—could result from statistical fluctuations, spectral variability, or limitations of the assumed power-law model. An alternative explanation may be time variability (P. Dave & I. Taboada 2023), which was not tested in this work due to the focus on steady emission scenarios. Future

⁷¹ <https://gcn.nasa.gov/circulars/31942>

⁷² <https://gcn.nasa.gov/circulars/33633>

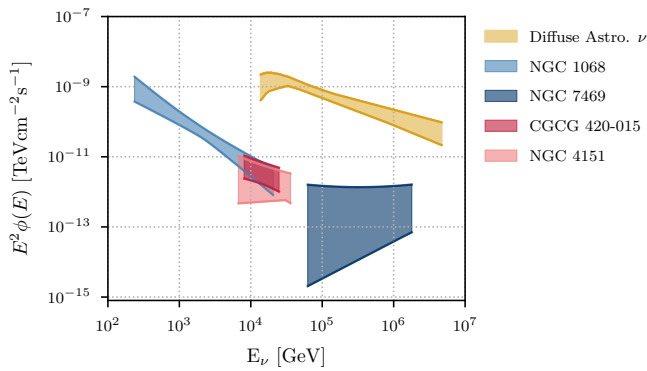


Figure 5. Best-fit neutrino all-flavor fluxes with their 1σ uncertainty and constrained in the 95% C.L. energy range of the top four sources in the list of X-ray-bright AGN. The yellow band represents the most recent measurement of the diffuse astrophysical neutrino flux. The energy range is constrained at the 90% C.L., while the flux normalization is shown with a 68% C.L. (R. Abbasi et al. 2025c).

time-dependent studies will be essential to explore this possibility. Improved spectral constraints, particularly below 100 GeV, will also be crucial for future observations.

Taken together, our findings suggest that X-ray-bright AGN are promising contributors to the observed diffuse extragalactic neutrino flux. This builds on a growing body of independent results: a 2.7σ binomial excess from NGC 4151 and CGCG 420–015 (R. Abbasi et al. 2025a), a 2.9σ excess from NGC 4151 in a hard X-ray AGN study (R. Abbasi et al. 2025b), and a 3.0σ signal from a stacking analysis of 14 Seyfert galaxies in the southern sky (S. Yu 2025). While current significance levels remain marginal, the emerging picture points to a new class of potential neutrino sources: X-ray-bright AGN. These results yield growing support for theoretical models of neutrino production in the coronal regions surrounding SMBHs. At the same time, the very high-energy events associated with NGC 7469 cannot be accommodated by the custom core–corona model (A. Kheirandish et al. 2021) tested in R. Abbasi et al. (2025a) and in this work (see Appendix D). This suggests that X-ray-bright AGN may not all share the same cosmic-ray acceleration or neutrino production conditions and that further measurements will be essential. Future data from IceCube and next-generation detectors (S. Adrián-Martínez et al. 2016; M. Agostini et al. 2020; M. G. Aartsen et al. 2021), combined with multi-wavelength observations, will be critical to confirm this connection and further unravel the components of the diffuse astrophysical neutrino flux.

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Appendix A IC 86 and IC 79 Geometry Comparison

The results presented in this Letter were obtained using a dataset comprising data recorded with two different detector configurations, as explained in Section 2.

In Figure 6, we show a top view of the IceCube array, highlighting the two configurations with different colors. Apart from two strings belonging to the DeepCore array, the only difference between the IC 86 and IC 79 arrays consists of five strings on one edge of the detector. Due to the minimal changes in the overall geometry, data recorded in this configuration have minimal losses in terms of reconstruction quality and resolution compared to data recorded using the full IceCube array. In the left panel of Figure 7, we display the angular resolution of the events in the sample for both configurations. Both the medians and the central 1σ quantile of the angular resolutions show minimal differences over the whole energy range, proving the comparable quality of IC 79 and IC 86 data.

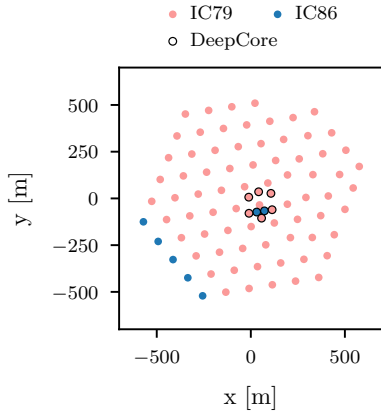


Figure 6. Disposition of the strings in the IceCube coordinate system. The blue dots mark the last seven strings added, which changed the detector from the IC 79 to the IC 86 configuration. For this work, only the five strings added to the edge of the detector are relevant, as the other two are part of the DeepCore detector (dots outlined in black), excluded from the analysis.

Appendix B Analysis Performance

As described in Section 3, the likelihood analysis used in this work incorporates an improved description of the spatial and energy pdfs, which are parameterized from simulations using the KDE method (A. Poluektov 2015). The right panel of Figure 7 shows two spatial pdfs obtained for different spectral indices. Because the signal pdfs are derived directly from simulation, they accurately capture the relationship between reconstructed event properties and the underlying signal hypothesis. In particular, as seen in the right panel of

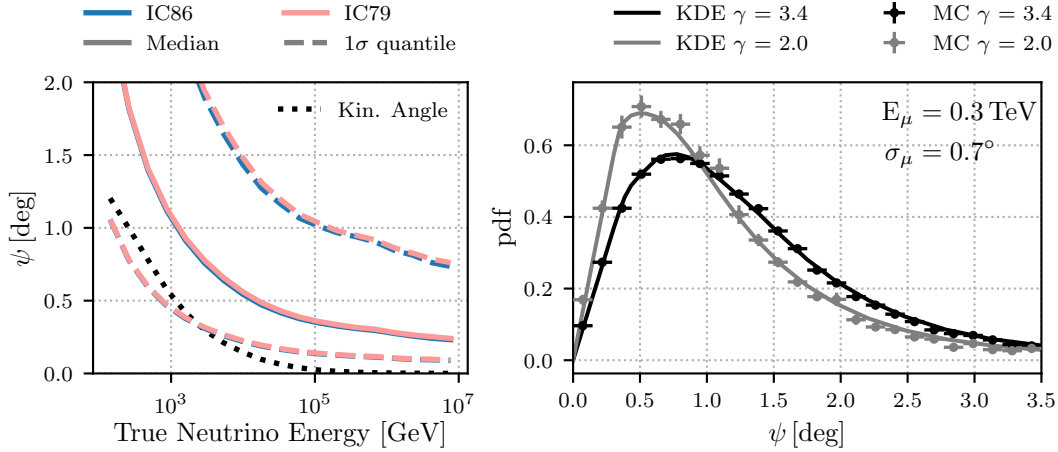


Figure 7. Left panel: angular resolution for both IC 86 (blue) and IC 79 (pink) data in the neutrino energy range employed in the analysis. The solid line represents the median, while the dashed lines denote the central 1σ quantiles (16% and 84% quantiles). The black dotted line marks the median kinematic angle, i.e., the median angular separation between the incoming neutrino and the muon produced in the interaction. Right panel: two examples of KDE pdfs for two spectral assumptions (solid lines) superimposed with the MC data (dots). This example shows the pdfs for a muon energy of 300 GeV and an uncertainty on the angular reconstruction of 0.7° .

Figure 7, events with the same reconstructed muon energy (E_μ) and angular uncertainty (σ_μ) can have different likelihoods of originating from a given source depending on their angular separation ψ and the given spectral assumption. This effect arises because lower-energy neutrinos, which dominate for softer spectra, produce muons with larger kinematic angles from the parent particle (see the left panel of Figure 7), leading to a broader spatial distribution. Hence, incorporating the full simulation-based dependence of ψ on energy and spectral shape improves our ability to accurately characterize the spectral emission of a source.

In addition to the improved description of the pdfs used in the likelihood analysis, this work also benefits from a significant increase in available statistics. The dataset analyzed here spans over 13.1 yr of uniformly processed IceCube data, resulting in a final event sample of approximately 1 million events, a $\sim 50\%$ increase compared to the IceCube Collaboration (2022) analysis. To quantify the analysis capability of observing an astrophysical signal, we use the *sensitivity* and the 5σ *discovery potential* neutrino fluxes. The sensitivity is defined as the average flux a point source would need to exceed the median TS value obtained from background-only simulations in 90% of the cases. The 5σ discovery potential is instead defined as the average flux needed to exceed the 5σ quantile of the background TS distribution in 50% of the cases. In Figure 8, we show the sensitivity and the 5σ discovery potential fluxes of this analysis, comparing them to those of IceCube Collaboration (2022). This work shows an increased sensitivity to astrophysical signals up to $\sim 30\%$ across the whole analyzed decl. range for both hard ($\gamma = 2.0$) and soft ($\gamma = 3.4$) spectral indices.

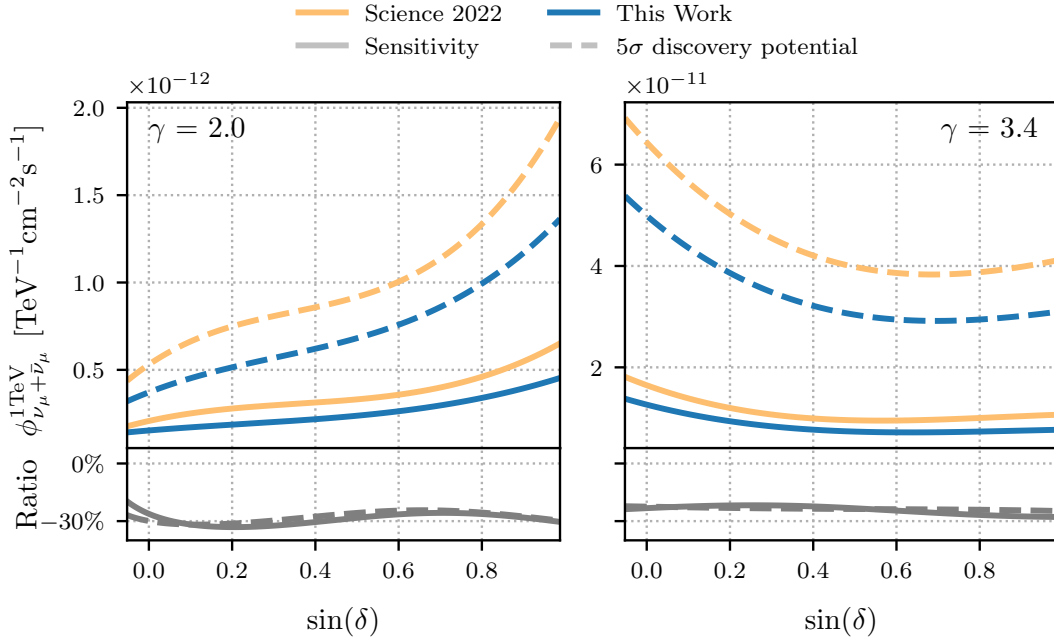


Figure 8. In the upper panels, we show the sensitivity (solid) and discovery potential (dashed) fluxes for a spectral index $\gamma = 2.0$ (left) and $\gamma = 3.4$ (right) for this analysis (blue) and IceCube Collaboration (2022; yellow). In the lower panels, we show the percentage ratio between the two. In both cases, the sensitivity and discovery potential improve by up to 30% compared to previous results.

Appendix C

Significance Calculation for the Likelihood Ratio Test

The calculation of the sensitivity, 5σ discovery potential, or significance of an observed TS value requires the determination of the TS distribution under the null hypothesis. Wilks's theorem states that, under suitable regularity conditions, the likelihood ratio TS asymptotically follows a chi-squared (χ^2) distribution (S. S. Wilks 1938). However, in the analysis presented in this work (see Section 3), the signal-strength parameter lies on the boundary of the parameter space under the null hypothesis ($n_s = 0$), thereby violating the required regularity conditions. In this situation, the asymptotic χ^2 approximation is formally invalid (see, e.g., R. Protasov et al. 2002). For this reason, no χ^2 distribution is used anywhere in this work to compute local or global significances for rejecting the null hypothesis. Instead, the TS distribution under the null hypothesis is empirically determined using a large ($\sim 10^6$, or more if needed) number of background-only pseudoexperiments, each analyzed with the same likelihood ratio maximization procedure used for the experimental data. Figure 9 shows the background TS distributions at four representative declinations (including that of NGC 1068). The red dashed curve indicates the half- χ^2 distribution with 2 degrees of freedom, corresponding to the difference in the number of free parameters between the signal (two) and background (zero) hypotheses, only for illustrative comparison; the clear discrepancies demonstrate explicitly that the χ^2 approximation does not accurately describe the null distribution in this analysis.

The local p -value (i.e., prior to accounting for the look-elsewhere effect) for each tested sky direction is defined as the survival probability of the observed TS with respect to the empirically determined background-only TS distribution at that direction. An example is shown in the upper left panel of Figure 9, where the observed TS at the location of NGC 1068 (TS = 27.2) is compared to the TS distribution obtained from $\sim 10^8$ pseudoexperiments.

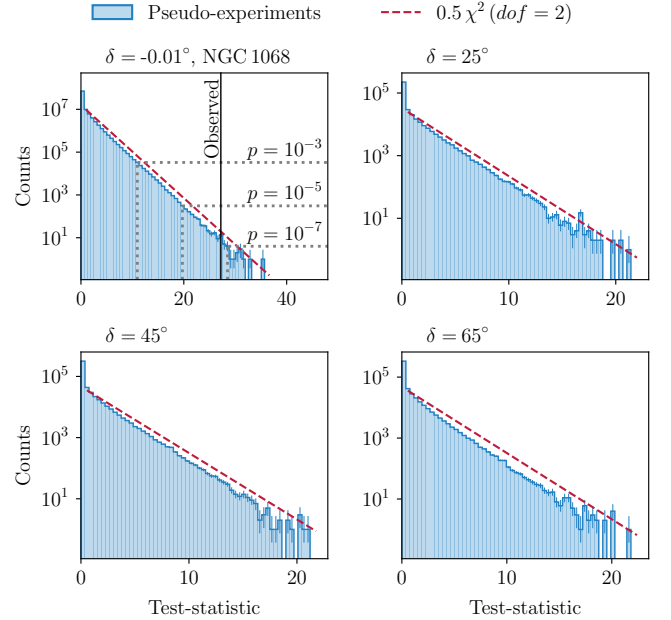


Figure 9. TS distributions under the null hypothesis at four representative declinations for the floating spectral index hypothesis. The blue histograms show the distributions obtained from background-only pseudoexperiments. The red dashed curve indicates the half- χ^2 with 2 degrees of freedom, shown only for reference. The upper left panel, in particular, shows the distribution of $\sim 10^8$ background-only TS values obtained at the location of NGC 1068 together with the observed TS (solid black line).

Global p -values (i.e., accounting for penalization due to the look-elsewhere effect) are likewise obtained from pseudoexperiments in both the all-sky scan and the binomial test. This is necessary because, in those two cases, the tested spectral hypotheses are correlated, and any analytical corrections to p -values become either inaccurate or overly conservative. In each result subsection of Appendix D, we specify the exact procedure used to evaluate the final global significance of the result.

Appendix D Additional Results

In addition to the results presented in Section 4, we conducted several other tests that did not yield statistically significant outcomes and are therefore not included in the main body of the Letter. Nevertheless, they have been taken into account to correct the global significance of our findings for the look-elsewhere effect. For completeness, we summarize these tests in the following subsections.

D.1. All-sky Searches with Fixed Spectral Index

Alongside the all-sky scan using a free spectral index, we performed two supplementary tests in which the spectral index γ was held fixed, as done in previous searches (IceCube Collaboration 2022) and to be consistent in the calculation of the global significance accounting for the look-elsewhere effect. We tested values of $\gamma = 2.5$ and $\gamma = 2.0$. The former was chosen due to its similarity to the best-fit astrophysical neutrino flux (R. Naab et al. 2023), while the latter reflects expectations for neutrino emission in scenarios governed by Fermi acceleration mechanisms.

The global significance of the result reported in Section 4.1 cannot be calculated analytically because the tested spectral hypotheses are highly correlated, as are adjacent tested locations in the sky. The correction is done in two steps. First, we account for testing multiple correlated locations in the sky by building the distribution of the minimum p -values from the pseudoexperiments for each spectral hypothesis. The fraction of realizations with smaller values than the observed p -values gives the scan-corrected significance. In the second step, the three scan-corrected p -values (one per all-sky scan) are compared with the distribution of the minimum p -values obtained by analyzing the same pseudodata under the three different signal hypotheses. The fraction of realizations resulting in a p -value smaller than the scan-corrected p -value yields the final globally corrected p -value.

Both fixed-index scans revealed hot-spot locations different from those found in the free-index case but with lower local significances. For the fixed $\gamma = 2$ case, the most significant hot spot was located at $(\alpha, \delta) = (77.01, 12.98)$, with a best-fit number of signal events of $\hat{n}_s = 16.8$ and a local significance of approximately 4.9σ . In the $\gamma = 2.5$ case, the hot spot appeared at $(\alpha, \delta) = (161.48, 27.32)$, with $\hat{n}_s = 34.3$ and a local significance of around 4.5σ . Neither of these hot spots is spatially compatible with any known source in the 4FGL-DR4 gamma-ray (J. Ballet et al. 2023) or eRASS1 X-ray (A. Merloni et al. 2024) catalogs. However, we note that the hottest spot from the $\gamma = 2.5$ scan is located 0.54° away from the 4FGL source NVSS J104516+275136.

As a cross-check of the statistical behavior of the sky scans, we verified that the distribution of the most significant p -values obtained from 2000 background-only pseudoexperiments follows the expectation from the extreme value theory (EVT; S. Coles 2001). For each spectral hypothesis, we constructed the distribution of the maximum $-\log_{10}(p)$ and fitted it with a Gumbel distribution, which is the limiting distribution for the maximum of a large number of independent random variables with exponentially decaying tails. An example fit for the floating spectral index hypothesis is shown in Figure 10. The Anderson–Darling goodness-of-fit test (T. W. Anderson & D. A. Darling 1954) shows no

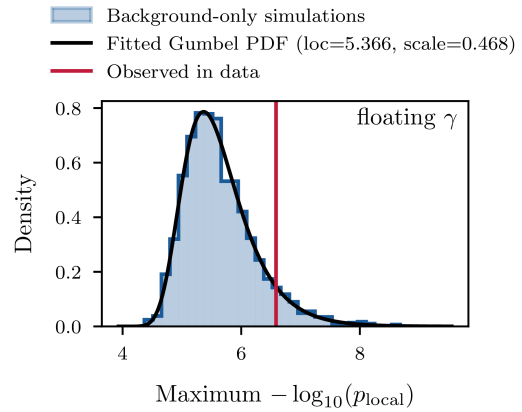


Figure 10. Distribution of the maximum local $-\log_{10}p$ -value from 2000 background-only sky scan simulations under the floating spectral index hypothesis. The histogram shows the distribution of the most significant fluctuation per pseudoexperiment, while the black curve represents the best-fit Gumbel pdf. The vertical red line indicates the maximum $-\log_{10}p$ observed in the experimental data. Equivalent agreement is found for the other spectral hypotheses.

evidence of deviations from the Gumbel hypothesis, with an A^2 statistic of 0.147 (the critical value to reject the Gumbel hypothesis at 95% C.L. is 0.754). Similar agreement is found for the other two spectral hypotheses. Scan-corrected p -values derived from the fitted Gumbel distributions are fully consistent with those obtained directly from the pseudoexperiments, confirming the expected asymptotic extreme-value behavior of the sky scan fluctuations. However, pseudoexperiments remain necessary for two reasons. First, the parameters of the Gumbel distribution cannot be determined analytically due to nontrivial correlations between adjacent pixels and must be calibrated from simulations. Second, the final global correction involves the minimum of only three strongly correlated, already scan-corrected p -values corresponding to the three tested spectral hypotheses. These global minima do not lie in an asymptotic extreme-value regime and are not described by a Gumbel distribution. The final multihypothesis correction can therefore be obtained exclusively from pseudoexperiments, while the EVT analysis serves as a validation of the scan-correction procedure.

D.2. Binomial Tests with Fixed Spectral Shapes

To preserve consistency with the analysis presented in IceCube Collaboration (2022), we also performed the binomial test for the fixed spectral index cases using our list of 110 gamma-ray emitters. As with the all-sky scans, these tests did not yield statistically significant results. For $\gamma = 2$, we found 13 sources (including NGC 1068, TXS 0506+056, and PKS 1424+240) out of 110 with a local p -value of approximately 0.007, while for $\gamma = 2.5$, the test returned a local p -value of about 0.001 for two sources (NGC 1068 and TXS 0506+056). Since the local p -values of these excesses are orders of magnitude lower than those obtained when performing the test with a floating spectral index, we did not further investigate these results.

For the list of 47 X-ray-bright AGN, we also applied the binomial test assuming the core–corona emission model (A. Kheirandish et al. 2021). The core–corona model predicts the shape of the neutrino flux for each tested source candidate based on the intrinsic X-ray luminosities as reported in the

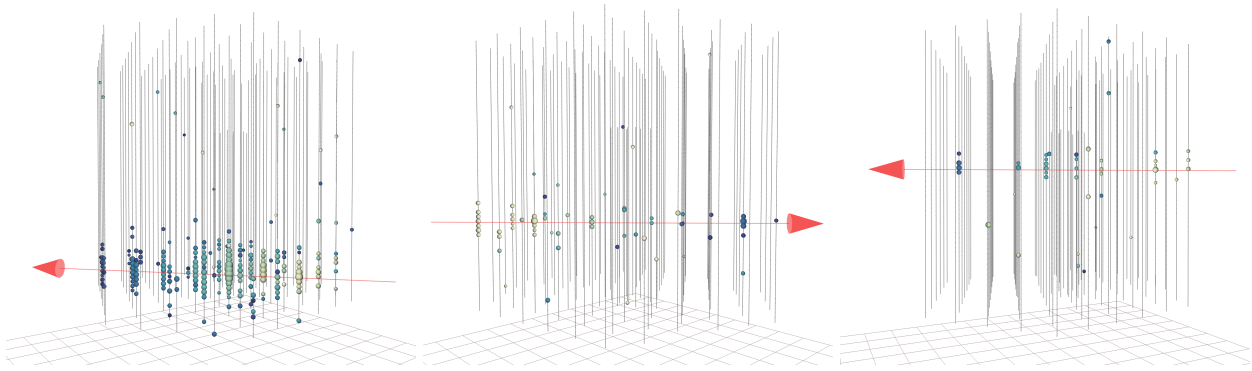


Figure 11. Event views of three of the events associated with the neutrino excess of NGC 1068. The leftmost event is the one that contributes the highest S/B ; in the middle panel, we show the event that contributes the most among those added in this new iteration of the sample; and in the right panel, we show the highest-contributing event from the IC 79 season. The colored blobs indicate the photons that hit each DOM: the bigger the blob, the higher the number of detected Cherenkov photons. The color scale indicates the time from the first hit in the detector (light) to the last (dark).

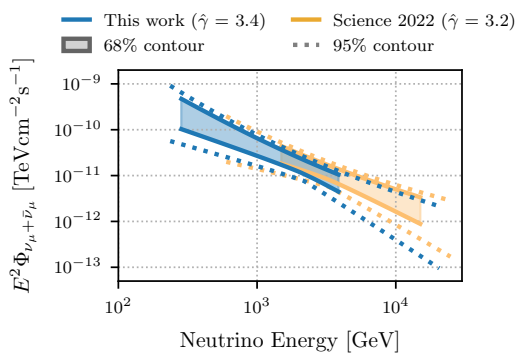


Figure 12. Comparison between the 68% (solid) and 95% (dashed) ν_μ flux uncertainties and energy ranges of NGC 1068’s measurement of this work (blue) and IceCube Collaboration (2022; yellow). The 68% energy interval obtained for this analysis ranges between 0.3 and 3.9 TeV, while the 95% one ranges between 0.2 and 20.6 TeV; the previous result spanned 1.5 and 15 TeV for the 68% contour and 0.6 and 24.8 TeV for the 95% one.

BASS catalog (M. J. Koss et al. 2022), where additional model parameters are fixed to match the NGC 1068 neutrino flux (R. Abbasi et al. 2025a). The f_S signal pdfs in the likelihood function (Equation (1)) were constructed using the same KDE method as in the power-law case but assuming the predicted flux shape. It results in n_s as the only free parameter in the likelihood ratio test (Equation (2)) that corresponds to the neutrino flux normalization. This is not the first time such a model has been tested using IceCube track-like events. R. Abbasi et al. (2025a) previously reported a 2.7σ excess from a binomial test applied to a list of 27 Seyfert galaxies. Given the substantial overlap between the two samples, 23 out of 47 sources (or 24 out of 48 when including NGC 1068), we applied the same model to our list. However, this analysis should not be interpreted as a direct follow-up to R. Abbasi et al. (2025a). The selection criteria used to compile the source list presented here differ from previous ones as described in Section 4.3.

The binomial test based on the core–corona model yielded a local p -value of approximately 0.001 for three sources out of 47. When NGC 1068 is included in the sample, the p -value further decreases to $\sim 10^{-5}$. Notably, two of the three sources contributing to the excess, NGC 4151 and CGCG 420–015, were also identified in R. Abbasi et al. (2025a). Moreover,

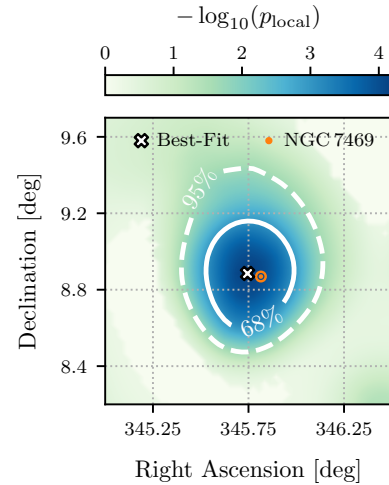


Figure 13. High-resolution scan and likelihood contours around the location of NGC 7469, a type I Seyfert galaxy found to be the most significant source in the list of 47 X-ray-bright AGN. Similarly to the right panel of Figure 2, the white cross represents the best-fit location, the solid (dashed) line represents the 68% (95%) uncertainty contour of the excess, and the orange dot and circle represent, respectively, the source location and its optical size (obtained from the NASA/IPAC Extragalactic Database).

these sources also contribute to the excess observed under the power-law spectral assumption, further supporting the hypothesis that they may be genuine neutrino emitters.

As discussed before, the global significance of the results cannot be calculated analytically. Therefore, we compute them from pseudoexperiments. Similarly to what is described in Appendix D.1, the a posteriori correction is obtained in two steps. First, we account for testing multiple significance thresholds by comparing the most significant excess with the distribution of the minimum p -values from background-only pseudoexperiments generated for each individual hypothesis. The fraction of realizations with smaller values than the observed result gives the threshold-corrected significance. In the second step, this threshold-corrected p -value is compared to the distribution of the minimum p -values obtained by analyzing the same pseudodata under the different signal hypotheses. The fraction of realizations resulting in a p -value smaller than the corrected p -value yields the final globally corrected p -value.

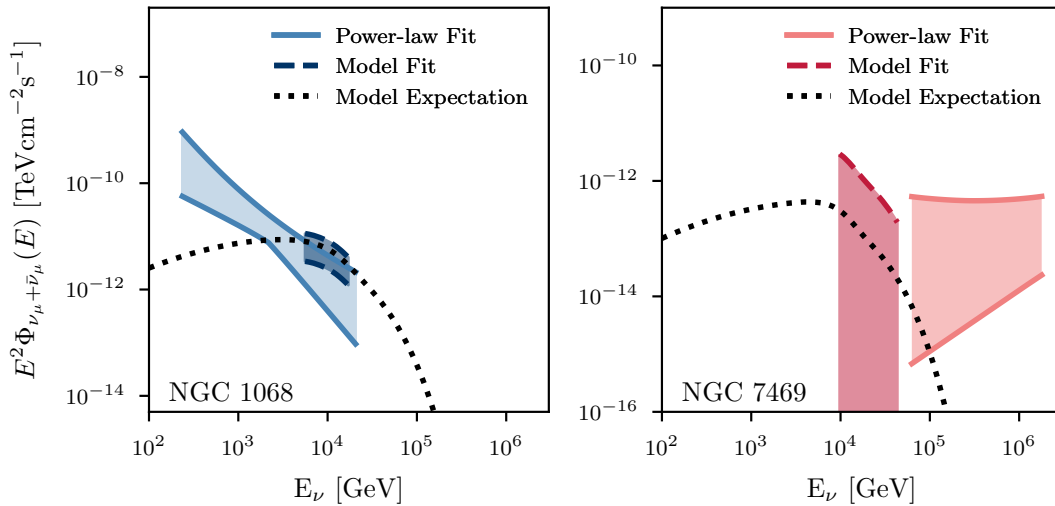


Figure 14. Comparison between the 95% uncertainties and sensitive energy ranges of the power-law and core-corona model spectra for NGC 1068 (left) and NGC 7469 (right). The power laws are shown in lighter colors with solid contours, while fluxes from the core-corona model are represented in darker colors with dashed contours. The dotted black lines represent the core-corona model expectations. The ν_μ flux from NGC 7469 in the core-corona model assumption is only partially constrained by the data.

D.3. Most Significant Sources

In Section 4.2, we present the result of a follow-up analysis of a list of 110 gamma-ray emitters, including NGC 1068. The mean number of signal events \hat{n}_s contributing to the neutrino excess increased from 79 to 102, demonstrating that, despite the slight decrease in global significance of the analysis, new events contribute to the excess. In Figure 11, we display three of the events contributing the most to the signal excess from NGC 1068. On the left, we show the event contributing the most to the observed TS value obtained when testing NGC 1068 for pointlike neutrino emission. In the middle, we show the most-contributing event among the additional data that were included in this round of the analysis, which is the ninth-most-contributing event. Finally, in the right panel, we display the most-contributing event in IC 79 data. This event is ranked 25th among those most contributing to the excess.

In Figure 12, we compare the newly measured spectrum of NGC 1068 with that of IceCube Collaboration (2022). The energy ranges are computed by evaluating the TS contribution of each event adding to the excess, selecting events in the simulations with similar reconstructed energy and angular error, and building a histogram of their true neutrino energy. The histograms are then summed, weighted by their TS contribution. The displayed energy ranges correspond to the 68% and 95% central intervals of such a histogram. The flux intervals are instead computed assuming Wilks’s theorem, which has been validated on MC simulations. As already noted in Section 4.2, this analysis reports a softer spectrum compared to the previous result (from $\hat{\gamma} = 3.2$ to $\hat{\gamma} = 3.4$) and a shift in the energy range toward lower energies. Despite the large overlap in the events contributing to NGC 1068’s neutrino excess in the two analyses, the new events are characterized by lower energies with respect to the ones driving the excess in IceCube Collaboration (2022), causing the observed change in energy range and best-fit spectral index. Nevertheless, the newly measured energy spectrum remains compatible with that of IceCube Collaboration (2022) within the 95% contours.

The search for neutrino emission from the list of 47 X-ray-bright, nonblazar AGN instead found NGC 7469 as the most significant, with 2.4σ after accounting for the look-elsewhere

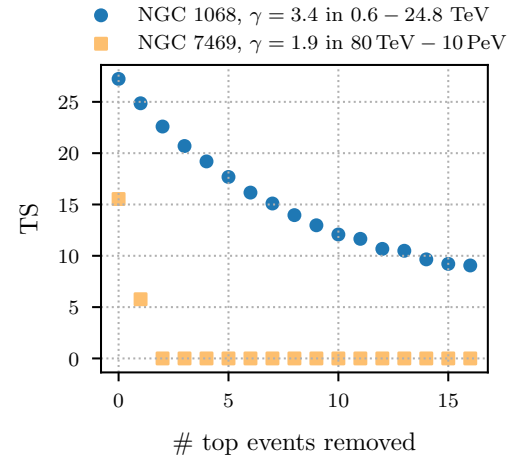


Figure 15. Dependency of the TS from the events with the highest S/B ratio for the two most significant sources found in this analysis. We calculate the TSs multiple times, removing the events with the highest S/B one by one. While NGC 1068’s TS decreases slowly by removing the highest-contributing events, NGC 7469’s TS rapidly drops to zero, as two high-energy events drive the excess.

effect. The likelihood fit for this source returned a very hard spectral index of $\hat{\gamma} = 1.9$ and a mean number of signal events of $\hat{n}_s = 5.5$. As shown in Figure 13, the likelihood scan in the vicinity of the source reveals that the best-fit position is well compatible with the cataloged coordinates. At the same time, the energy spectrum of this source, as opposed to NGC 1068, spans energies of the order of 10–100 TeV (see Figure 5). In fact, in this case, the neutrino excess is driven by only two very high-energy ($E_\nu > 100$ TeV) neutrino alert events (see also G. Sommani et al. 2025).

The emergence of a source with a markedly different spectrum than the well-known soft-spectrum source NGC 1068 demonstrates the flexibility of the likelihood method. The custom core-corona model, which was also tested in this context, peaks at $\lesssim 10$ TeV with a sharp cutoff at higher energies (R. Abbasi et al. 2025a). By construction, it penalizes any high-energy emission from the target source, as the right panel of Figure 14 shows. Figure 15 shows how the

TS depends on the events with the highest signal-over-background (S/B) ratio. For NGC 1068, removing the most significant events one by one leads to a gradual reduction in the TS, which eventually reaches zero. In contrast, for NGC 7469, the TS drops to zero immediately after removing the two most significant events. This highlights the ability of the analysis to recover signals in both scenarios: a steady accumulation of many low-energy events from a soft-spectrum source like NGC 1068 and a few high-energy events from a hard-spectrum source like NGC 7469. The starkly different behavior of these two sources also suggests that not all X-ray-bright AGN share the same cosmic-ray acceleration and/or neutrino production mechanisms.

D.4. Candidate Source Lists

Unlike the all-sky searches and binomial tests, the catalog search on the list of 110 gamma-ray emitters was performed only under the assumption of a floating spectral index. This is the only case in which the global significance is computed analytically using the well-known Šidák correction (Z. Šidák 1967), $p_{\text{post}} = 1 - (1 - p_{\text{local}})^N$, where N , in this case, is 110, i.e., the number of tested sources. No correlation between tests is included in the global significance calculation, since all sources are sufficiently spatially separated to be treated as independent.

The only catalog search not discussed in the main text is the one carried out on the list of 47 bright X-ray AGN, assuming neutrino emission follows the core–corona model (A. Kheirandish

et al. 2021). In this case, the global significance cannot be obtained with the Šidák correction alone, since we tested two spectral assumptions that might share some level of correlation. Therefore, after correcting for having tested 47 sources, we calculate the distribution of the minimum p -values out of pseudoexperiments analyzed under both spectral hypotheses and define the globally corrected p -value as the fraction of pseudoexperiments resulting in a p -value smaller than the source-corrected p -value.

As in the case of the binomial test, the catalog search yielded a lower significance under the core–corona hypothesis compared to the power-law assumption. This outcome reinforces the conclusion that, given our current understanding of neutrino emission, the simple yet flexible assumption of an unbroken power-law spectrum remains the most consistent with the data.

In agreement with R. Abbasi et al. (2025a), CGCG 420–015 emerged as the most significant source in the list under the core–corona model, with a fitted signal strength of $\hat{n}_s = 33.3$ and a local significance of 3.8σ . NGC 1068, the most prominent neutrino emitter in the sky, is always evaluated a posteriori. In this case, it was found to exhibit an excess of 48 signal events with a local significance of 5σ .

A complete summary of all fit results is provided in Tables 2 and 3. These include the likelihood fit values, local significances, and 90% upper limits for both $\gamma = 2$ and $\gamma = 3$ for all tested sources. In the second table, which presents results for the X-ray-bright AGN sample, we additionally report the outcomes obtained under the core–corona model.

Table 2
List of Gamma-Ray Emitters Selected as Candidate Neutrino Sources

Source Name	R.A.	Decl.	\hat{n}_s	$\hat{\gamma}$	$-\log_{10} p_{\text{local}}$	$\phi_{90\%}^{E^{-2}}$	$\phi_{90\%}^{E^{-3}}$
NGC 1068	40.67	−0.01	102	3.4	6.6 (5.0 σ)	6.6	385.6
PKS 1424+240	216.76	23.80	96	3.6	4.1 (3.8 σ)	8.3	227.2
TXS 0506+056	77.36	5.70	5	1.9	3.4 (3.3 σ)	5.2	249.5
GB6 J1542+6129	235.75	61.50	27	3.2	1.9 (2.2 σ)	9.3	145.8
LQAC 284+003	284.48	3.22	14	2.6	1.6 (1.9 σ)	3.4	178.5
S3 0458−02	75.30	−1.97	34	4.4	1.4 (1.8 σ)	2.9	180.2
OJ 014	122.86	1.78	49	4.0	1.4 (1.8 σ)	3.0	173.9
MITG J201534+3710	303.89	37.18	45	3.8	1.3 (1.7 σ)	5.6	119.4
B2 0619+33	95.73	33.43	48	4.4	1.3 (1.6 σ)	5.2	117.5
MGR0 J1908+06	286.91	6.32	2	1.8	1.3 (1.6 σ)	3.3	152.0
NGC 5380	209.33	37.50	12	2.7	1.2 (1.6 σ)	5.4	115.2
B2 1215+30	184.48	30.12	20	3.0	1.2 (1.6 σ)	4.9	117.6
NGC 2146	94.53	78.33	2	1.3	1.2 (1.6 σ)	9.7	138.6
B3 0609+413	93.22	41.37	9	2.4	1.2 (1.5 σ)	5.5	110.9
S2 0109+22	18.03	22.75	23	3.0	1.1 (1.4 σ)	4.1	117.3
3C 454.3	343.50	16.15	2	1.6	1.1 (1.4 σ)	3.7	124.9
B2 2234+28A	339.10	28.48	28	3.3	1.0 (1.3 σ)	4.5	109.5
TXS 0603+476	91.86	47.66	35	4.4	1.0 (1.3 σ)	5.5	102.7
S5 1044+71	162.11	71.73	45	4.4	0.9 (1.2 σ)	8.6	111.1
M31	10.82	41.24	20	3.5	0.9 (1.2 σ)	4.8	96.1
PKS 1441+25	220.99	25.03	6	2.2	0.9 (1.2 σ)	3.9	106.1
3C 273	187.27	2.05	34	4.4	0.9 (1.1 σ)	2.4	129.0
7C 2010+4619	303.02	46.49	4	2.0	0.8 (1.1 σ)	4.8	90.1
PKS 1502+106	226.10	10.50	7	2.3	0.8 (1.0 σ)	2.9	114.5
B2 2308+34	347.77	34.42	27	3.8	0.8 (1.0 σ)	4.2	90.3
TXS 0518+211	80.44	21.21	14	2.9	0.8 (0.9 σ)	3.3	93.3
PKS 1717+177	259.81	17.75	28	3.7	0.8 (0.9 σ)	3.2	99.2
PMN J0948+0022	147.24	0.37	23	4.4	0.7 (0.9 σ)	2.1	120.0
TXS 1902+556	285.81	55.68	18	3.7	0.7 (0.9 σ)	4.7	80.6
IC 678	168.56	6.63	15	2.8	0.7 (0.8 σ)	2.5	106.6
4C +55.17	149.42	55.38	18	3.4	0.7 (0.8 σ)	4.7	81.5
4C +38.41	248.82	38.14	4	2.4	0.7 (0.8 σ)	3.9	79.3

Table 2
(Continued)

Source Name	R.A.	Decl.	\hat{n}_s	$\hat{\gamma}$	$-\log_{10} p_{\text{local}}$	$\phi_{90\%}^{E^{-2}}$	$\phi_{90\%}^{E^{-3}}$
PKS 0215+015	34.46	1.73	24	4.0	0.7 (0.8 σ)	2.1	111.7
Mrk 421	166.12	38.21	12	2.8	0.7 (0.8 σ)	3.9	78.9
NVSS J141826–023336	214.61	–2.56	11	4.4	0.6 (0.8 σ)	2.0	120.9
OJ 287	133.71	20.12	24	4.4	0.6 (0.7 σ)	3.0	84.5
SBS 0846+513	132.51	51.14	7	3.1	0.6 (0.7 σ)	4.2	72.1
IES 0647+250	102.70	25.05	23	4.4	0.6 (0.6 σ)	3.1	81.7
PKS B1130+008	173.20	0.57	17	4.4	0.6 (0.6 σ)	1.9	102.1
IES 1959+650	300.01	65.15	6	2.8	0.6 (0.6 σ)	5.6	77.1
BL Lac	330.69	42.28	16	3.7	0.6 (0.6 σ)	3.7	72.1
PKS 0235+164	39.67	16.62	20	3.8	0.5 (0.5 σ)	2.6	79.6
B2 0218+357	35.28	35.94	14	4.4	0.4 (0.4 σ)	3.2	61.6
MITG J200112+4352	300.30	43.89	7	3.0	0.4 (0.4 σ)	3.4	62.6
Ton 599	179.88	29.24	14	4.4	0.4 (0.3 σ)	2.9	65.0
NGC 1275	49.96	41.51	7	3.2	0.4 (0.3 σ)	3.2	59.7
TXS 1055+567	164.67	56.46	5	3.5	0.4 (0.3 σ)	3.7	57.3
RX J1931.1+0937	292.78	9.63	12	4.4	0.4 (0.3 σ)	2.1	76.6
4C +14.23	111.32	14.42	12	3.8	0.4 (0.3 σ)	2.2	69.5
Arp 299	172.07	58.52	8	4.4	0.4 (0.3 σ)	3.7	55.5
Mrk 501	253.47	39.76	9	4.4	0.4 (0.2 σ)	3.0	55.1
S4 1250+53	193.31	53.02	5	4.4	0.4 (0.2 σ)	3.4	51.2
4C +15.54	241.77	15.84	10	4.4	0.4 (0.2 σ)	2.1	62.3
NVSS J184425+154646	281.12	15.79	7	4.4	0.3 (0.1 σ)	2.1	61.0
B2 2114+33	319.06	33.66	3	2.8	0.3 (0.1 σ)	2.7	50.6
4C +28.07	39.47	28.80	3	2.9	0.3 (0.0 σ)	2.5	51.1
B3 1343+451	206.39	44.88	2	3.0	0.3 (0.0 σ)	2.9	47.6
4C +41.11	65.98	41.83	1	3.0	0.3 (0.0 σ)	2.8	45.5
PKS 0735+17	114.54	17.71	3	3.0	0.3 (0.0 σ)	2.1	54.7
OX 169	325.89	17.73	4	4.4	0.3 (0.0 σ)	2.0	53.1
M82	148.95	69.67	4	4.4	0.3 (0.0 σ)	4.1	47.0
MG2 J043337+2905	68.41	29.10	1	4.4	0.3 (0.0 σ)	2.3	43.5
W Comae	185.38	28.24	0	...	0.3 (0.0 σ)	2.2	44.5
S5 1803+784	270.17	78.47	0	...	0.0 (0.0 σ)	6.1	80.9
PMN J0709–0255	107.45	–2.93	0	...	0.0 (0.0 σ)	1.6	80.9
PKS 1216–010	184.64	–1.33	0	...	0.0 (0.0 σ)	1.4	66.7
PKS 0422+00	66.19	0.60	0	...	0.0 (0.0 σ)	1.4	67.1
PKS 0420–01	65.83	–1.33	0	...	0.0 (0.0 σ)	1.4	68.1
OT 081	267.88	9.65	0	...	0.0 (0.0 σ)	1.8	58.6
CTA 102	338.15	11.73	0	...	0.0 (0.0 σ)	1.7	53.1
PKS 0336–01	54.88	–1.78	0	...	0.0 (0.0 σ)	1.5	74.3
PKS 0736+01	114.82	1.62	0	...	0.0 (0.0 σ)	1.4	62.1
PKS 0440–00	70.66	–0.30	0	...	0.0 (0.0 σ)	1.4	62.5
1H 1720+117	261.27	11.87	0	...	0.0 (0.0 σ)	1.7	52.2
4C +01.28	164.62	1.56	0	...	0.0 (0.0 σ)	1.4	62.9
M87	187.71	12.39	0	...	0.0 (0.0 σ)	1.8	54.4
4C +01.02	17.17	1.58	0	...	0.0 (0.0 σ)	1.4	64.5
OG 050	83.17	7.55	0	...	0.0 (0.0 σ)	1.6	58.9
TXS 0141+268	26.15	27.09	0	...	0.0 (0.0 σ)	2.2	47.1
RGB J2243+203	340.99	20.36	0	...	0.0 (0.0 σ)	2.0	48.3
PKS 0829+046	127.97	4.49	0	...	0.0 (0.0 σ)	1.6	61.3
MITG J123931+0443	189.89	4.73	0	...	0.0 (0.0 σ)	1.5	59.2
PG 1553+113	238.93	11.19	0	...	0.0 (0.0 σ)	1.7	53.4
PKS 2032+107	308.85	10.94	0	...	0.0 (0.0 σ)	1.7	53.4
4C +21.35	186.23	21.38	0	...	0.0 (0.0 σ)	1.9	47.0
PG 1218+304	185.34	30.17	0	...	0.0 (0.0 σ)	2.2	45.5
RX J1754.1+3212	268.55	32.20	0	...	0.0 (0.0 σ)	2.3	45.8
Arp 220	233.70	23.53	0	...	0.0 (0.0 σ)	2.1	46.8
MITG J021114+1051	32.81	10.86	0	...	0.0 (0.0 σ)	1.7	52.7
87GB 194024.3+102612	295.70	10.56	0	...	0.0 (0.0 σ)	1.7	54.4
PKS 0502+049	76.34	5.00	0	...	0.0 (0.0 σ)	1.5	56.7
IES 0033+595	8.98	59.83	0	...	0.0 (0.0 σ)	3.6	49.2
NGC 3424	162.91	32.89	0	...	0.0 (0.0 σ)	2.3	45.5
GB6 J1037+5711	159.43	57.19	0	...	0.0 (0.0 σ)	3.3	46.8
ON 246	187.56	25.30	0	...	0.0 (0.0 σ)	2.1	43.4
PKS 1502+036	226.27	3.45	0	...	0.0 (0.0 σ)	1.5	59.0

Table 2
(Continued)

Source Name	R.A.	Decl.	\hat{n}_s	$\hat{\gamma}$	$-\log_{10}P_{\text{local}}$	$\phi_{90\%}^{E^{-2}}$	$\phi_{90\%}^{E^{-3}}$
PKS 0507+17	77.52	18.01	0	...	0.0 (0.0 σ)	1.8	47.0
B2 1520+31	230.55	31.74	0	...	0.0 (0.0 σ)	2.3	45.2
PG 1246+586	192.08	58.34	0	...	0.0 (0.0 σ)	3.3	46.3
IES 0806+524	122.46	52.31	0	...	0.0 (0.0 σ)	3.1	45.8
PKS 0019+058	5.64	6.13	0	...	0.0 (0.0 σ)	1.5	56.0
TXS 2241+406	341.06	40.96	0	...	0.0 (0.0 σ)	2.6	44.6
TXS 1452+516	223.62	51.41	0	...	0.0 (0.0 σ)	3.0	43.7
1H 1013+498	153.77	49.43	0	...	0.0 (0.0 σ)	2.9	44.8
B3 0133+388	24.14	39.10	0	...	0.0 (0.0 σ)	2.5	44.0
S4 0814+42	124.56	42.38	0	...	0.0 (0.0 σ)	2.6	42.4
S4 0917+44	140.23	44.70	0	...	0.0 (0.0 σ)	2.6	40.7
3C 66A	35.67	43.04	0	...	0.0 (0.0 σ)	2.6	43.0
S4 1749+70	267.16	70.10	0	...	0.0 (0.0 σ)	3.7	41.6
S5 0716+71	110.49	71.34	0	...	0.0 (0.0 σ)	3.5	37.9

Note. Sources are ordered by descending significance. For each source, we list the name, the equatorial coordinates (J2000 equinox) from the 4FGL-DR2 catalog (S. Abdollahi et al. 2020), and the likelihood search results: number of signal events \hat{n}_s , spectral index $\hat{\gamma}$, $-\log_{10}P_{\text{local}}$ and corresponding significance in brackets, and 90% C.L. astrophysical flux upper limits ($\phi_{90\%}$) with $\phi_{\nu_\mu+\bar{\nu}_\mu} = \phi_{90\%}(E_\nu/1 \text{ TeV})^{-\gamma} \times 10^{-13} \text{ TeV}^{-1} \text{ cm}^{-2} \text{ s}^{-1}$ for $\gamma = 2.0$ and 3.0.

Table 3
List of X-Ray-bright AGN Selected as Candidate Neutrino Sources



Source Name	R.A.	Decl.	$F_{20-50 \text{ keV}}^{\text{intr}}$	Power-law Fit Results					Core-Corona Model Fit Results		
				\hat{n}_s	$\hat{\gamma}$	$-\log_{10}P_{\text{local}}$	$\phi_{90\%}^{E^{-2}}$	$\phi_{90\%}^{E^{-3}}$	\hat{n}_s	$-\log_{10}P_{\text{local}}$	$n_{90\%}^{\text{Model}}$
NGC 1068*	40.67	-0.01	7.72	102	3.4	6.6 (5.0 σ)	6.6	385.6	48	5.8 (4.7 σ)	73.0
NGC 7469	345.82	8.87	2.69	5	1.9	4.1 (3.8 σ)	6.3	265.7	7	0.9 (1.1 σ)	58.0
NGC 4151*	182.64	39.41	18.09	28	2.7	2.9 (3.1 σ)	8.5	178.1	23	3.0 (3.1 σ)	45.5
CGCG 420-015*	73.36	4.06	1.77	35	2.7	2.4 (2.7 σ)	4.3	222.1	33	3.4 (3.4 σ)	46.3
Cygnus A*	299.87	40.73	4.93	3	1.6	2.2 (2.5 σ)	7.5	153.8	1	0.5 (0.5 σ)	34.2
LEDA 166445	42.68	54.70	1.61	57	4.4	1.8 (2.1 σ)	7.6	136.9	0	0.0 (0.0 σ)	30.4
NGC 4992	197.27	11.63	2.34	27	2.9	1.6 (2.0 σ)	4.1	162.0	22	2.5 (2.7 σ)	35.7
NGC 1194*	45.95	-1.10	3.87	43	4.4	1.5 (1.8 σ)	3.0	182.0	1	0.4 (0.3 σ)	32.8
Mrk 1498	247.02	51.78	1.86	40	3.6	1.4 (1.7 σ)	6.5	120.5	8	1.0 (1.3 σ)	25.3
MCG +4-48-2*	307.15	25.73	4.32	37	3.2	1.4 (1.7 σ)	4.9	130.3	17	1.6 (2.0 σ)	31.0
NGC 3079	150.49	55.68	3.33	34	3.6	1.3 (1.7 σ)	6.6	117.2	16	1.3 (1.6 σ)	36.9
Mrk 417	162.38	22.96	1.73	4	1.9	1.3 (1.6 σ)	4.4	127.6	5	1.0 (1.3 σ)	28.3
Q0241+622	41.24	62.47	3.45	17	2.8	1.2 (1.6 σ)	7.8	118.7	14	1.9 (2.3 σ)	23.6
LEDA 138501	32.41	52.44	1.95	34	4.4	1.0 (1.2 σ)	5.4	97.5	0	0.0 (0.0 σ)	21.2
LEDA 86269	71.04	28.22	1.76	39	4.4	1.0 (1.2 σ)	4.3	104.5	0	0.0 (0.0 σ)	26.3
NGC 5252	204.57	4.54	3.65	32	3.6	0.9 (1.1 σ)	2.6	125.4	7	0.7 (0.9 σ)	28.1
3C 382*	278.76	32.70	2.62	33	4.4	0.9 (1.1 σ)	4.2	93.7	0	0.0 (0.0 σ)	23.3
NGC 4388*	186.44	12.66	10.79	2	2.0	0.8 (1.0 σ)	3.0	111.2	0	0.0 (0.0 σ)	26.0
LEDA 168563	73.02	49.55	2.15	2	1.8	0.8 (1.0 σ)	4.9	88.3	2	0.6 (0.7 σ)	19.5
Ark 120	79.05	-0.15	2.75	27	4.4	0.8 (0.9 σ)	2.2	124.4	0	0.0 (0.0 σ)	24.3
Z164-19*	221.40	27.03	8.81	3	1.9	0.7 (0.9 σ)	3.6	89.5	0	0.0 (0.0 σ)	22.5
3C 445	335.96	-2.10	2.02	14	4.4	0.7 (0.8 σ)	2.0	124.1	2	0.5 (0.5 σ)	20.4
NGC 5548	214.50	25.14	2.70	17	3.2	0.7 (0.8 σ)	3.4	85.8	7	0.9 (1.2 σ)	22.1
Mrk 6	103.05	74.43	2.15	10	2.8	0.6 (0.7 σ)	6.8	88.2	9	1.1 (1.4 σ)	19.1
NGC 3516*	166.70	72.57	4.17	34	4.4	0.6 (0.6 σ)	6.8	86.7	0	0.0 (0.0 σ)	20.6
UGC 11910*	331.76	10.23	2.20	21	4.4	0.6 (0.6 σ)	2.4	87.9	6	0.7 (0.9 σ)	21.6
4C +50.55*	321.16	50.97	7.73	9	3.0	0.6 (0.6 σ)	4.1	67.8	8	1.0 (1.2 σ)	16.2
IGR J21277+5656	321.94	56.94	1.67	4	2.4	0.6 (0.6 σ)	4.3	67.9	6	0.9 (1.1 σ)	16.7
Mrk 1040*	37.06	31.31	2.37	16	4.4	0.5 (0.5 σ)	3.1	69.7	3	0.6 (0.7 σ)	17.3
3C 111*	64.59	38.03	4.13	12	4.4	0.4 (0.3 σ)	3.1	59.3	0	0.0 (0.0 σ)	15.3
Mrk 1210*	121.02	5.11	2.36	9	4.4	0.4 (0.1 σ)	1.7	72.5	0	0.0 (0.0 σ)	18.4
NGC 1142	43.80	-0.18	4.05	6	3.7	0.3 (0.1 σ)	1.5	75.7	2	0.4 (0.3 σ)	16.0
NGC 7682*	352.27	3.53	1.99	1	2.8	0.3 (0.1 σ)	1.6	69.6	2	0.6 (0.7 σ)	16.9
NGC 7603	349.74	0.24	1.87	2	4.4	0.3 (0.0 σ)	1.5	68.4	0	0.0 (0.0 σ)	15.1
3C 390.3*	280.54	79.77	3.66	0	...	0.0 (0.0 σ)	6.7	93.7	0	0.0 (0.0 σ)	13.4
4C +74.26	310.66	75.13	2.00	0	...	0.0 (0.0 σ)	4.5	49.8	0	0.0 (0.0 σ)	12.1
NGC 6240*	253.25	2.40	12.89	0	...	0.0 (0.0 σ)	1.4	57.4	0	0.0 (0.0 σ)	14.8


Table 3
(Continued)




Source Name	R.A.	Decl.	$F_{20-50\text{ keV}}^{\text{intr}}$	Power-law Fit Results					Core-Corona Model Fit Results		
				\hat{n}_s	$\hat{\gamma}$	$-\log_{10}P_{\text{local}}$	$\phi_{90\%}^{E^{-2}}$	$\phi_{90\%}^{E^{-3}}$	\hat{n}_s	$-\log_{10}P_{\text{local}}$	$n_{s90\%}^{\text{Model}}$
NGC 3227*	155.88	19.87	4.16	0	...	0.0 (0.0 σ)	1.9	47.7	0	0.0 (0.0 σ)	15.4
IRAS 05589+2828	90.54	28.47	2.64	0	...	0.0 (0.0 σ)	2.2	42.8	0	0.0 (0.0 σ)	12.6
Mrk 79	115.64	49.81	1.82	0	...	0.0 (0.0 σ)	3.0	44.8	0	0.0 (0.0 σ)	12.1
2MASX J20145928+2523010*	303.75	25.38	2.70	0	...	0.0 (0.0 σ)	2.0	42.8	0	0.0 (0.0 σ)	13.1
IRAS 05078+1626*	77.69	16.50	3.37	0	...	0.0 (0.0 σ)	1.8	46.0	0	0.0 (0.0 σ)	13.2
Mrk 110*	141.30	52.29	2.12	0	...	0.0 (0.0 σ)	3.0	44.6	0	0.0 (0.0 σ)	11.7
NGC 4102	181.60	52.71	2.24	0	...	0.0 (0.0 σ)	3.1	46.5	0	0.0 (0.0 σ)	15.2
NGC 7319	339.01	33.98	1.70	0	...	0.0 (0.0 σ)	2.3	41.2	1	0.5 (0.4 σ)	12.7
NGC 4051	180.79	44.53	1.75	0	...	0.0 (0.0 σ)	2.7	43.1	0	0.0 (0.0 σ)	16.6
UGC 3374*	88.72	46.44	4.94	0	...	0.0 (0.0 σ)	2.7	42.0	0	0.0 (0.0 σ)	11.7
Mrk 3*	93.90	71.04	8.98	0	...	0.0 (0.0 σ)	3.5	39.7	0	0.0 (0.0 σ)	11.3




























































Note. Sources are ordered by descending significance under the power-law spectral assumption. For each source, we list the name, the equatorial coordinates (J2000 equinox) from the BASS catalog (M. J. Koss et al. 2022), and the intrinsic 20–50 keV X-ray flux ($F_{20-50\text{ keV}}^{\text{intr}} \times 10^{-11} \text{ erg cm}^{-2}\text{s}^{-1}$); the likelihood search results under the power-law hypothesis: \hat{n}_s , spectral index $\hat{\gamma}$, $-\log_{10}P_{\text{local}}$ and corresponding significance in brackets, and 90% C.L. astrophysical flux upper limits ($\phi_{90\%}$) with $\phi_{\nu_1+\nu_2} = \phi_{90\%}(E_{\nu}/1 \text{ TeV})^{-\gamma} \times 10^{-13} \text{ TeV}^{-1}\text{cm}^{-2}\text{s}^{-1}$ for $\gamma = 2.0$ and 3.0 ; and the results for the core–corona model hypothesis: number of signal events \hat{n}_s , $-\log_{10}P_{\text{local}}$ and corresponding significance in brackets, and 90% C.L. astrophysical flux upper limits given as the mean number of signal events. The sources highlighted with an asterisk were included in the R. Abbasi et al. (2025a) selection.

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